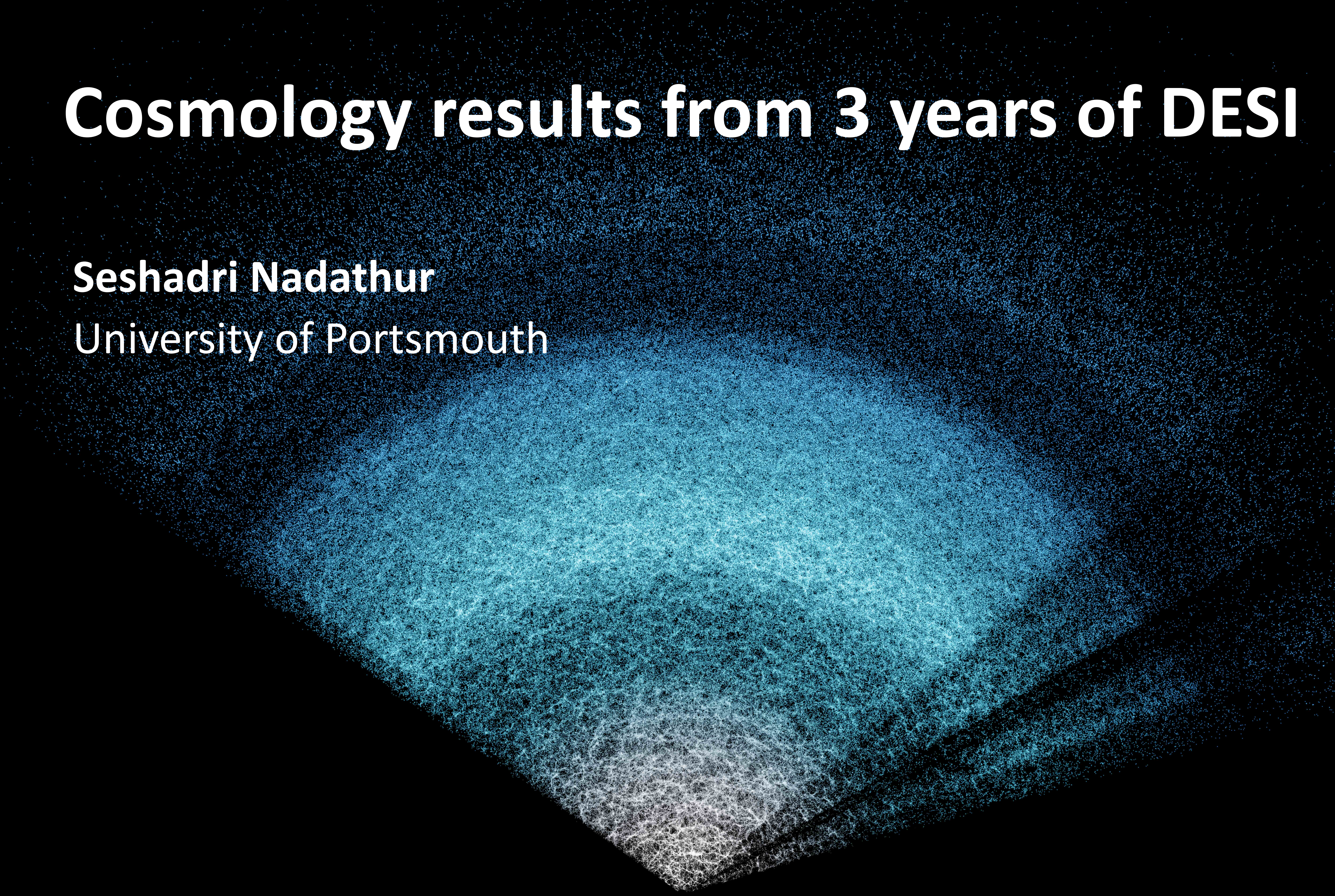
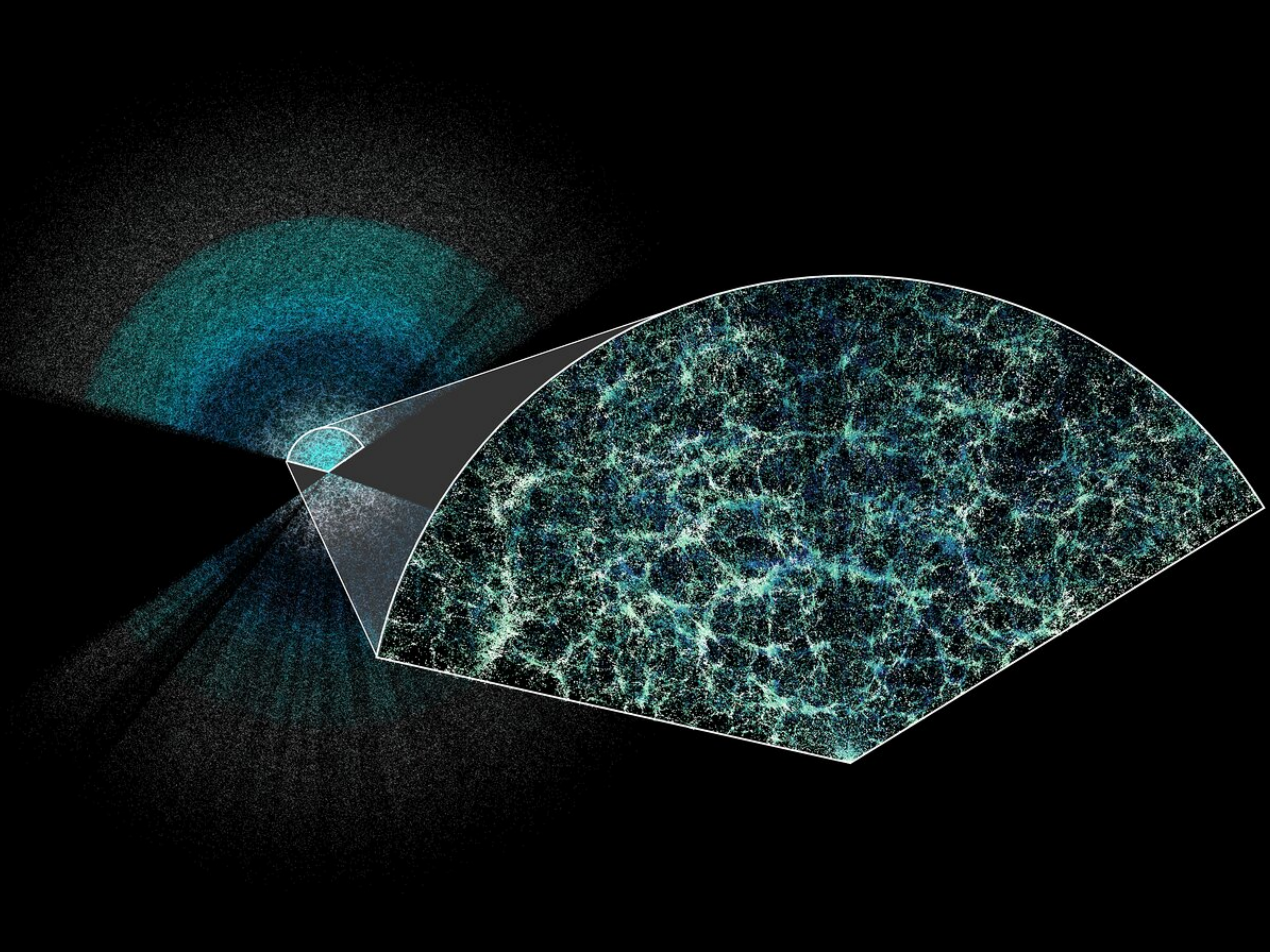
Cosmology results from 3 years of DESI



Seshadri Nadathur
University of Portsmouth

Karlsruhe, July 2025



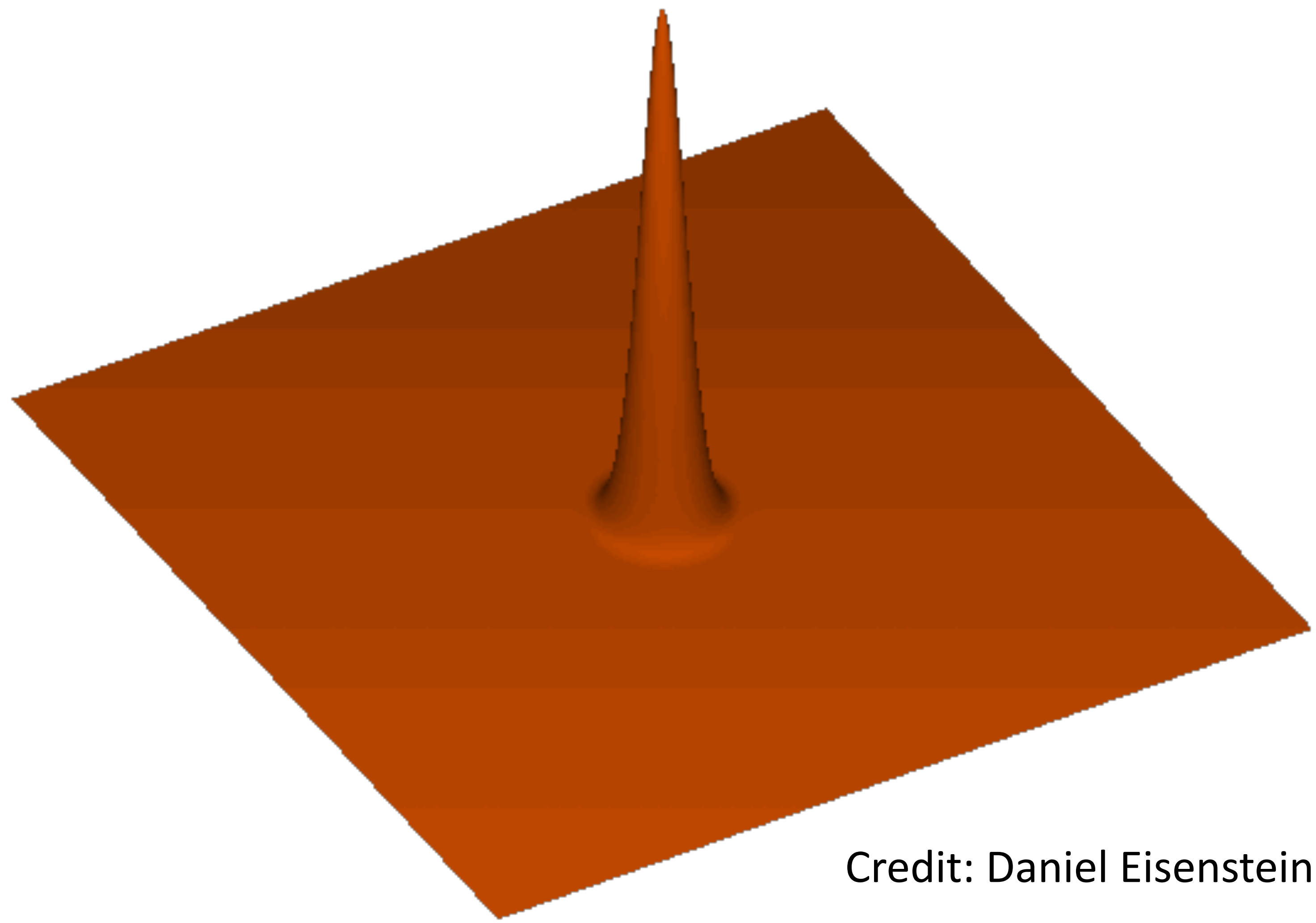
DARK ENERGY SPECTROSCOPIC INSTRUMENT
U.S. Department of Energy Office of Science

Baryon Acoustic Oscillations (BAO)

Gravity + pressure in primordial plasma → sound waves

When baryons and photons decouple ($z \sim 1100$), sound waves stall

Characteristic scale imprinted in matter distribution at sound horizon scale, $r_{\rm d} \sim 150$ Mpc



Credit: Daniel Eisenstein

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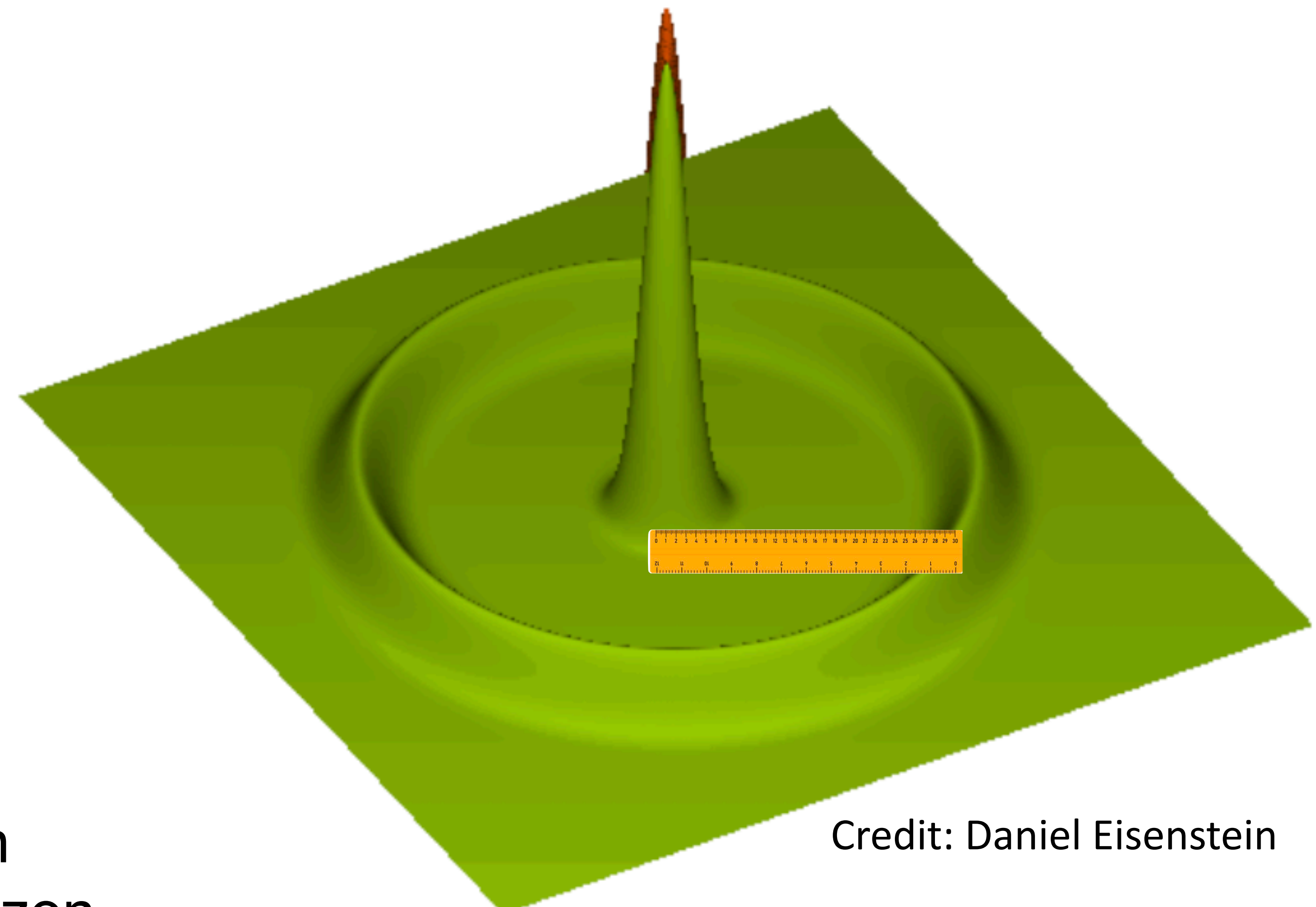

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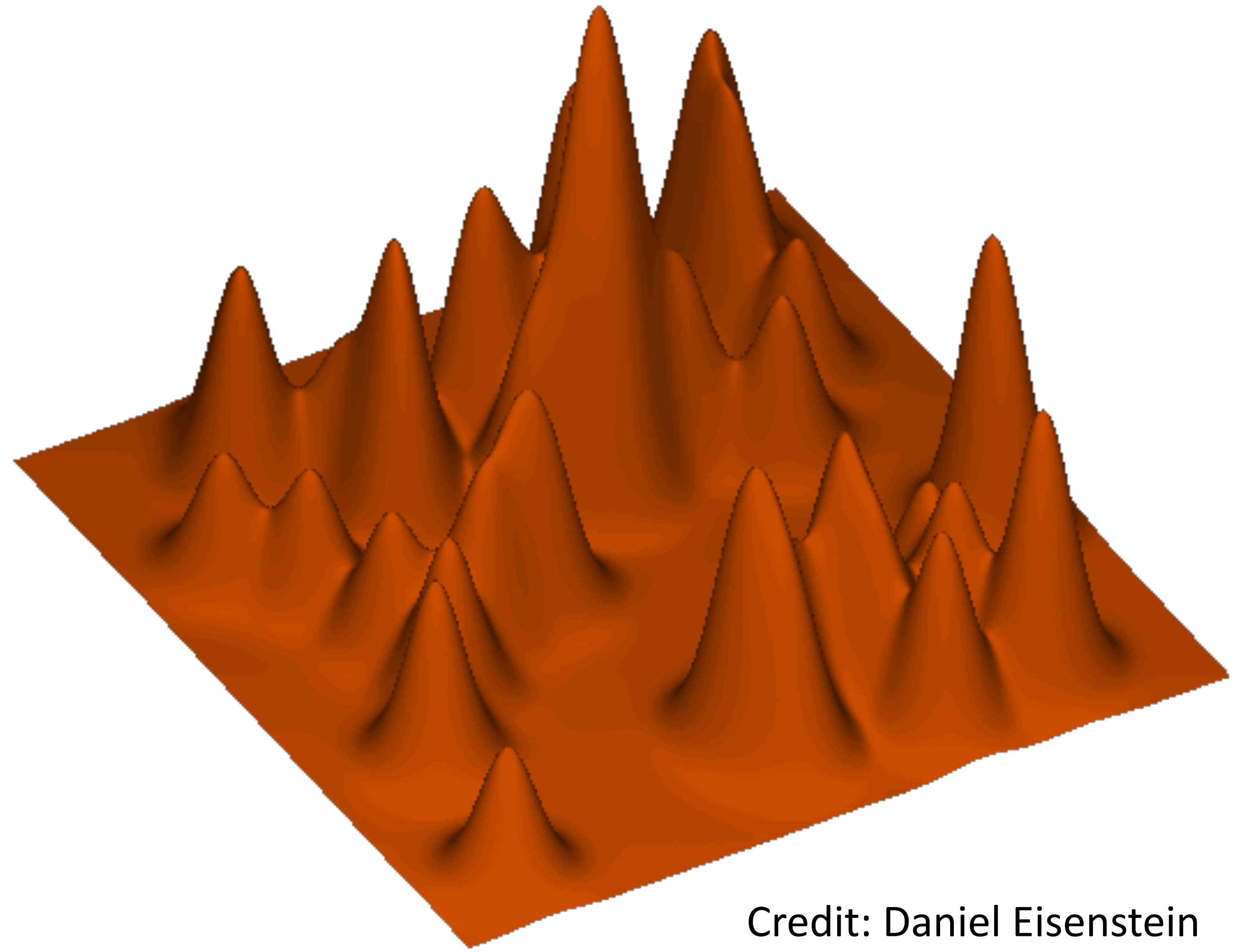
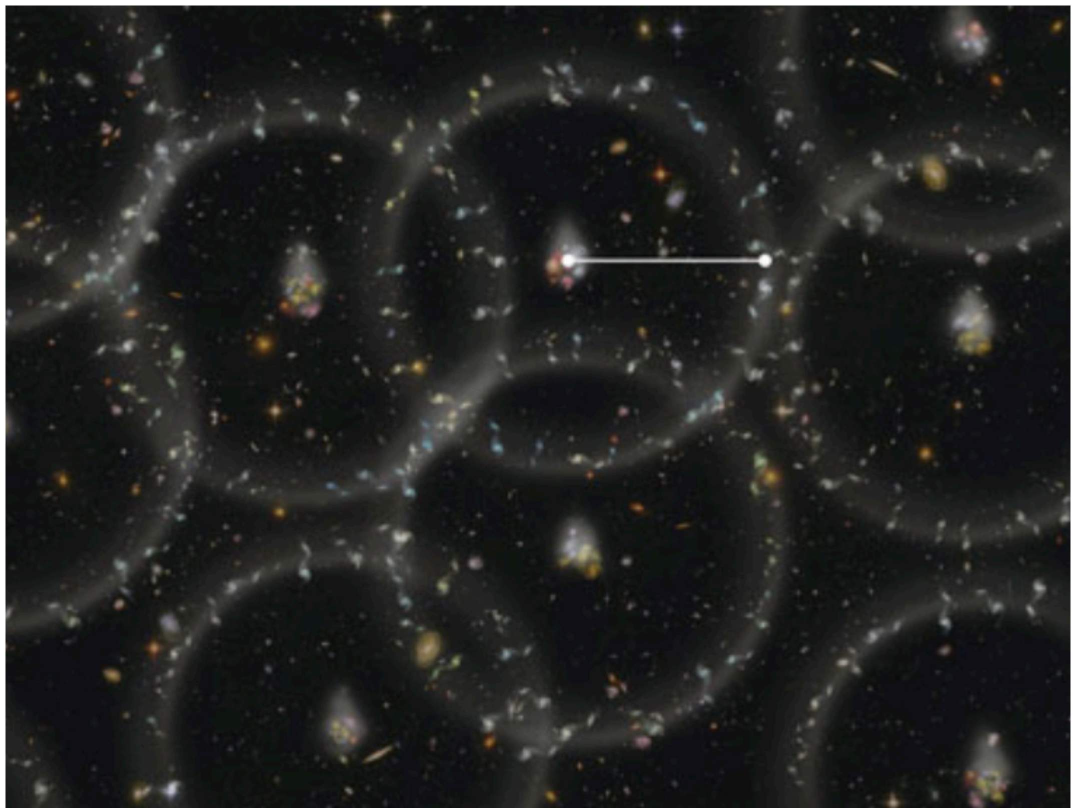
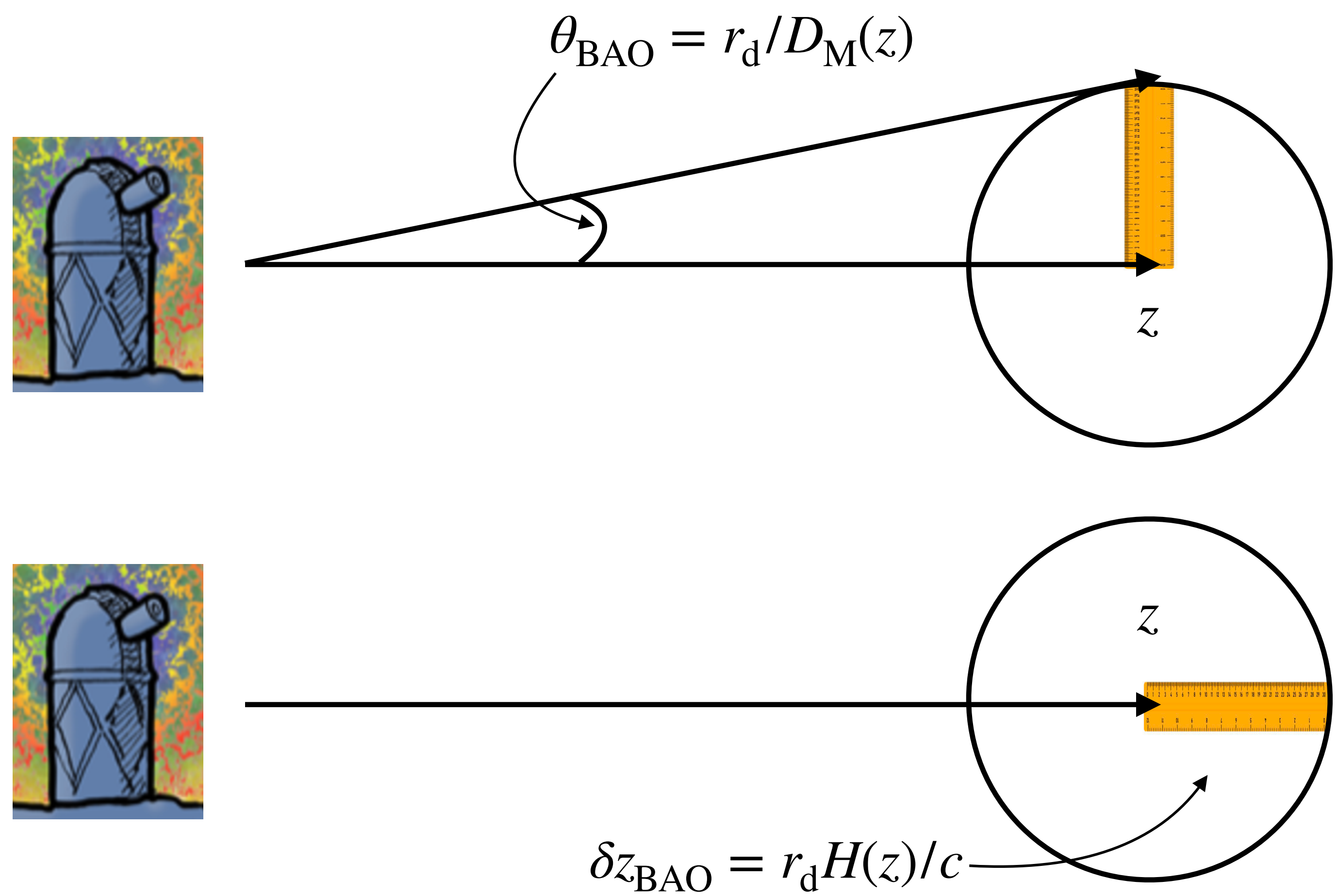


Credit: Daniel Eisenstein

Baryon Acoustic Oscillations (BAO)





$D_{\rm M}(z)$ and $H(z)$ encode expansion history of the Universe

$$D_{\rm M} = \frac{c}{H_0\sqrt{\Omega_{\rm K}}}\sinh\left[\sqrt{\Omega_{\rm K}}\int_0^z \frac{dz'}{H(z')/H_0}\right]$$

$$H^2 = H_0^2\left(\Omega_{\rm m}a^{-3} + \Omega_{\rm K}a^{-2} + \Omega_{\rm r}a^{-4} + \Omega_\Lambda\right)$$

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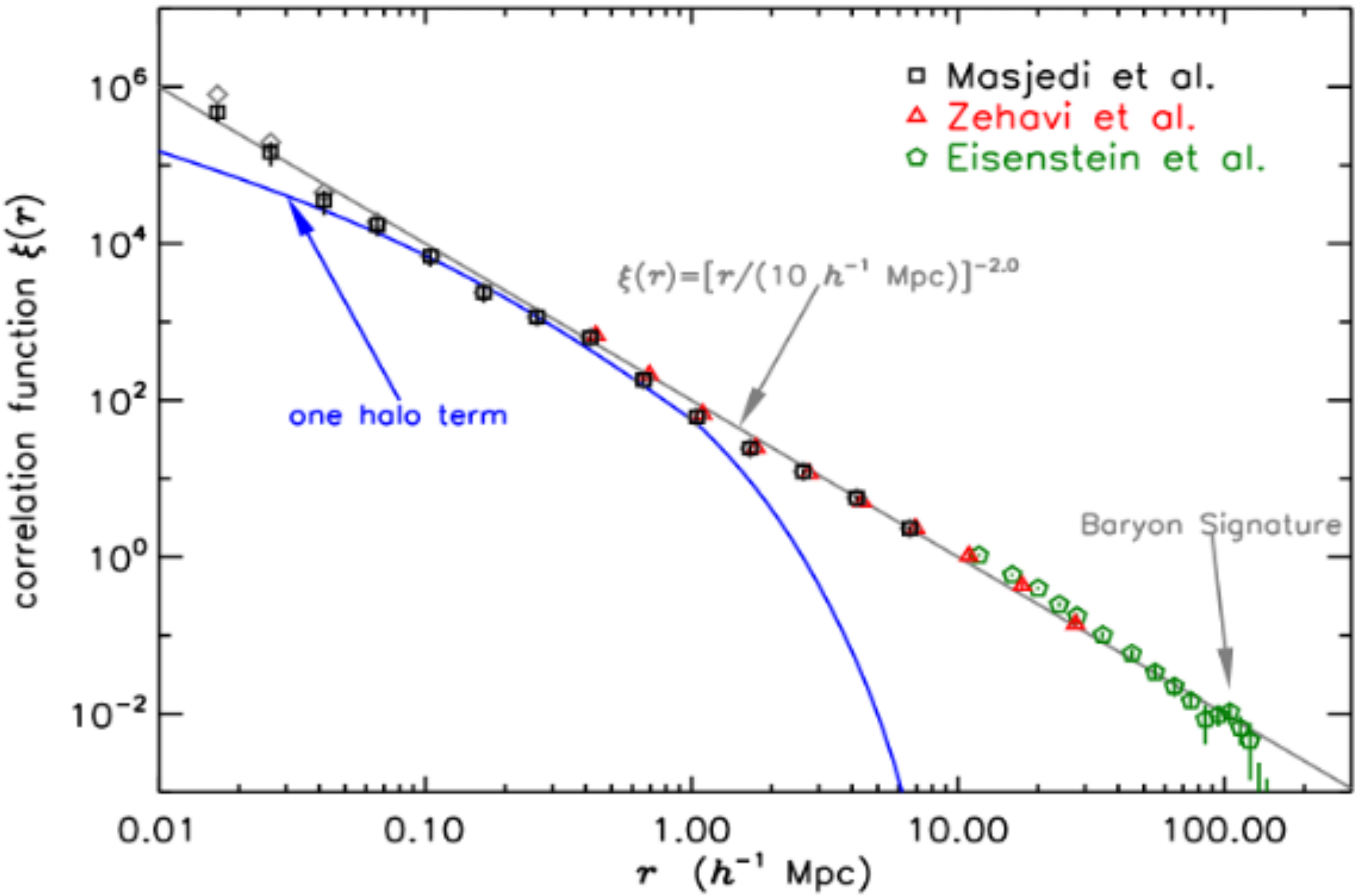
Baryon Acoustic Oscillations (BAO)



