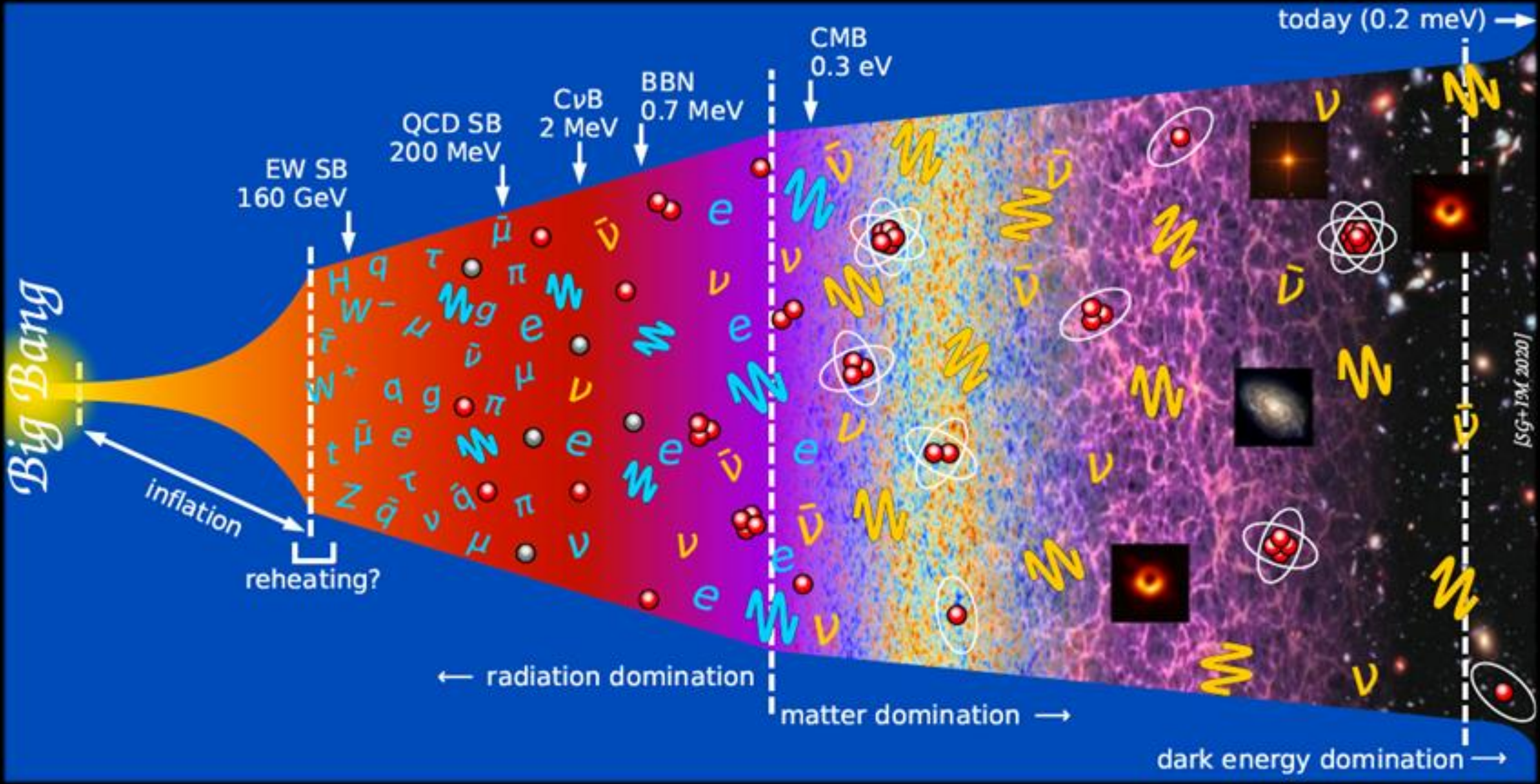
Current constraints on cosmological scenarios with very low reheating temperatures

Sergio Pastor
(IFIC Valencia)

