

Small Primordial Black Holes as Window on Quantum Gravity

Sebastian Zell

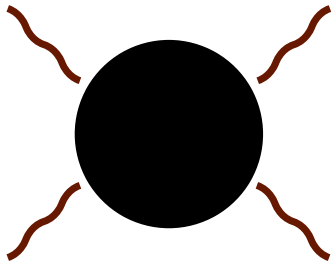
UCLouvain

Work¹ with Gia Dvali, Lukas Eisemann and Marco Michel

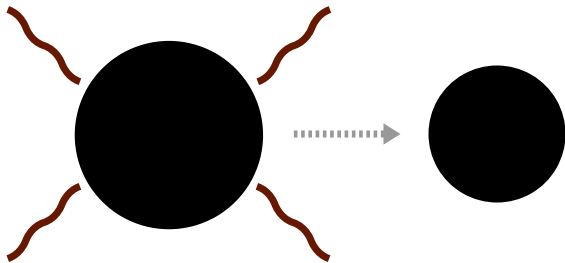
20th June 2023

¹M. M., S. Z., *The Timescales of Quantum Breaking*, arXiv:2306.09410.
G. D., L. E., M. M., S. Z., *Black Hole Metamorphosis and Stabilization by Memory Burden*, Phys. Rev. D **102** (2020), arXiv:2006.00011.

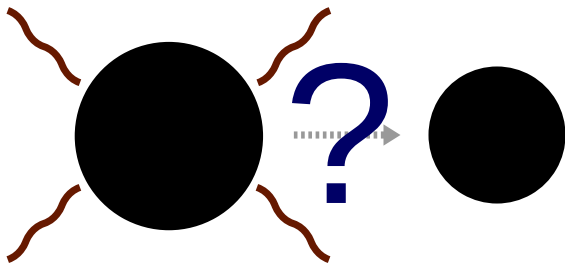
Black hole evaporation



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What happens to a black hole as it evaporates?

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- 1 Why it is an open question
- 2 Searching for small primordial black holes
- 3 Hints from analogue models

Scales of a black hole

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- ▶ Naive timescale of (half) evaporation

$$t_{1/2} \sim S r_g$$

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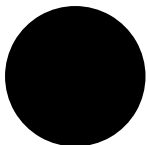
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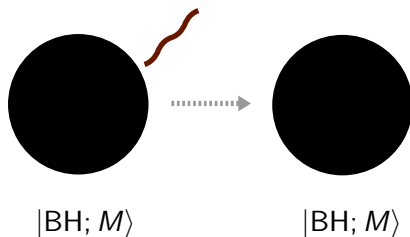
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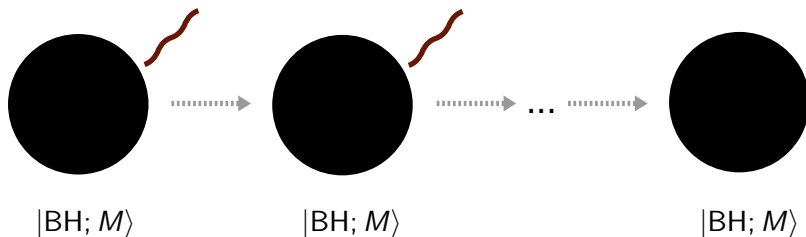


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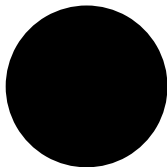
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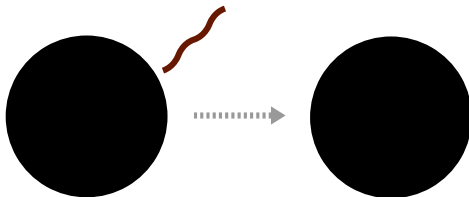


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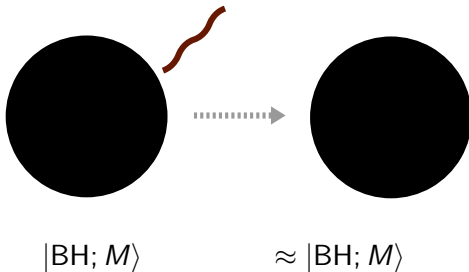