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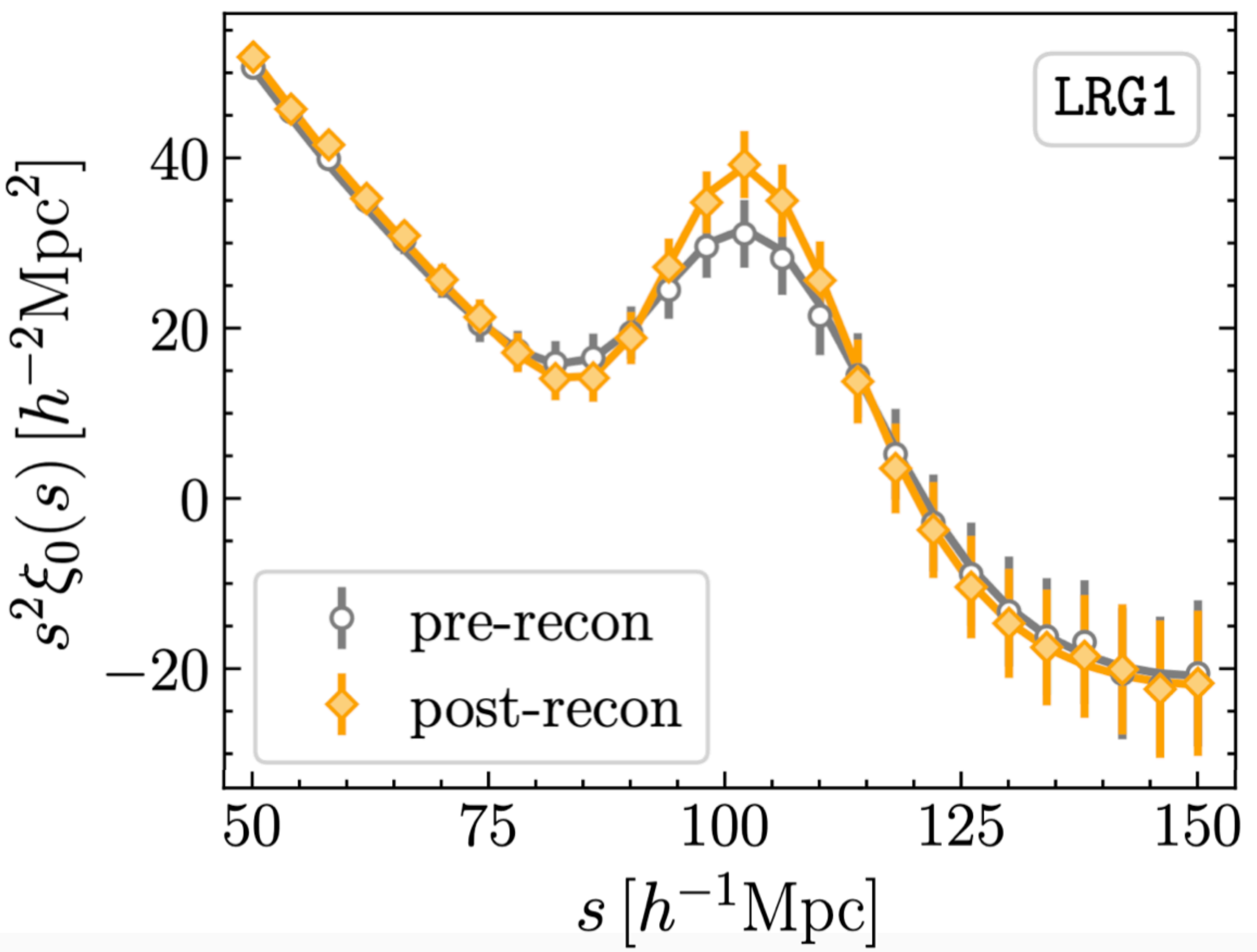
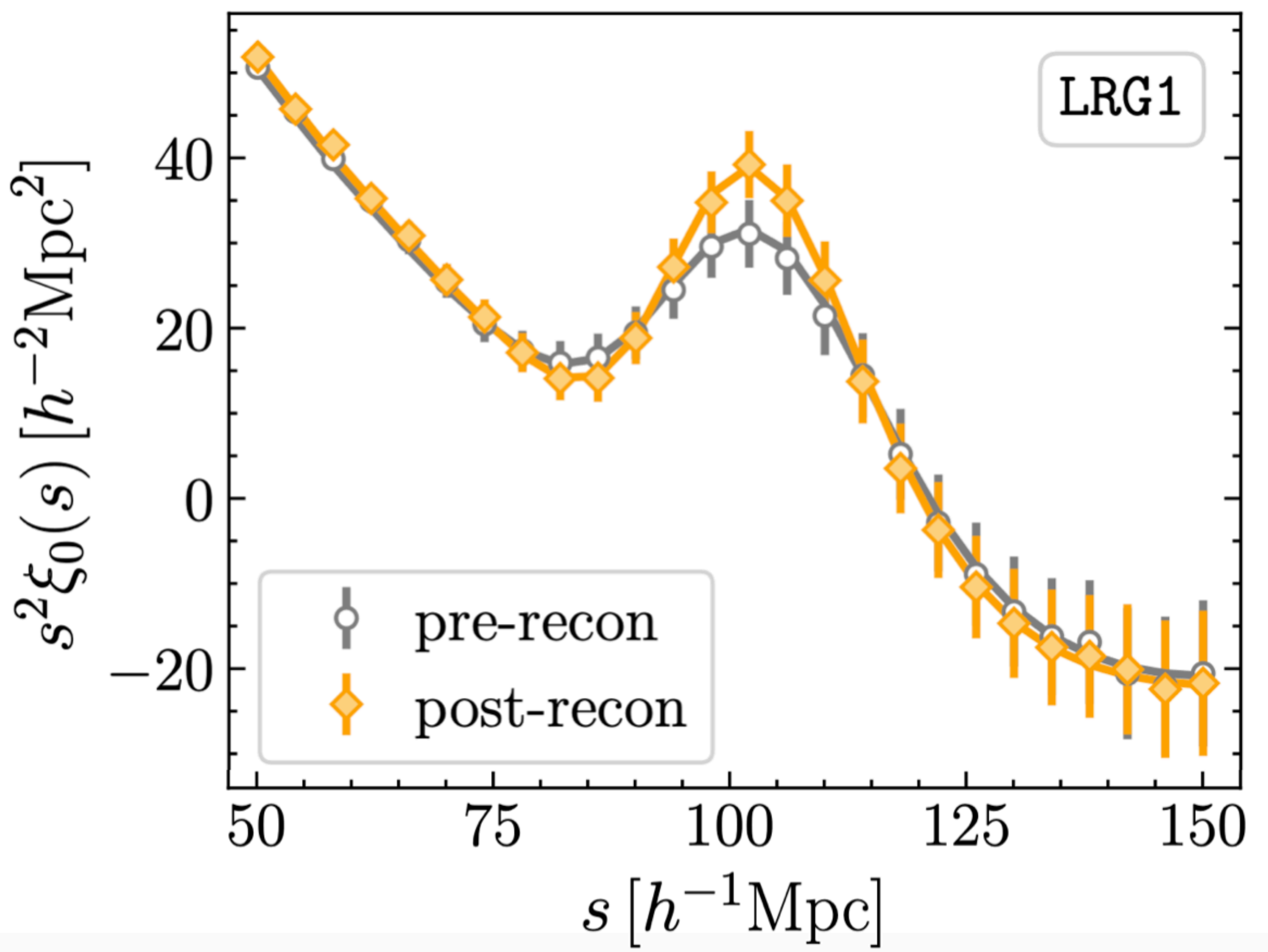
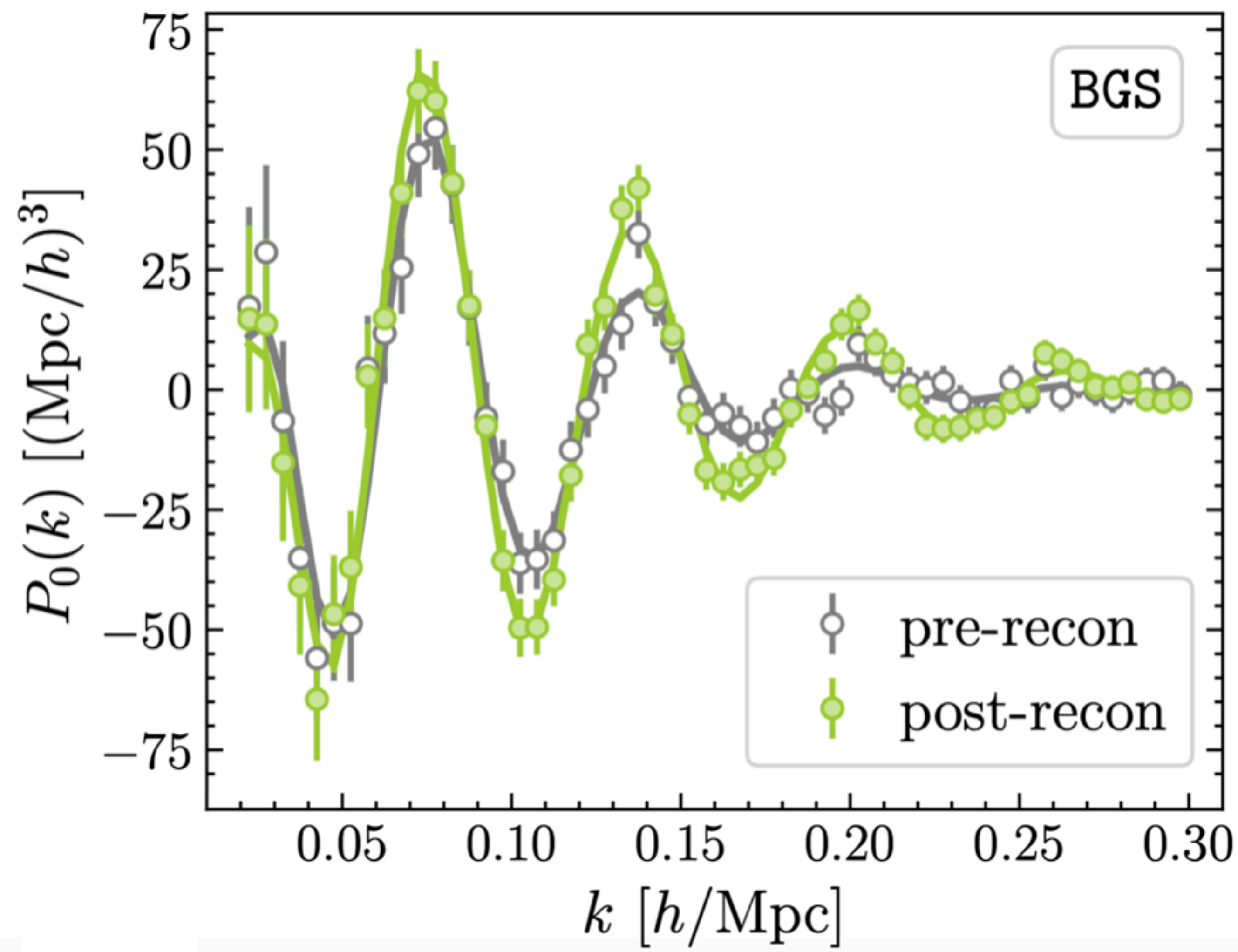


A peak in the correlation function …

Baryon Acoustic Oscillations (BAO)

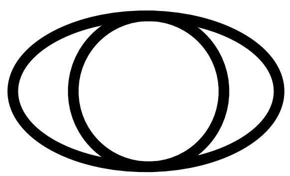




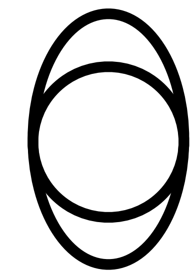
A peak in the correlation function … or wiggles in the power spectrum

BAO measurements

"Scaling parameters":



$$\alpha_{\perp} = \frac{D_{\mathrm{M}}}{r_{\mathrm{d}}} \frac{r_{\mathrm{d}}^{\mathrm{fid}}}{D_{\mathrm{M}}^{\mathrm{fid}}} \quad \text{and} \quad \alpha_{||} = \frac{D_{\mathrm{H}}}{r_{\mathrm{d}}} \frac{r_{\mathrm{d}}^{\mathrm{fid}}}{D_{\mathrm{H}}^{\mathrm{fid}}}$$



perpendicular std ruler size

line-of-sight std ruler size

BAO measurements

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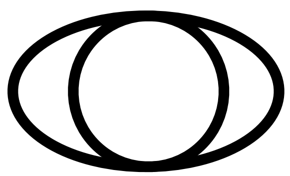

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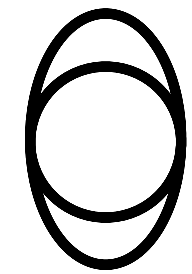
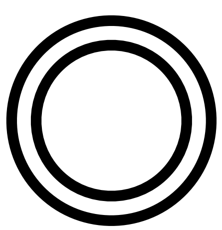
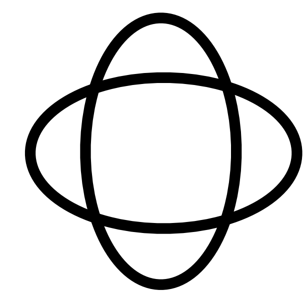


perpendicular std ruler size — line-of-sight std ruler size

OR



$$\alpha_{\mathrm{iso}} = \left(\alpha_{\perp}^{2} \alpha_{||}\right)^{1/3} \quad \text{and} \quad \alpha_{\mathrm{AP}} = \frac{D_{\mathrm{H}}}{D_{\mathrm{M}}} \frac{D_{\mathrm{M}}^{\mathrm{fid}}}{D_{\mathrm{H}}^{\mathrm{fid}}}$$



overall scale of std ruler — anisotropy of std ruler

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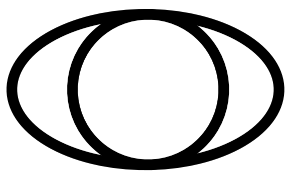
BAO measurements

"Scaling parameters":



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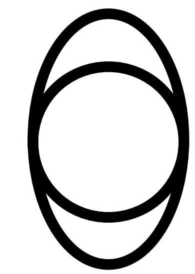
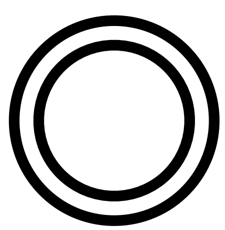


perpendicular std ruler size

line-of-sight std ruler size

OR



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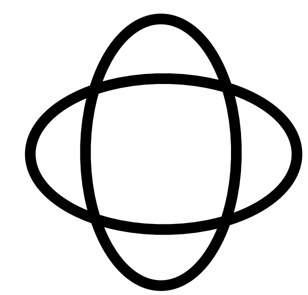


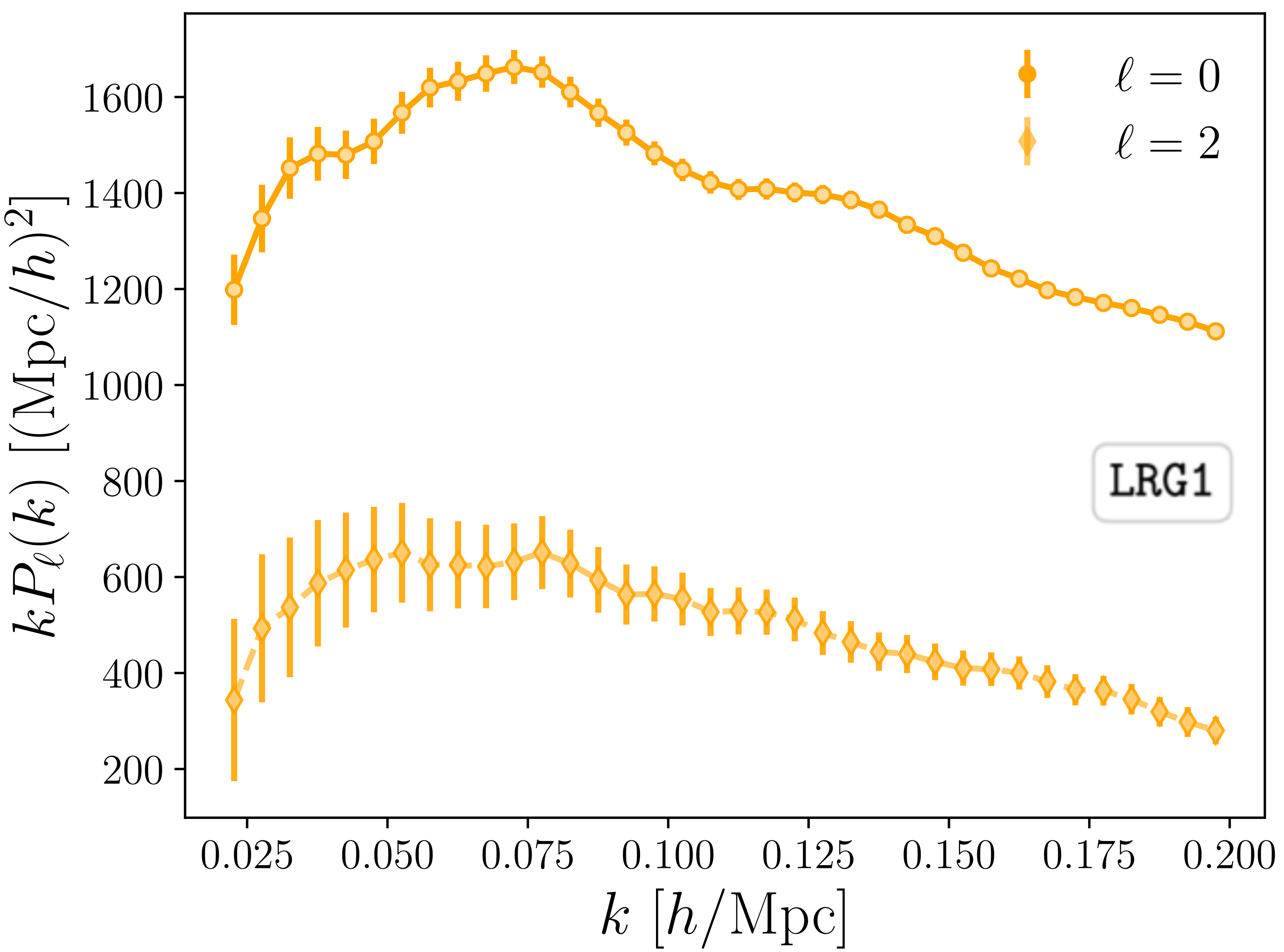
overall scale of std ruler

anisotropy of std ruler

OR

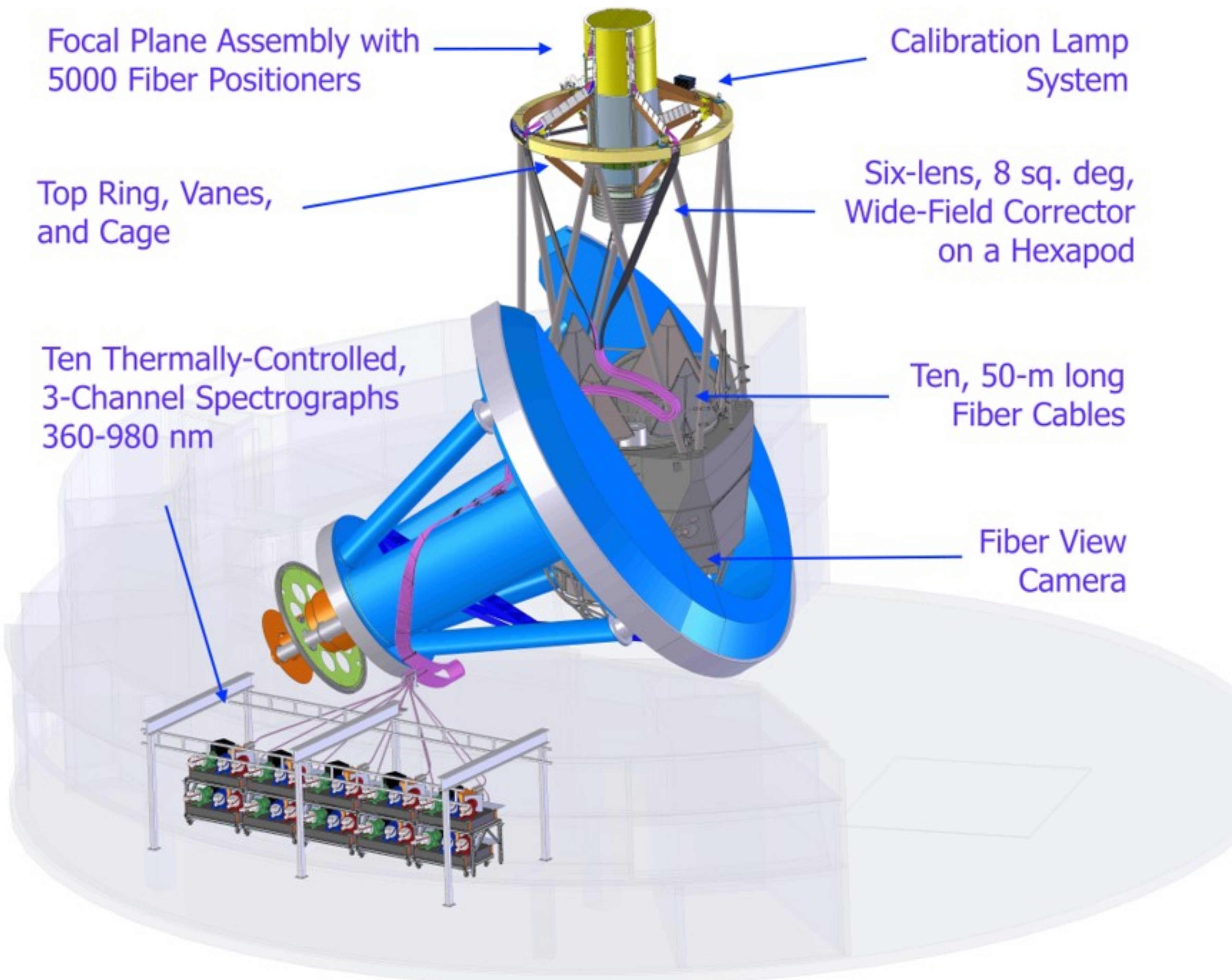
just $\alpha_\mathrm{iso} = \left(\alpha_\perp^2 \alpha_{||}\right)^{1/3}$ (if SNR is low)

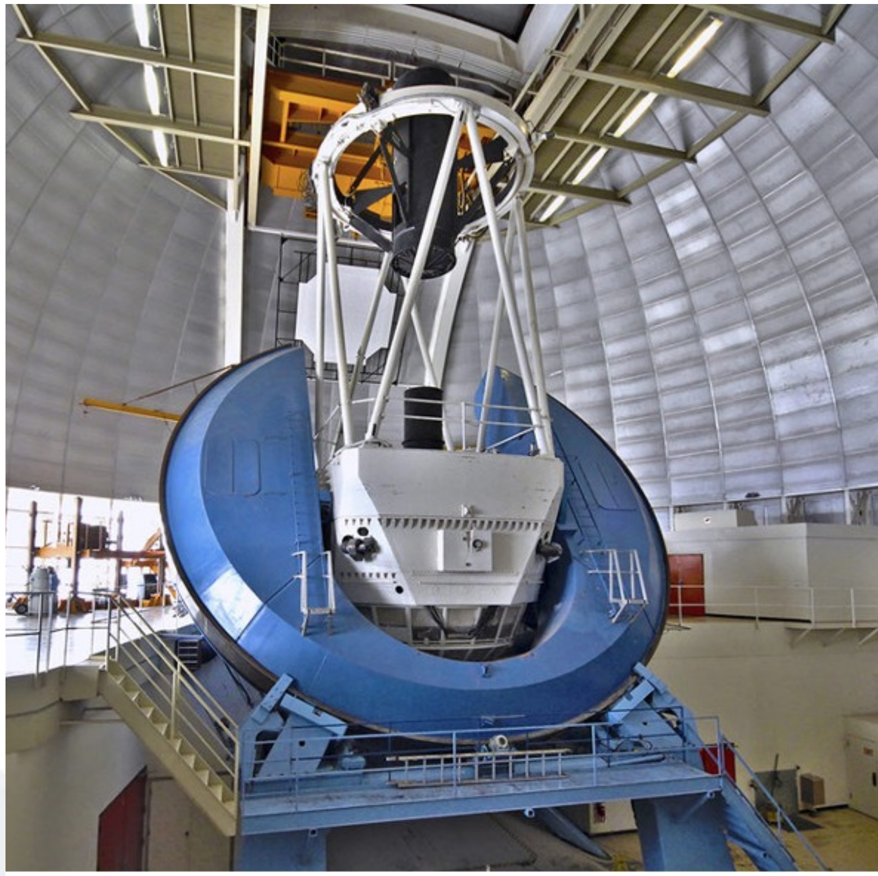
Full shape power spectrum



$$P_{s,g}(\boldsymbol{k}) = P^{PT}_{s,g}(\boldsymbol{k}) + (b + f\mu^2)(b\alpha_0 + f\alpha_2\mu^2 + f\alpha_4\mu^4)k^2 P_{\mathrm{s},b_1^2}(\boldsymbol{k}) + (\mathrm{SN}_0 + \mathrm{SN}_2 k^2\mu^2 + \mathrm{SN}_4 k^4\mu^4),$$

DESI: the instrument





in situ at KPNO

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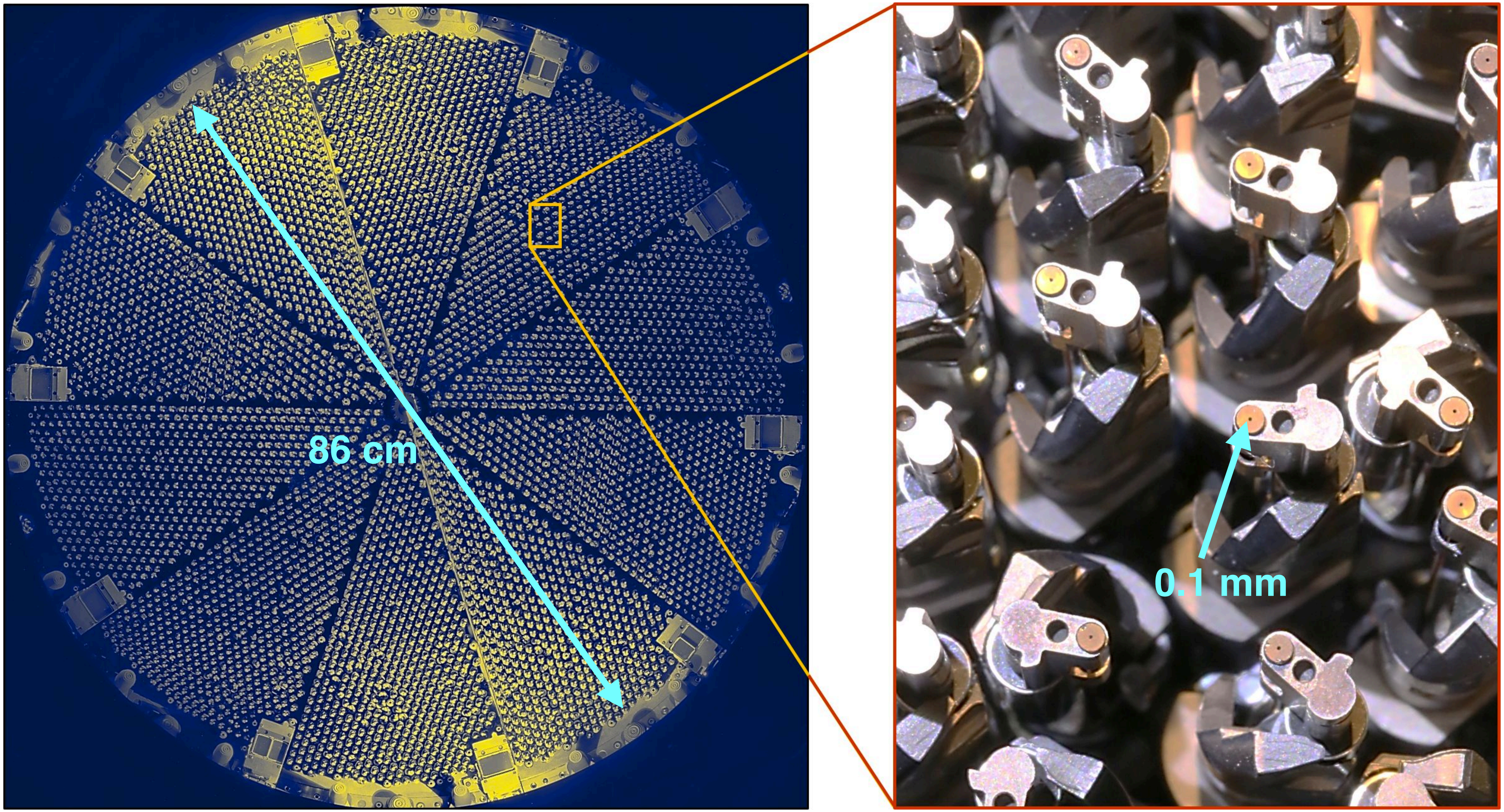
DESI: the instrument