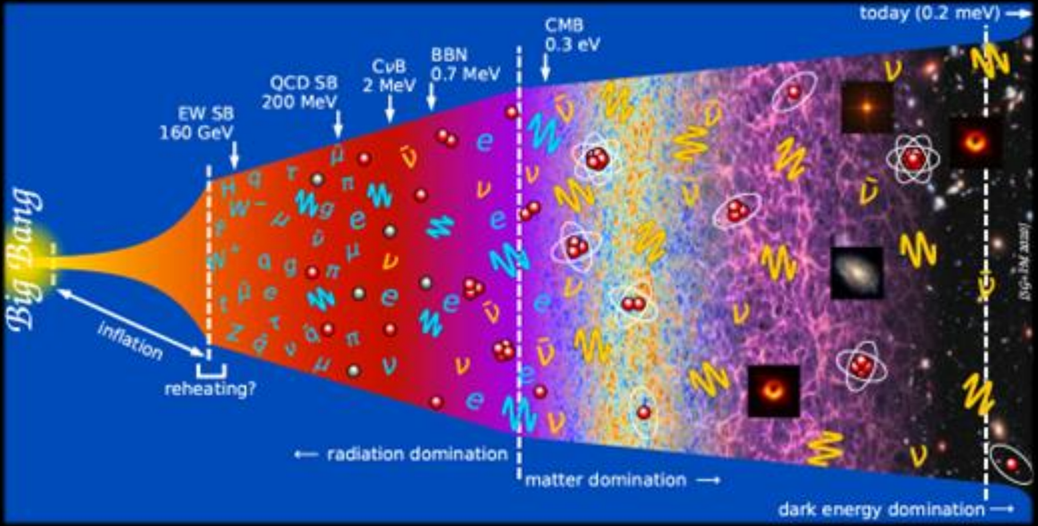
based on arXiv:2501.01369 with N Barbieri, T Brinckmann, S Gariazzo, M Lattanzi & O Pisanti