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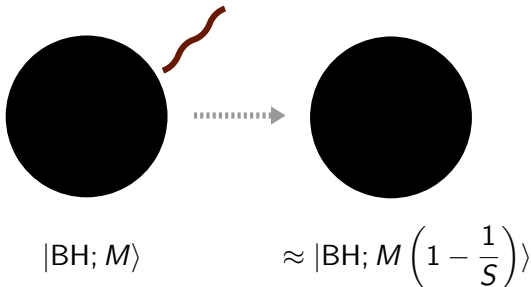
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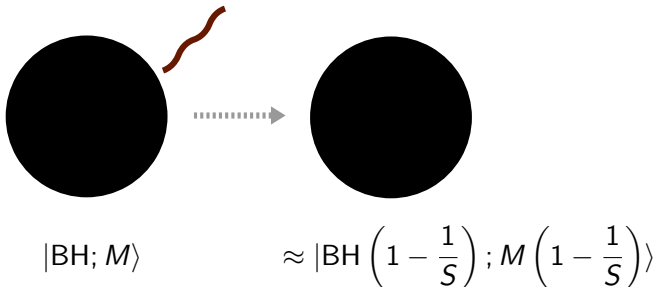
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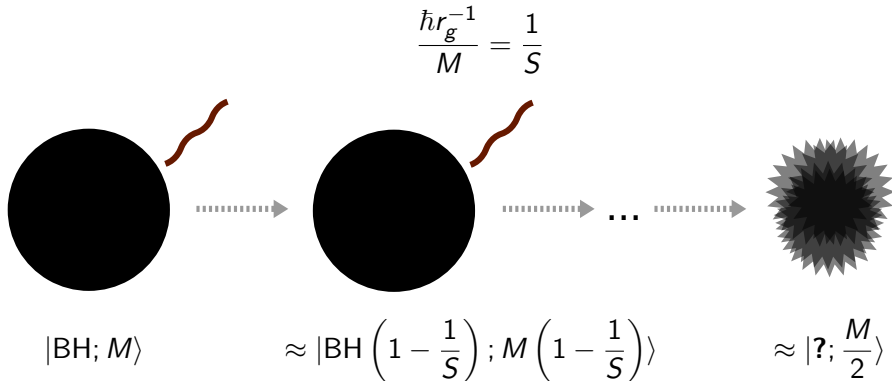
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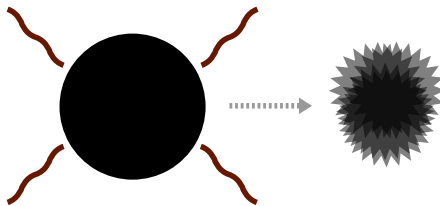


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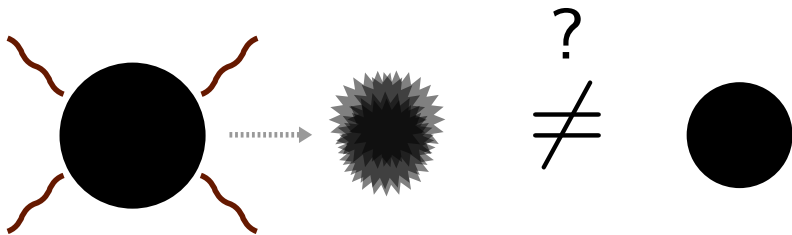
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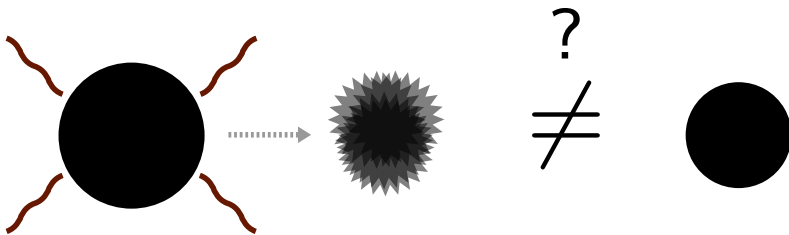
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Open questions:

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- ② What happens after a potential breakdown?