All 5000 fibres can be positioned to accuracy of <5 μm RMS in <120s

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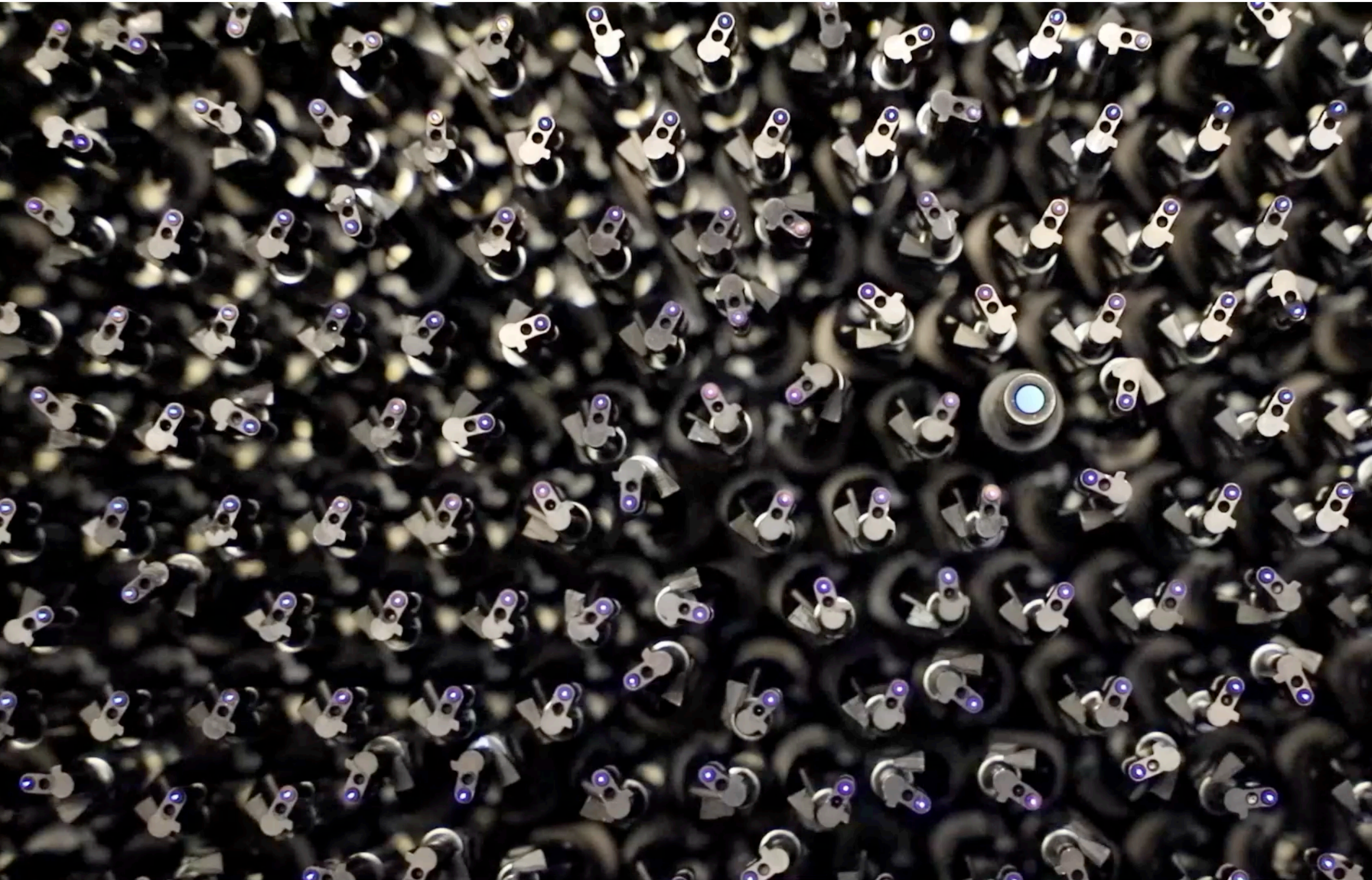
DESI: the instrument



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DESI: the survey

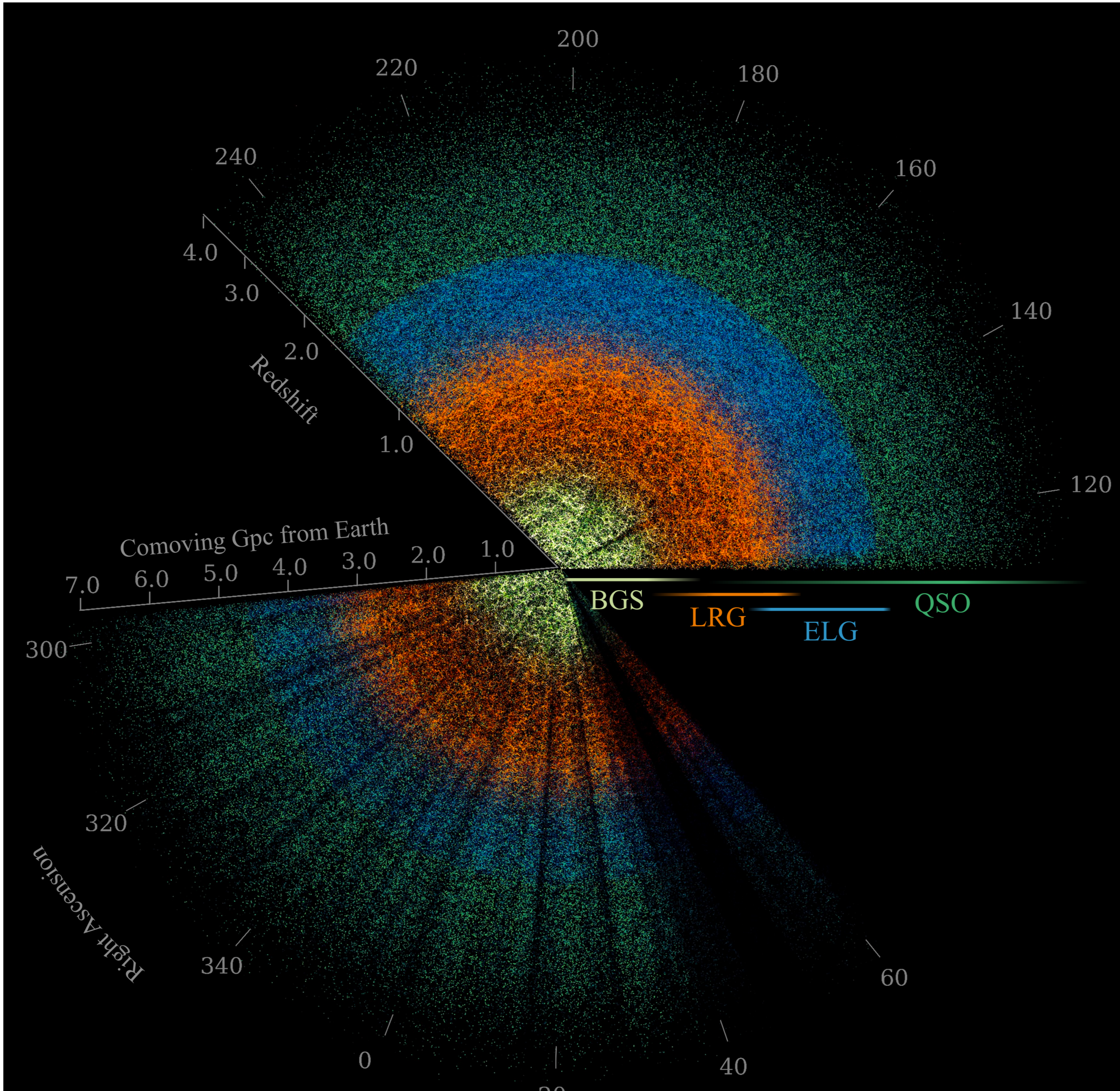
~~30~~ ~~40~~ **50+ million** extragalactic redshifts in ~~5~~ **8** years

Bright Galaxies (BGS)
$0 < z < 0.4$

Luminous Red Galaxies (LRG)
$0.4 < z < 1.1$

Emission Line Galaxies (ELG)
$0.8 < z < 1.6$

Quasars (QSO)
tracers: $0.9 < z < 2.1$
Lya: $z > 2.1$



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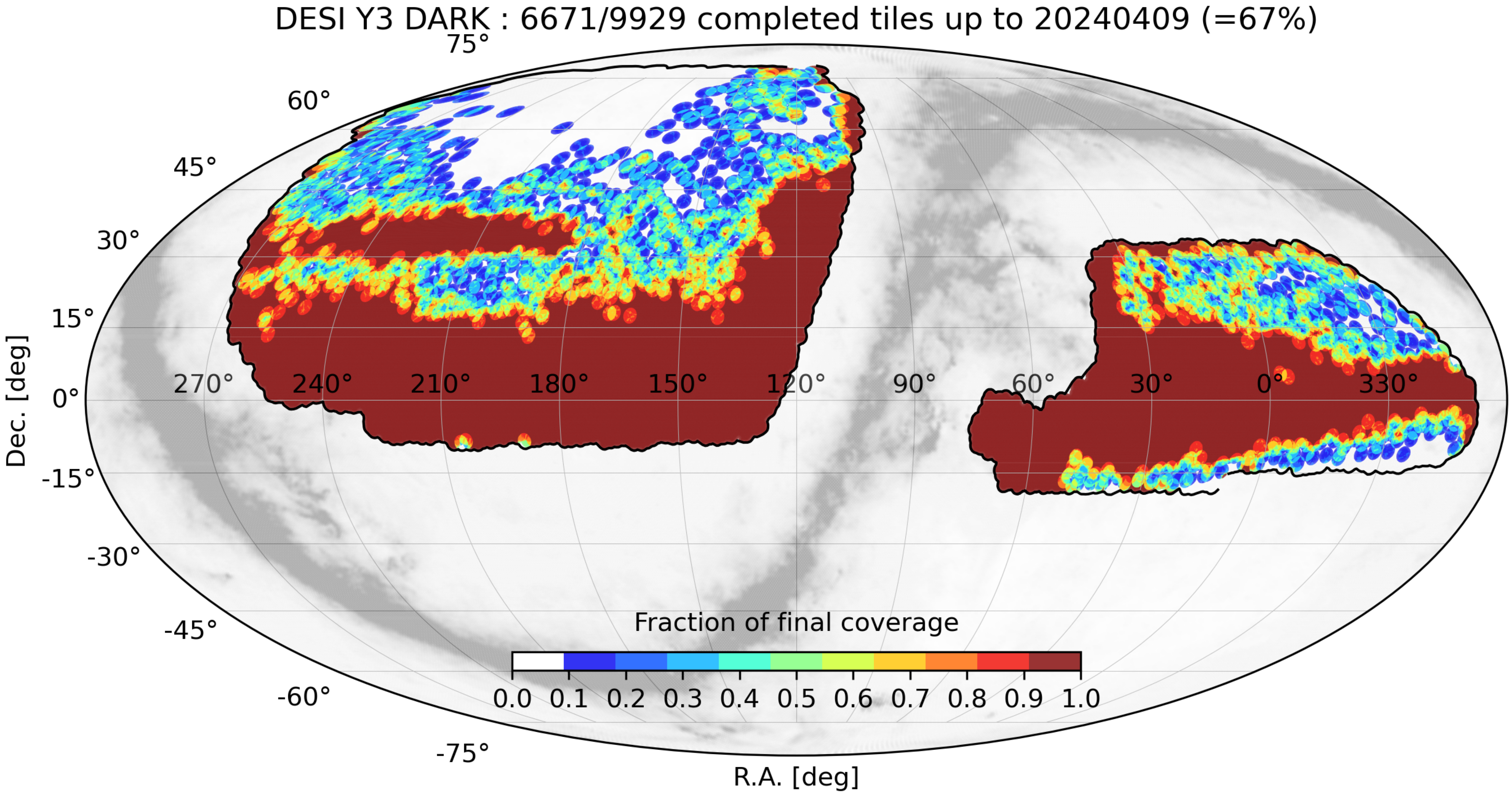
DR2 – 3 years of data

DR2 BAO results: March 2025



>30 million galaxies and quasars, ~14 million of which are used for cosmology

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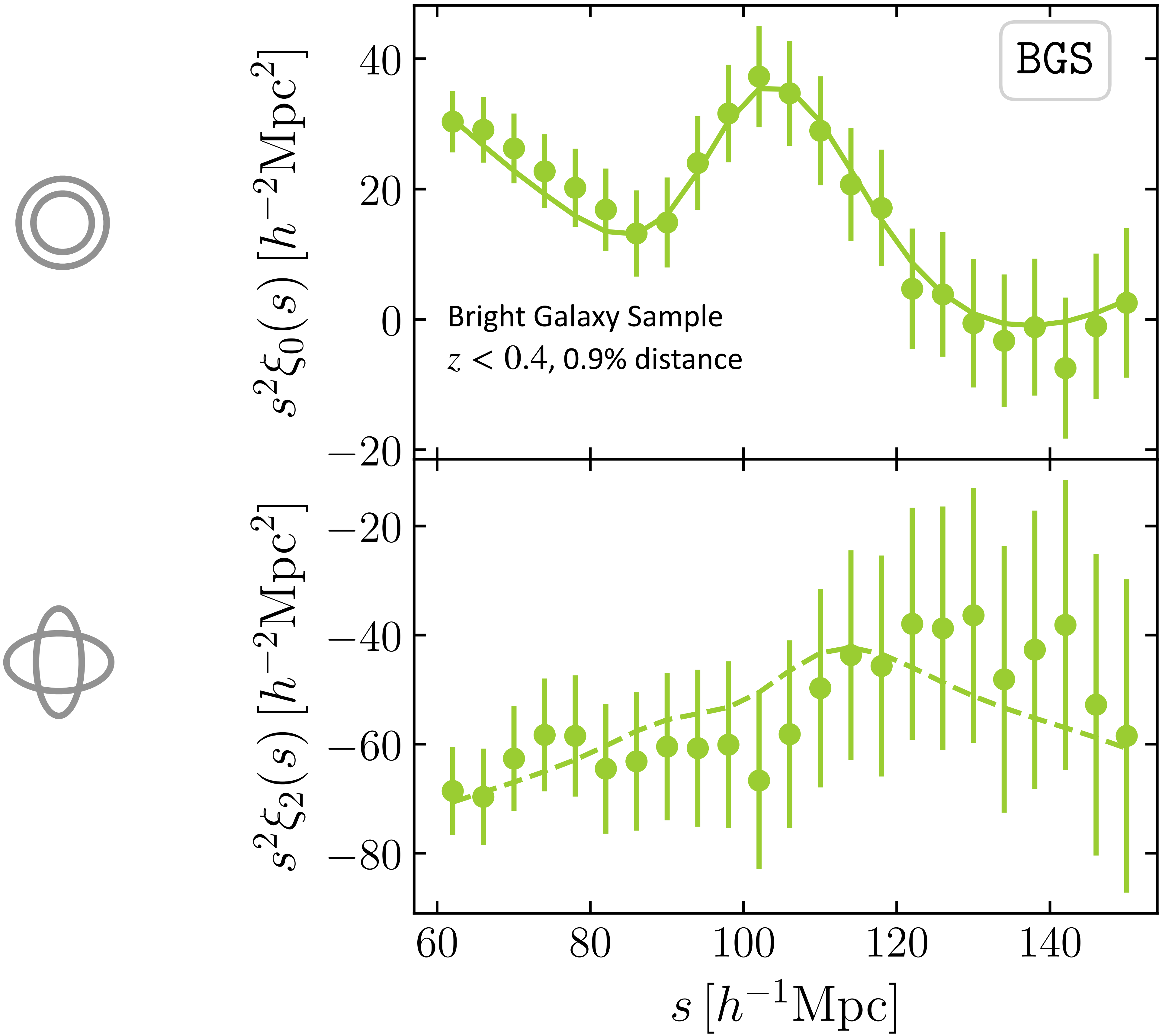
BAO measurements in DR2



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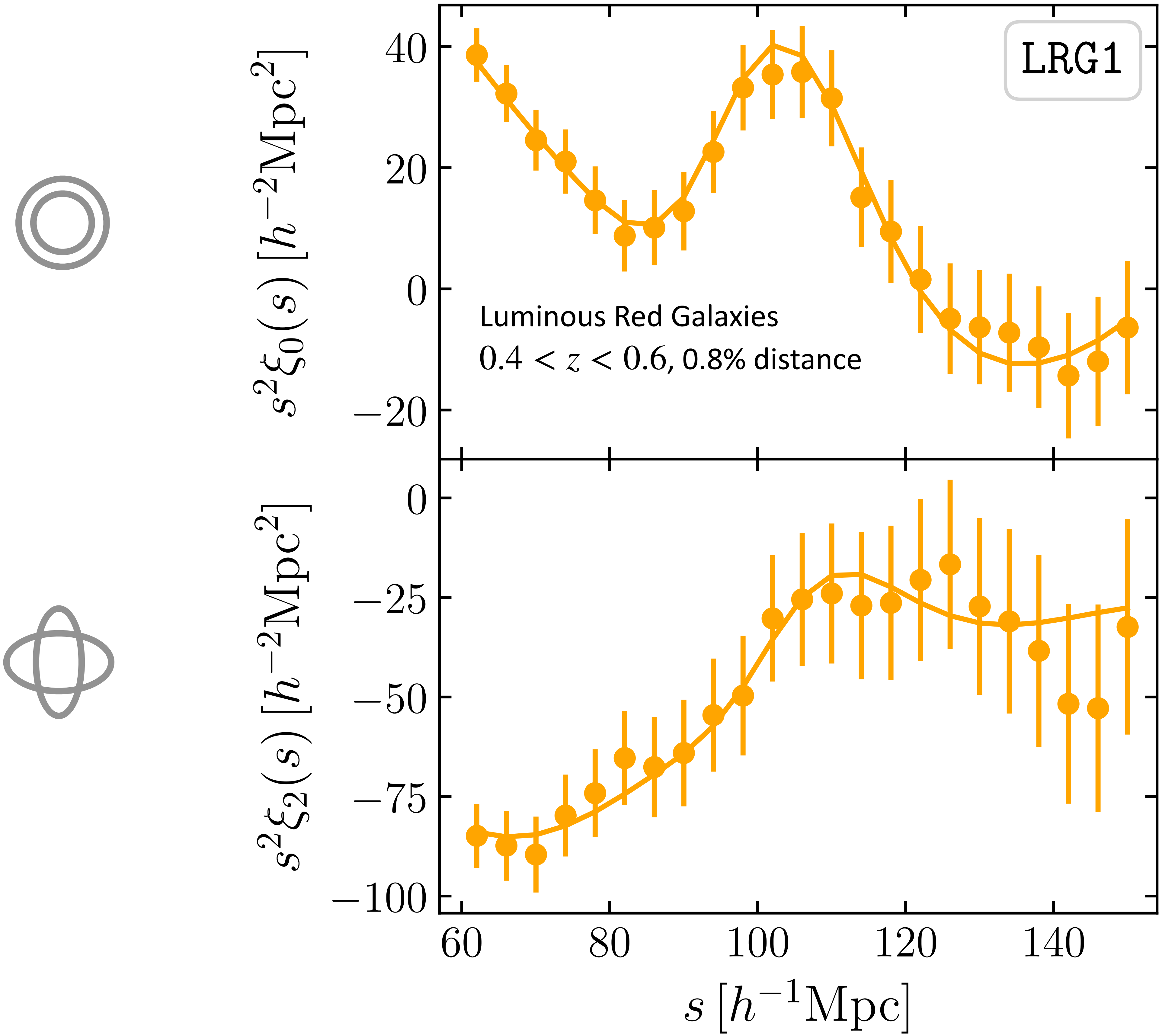

BAO measurements in DR2

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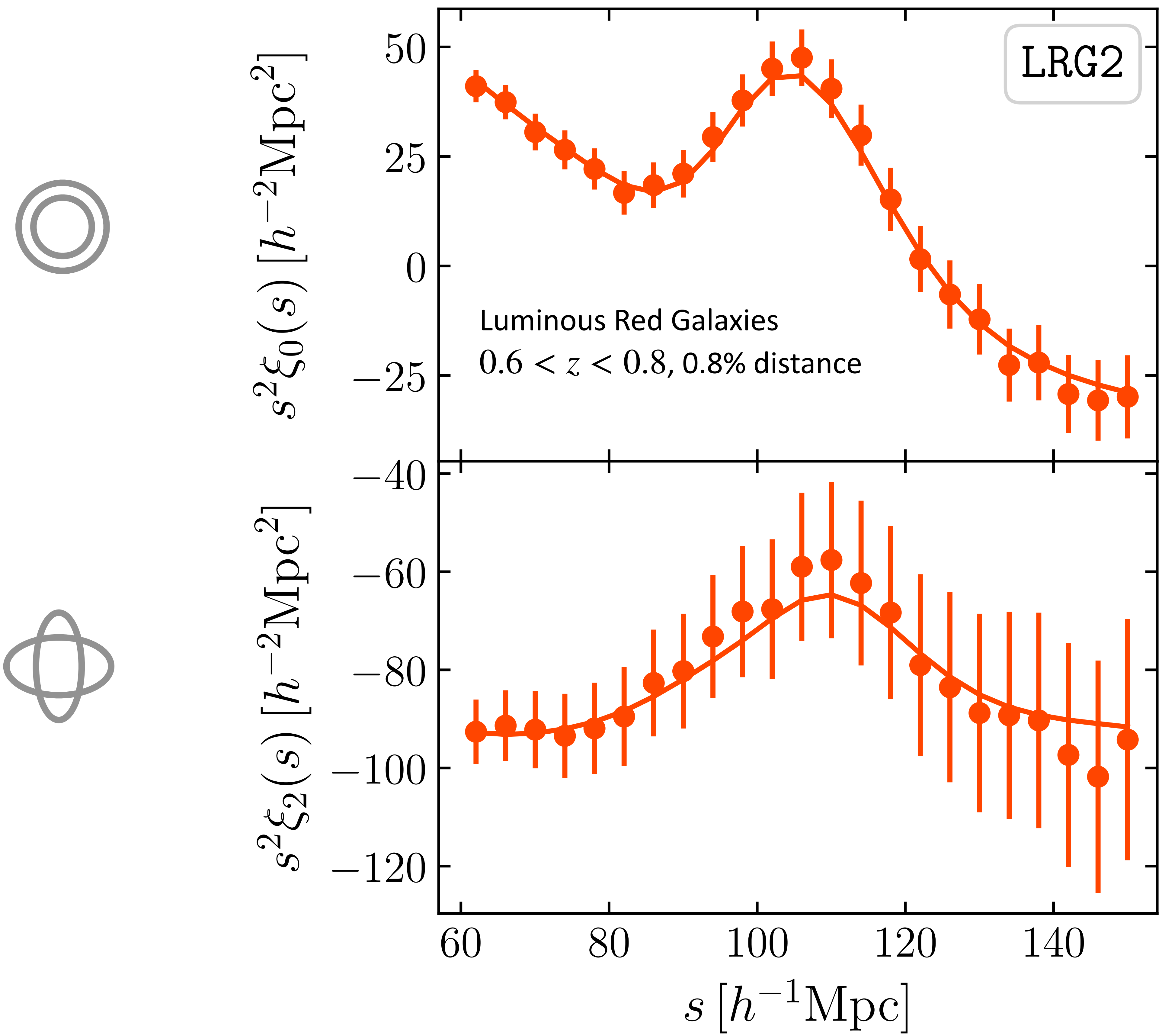

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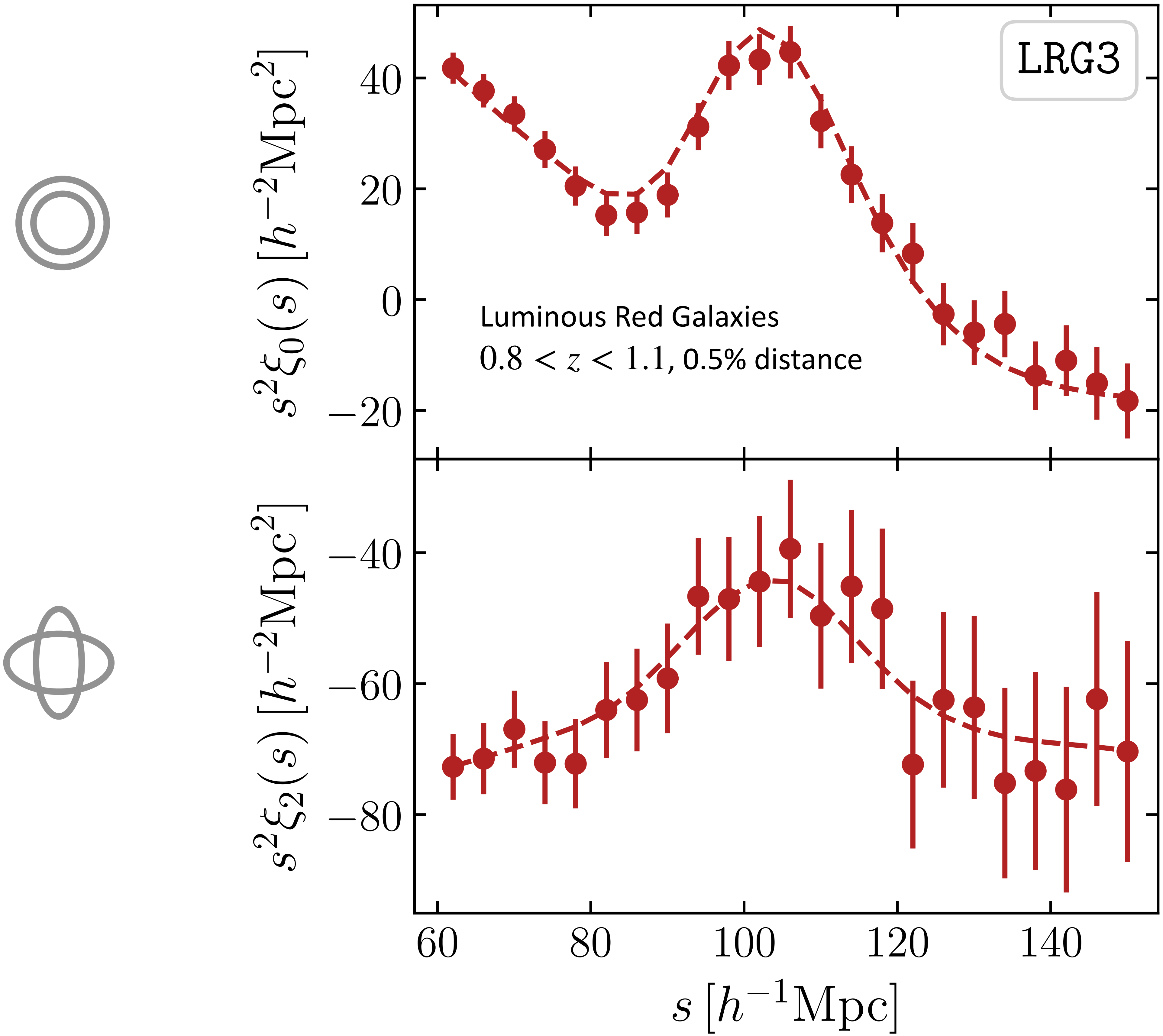