PASCOS 2025
Durham, 21-25 July

Evolution of the universe



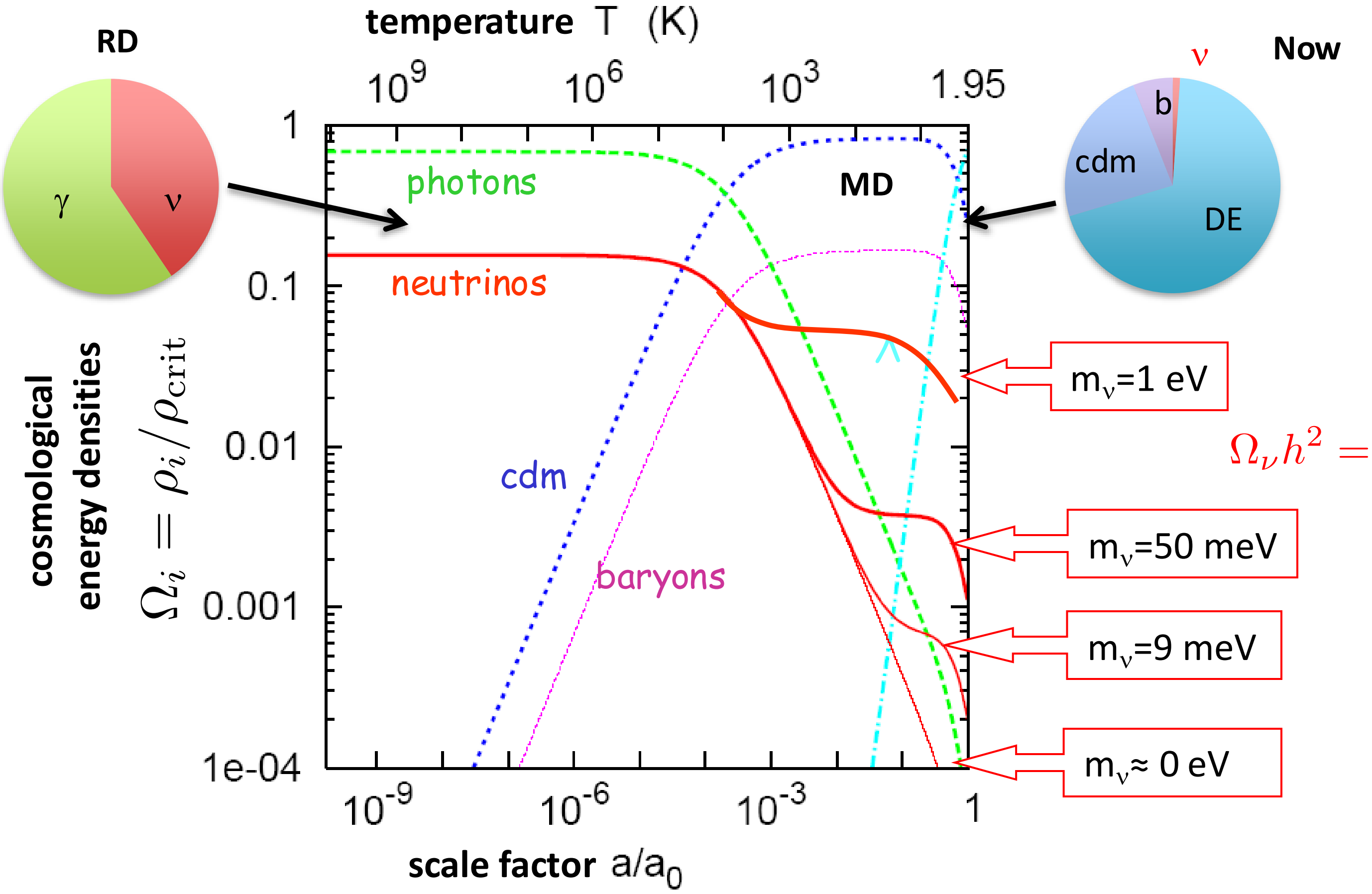
Evolution of the universe



energy density: $\rho(a) = a^{-3(1+w)}$

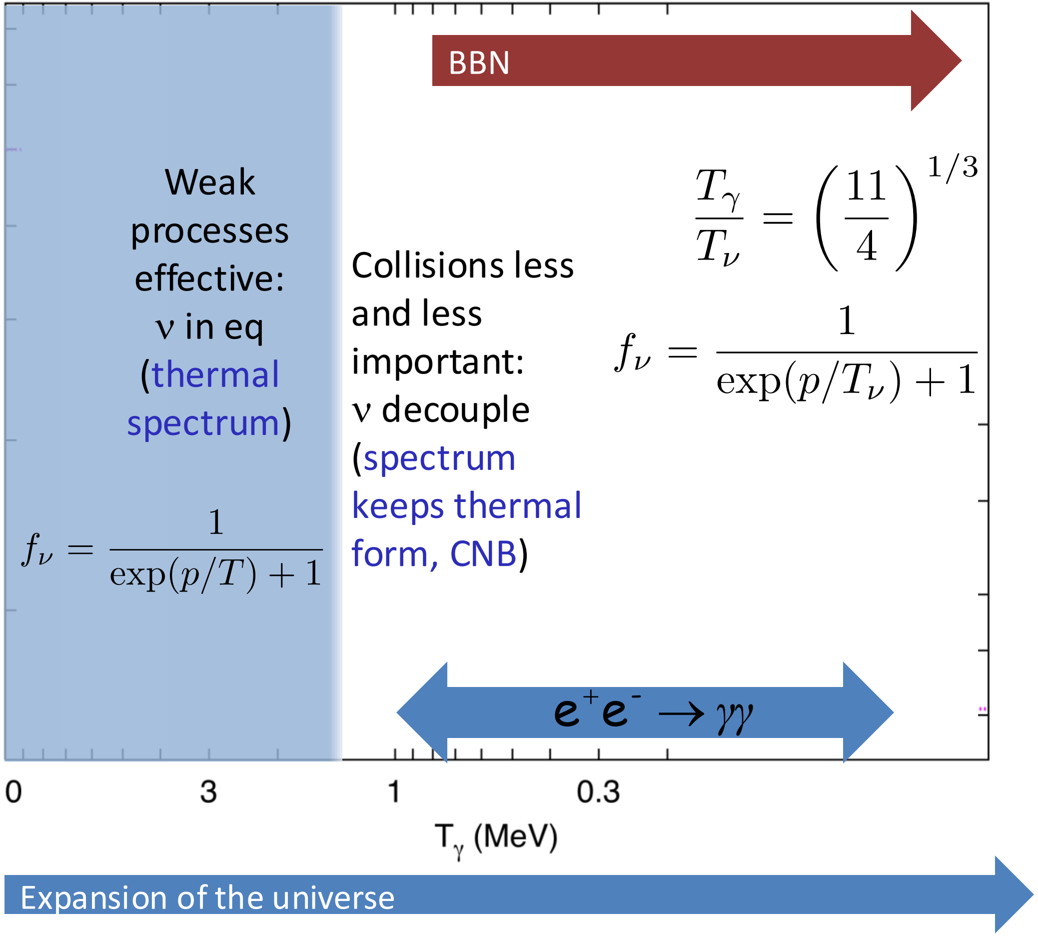
$$\rho_R \sim a^{-4} \quad , \qquad w = 1/3 \qquad \text{(Radiation)}$$
$$\rho_M \sim a^{-3} \quad , \qquad w = 0 \qquad \text{(Matter)}$$
$$\rho_\Lambda \sim \text{const.} \quad , \qquad w = -1 \qquad \text{(Cosmological constant)}$$

Evolution of the background densities: 1 MeV → now



$$\Omega_\nu h^2 = \frac{\sum m_{\nu_i}}{93.2\,\mathrm{eV}}$$

Neutrino decoupling and $e^{\pm}$ annihilation



Relativistic particles in the universe

At $T<m_e$, the radiation content of the Universe is

$$\rho_{\rm rad} = \rho_\gamma + \rho_\nu = \rho_\gamma \left[1 + \frac{7}{8}\left(\frac{4}{11}\right)^{4/3} \times 3\right]$$

Valid for standard neutrinos in the
instantaneous decoupling approximation

Relativistic particles in the universe

At $T<m_e$, the radiation content of the Universe is

$$\rho_{\rm rad} = \rho_\gamma + \rho_\nu + \rho_x = \rho_\gamma \left[1 + \frac{7}{8}\left(\frac{4}{11}\right)^{4/3} N_{\rm eff}\right]$$

effective number of relativistic neutrino species

(effective number of neutrinos)

$N_{\rm eff}$ is a way to measure the ratio $\dfrac{\rho_\nu + \rho_x}{\rho_\gamma}$