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Primordial black holes as dark matter

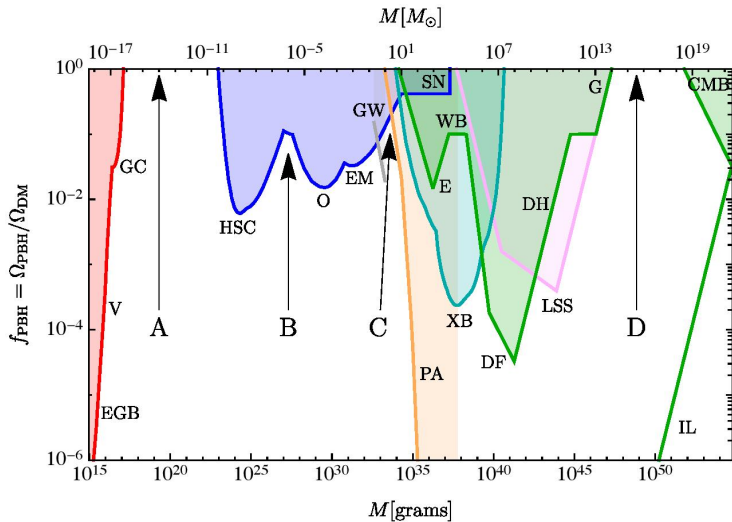


Figure from: B. Carr, F. Kühnel, *Primordial Black Holes as Dark Matter: Recent Developments*, arXiv:2006.02838.

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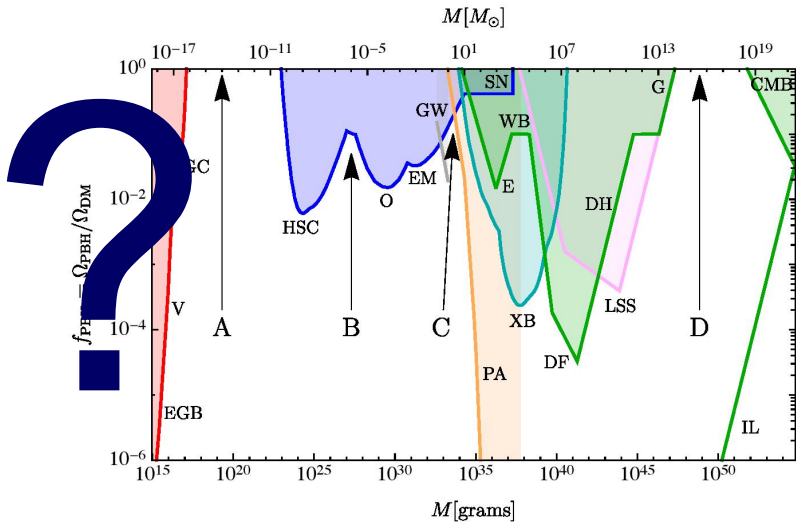


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Analogue quantum systems

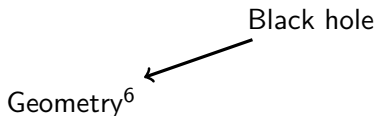
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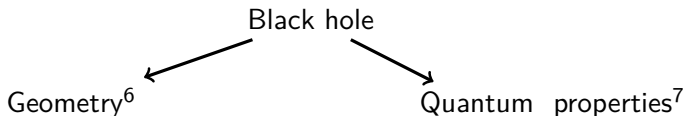
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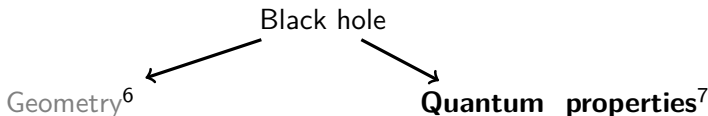


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- ▶ Analogue models: indications for early slowdown

Key property: entropy

► Entropy¹¹

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$$S = \hbar^{-1} r_g M$$

- ▶ Different microstates

$$\# = \exp(S)$$

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Prototype model¹²

- Use S modes $\hat{a}_1^\dagger, \dots, \hat{a}_S^\dagger$

$$\frac{\hat{\mathcal{H}}_S}{r_g^{-1}} = \sqrt{S} \sum_{k=1}^S \underbrace{\hat{n}_k}_{\hat{a}_k^\dagger \hat{a}_k}$$

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$$\Delta E_k \approx \sqrt{S} r_g^{-1} \left(1 - \frac{\langle \hat{n}_0 \rangle}{S} \right)$$