DARK ENERGY SPECTROSCOPIC INSTRUMENT

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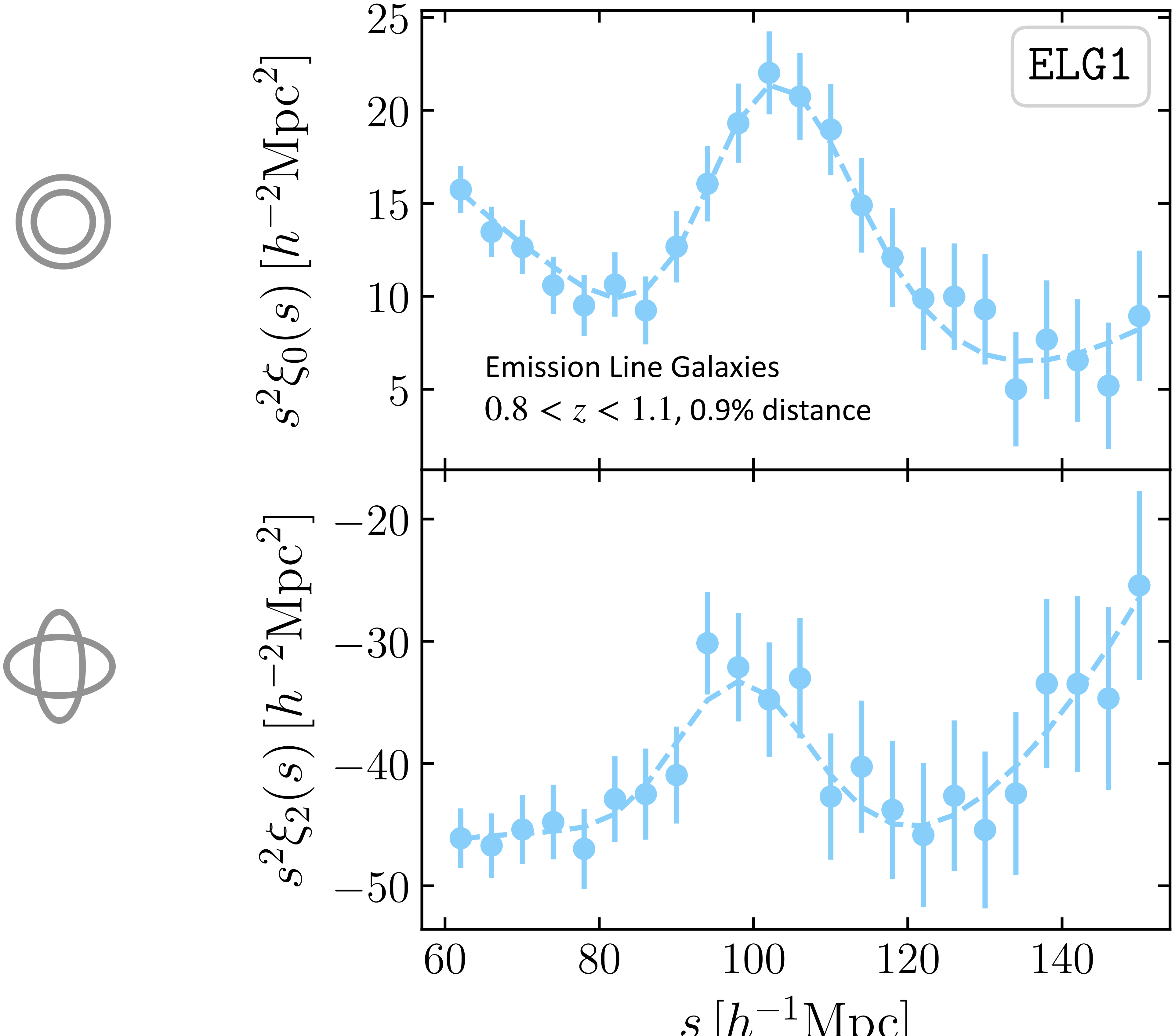
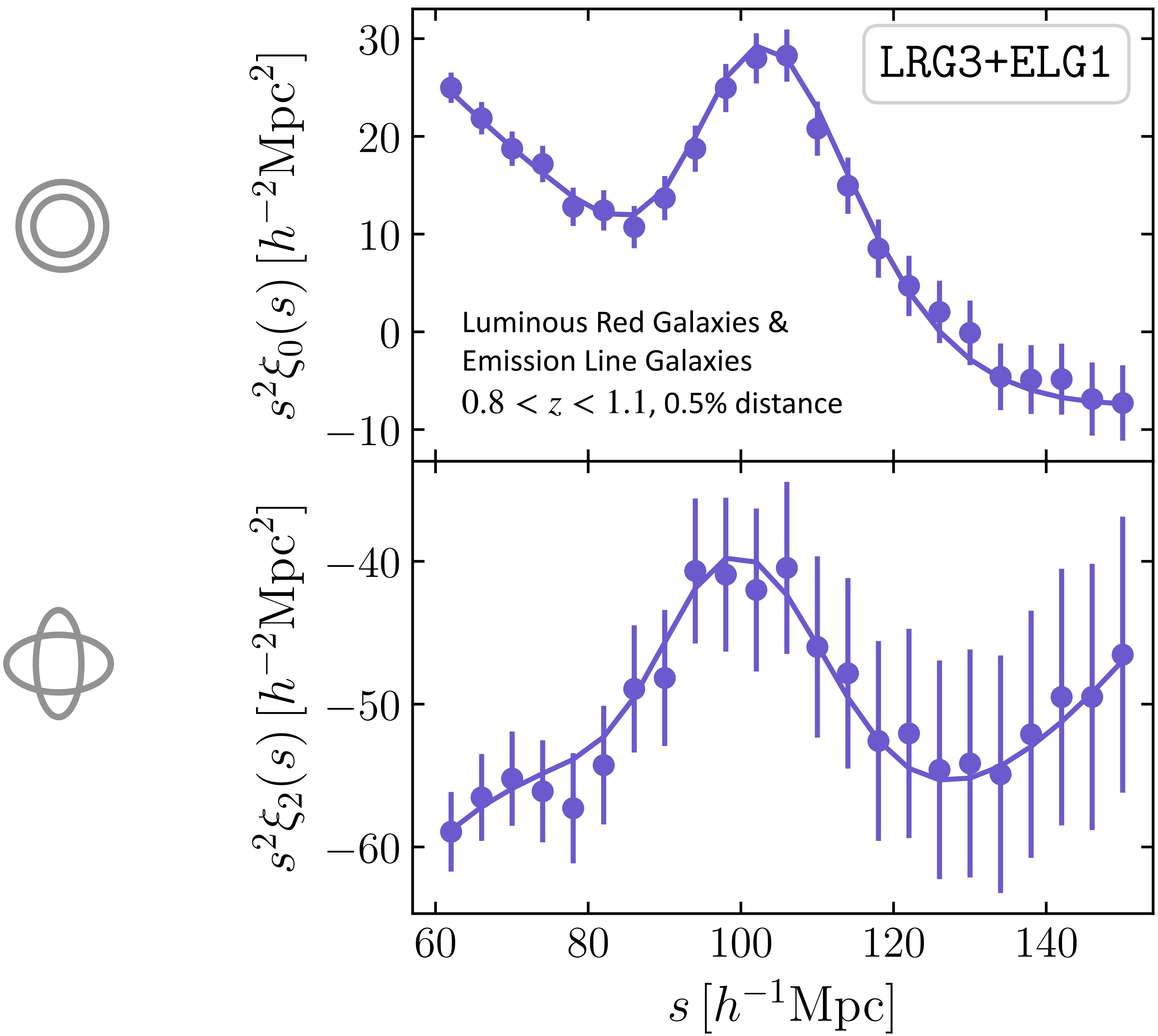
BAO measurements in DR2



BAO measurements in DR2



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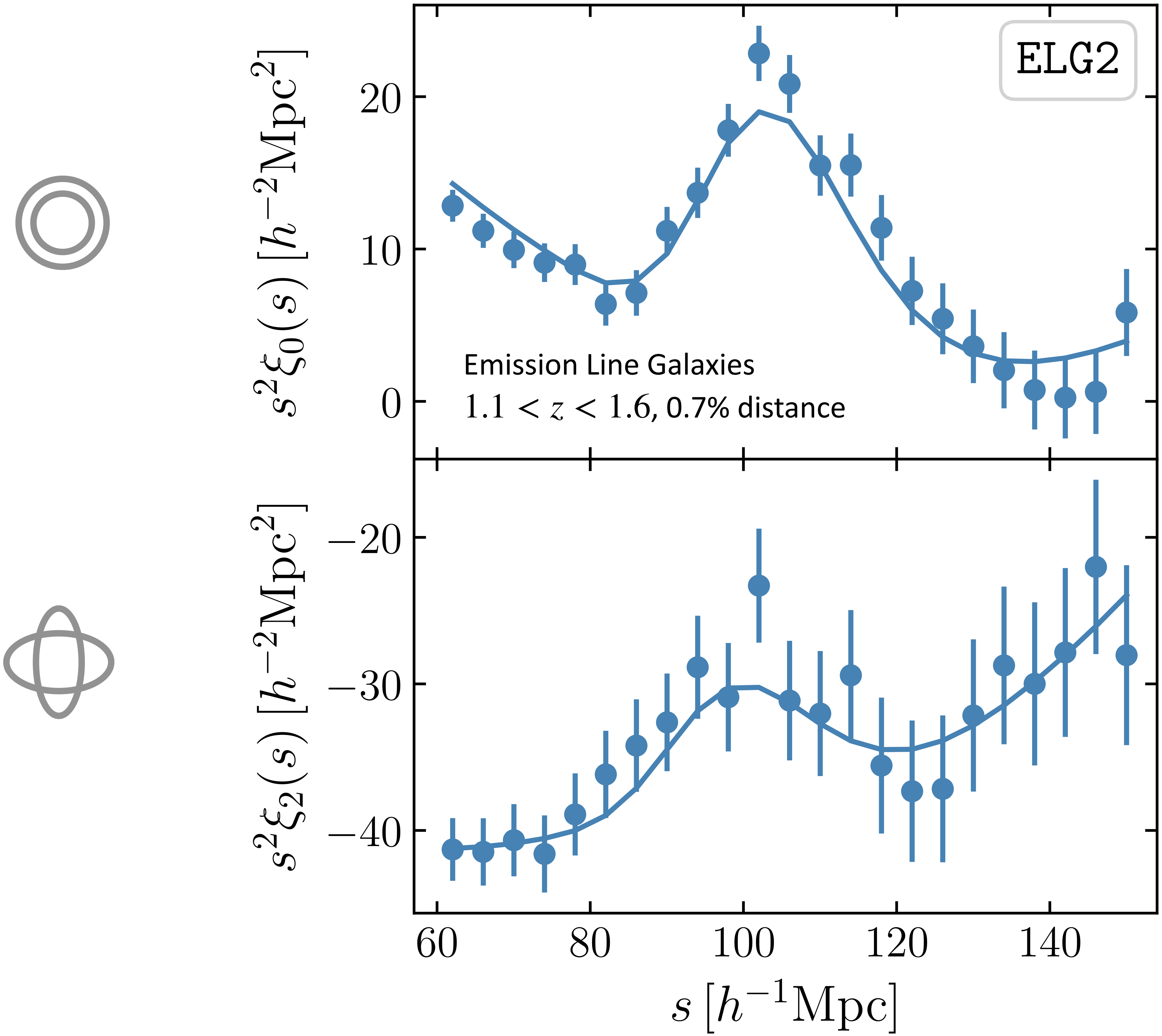

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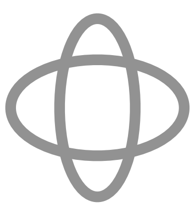

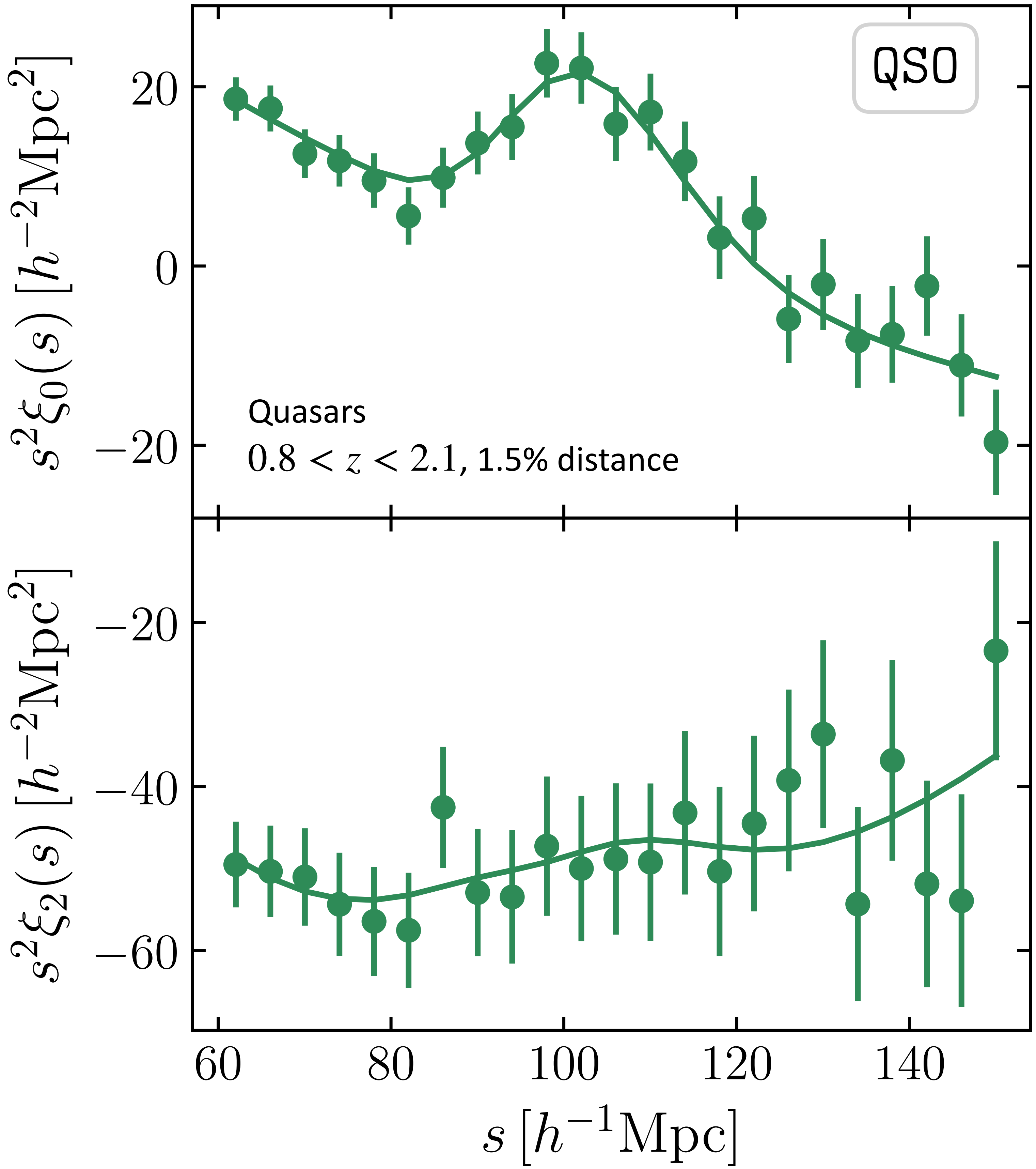

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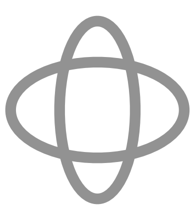

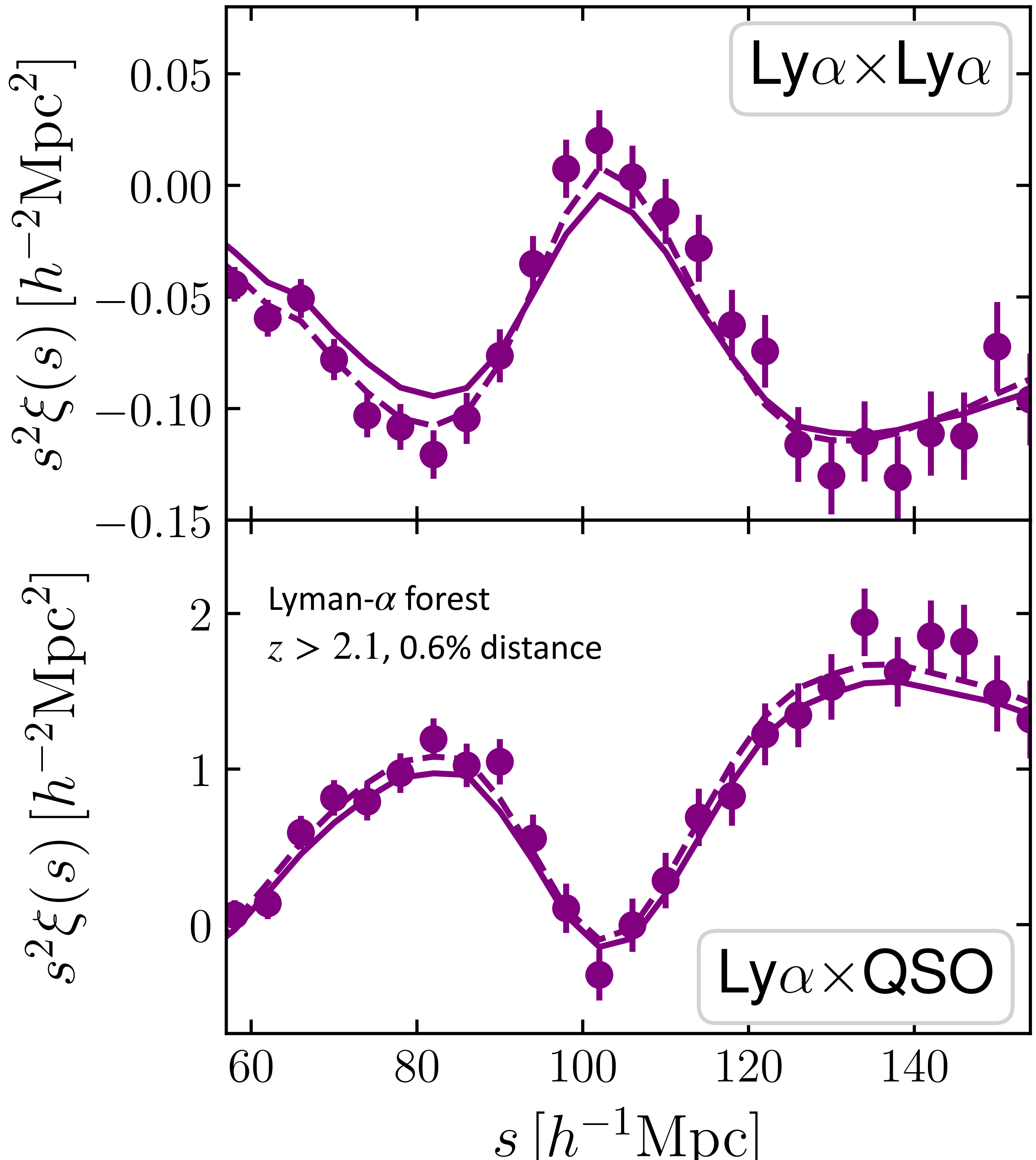

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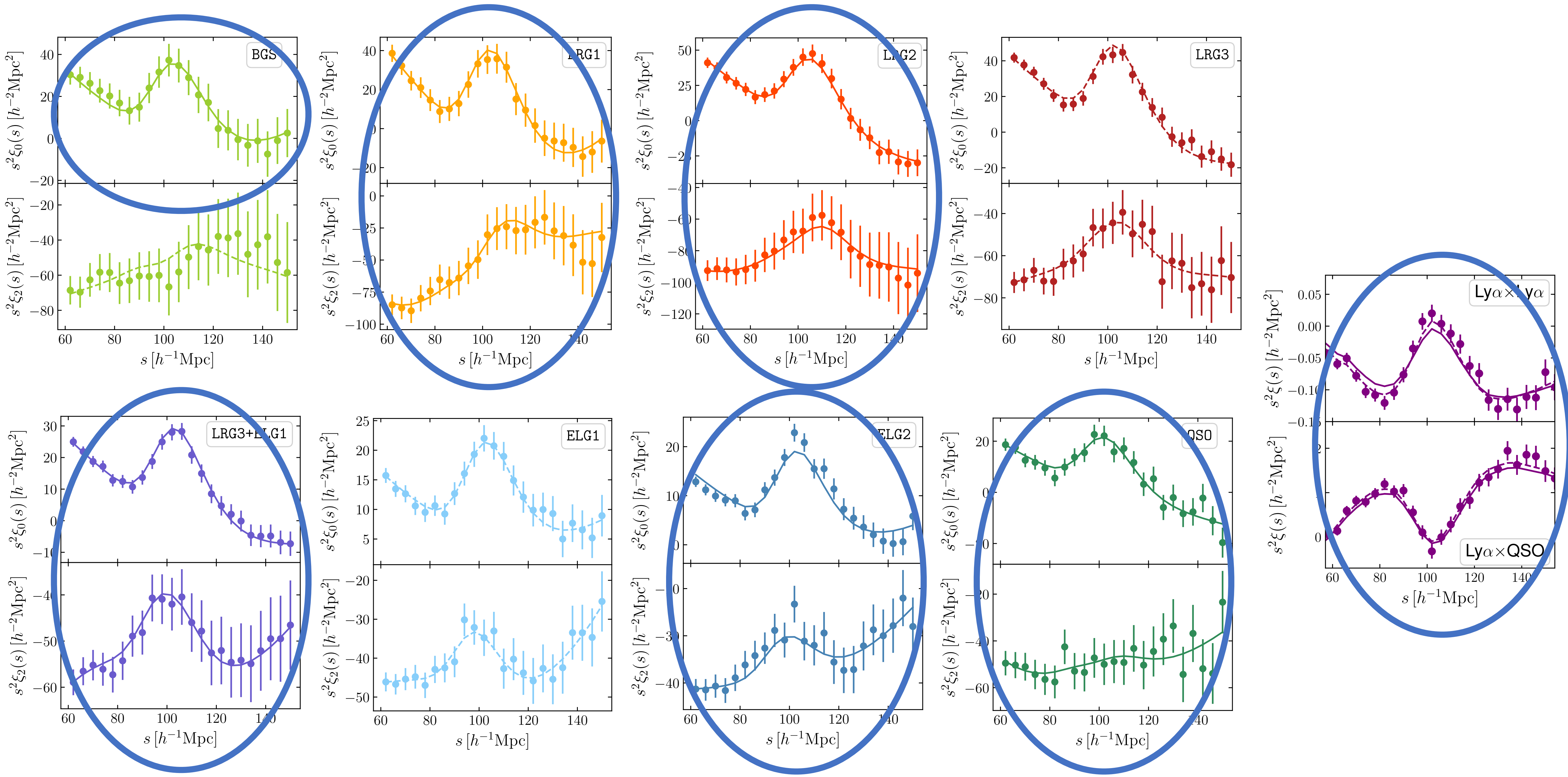
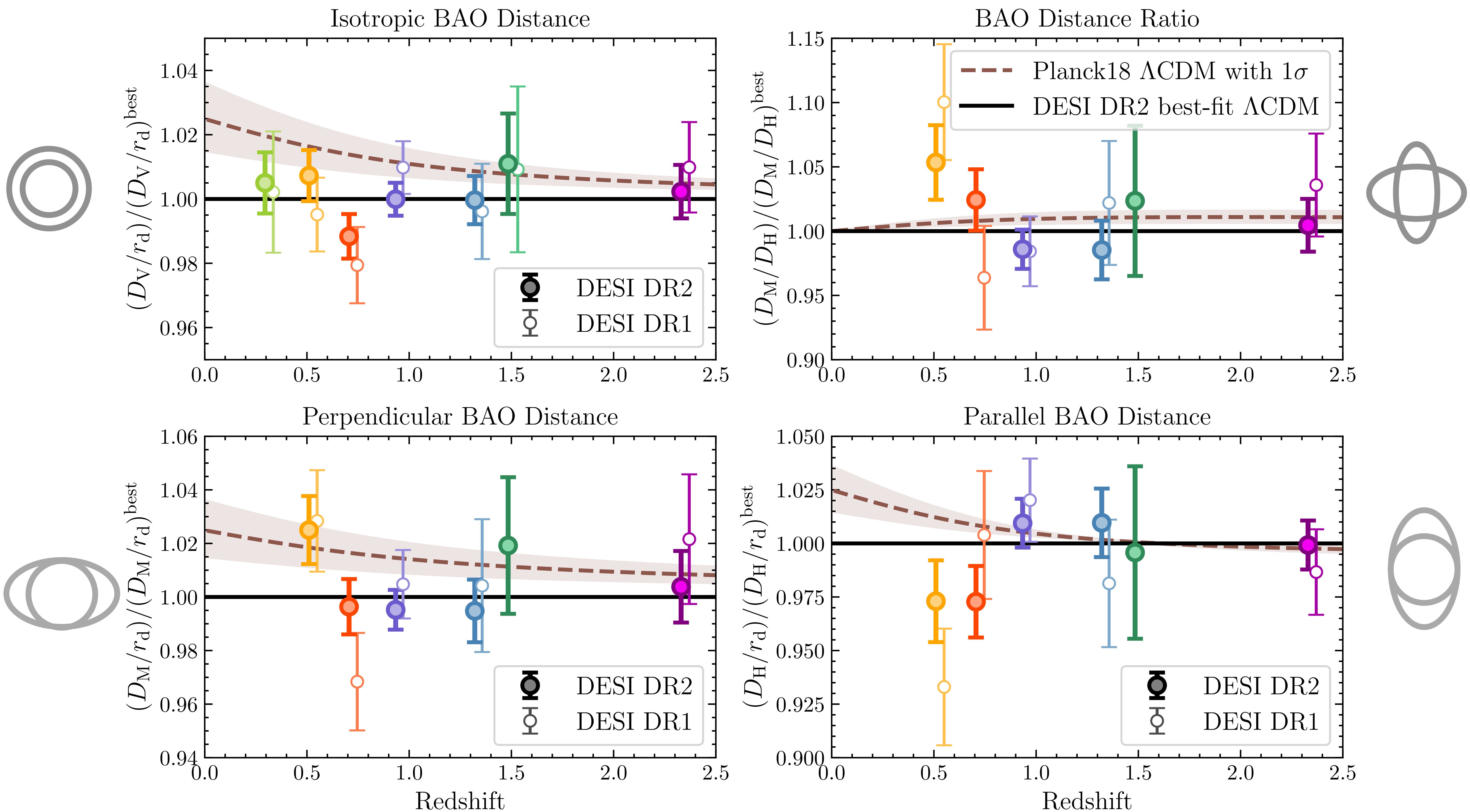
BAO measurements in DR2



= used for cosmology inference

Distances from BAO



Flat ΛCDM

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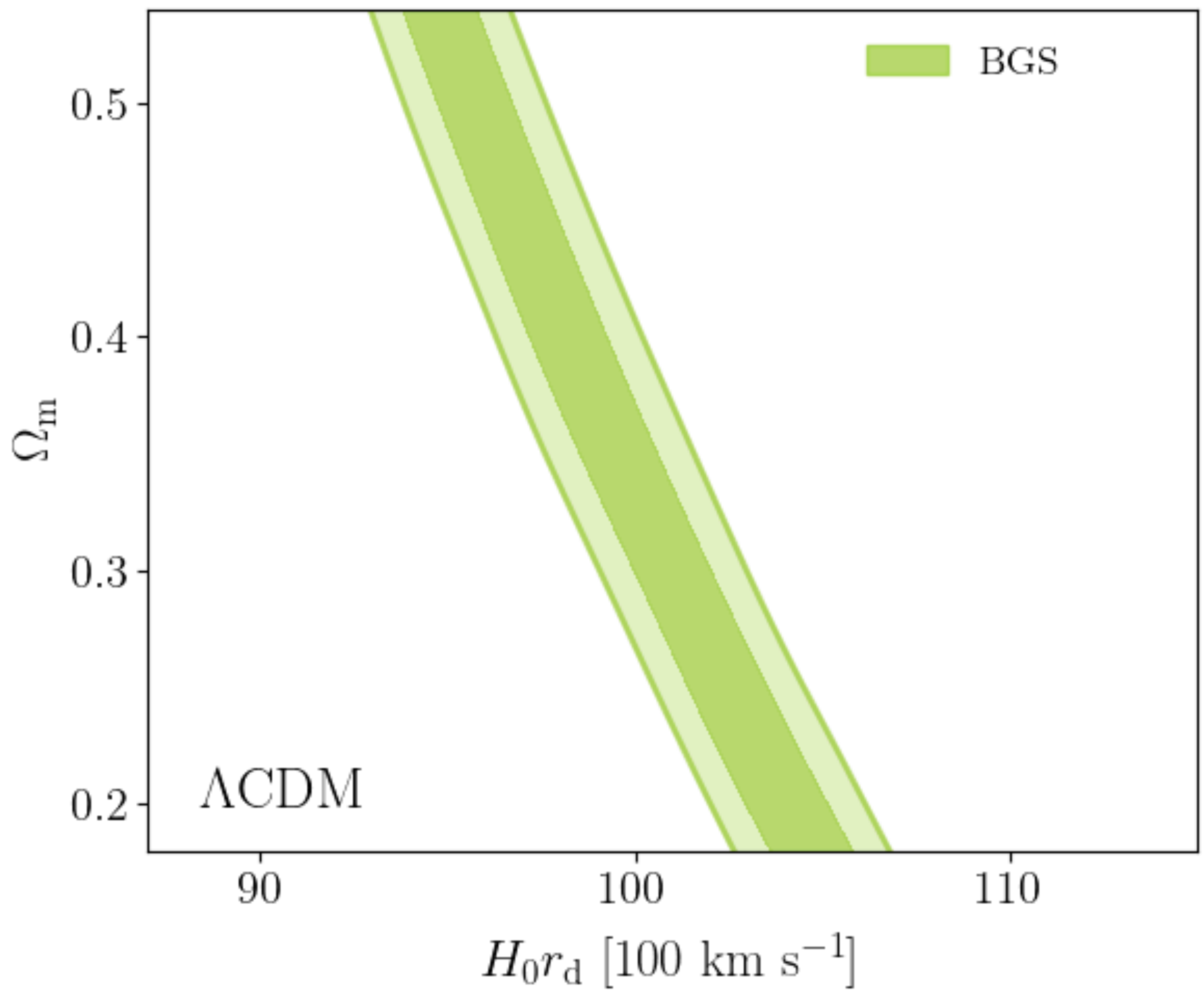
Flat ΛCDM model



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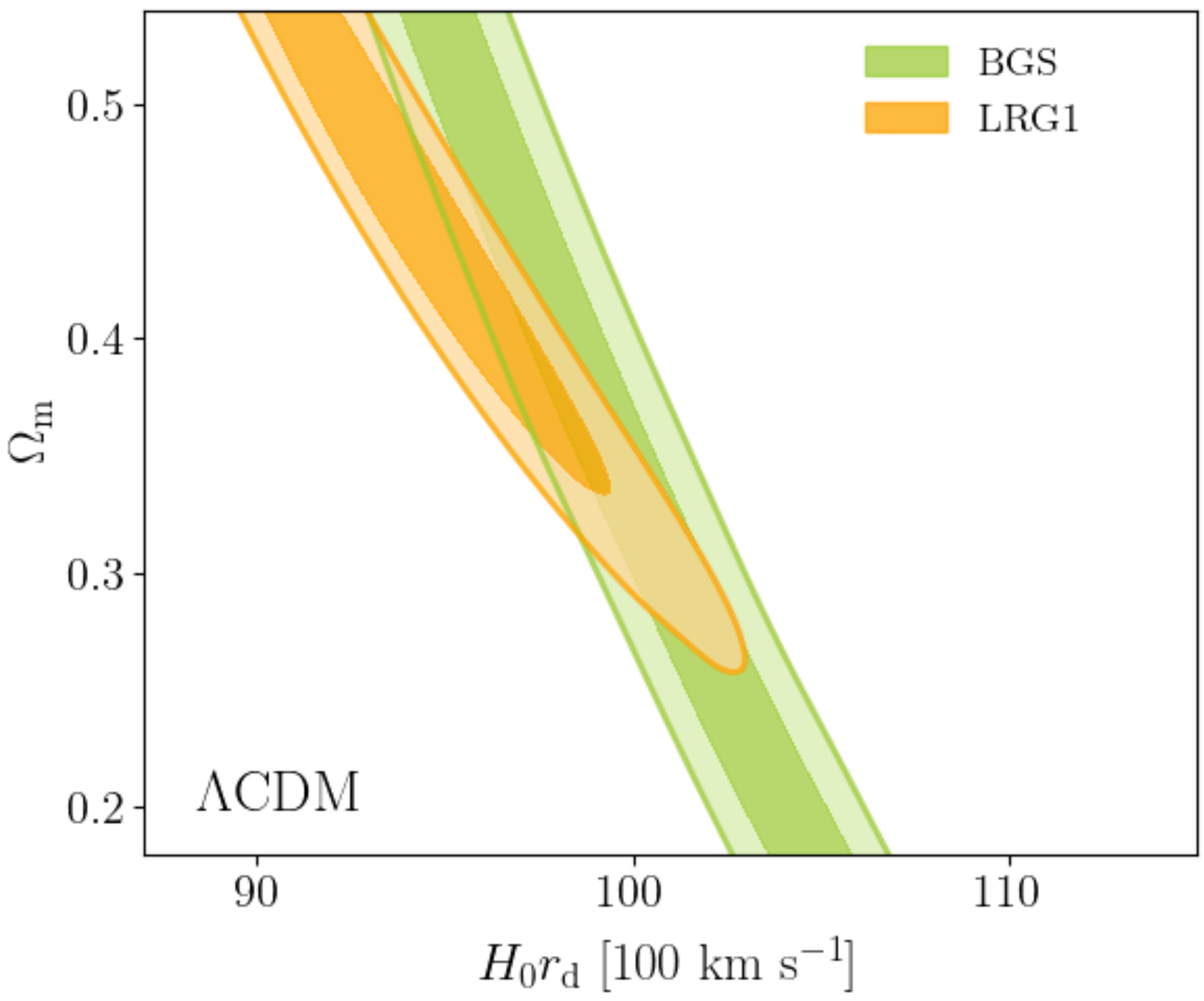