K Akita & M Yamaguchi, JCAP 08 (2020) 012

J Froustey, C Pitrou & MC Volpe, JCAP 12 (2020) 015

J J Bennett et al, JCAP 04 (2021) 073

$$N_{\rm eff} = 3.044$$

(3.0440 ± 0.0002)

standard value

Relativistic particles in the universe

At T<m$_e$, the radiation content of the Universe is

$$\rho_{\rm rad} = \rho_\gamma + \rho_\nu + \rho_x = \rho_\gamma \left[1 + \frac{7}{8}\left(\frac{4}{11}\right)^{4/3} N_{\rm eff}\right]$$

effective number of relativistic neutrino species

(effective number of neutrinos)

$N_{\rm eff} \neq 3$

additional relativistic particles (scalars, pseudoscalars, decay products of heavy particles,…)

non-standard neutrino physics (**primordial neutrino asymmetries**, totally or partially thermalised light sterile neutrinos, non-standard interactions with electrons,…)

$N_{\rm eff} = 2.99^{+0.34}_{-0.33}$ **(2018) Planck**

(95%, TT,TE,EE+lowE+lensing+BAO)

Very low reheating scenarios

Cosmological scenarios with low reheating temperatures (T_{RH})

REHEATING (standard picture): phase ending **inflation**

during **inflation**, a **non-relativistic scalar** (inflaton) dominates the energy density

during **reheating**, the **NR scalar** ϕ decays into standard particles

→ photons, $e^{\pm}$, etc are populated directly

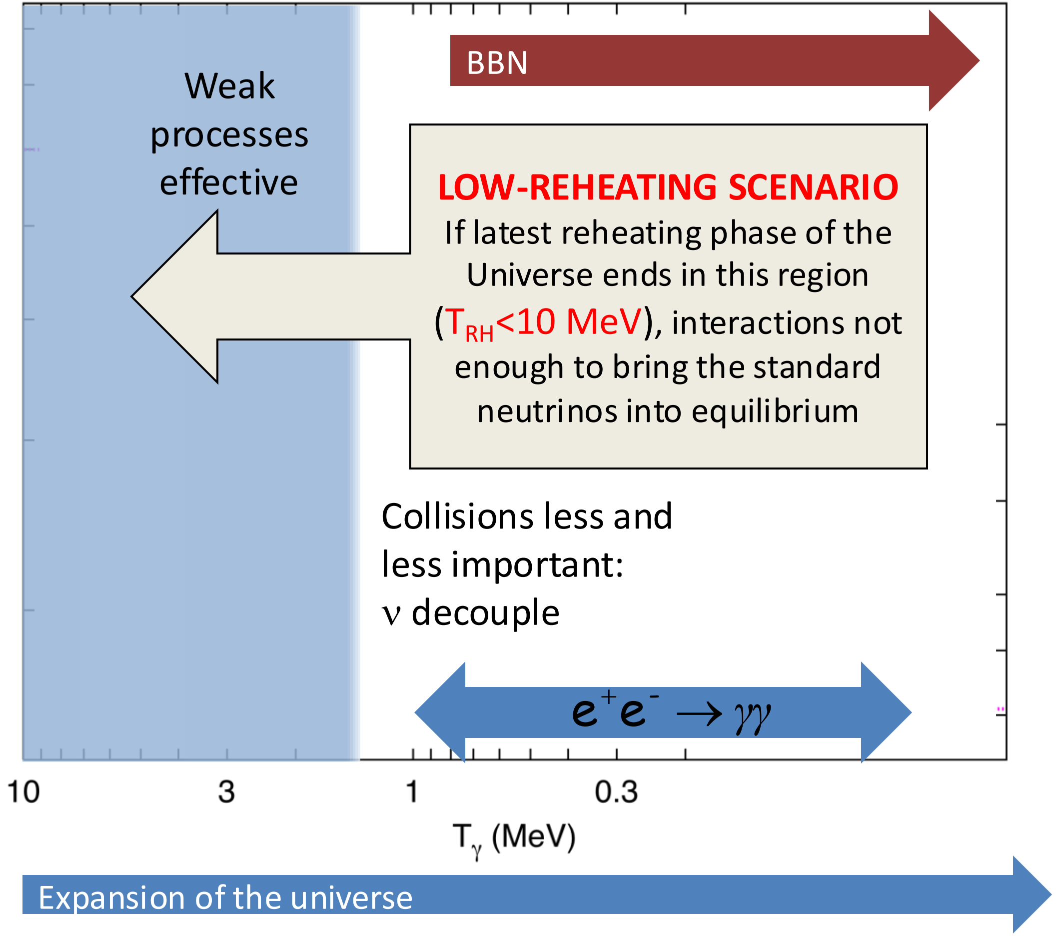
Radiation Domination (RD) begins after **reheating**