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- ▶ Dictionary

\hat{n}_0 : carries mass

$\langle \hat{n}_0 \rangle = S$: black hole state

\hat{n}_k : carry entropy

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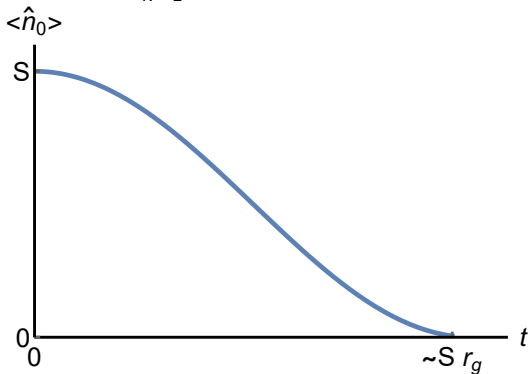
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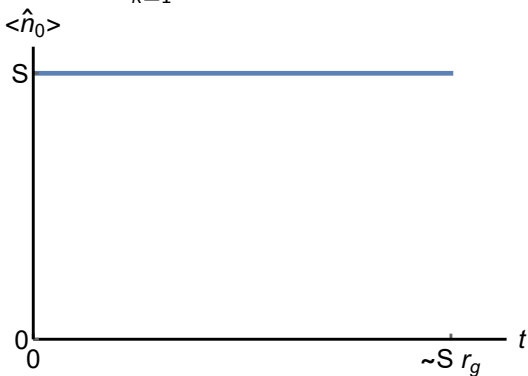
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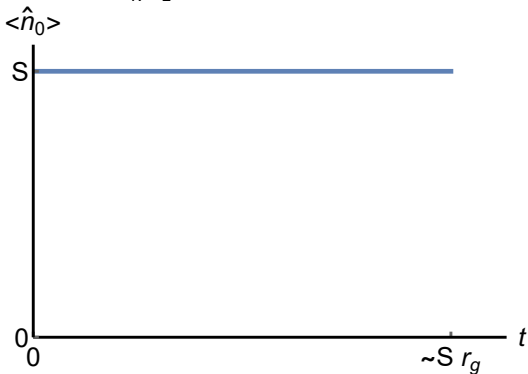
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Memory burden:¹³ entropy prevents evaporation

¹³G. Dvali, *A Microscopic Model of Holography: Survival by the Burden of Memory*, arXiv:1810.02336.

Full model¹⁶

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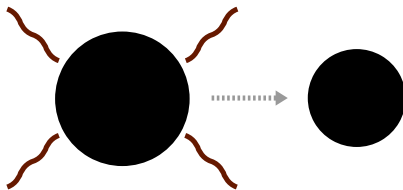
Full model¹⁶

$$\hat{\mathcal{H}} = \hat{\mathcal{H}}_{S_{>}} + \hat{n}_b + \frac{1}{S} \left(\hat{a}_0^\dagger \hat{b} + \text{h.c.} \right) + \hat{\mathcal{H}}_{S_{<}} + \text{interactions}$$

¹⁴ G. Dvali, L. Eisemann, M. Michel, S. Zell, *Black Hole Metamorphosis and Stabilization by Memory Burden*, arXiv:2006.00011.

Full model¹⁶

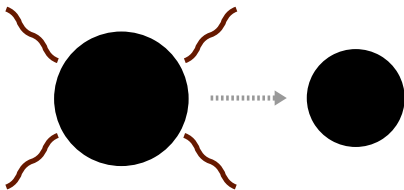
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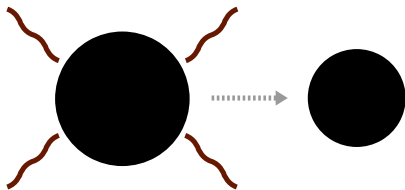


$$\langle \hat{n}_0 \rangle = S_> \longrightarrow \langle \hat{n}_0 \rangle = S_<$$

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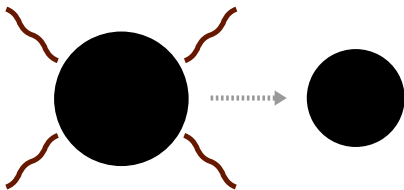
- Exact time evolution:¹⁷ transition suppressed dynamically

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¹⁵ M. Michel, S. Zell, *TimeEvolver: A Program for Time Evolution With Improved Error Bound*, arXiv:2205.15346.

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- ▶ Exact time evolution:¹⁷ transition suppressed dynamically
- ▶ Slowdown at the latest after half evaporation [back](#)

¹⁴ G. Dvali, L. Eisemann, M. Michel, S. Zell, *Black Hole Metamorphosis and Stabilization by Memory Burden*, arXiv:2006.00011.

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