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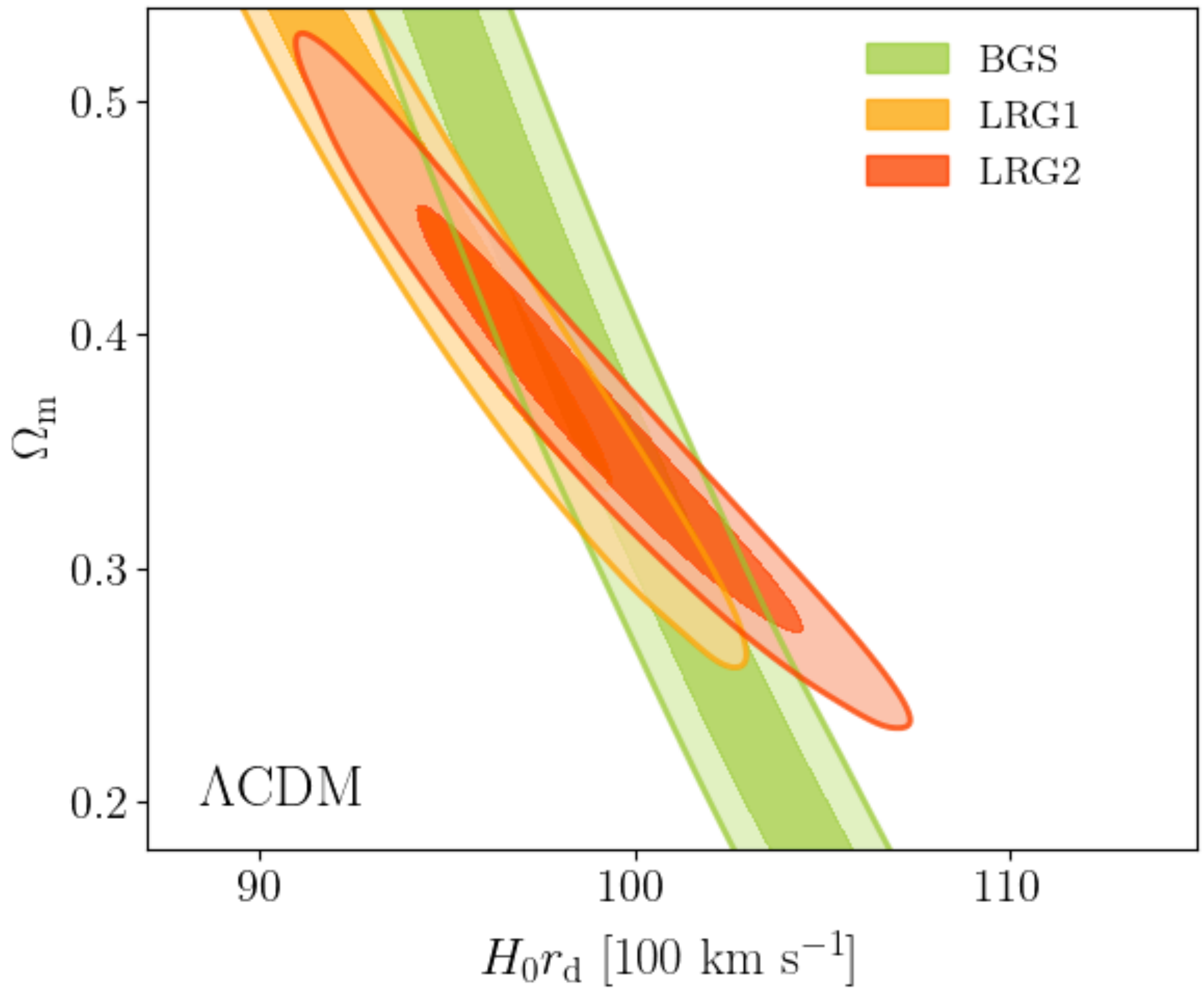

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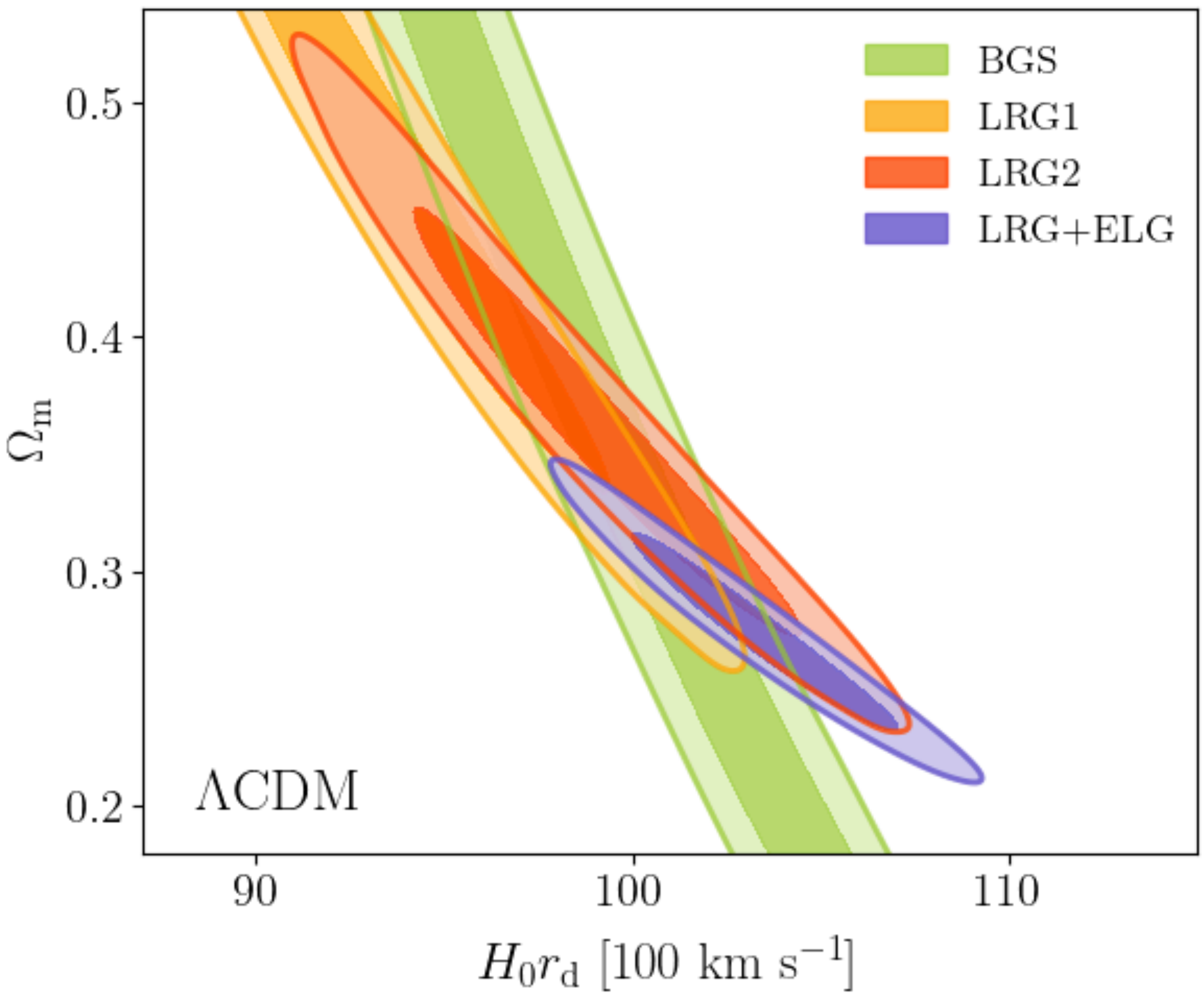

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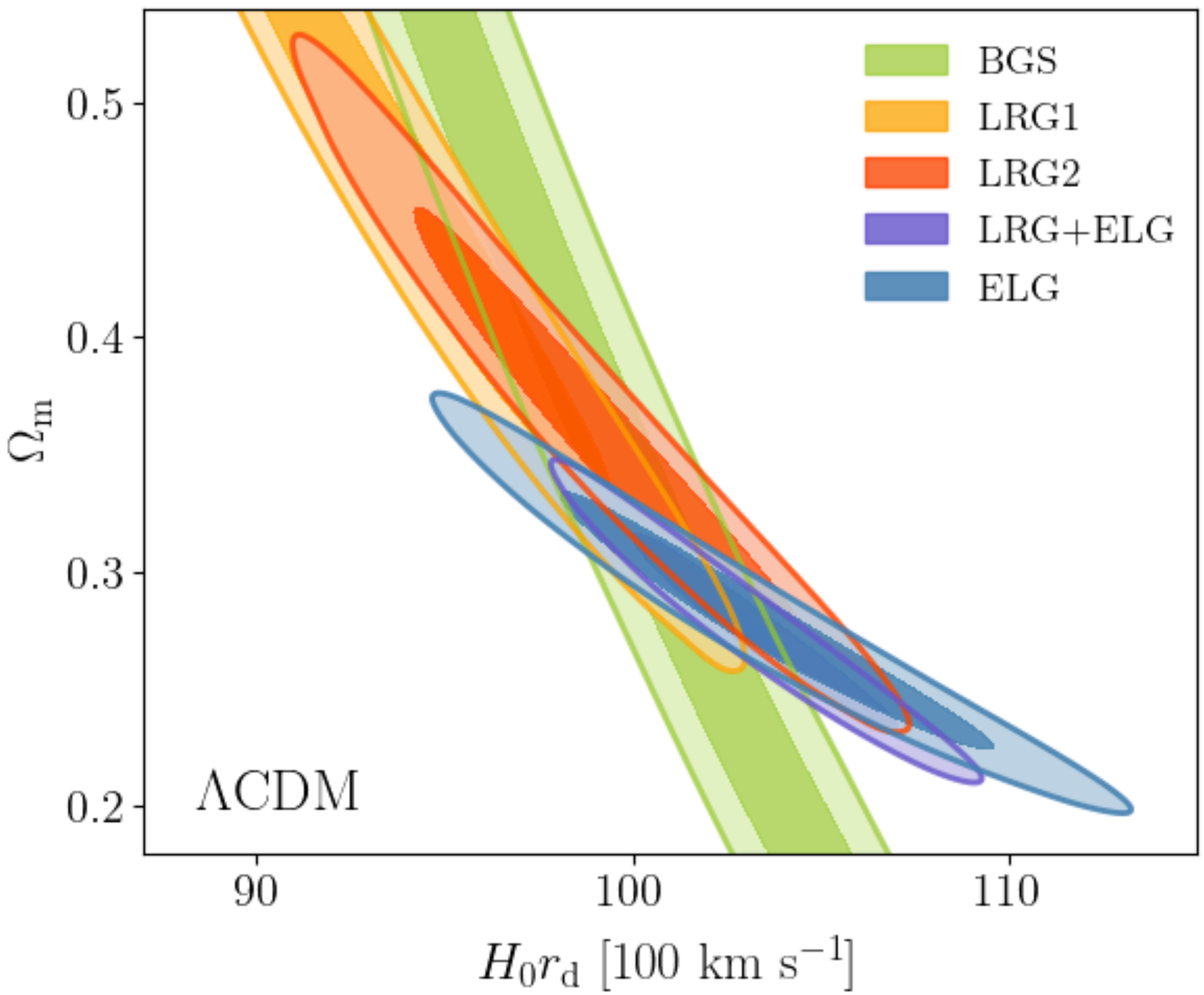

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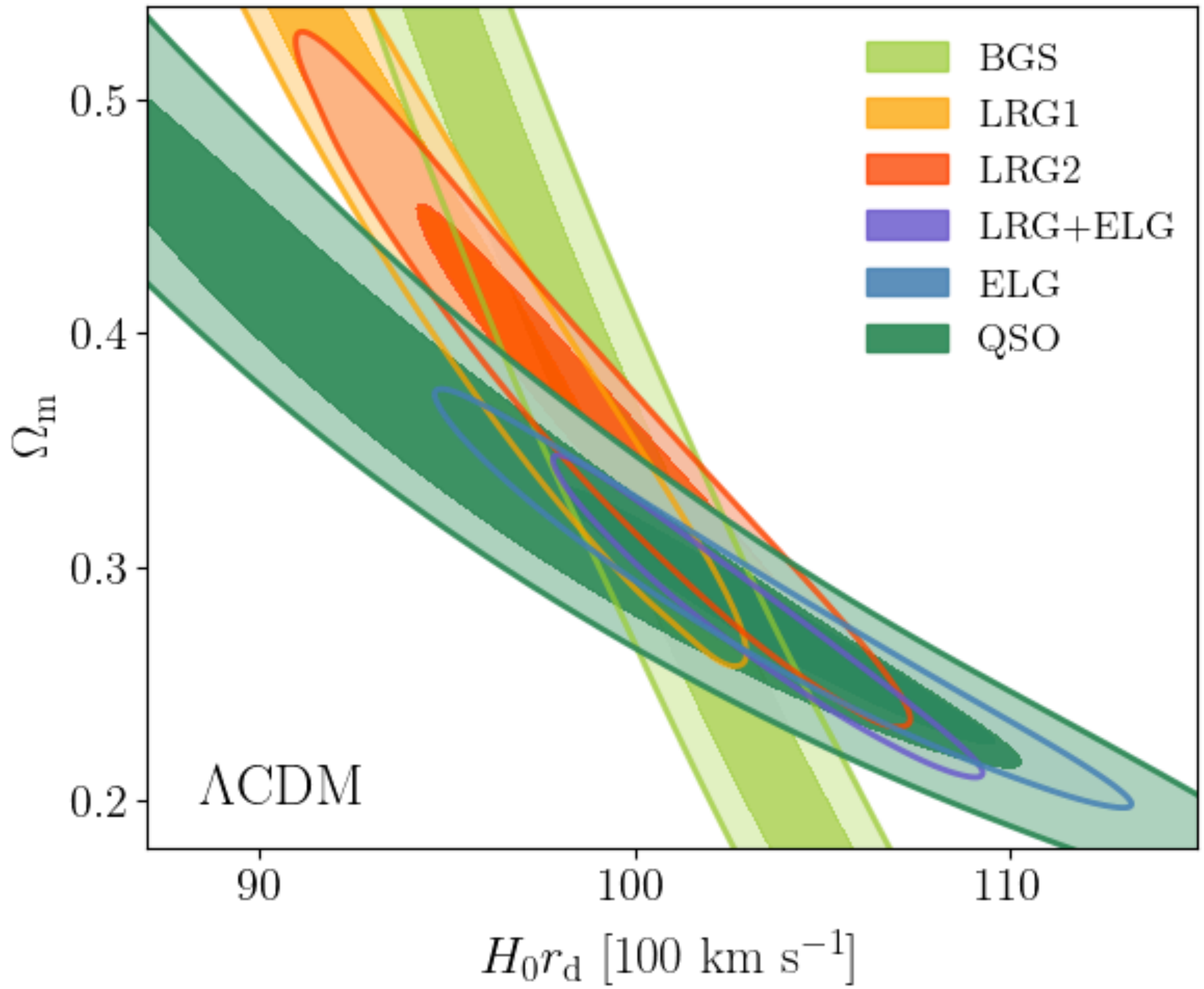

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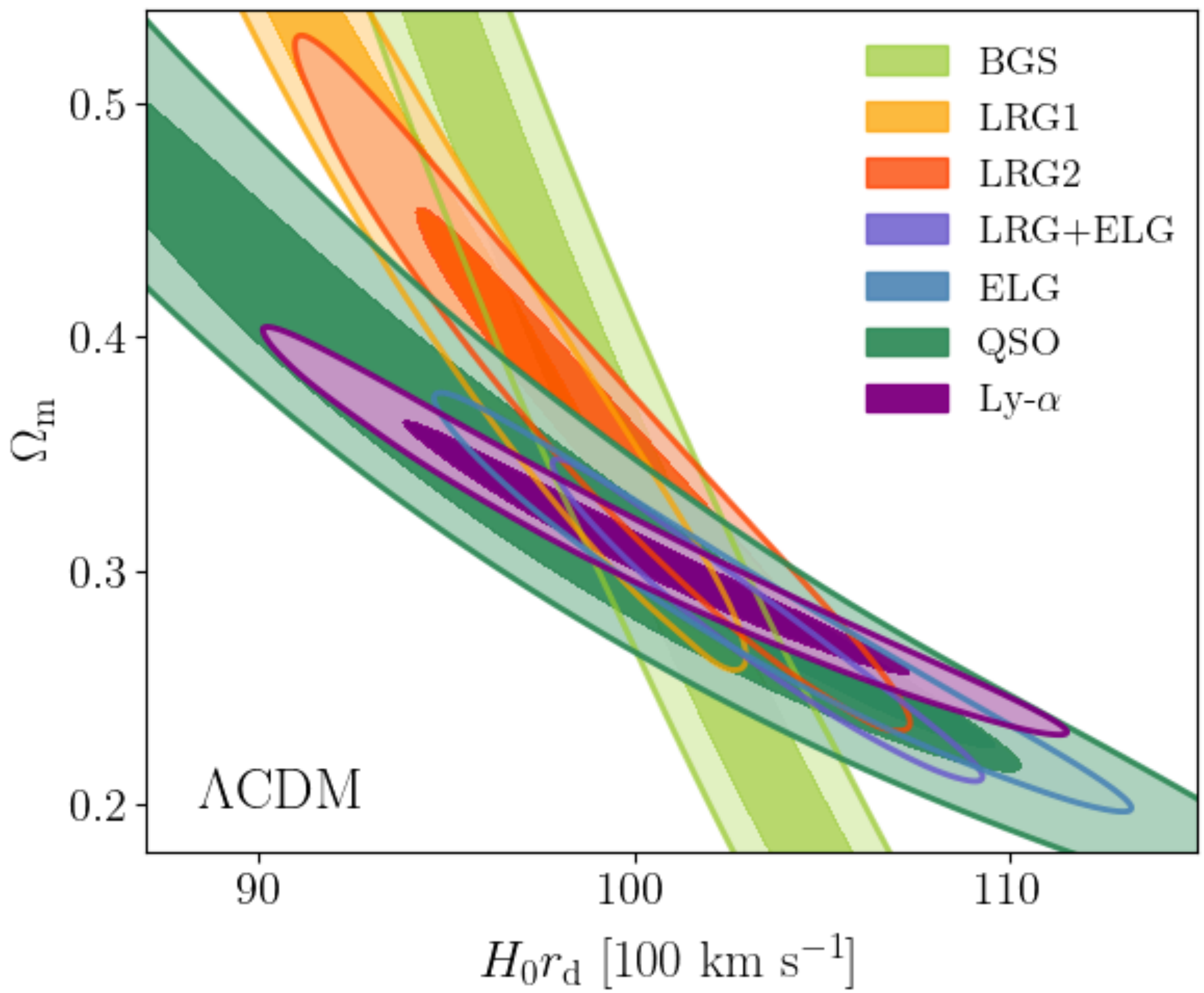
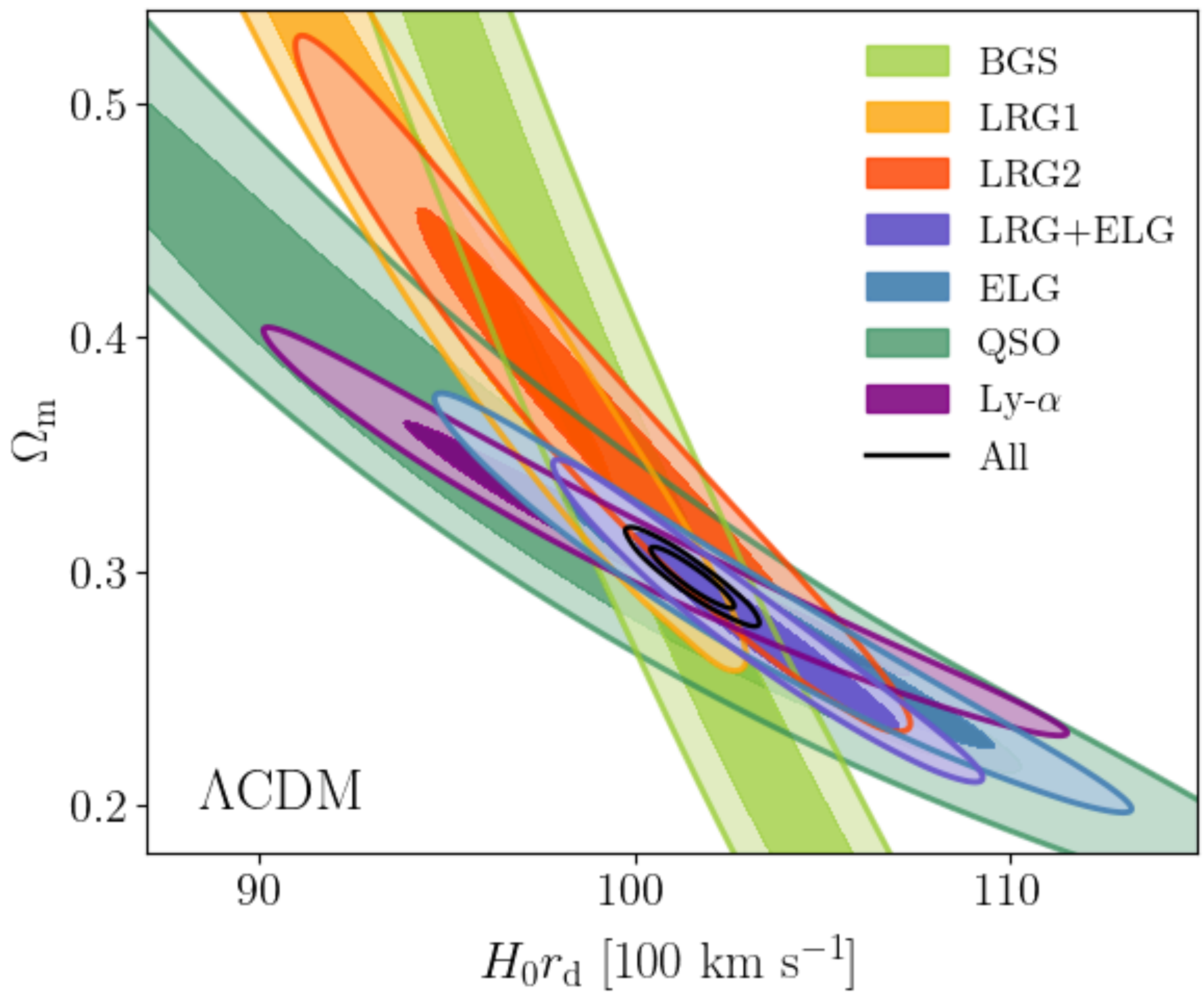
Flat ΛCDM model



Flat ΛCDM model



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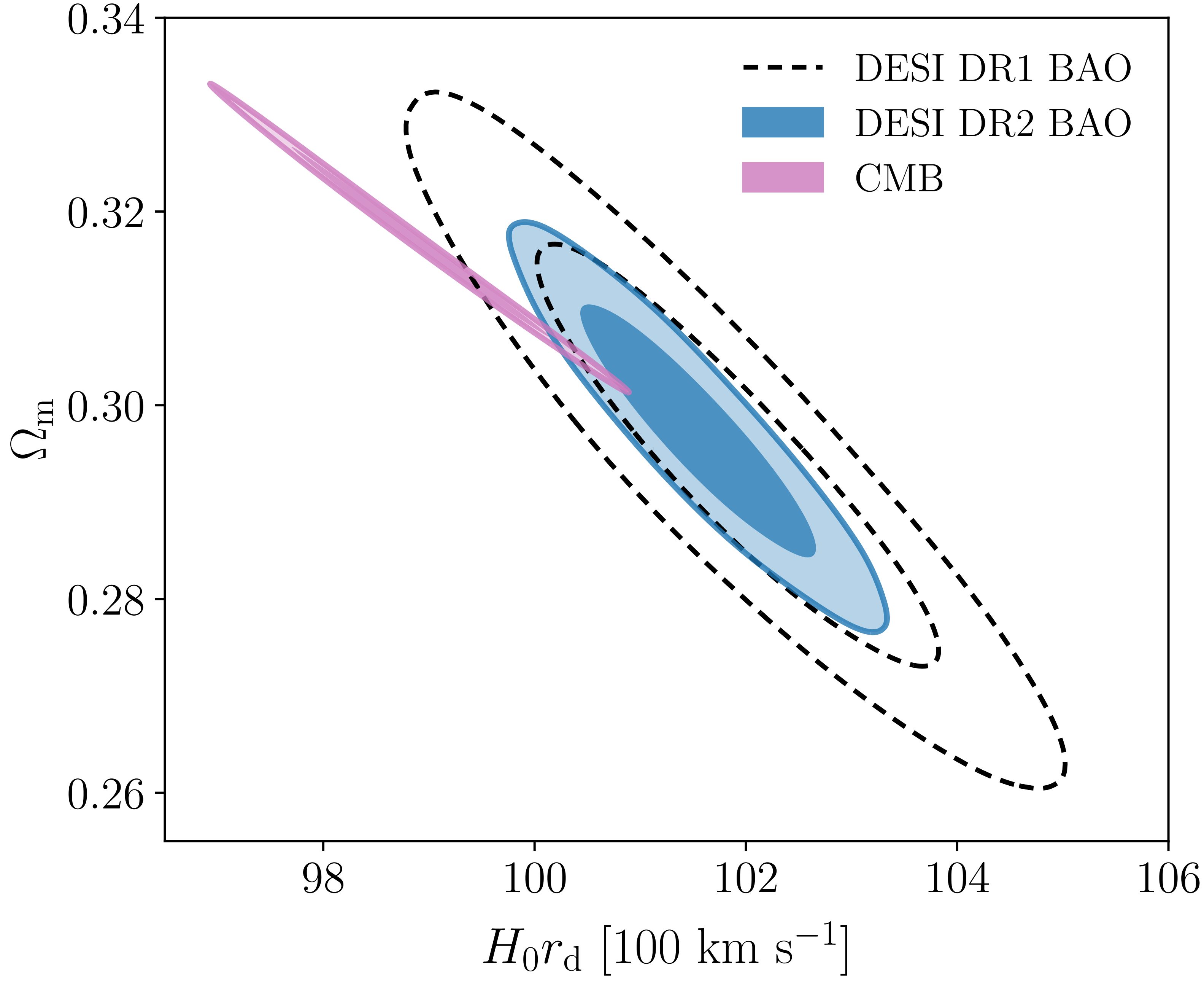