---

neutrinos are populated via weak interactions with **charged leptons**

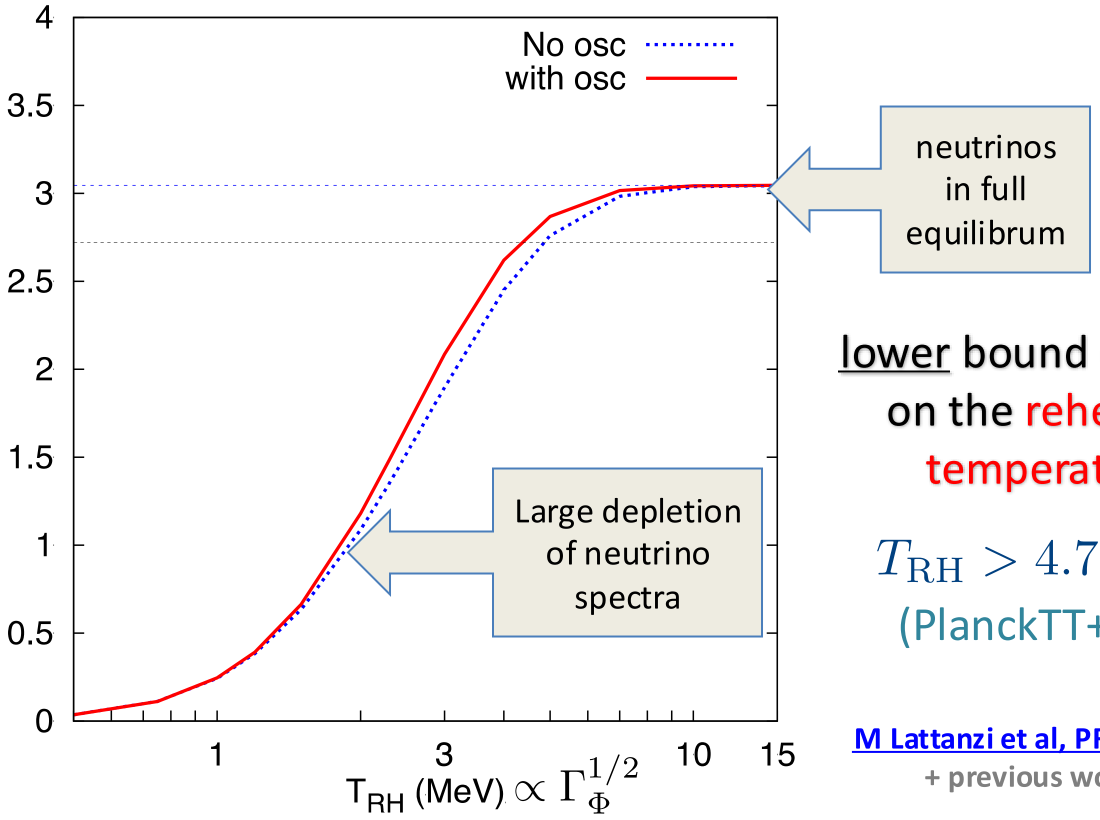
If (last period of) **reheating** occurs **too late:** $T_{\rm RH} \lesssim 10\,{\rm MeV} \rightarrow N_{\rm eff} < 3$

$N_{eff} < 3$?


BBN
Weak processes effective
LOW-REHEATING SCENARIO
If latest reheating phase of the Universe ends in this region (T_{RH}<10 MeV), interactions not enough to bring the standard neutrinos into equilibrium
Collisions less and less important: ν decouple
$e^+e^- \rightarrow \gamma\gamma$
10
3
1
0.3
T_γ (MeV)
Expansion of the universe

3ν in very low-reheating scenarios

$$N_{\rm eff} = \frac{8}{7}\left(\frac{11}{4}\right)^{4/3} \frac{\rho_\nu}{\rho_\gamma}$$



<u>lower</u> bound (95%CL) on the reheating temperature

$$T_{\rm RH} > 4.7\ {\rm MeV}$$

(PlanckTT+ lowP)

M Lattanzi et al, PRD 92 (2015) 123534

+ previous works since 1998

Our work: a more precise calculation of neutrino evolution + BBN production

Boltzmann evolution equations (matrix form)

$$(\partial_t - H p\, \partial_p)\, \varrho_p(t) = -i \left[\left(\frac{1}{2p} \mathbb{M}_\mathrm{F} - \frac{8\sqrt{2} G_\mathrm{F} p}{3 m_\mathrm{W}^2} \mathbb{E} \right), \varrho_p(t) \right] + \mathcal{I}\left[\varrho_p(t)\right]$$

vacuum osc. term

matter potential term

collision integrals (α G_F^2)

take into account neutrino-electron scattering and pair annihilation

2D integrals over the momentum, take most of the computation time

+ continuity equation

$$\dot{\rho} = -3H(\rho + P)$$

+ evolution of scalars

$$\dot{\rho}_\phi = -(3H + \Gamma_\phi)\rho_\phi$$

Code: FORTran-Evolved PrimordIAl Neutrino Oscillations (**FortEPiaNO**)