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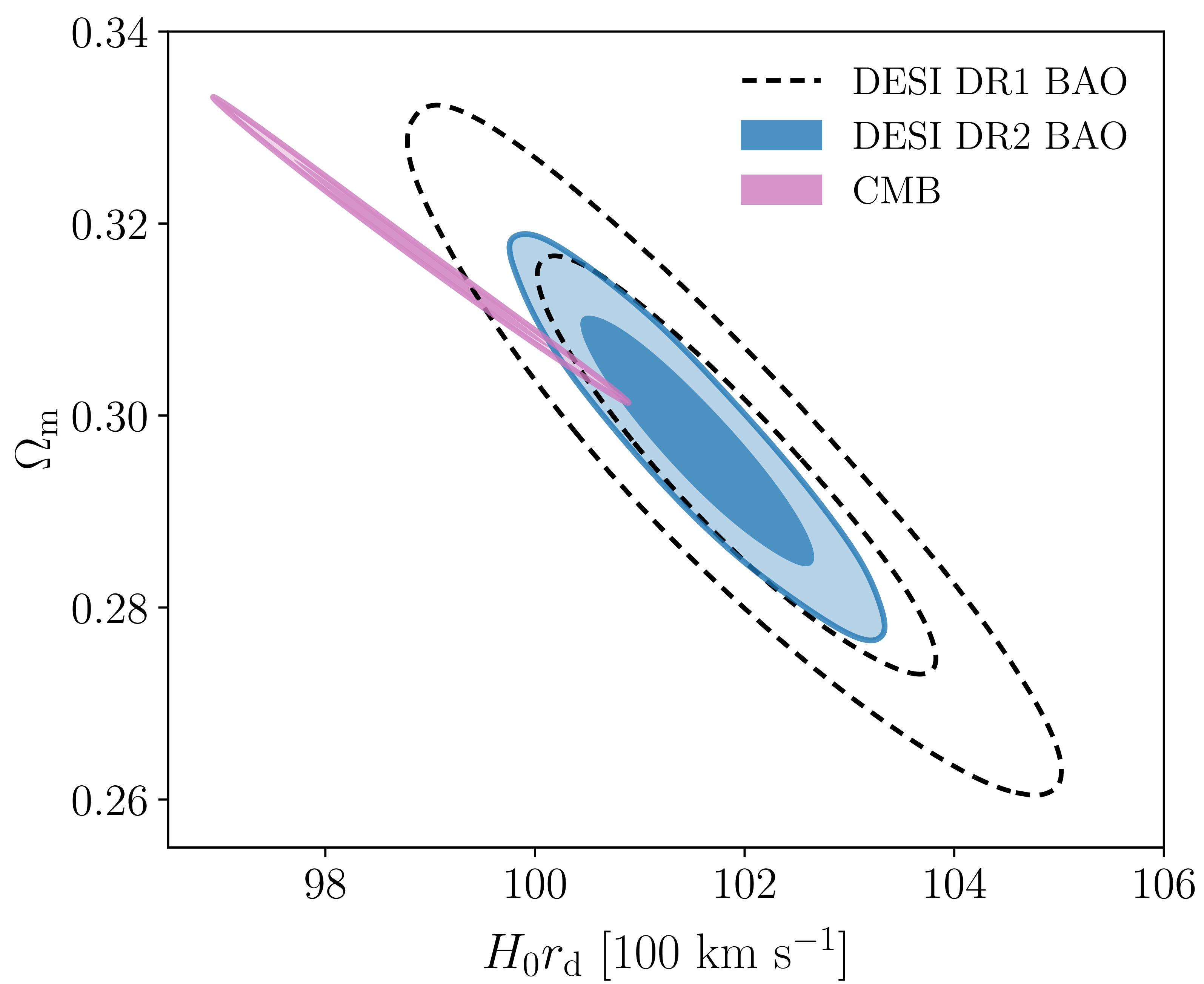
Flat ΛCDM model

- DR2 agrees beautifully with DR1
- DR2 is 40% more precise
- There is a $\sim 2.3\sigma$ **discrepancy** between DESI BAO and CMB
- The discrepancy has **increased** in DR2
- DESI BAO and CMB still agree very well on the acoustic angular scale θ_*
- The discrepancy is roughly in $\Omega_m h^2$
- Discrepancy increases to $\sim 2.8\sigma$ with latest SPT data



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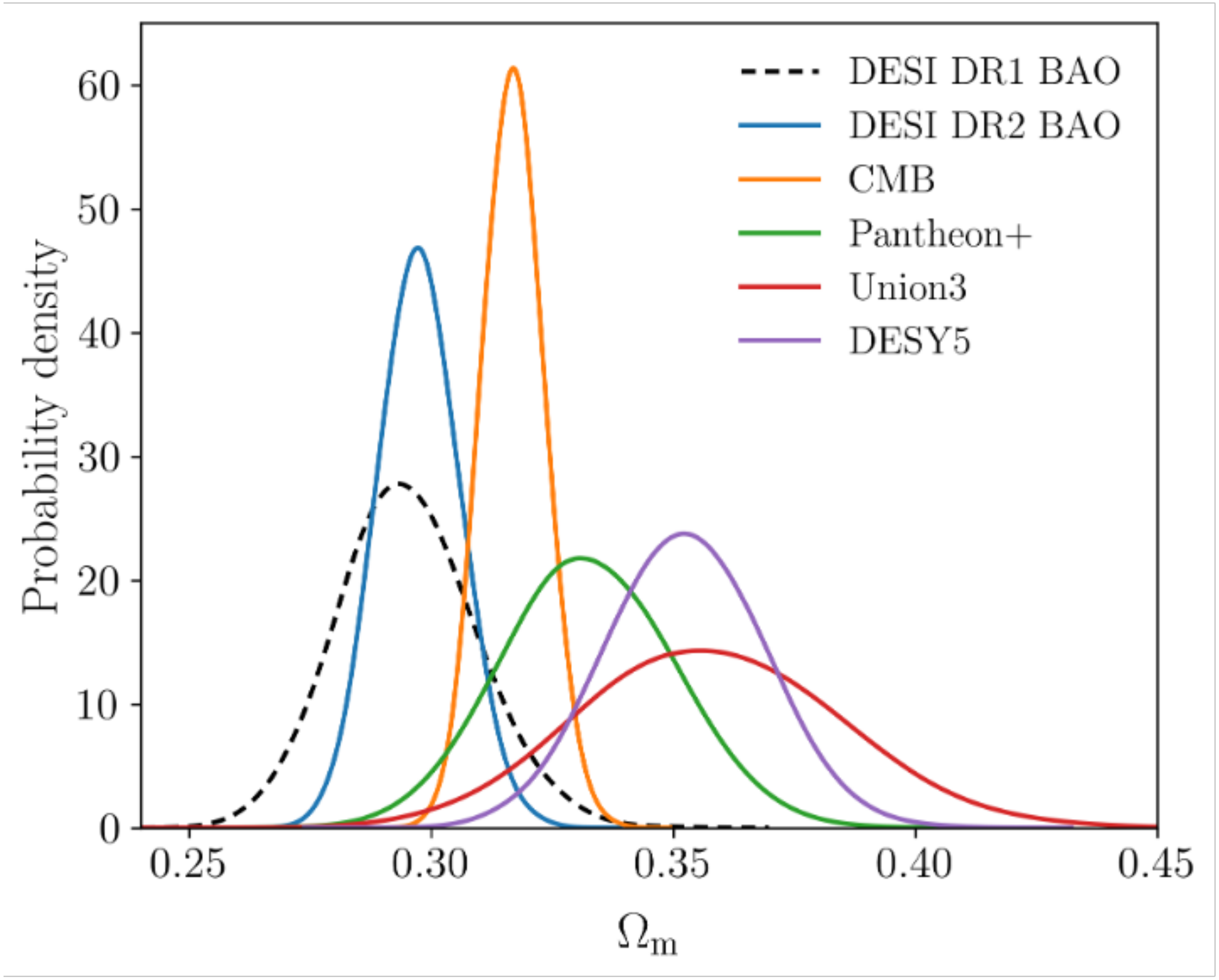
Flat ΛCDM model



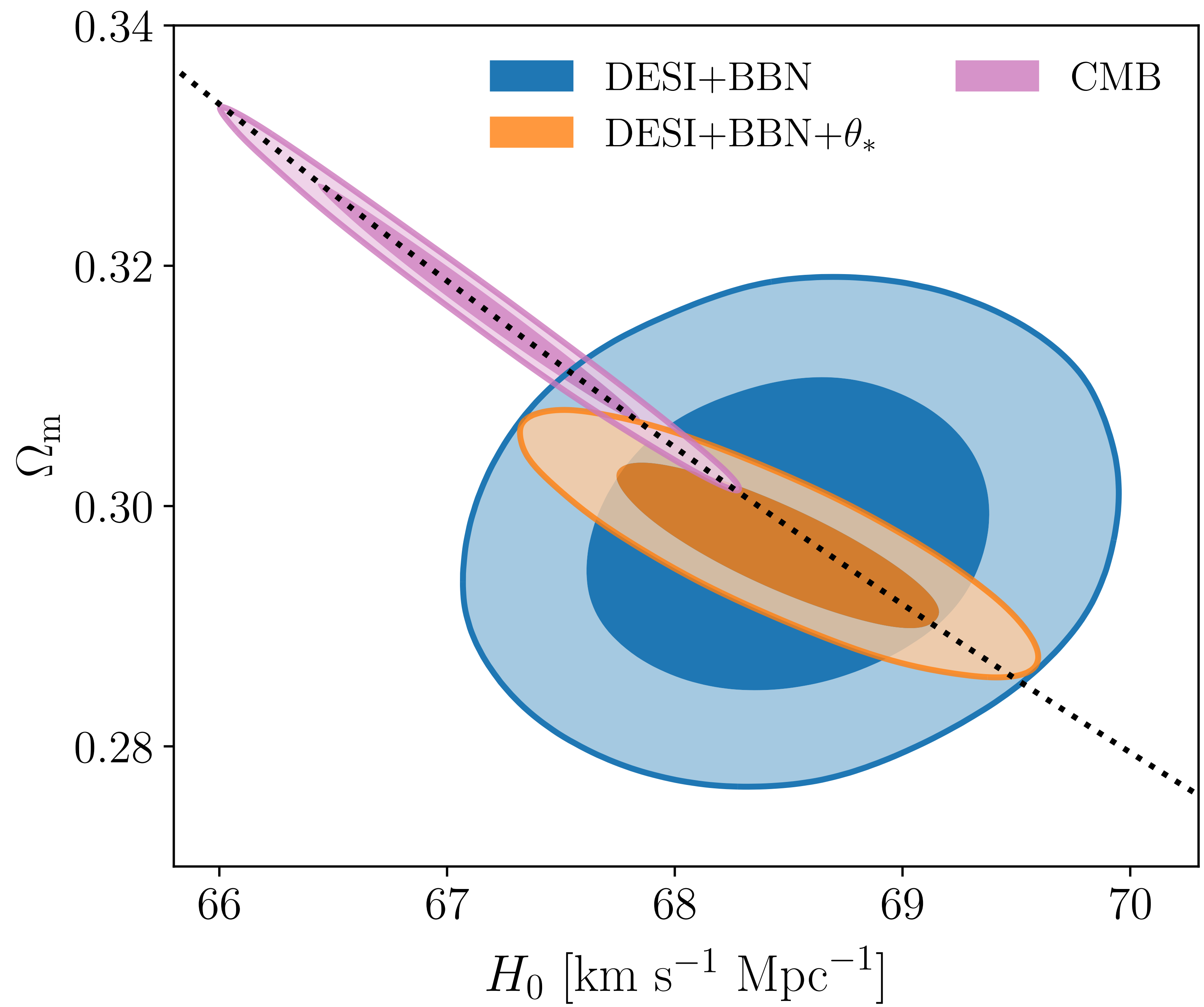
Which flat ΛCDM?

Flat ΛCDM model

Can calibrate sound horizon r_{d} with a BBN prior

Calibrated BAO then prefers slightly larger H_0 than Planck,
$H_0 = 68.5 \pm 0.6\ \mathrm{km\,s^{-1}\,Mpc^{-1}}$

(again, very consistent in θ_* but slightly shifted in $\Omega_{\mathrm{m}}h^2$ …)



$$\Sigma m_\nu$$

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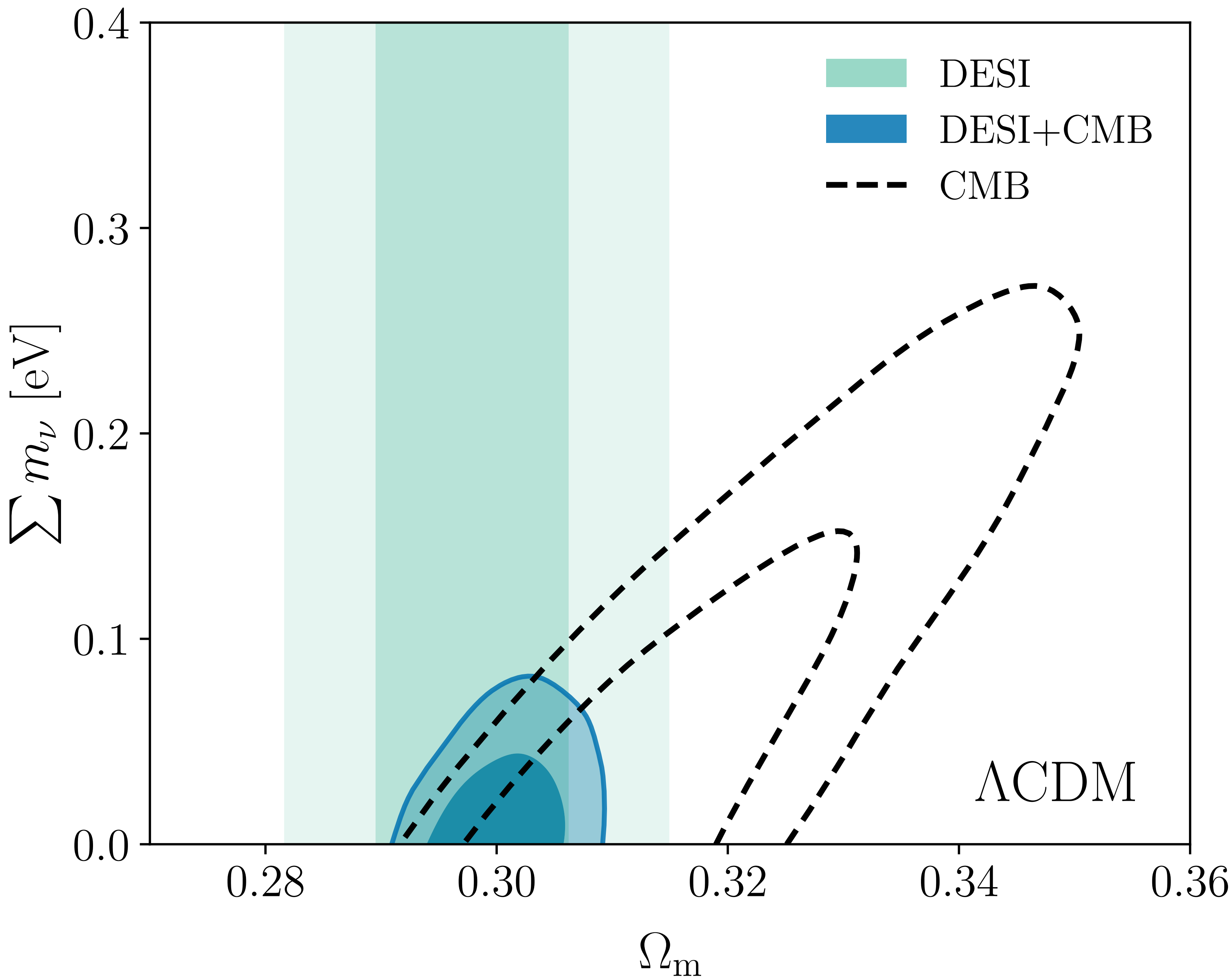
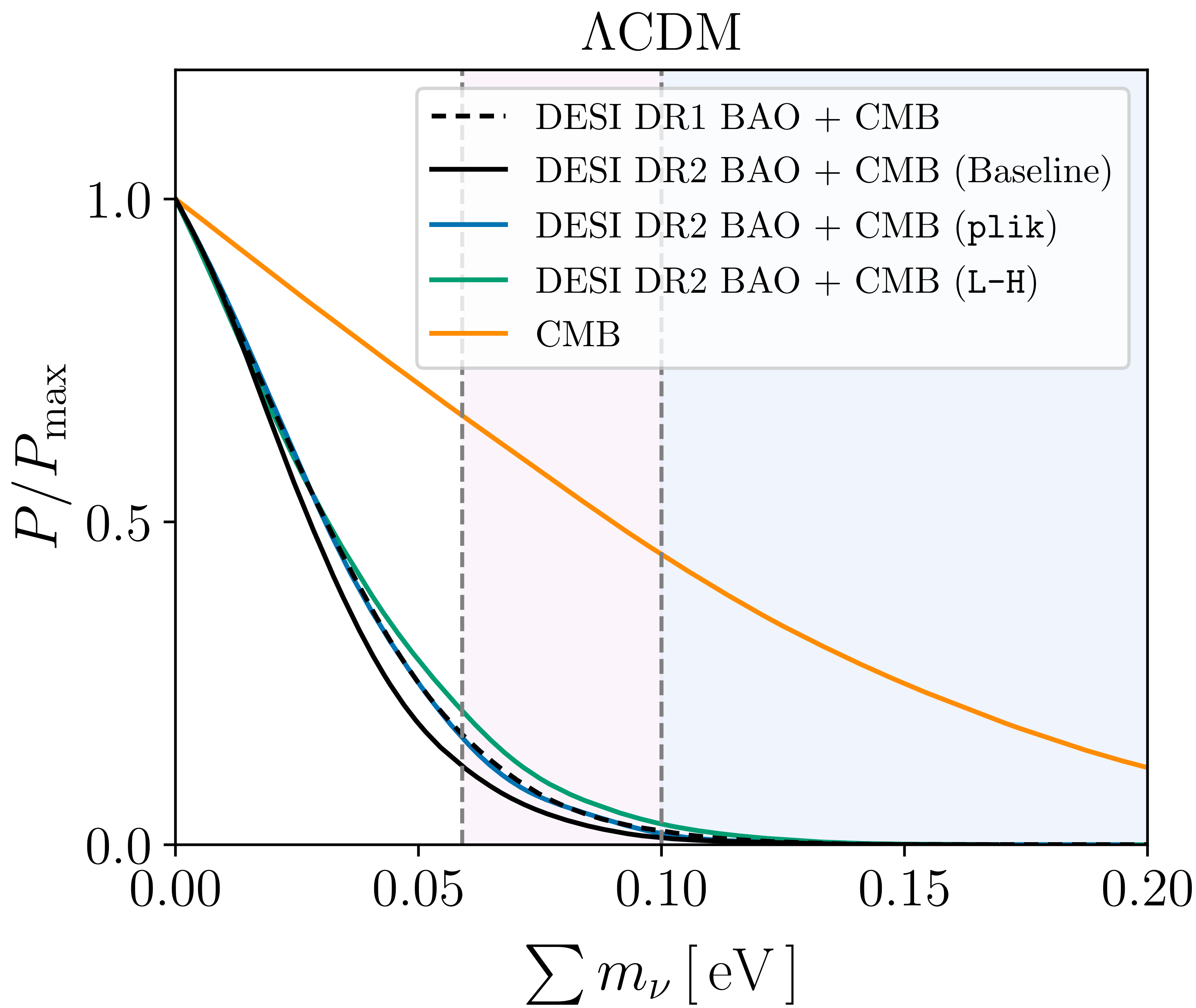
DESI and the neutrino mass scale



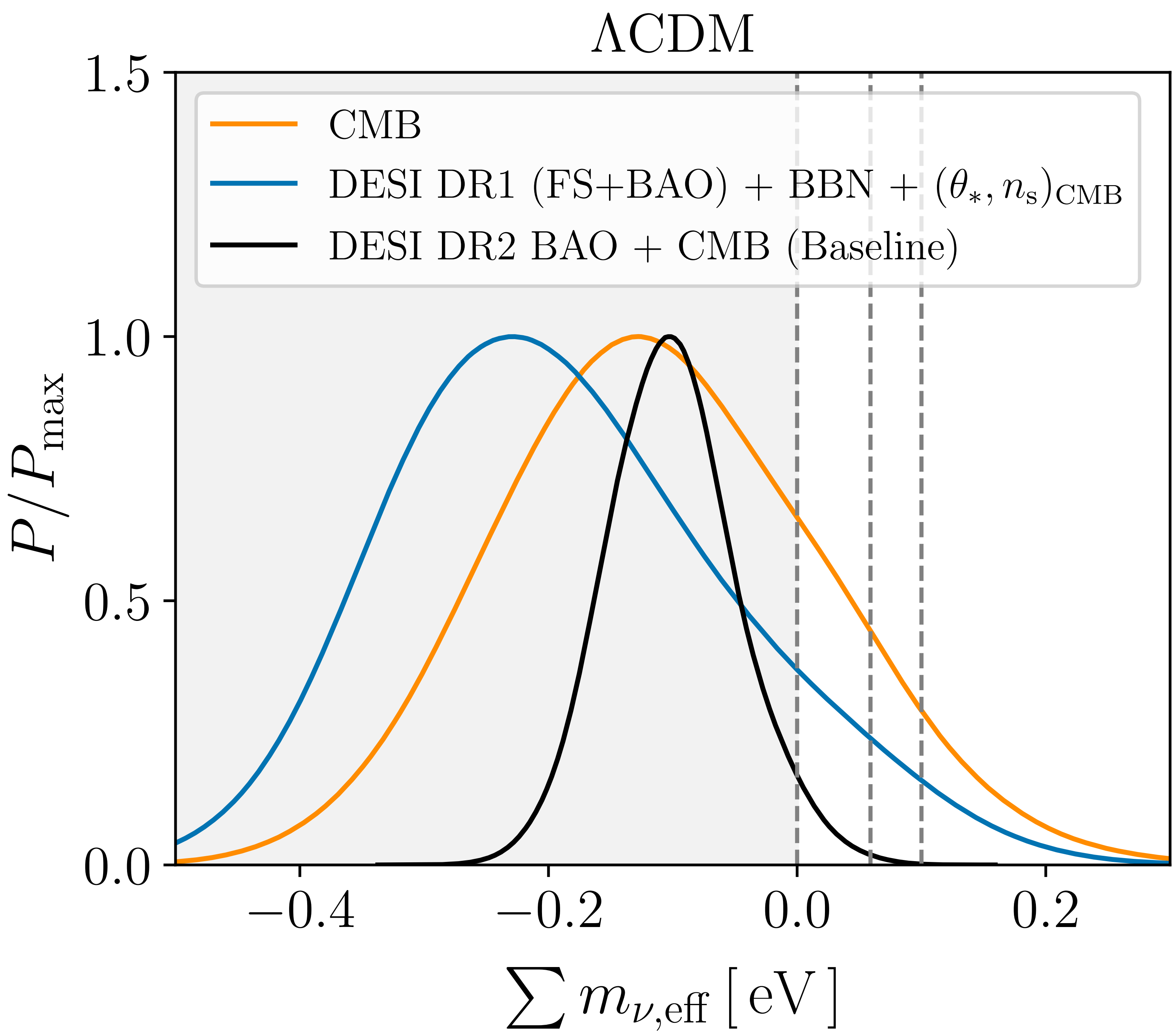
$\Sigma m_\nu < 0.064$ eV (95 % CL)

DESI and the neutrino mass scale



Uncomfortably close to terrestrial lower bound?

DESI and the neutrino mass scale



Elbers et al (2025)

If allowed, data would prefer even smaller (negative!) **effective** masses …

Dark energy $w(z)$

DESI BAO and Dark Energy

"Equation of state" governs effect of DE on expansion:

$$w = \frac{P_{\rm DE}}{c^2 \rho_{\rm DE}}$$

Allow equation of state to vary with time, $w = w(a)$. Baseline analysis uses the "**CPL parametrisation**"

$$w(a) = w_0 + w_a(1 - a)$$

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$$w(a) = w_0 + w_a(1 - a)$$

Remember: CPL is useful as it can match many physical models, but it is not to be taken too literally!

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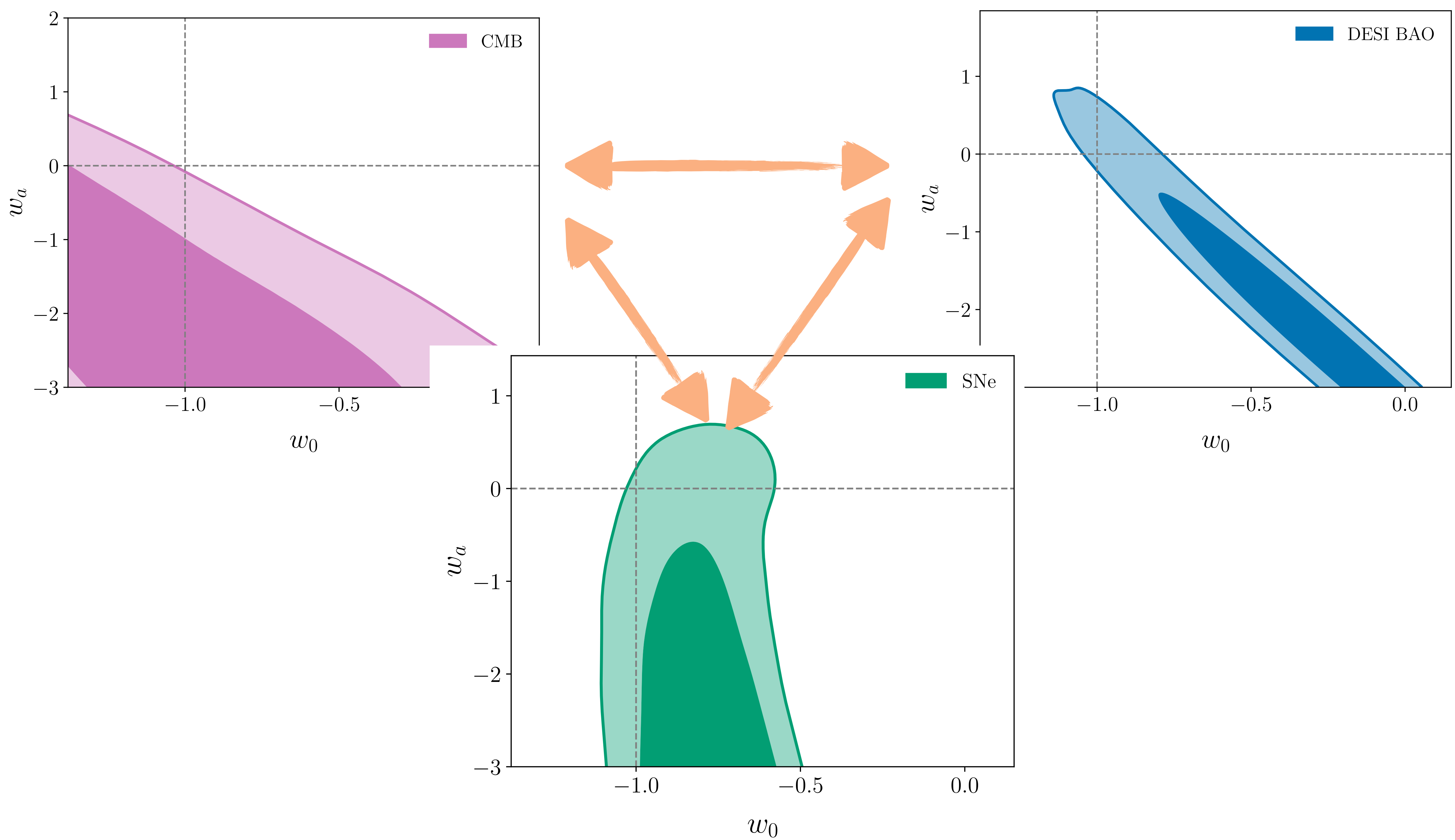
DESI BAO and Dark Energy



Very hard to say much with 1 probe alone – combinations essential!

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DESI BAO and Dark Energy

Level of discrepancy with ΛCDM cosmological constant DE from

DESI + CMB: 3.1σ

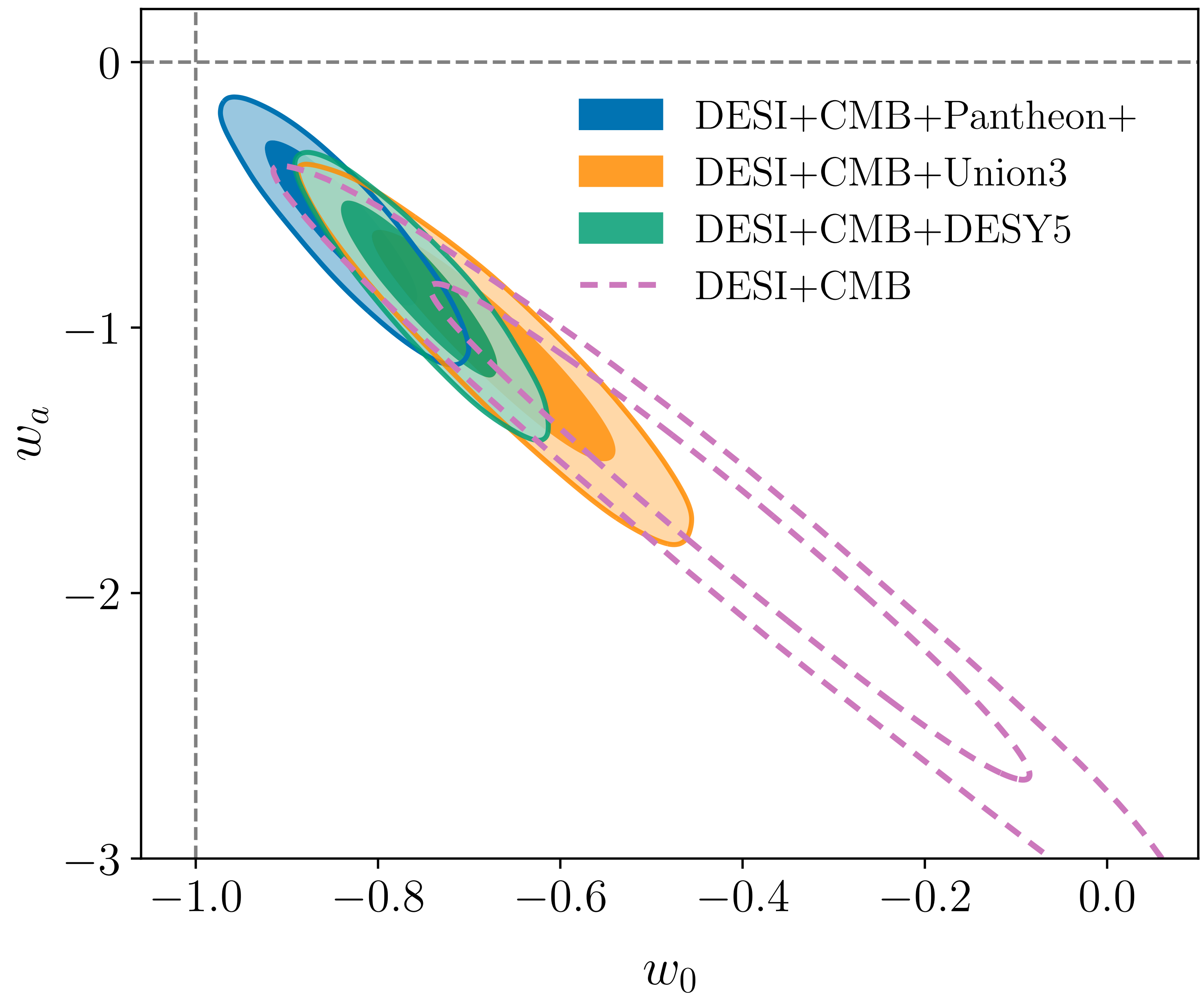
+ Pantheon+: 2.8σ

+ Union3: 3.8σ

+ DES-SN5YR: 4.2σ

*deliberately frequentist statistic! Other metrics in paper

Hints of dark energy **weakening**



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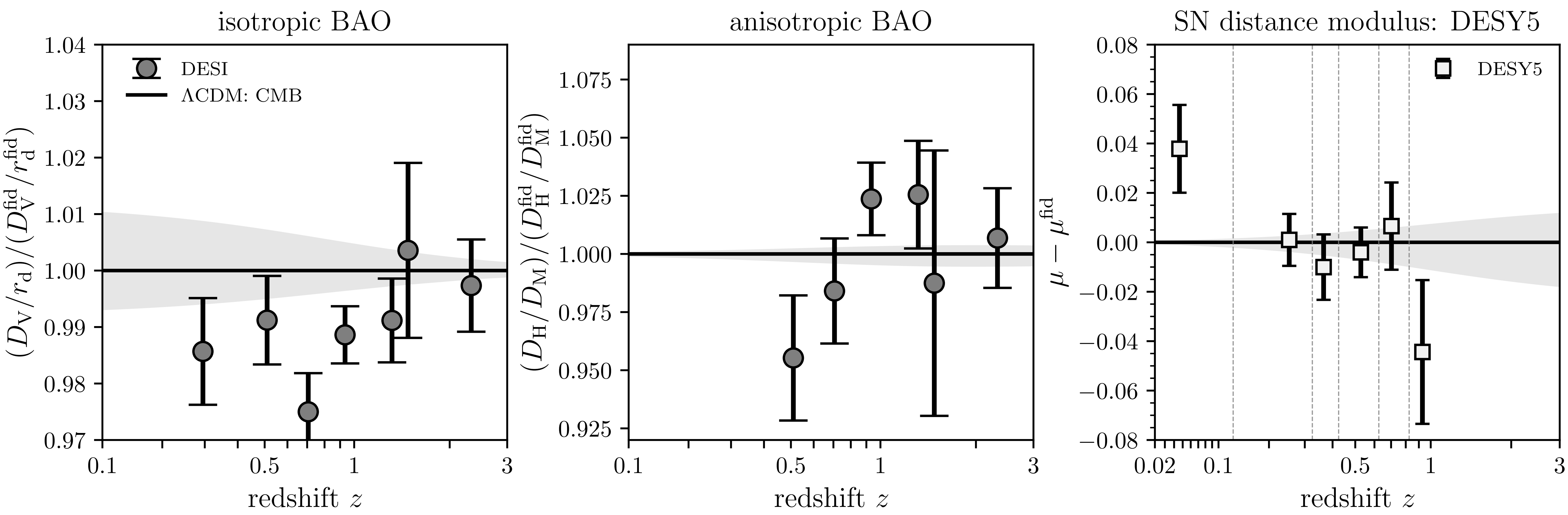
DESI BAO and Dark Energy

Understanding **why** the data pulls away from ΛCDM:



CMB can be encapsulated in 3 parameters: $(\theta_*, \omega_{bc}, \omega_b)$

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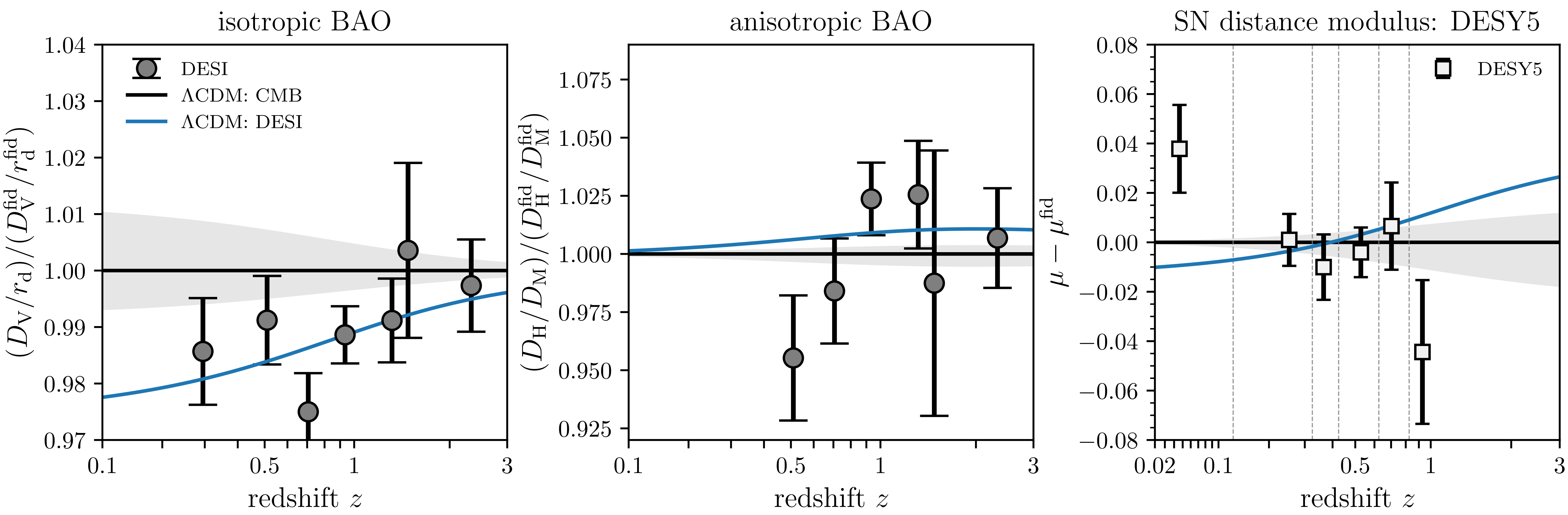
DESI BAO and Dark Energy

Understanding **why** the data pulls away from $\Lambda\mathrm{CDM}$:



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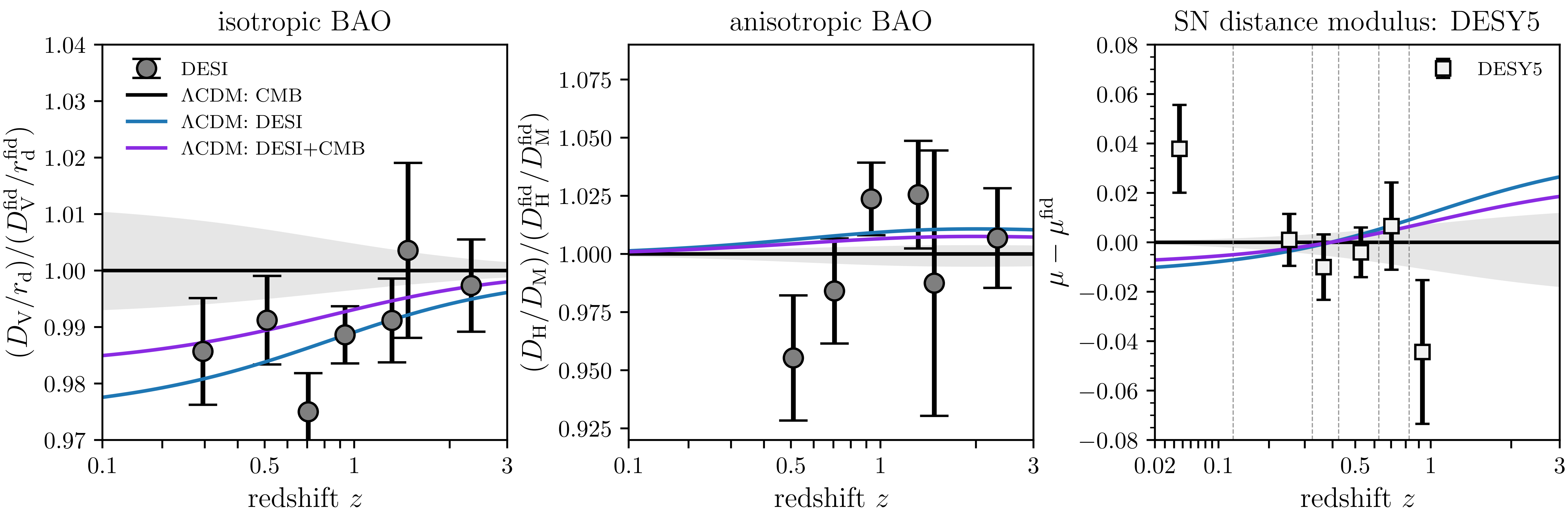

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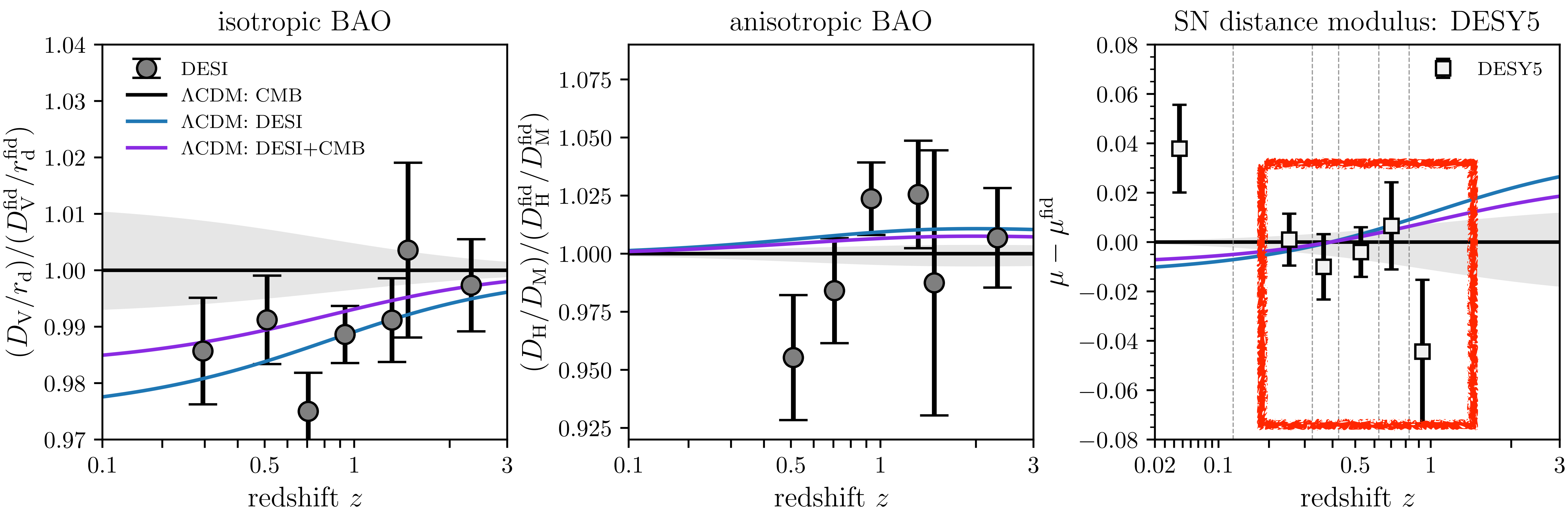
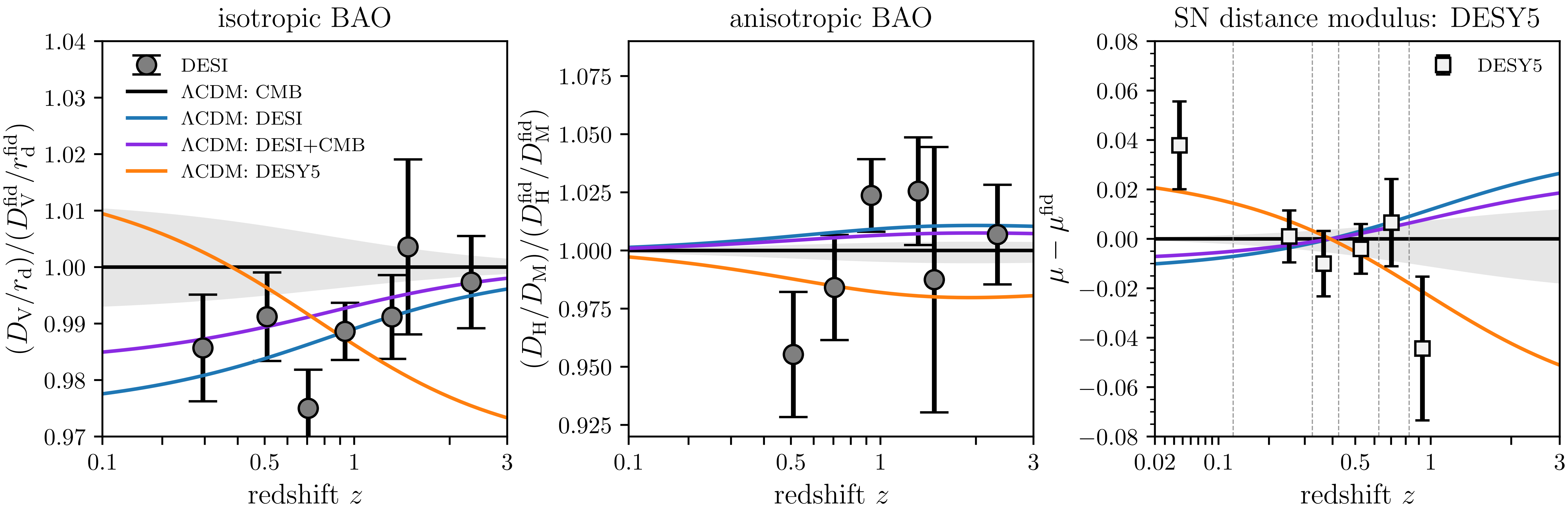


BAO and SN distances agree in the range of redshift overlap!

CMB can be encapsulated in 3 parameters: $(\theta_*, \omega_{bc}, \omega_b)$

DESI BAO and Dark Energy

Understanding **why** the data pulls away from ΛCDM:



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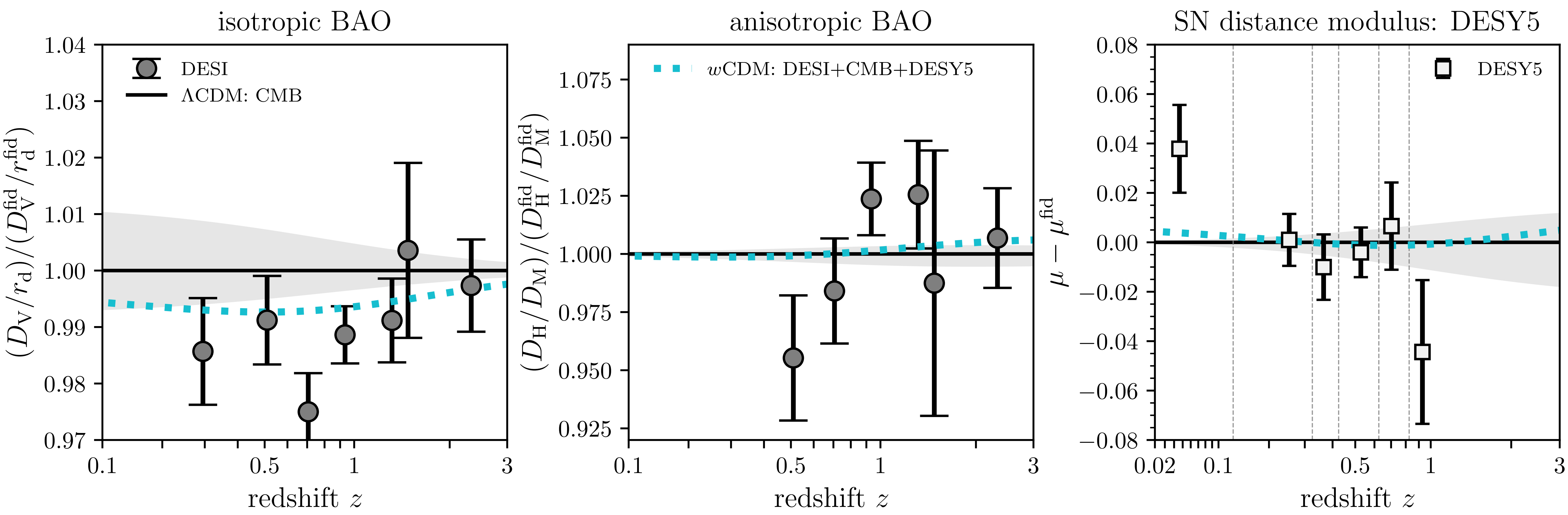

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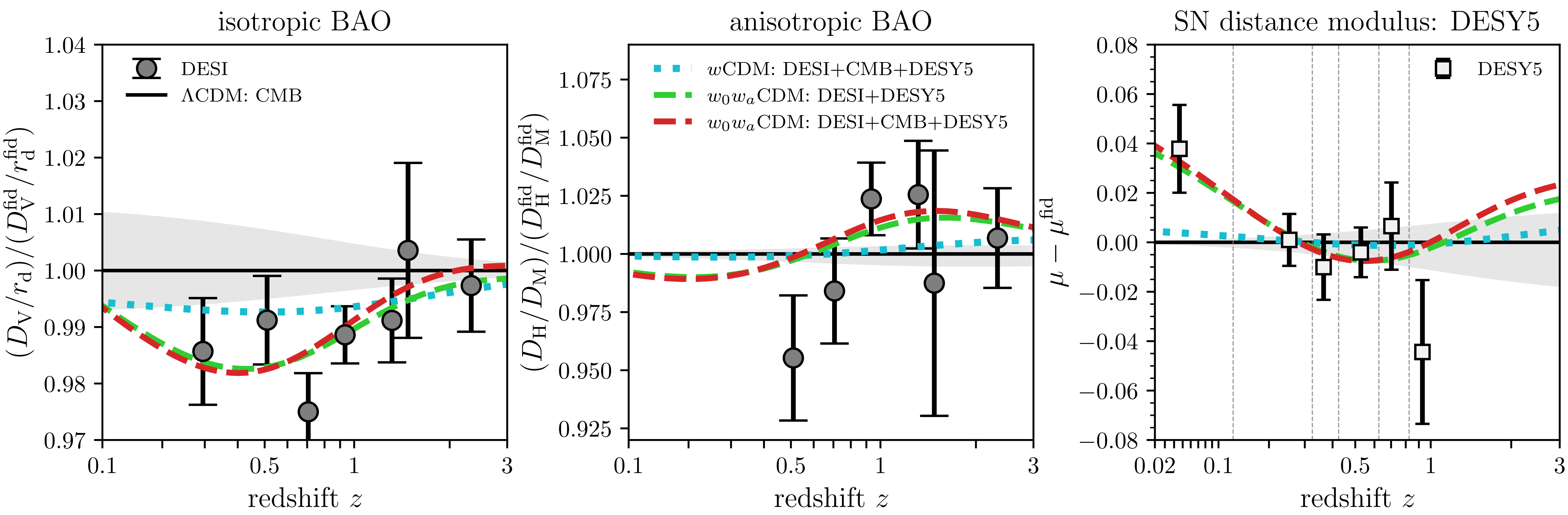
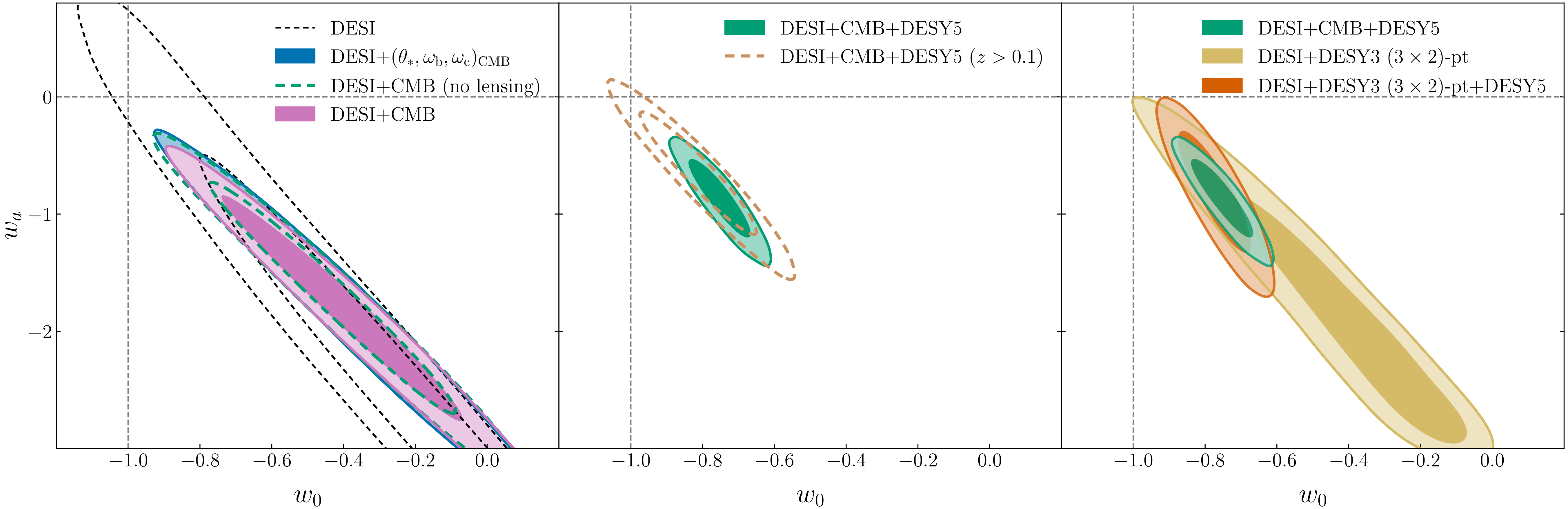


CMB can be encapsulated in 3 parameters: $(\theta_*, \omega_{bc}, \omega_b)$

How robust is this result?



Swapping out different datasets changes the significance, but *the results stay in the same part of parameter space*

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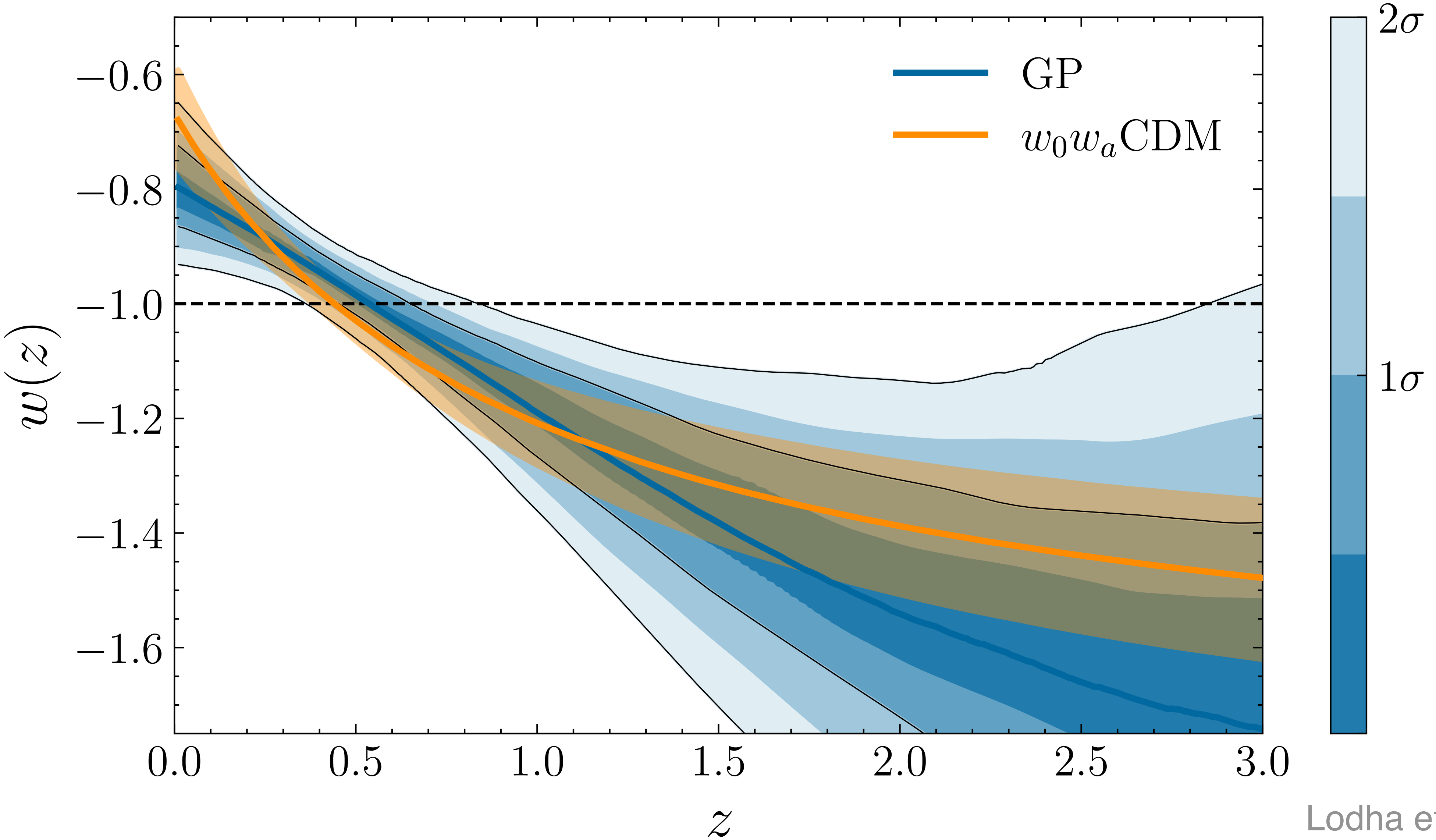
How robust is this result?



Lodha et al (2025)

It is not specific to the CPL parametrisation

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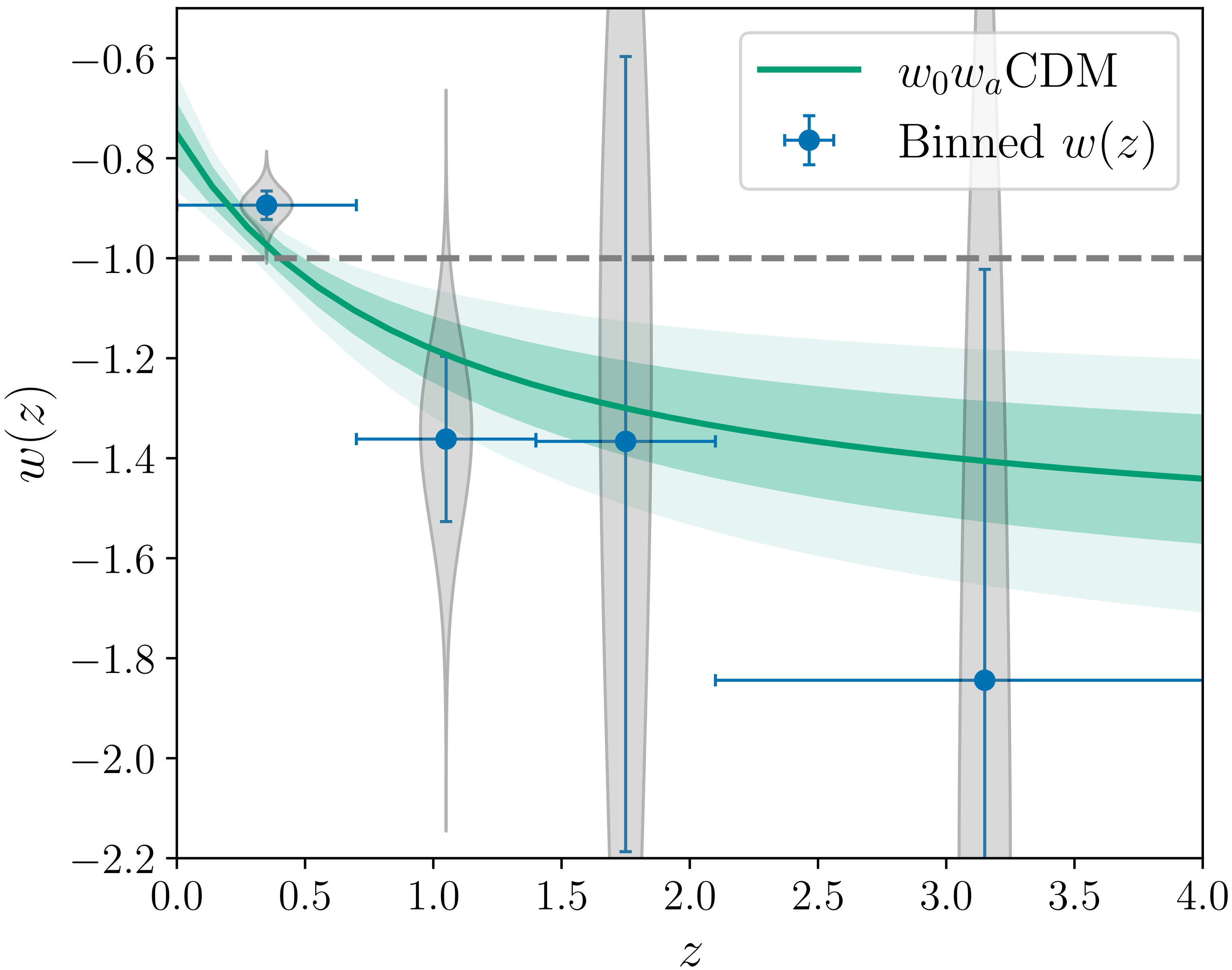
How robust is this result?



It is not specific to the CPL parametrisation

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Can we avoid DE evolution?