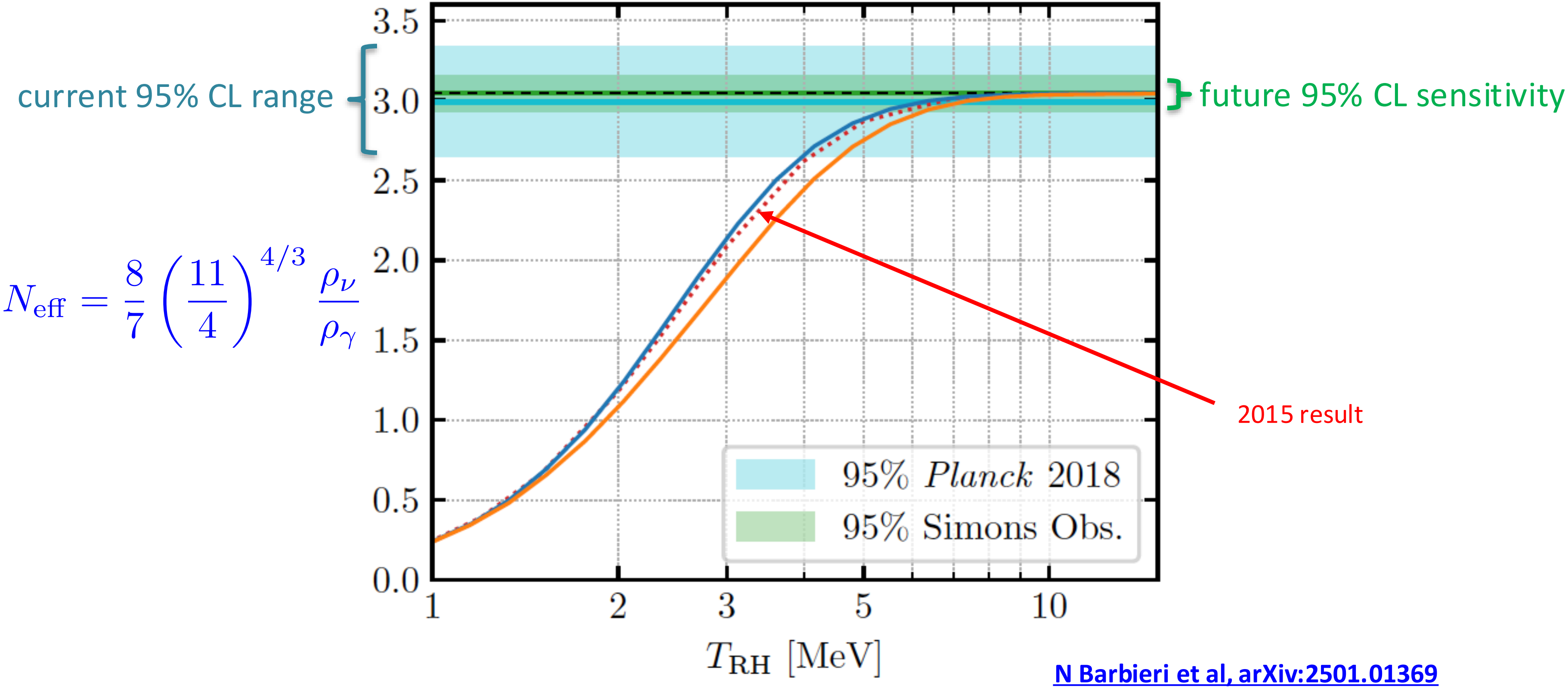
→ $f_{\nu_\alpha}(p, t)$ **Effect on BBN (PArthENoPE code)**

↓ $^4\mathrm{He}$ abundance

→ $\mathrm{Final}\ f_{\nu_i}(p)$ **Analysis of CMB+BAO data**

Final value of N_{eff} (T_{RH})

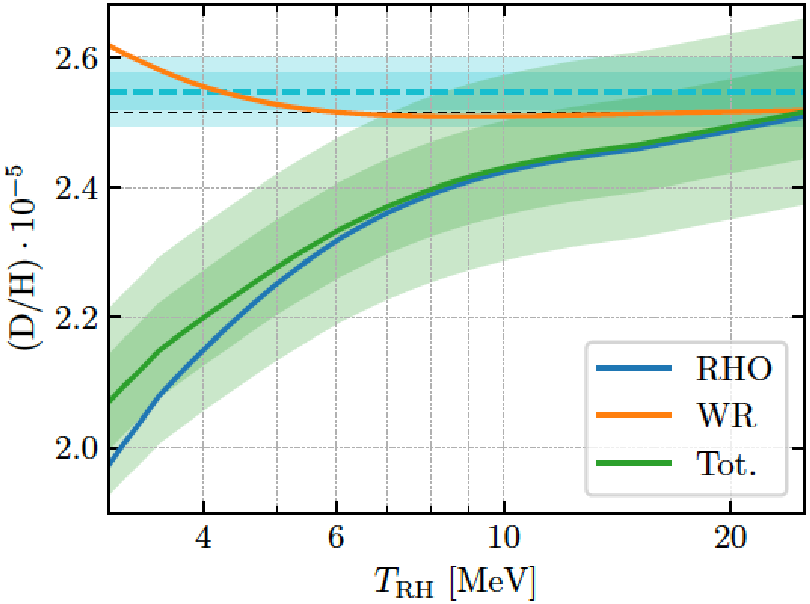
$$N_{\rm eff} = \frac{8}{7}\left(\frac{11}{4}\right)^{4/3}\frac{\rho_\nu}{\rho_\gamma}$$