- Hard to see any effects that could shift DESI BAO results
- CMB $\Omega_\mathrm{m}h^2$: ACT and SPT recently confirmed Planck results, hard to see much room for shifts
- SN data certainly could have some offsets (especially calibration across surveys and historical low-z samples)
  - big new SN data efforts on the horizon, esp. Rubin/LSST

- Perhaps some room for non-flat curvature? Dinda & Maartens (2025), Chen & Zaldarriaga
- Shift in optical depth to reionization, τ, could help ΛCDM? Sailer et al (2025)
- Modification of recombination history? Mirpoorian et al (2025)
- Interaction between DE and DM components? Khoury et al (2025)
- Non-minimal DE coupling to gravity? Wolf et al (2025)

What's next? – other analyses

From DESI:

- Adding full-shape information to BAO (for DR2 and beyond)
- Higher-order clustering statistics (3-pt, 4-pt, …)
- Simulation-based / emulator approaches
- Lensing and cross-correlations
- …

From other surveys:

- Euclid – DR1 in late 2026, DR3 in 2031 (ask me about Euclid BAO)
- ZTF and LSST SNe
- LSST and Euclid weak lensing (and DES Y6 soon)
- Roman, SO, SphereX, …


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Science & Technology Facilities Council

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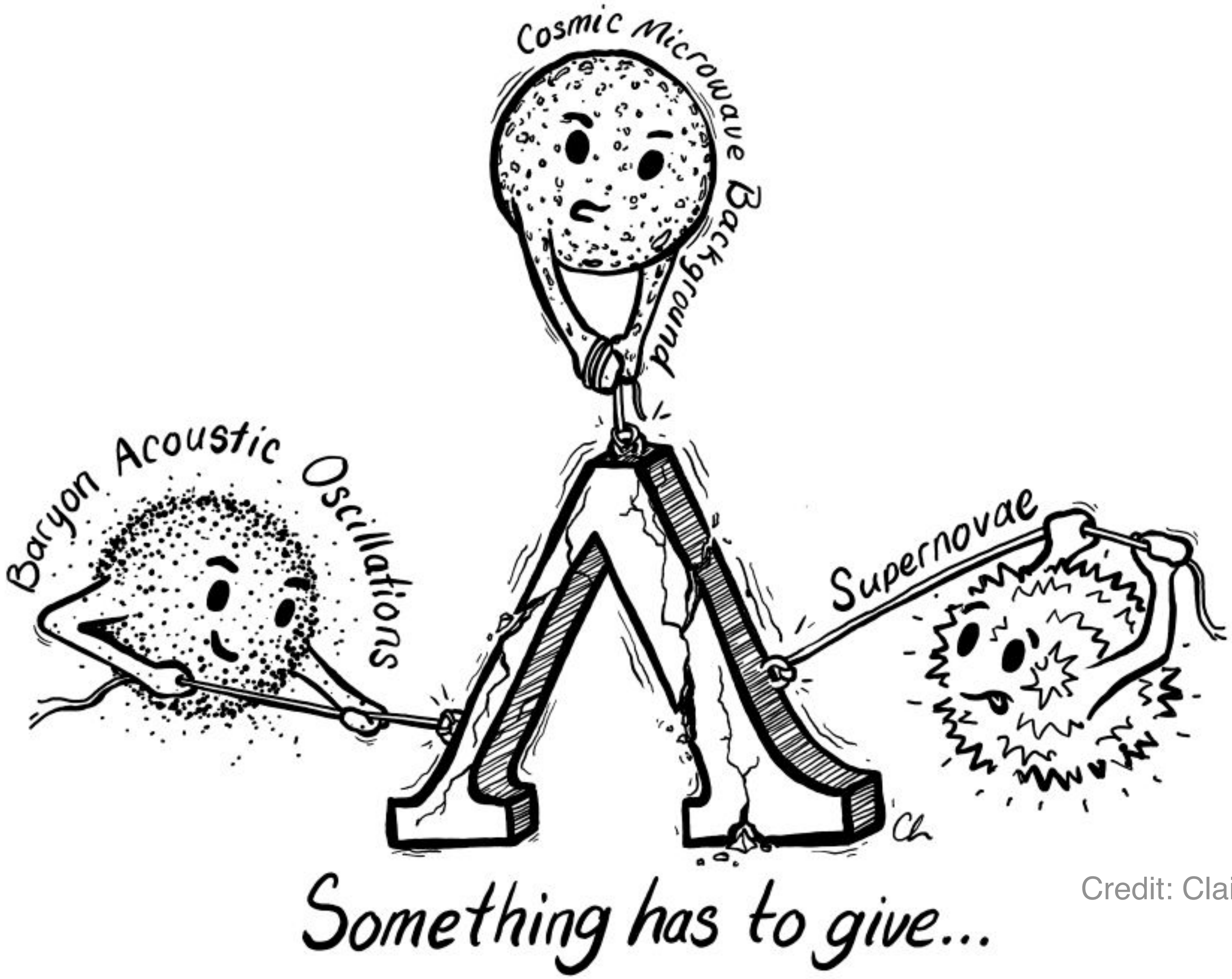
What's next?



Credit: Claire Lamman

Frequently Asked Questions for DR2

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Frequently asked questions for DR2

1. Are BAO and SN distance measurements in conflict?
2. Are DESI and DES parameter values inconsistent (in $w_0 w_a$CDM)?
3. Are there any BAO outliers?
4. What happened to tension between DESI and SDSS BAO?
5. I heard DESI DR2 is actually *more* consistent with Planck than DR1??
6. Why does the data give $w \simeq -1$ in fixed wCDM?
7. Does it matter which CMB likelihood you use?
8. Why do you use only a 1D BAO fit to BGS at $z = 0.3$?

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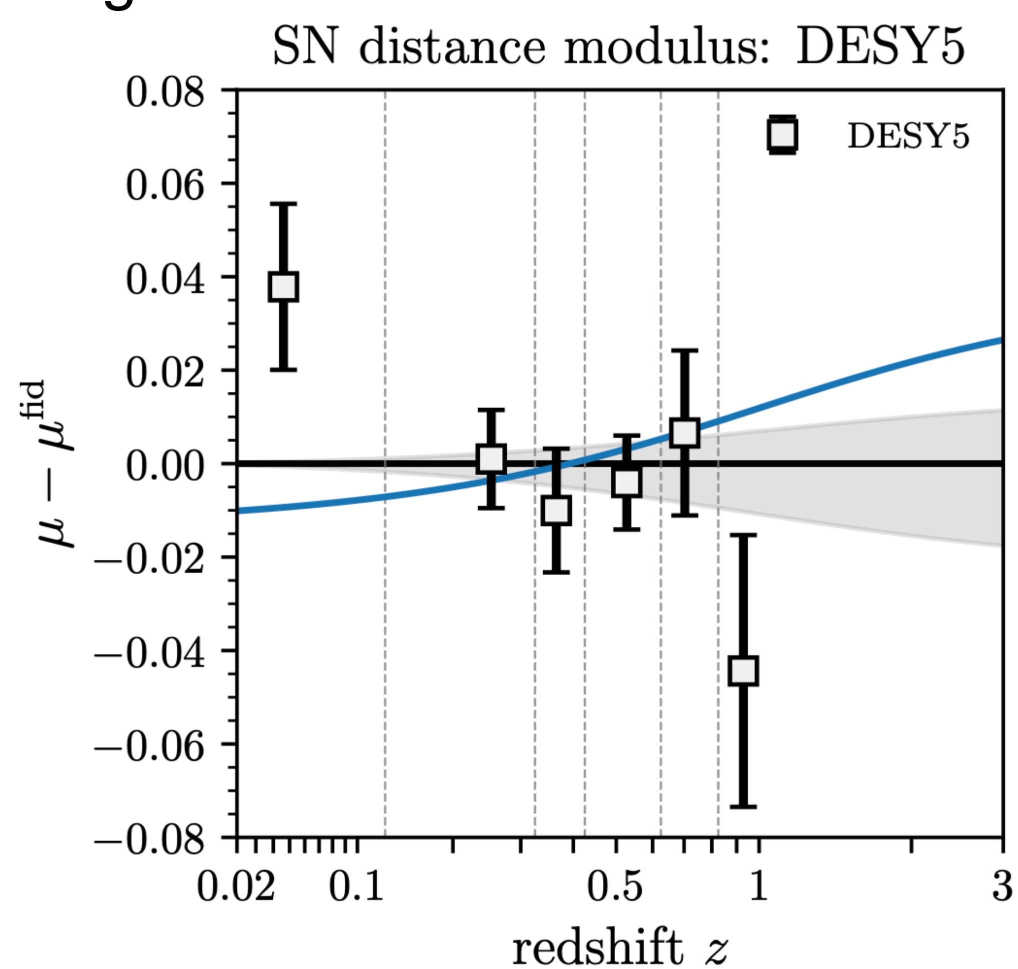
Q: Do BAO and SN give different distances?

A: **No**, in the overlapping redshift range they are very consistent!

From the paper:

pernovae. For supernovae at $z > 0.1$, which partially overlap the redshift range of DESI, the ΛCDM model that best fits the DESI data is also a good fit to the SNe data. Relative to models that best fit each of the DESY5, Union3 and Pantheon+ SNe samples alone, over the full redshift range, the DESI best-fit model gives only small shifts in the quality of the fit to the SNe data, with $\Delta\chi^2 = -1.2$, 1.5 and 2.3 respectively. Unfortunately, no

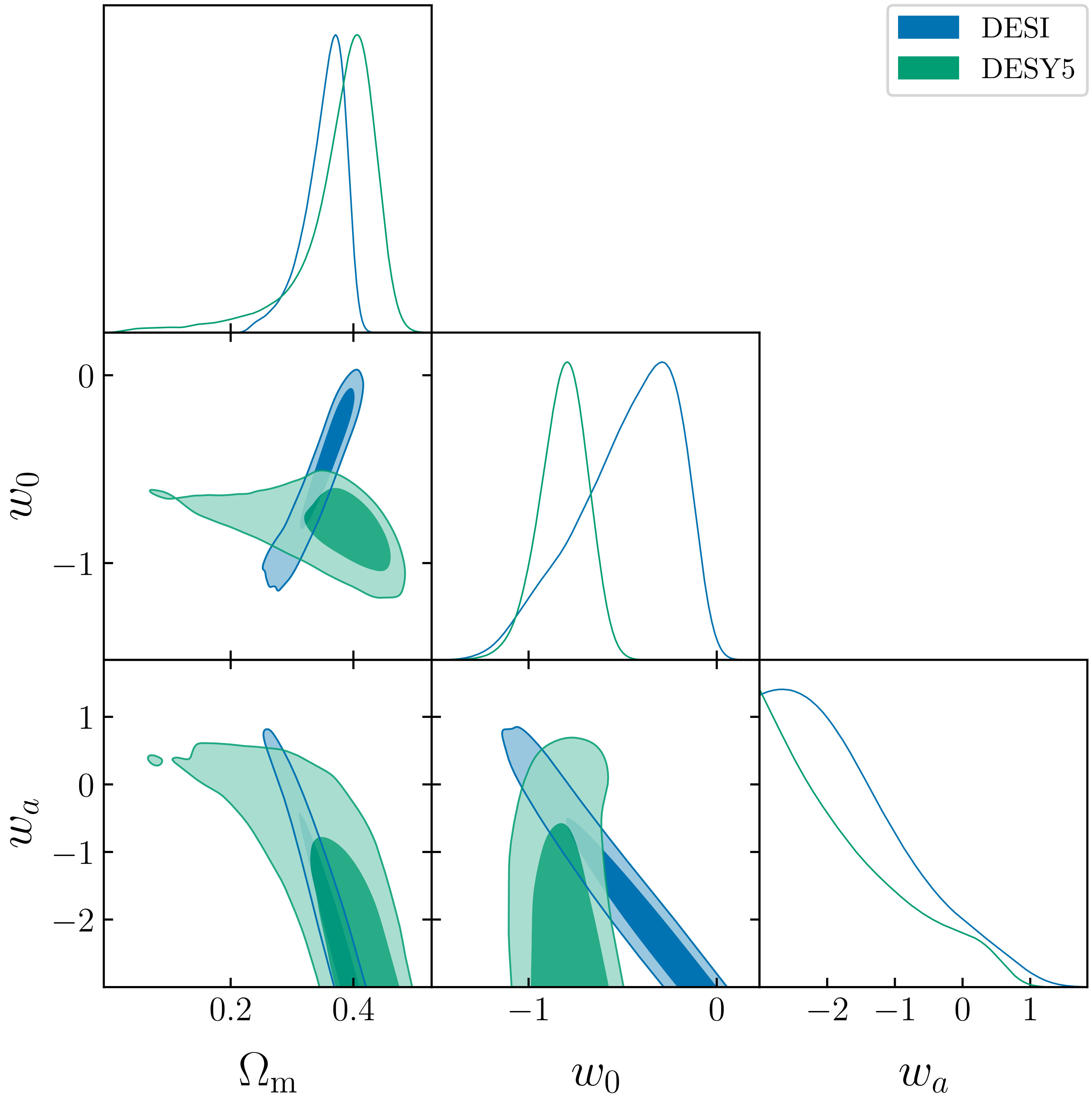
E.g.



Note: Some mistaken claims in the literature come because they compare *calibrated* SN (using SH0ES H_0) to *calibrated* BAO (using Planck H_0) – this is just the Hubble tension again.

Q: Are BAO and DES SN in conflict in $w_0 w_a$?

A: **No**, we don't think so!



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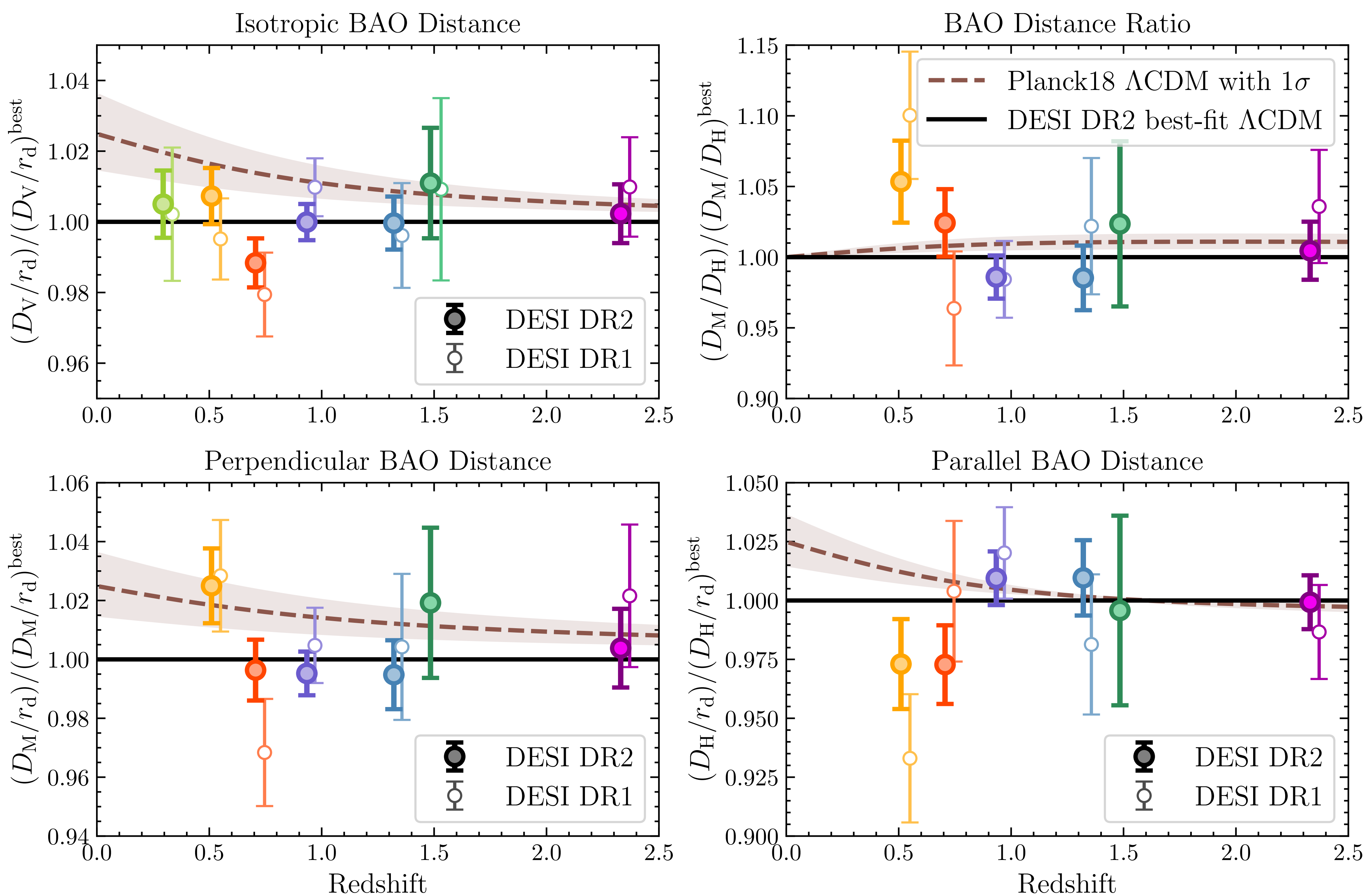
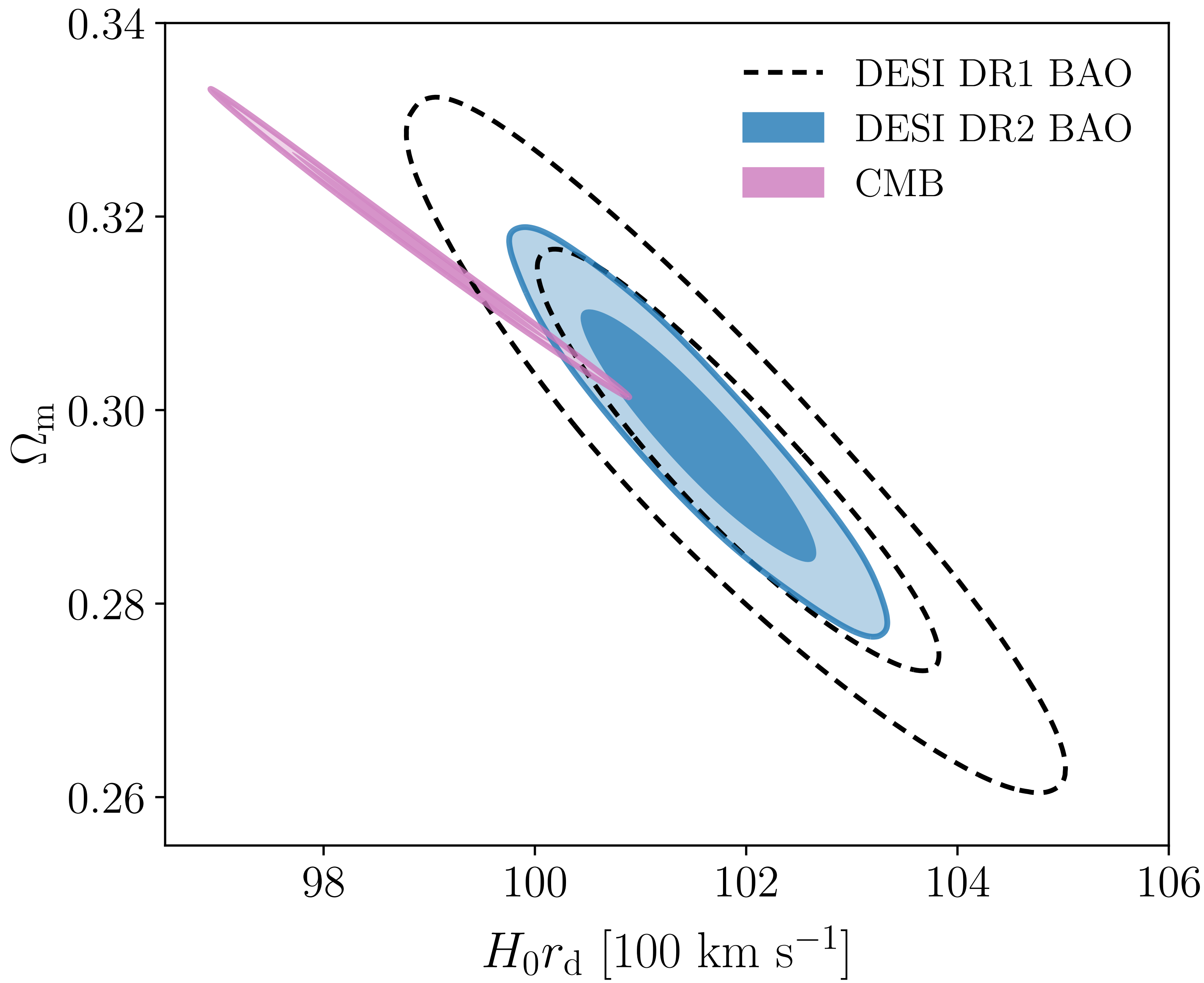
Q: BAO outliers/tensions with SDSS?

A: **No evidence** of unusual outliers; discrepancy between DESI and SDSS at $z = 0.71$ has **decreased** (~1.5 to 2.5σ) in DR2



Q: Is DR2 closer to Planck than DR1??

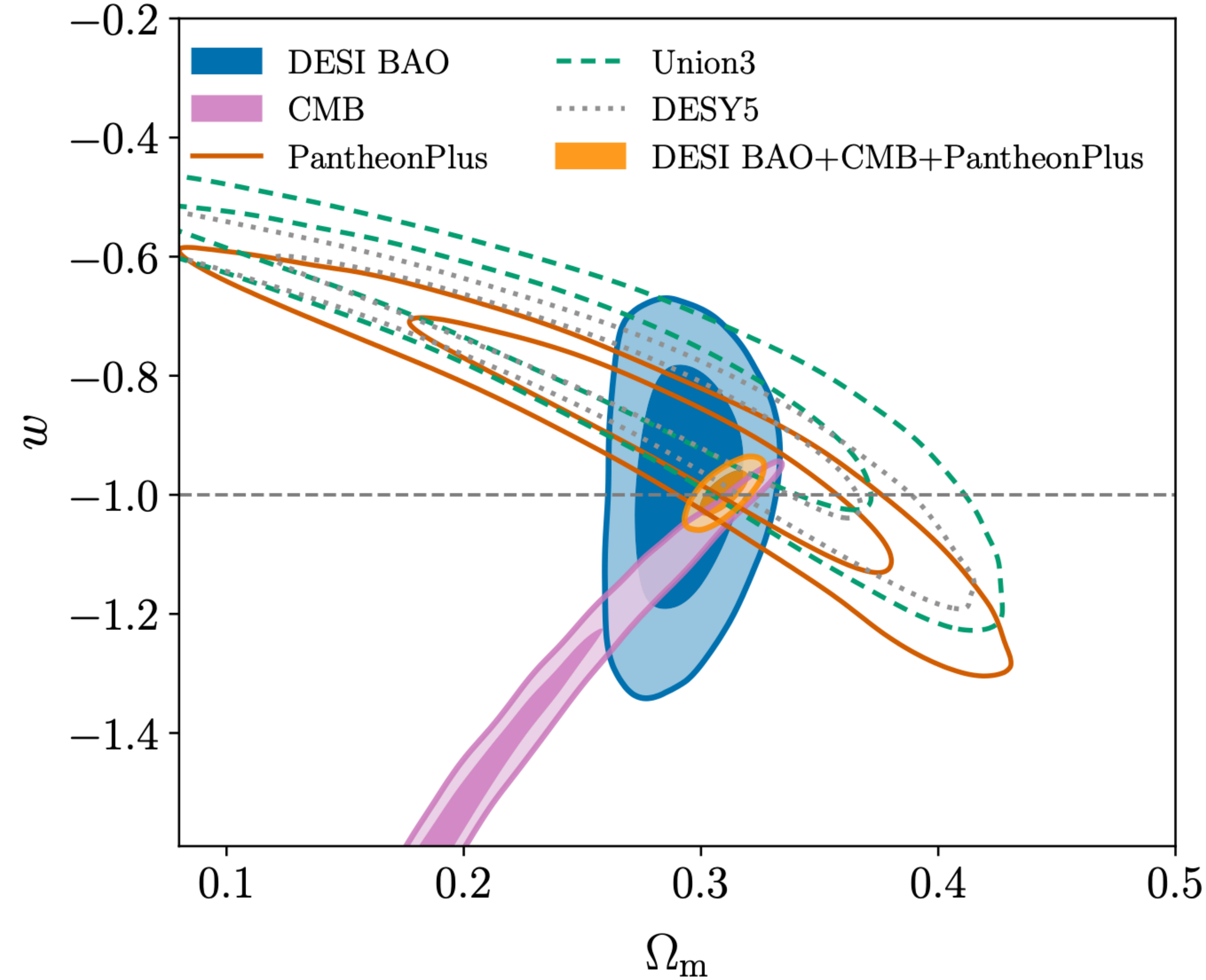
A: Judge for yourself!



Q: Why $w = -1$ in fixed wCDM?

Related Q: Why do you find $w(z) \simeq -1$ at the pivot redshift?

A #1: wCDM gives a (poor) compromise b/w high-z CMB and low-z SNe



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Q: Why $w = -1$ in fixed wCDM?

Related Q: Why do you find $w(z) \simeq -1$ at the pivot redshift?

https://arxiv.org/pdf/0708.0024.pdf

The Mirage of $w = -1$

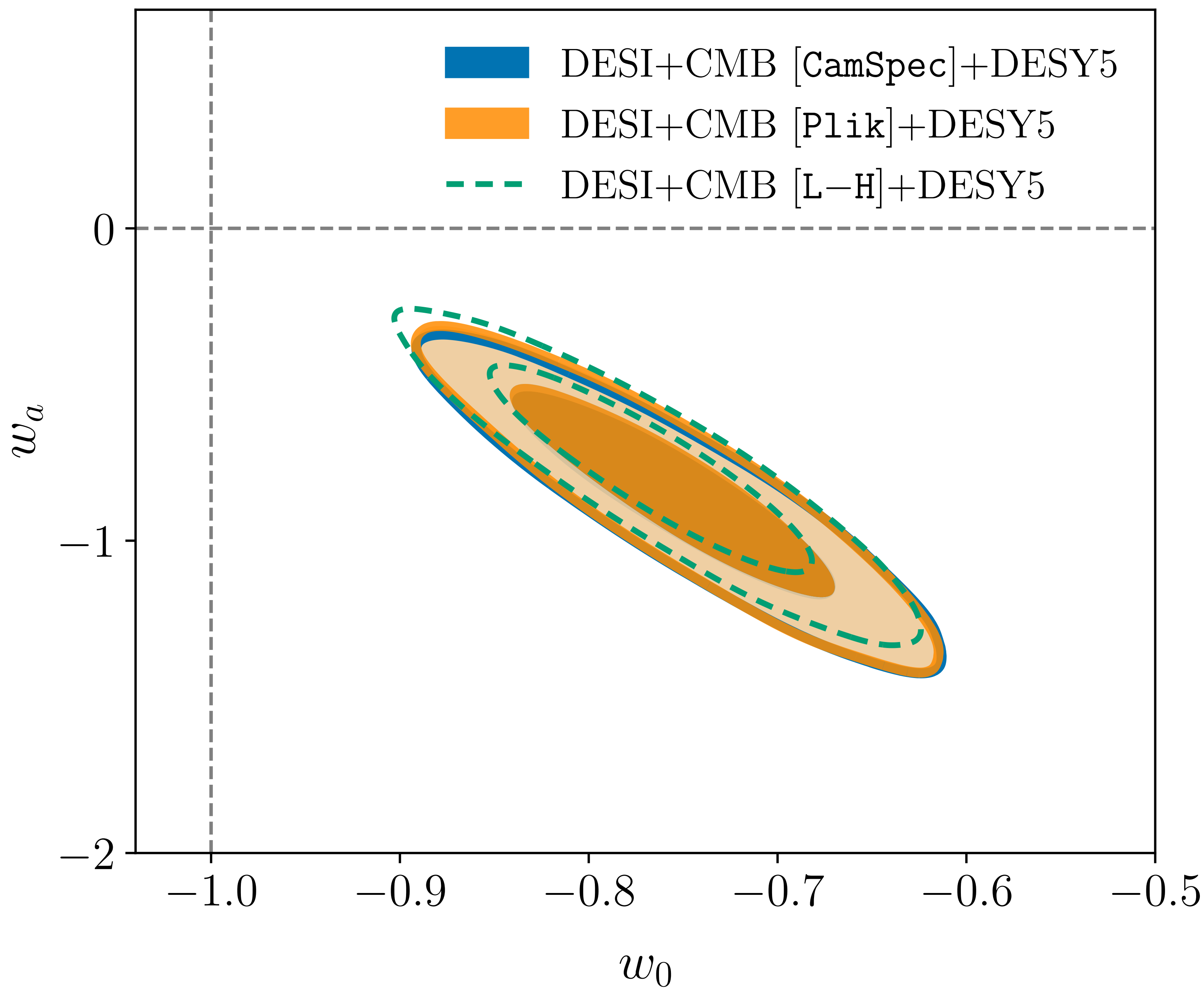
Eric V. Linder
Berkeley Lab, University of California, Berkeley, CA 94720, USA

Thus a high redshift distance measurement consistent with LCDM virtually forces (within the picture so far) the value $w(z \approx 0.4) = -1$, irrespective of true time variation. However, low redshift measurements insufficiently sensitive to time variation measure only an averaged EOS that corresponds strongly to the value at a sweet spot or "pivot" redshift with the pivot near $z \approx 0.4$. That is,

A #2: This was answered in 2007!

Q: Does it matter which CMB likelihood?

A: **Not important for DE:**



Q: Why only 1D BAO for BGS?

A: We made a **conservative** decision before unblinding

Factors considered:

- Blinded posterior for $\alpha_{\rm AP}$ suggested possible non-Gaussianity
- $\sim 5\,\%$ precision on $\alpha_{\rm AP}$ at BGS redshift adds very little cosmological information anyway
- Wanted to avoid any changes after unblinding