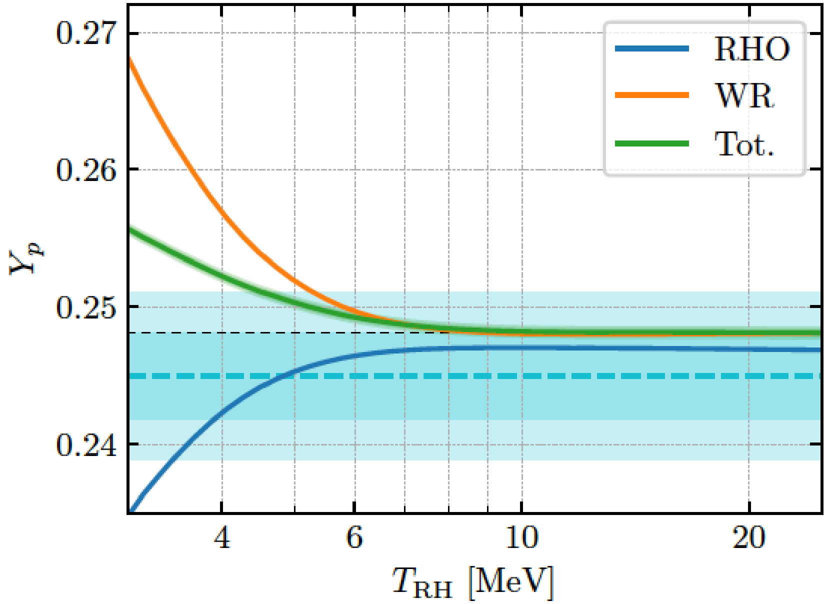


N Barbieri et al, arXiv:2501.01369

Effect on Primordial Nucleosynthesis: PArthENoPE code

current 68/95% CL regions from measurements of primordial D and ^{4}He (PDG)

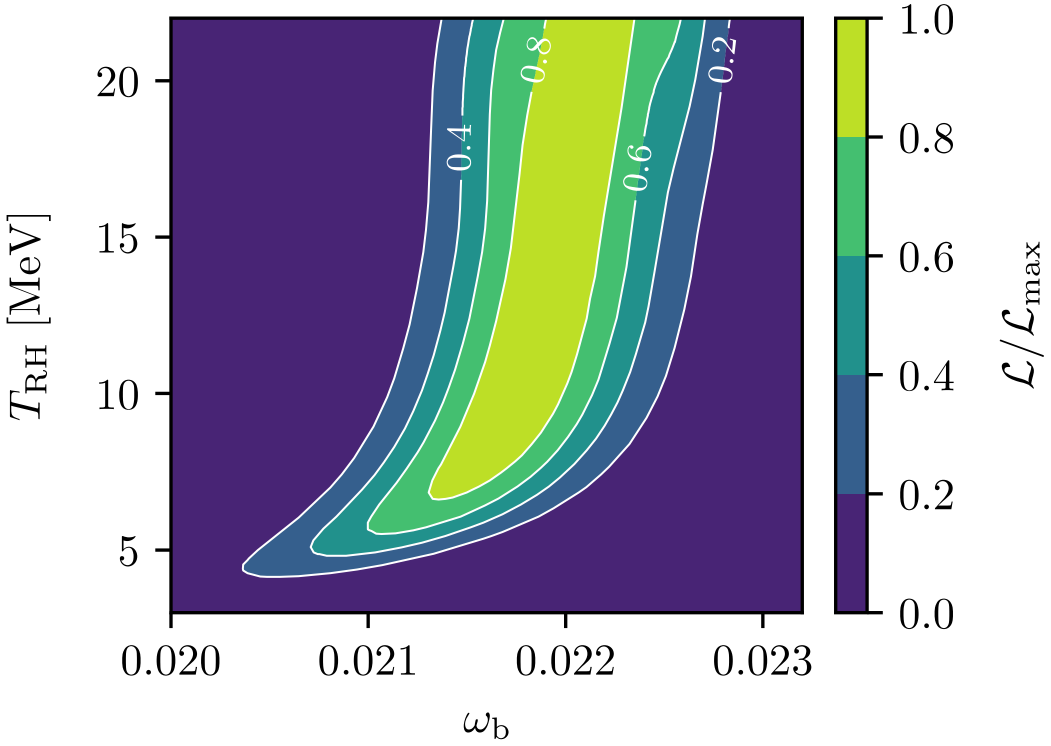




- RHO : effect of N_{eff} only
- WR : effect of $f(\nu_e)$ on weak rates only
- Tot. : **TOTAL effect:** D/H decreases and Y_p increases for **smaller T_{RH}**

N Barbieri et al, arXiv:2501.01369

BBN likelihood function (T_{RH})

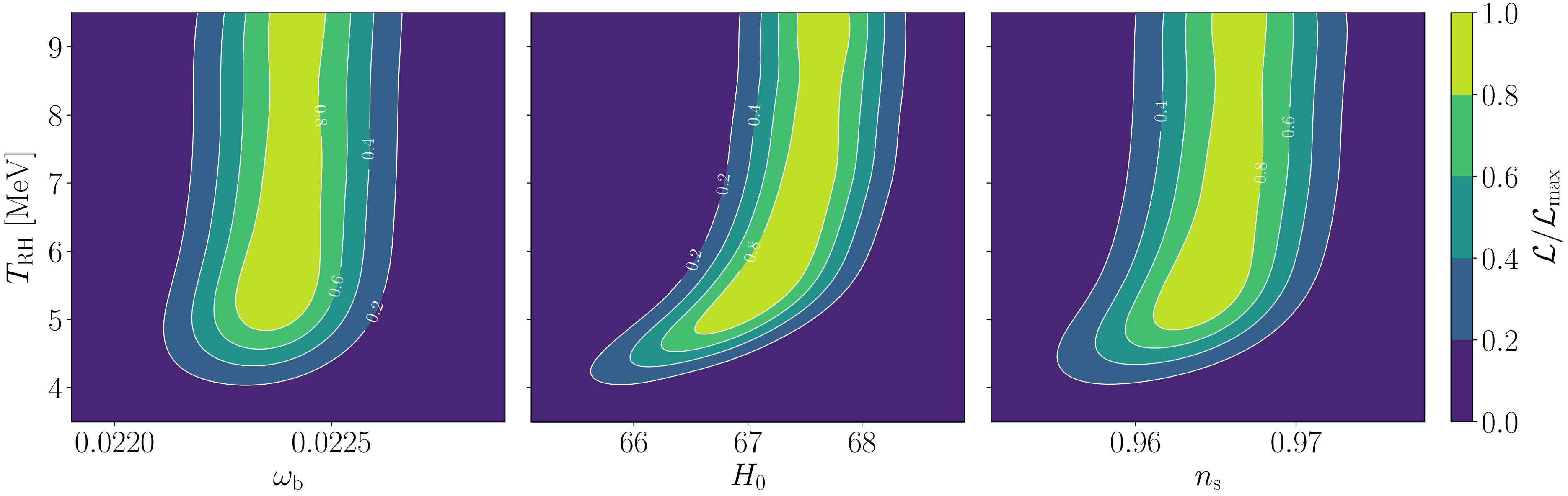


BBN only

$T_{\rm RH} > 3.67\,{\rm MeV}$ (95% CL)

N Barbieri et al, arXiv:2501.01369

Planck+lensing+DESI likelihood function (T_{RH})

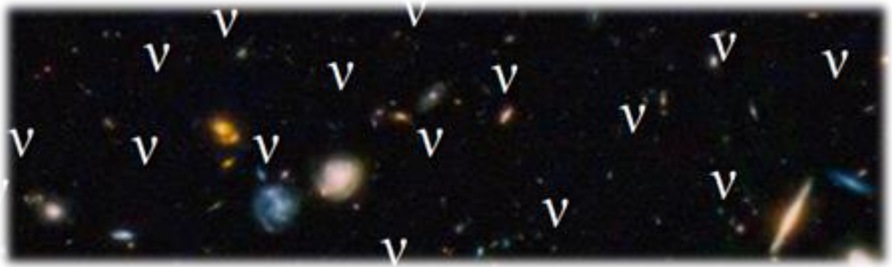


$T_{\rm RH} > 3.79$ MeV (Planck+lensing+DESI)

$T_{\rm RH} > 5.96$ MeV (BBN+Planck+lensing+DESI)

N Barbieri et al, arXiv:2501.01369

Conclusions



✓ We solved the momentum-dependent kinetic equations for neutrinos in the early universe, including flavour oscillations, in a **very low reheating scenario**: neutrino spectra are depleted ($N_{eff} < 3$)

✓ A consistent BBN+Planck+lensing+DESI analysis leads to the **most stringent bound** to date on the reheating temperature: $T_{\rm RH} > 5.96\,{\rm MeV}\ (95\%\ {\rm CL})$