

Right-Handed Sneutrino Modification of D-term Hybrid Inflation in Non-minimal Supergravity

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Topics

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- ▶ What is D-term inflation?
- ▶ What are the problems with D-term Inflation?
- ▶ Provide a solution.

Inflation

Once upon a time, the Universe was dominated by some mysterious vacuum energy:

$$V(\phi) = 3H^2 M_{Pl}^2 \approx \rho_\phi \approx \text{const.}$$

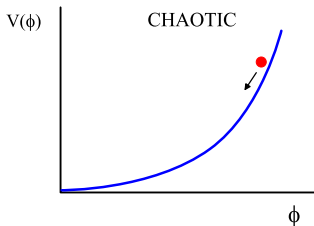
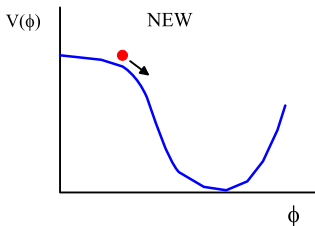
$$\therefore \dot{\rho} \approx 0 \Rightarrow H \equiv \frac{\dot{a}}{a} \approx \text{const.}$$

$$\frac{da}{a} = H dt \Rightarrow a \propto e^{Ht}$$

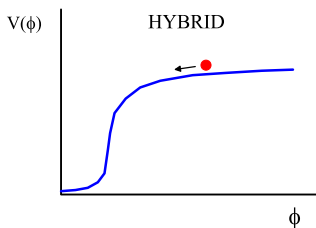
People named this dramatic expansion “Inflation”. The problem is, what in the world is $V(\phi)$ after all?!

Inflation

So people start to play the game of Inflation by inventing all kinds of potentials...



Inflation



Inflation

Slow-roll inflation:

A scalar potential of the field(s) $V(\phi, \dots)$ tells you almost everything! $\eta = M^2 V''/V$ and $\epsilon = (M^2/2)(V'/V)^2$ Spectral index:

$$n = 1 + 2\eta - 6\epsilon$$

Curvature perturbation:

$$P_\zeta = \frac{1}{12\pi^2 M^6} \frac{V^3}{V'^2}$$

etc.

Supersymmetry (Boson \leftrightarrow Fermion)

Every particle corresponds to its superpartner, which must be HEAVY due to SUSY breaking.



Supersymmetry

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- ▶ etc.

Supersymmetry

Scalar potential

$$V = V_F + V_D$$

SUSY:

$$V_F = |F|^2 \equiv |W_m|^2$$

$$V_D = \frac{1}{2} \left(\frac{g}{2} \sum_n q_n |\phi_n|^2 + \xi \right)^2$$

SUGRA:

$$V_F = e^{K/M^2} \left[\left(W_m + \frac{WK_m}{M^2} \right) K^{m\dagger n} \left(W_n + \frac{WK_n}{M^2} \right) - \frac{3|W|^2}{M^2} \right]$$

$$V_D = \frac{1}{2} (\text{Ref})^{-1} g^2 (q_n K_n \phi^n + \xi)^2$$

Conventional D-term Inflation

The superpotential of D-term hybrid inflation is given by

$$W_D = \lambda S \Phi_+ \Phi_-$$

One-loop potential:

$$V(S) = V_0 + \frac{g^4 \xi^2}{16\pi^2} \ln\left(\frac{|S|^2}{\Lambda^2}\right)$$

Problems of conventional D-term Inflation

From WMAP

$$n = 0.958 \pm 0.016$$

However,

$$n = 0.983$$

From cosmic string study (for example astro-ph:0702223):

$$\xi^{1/2} < 3.9 \times 10^{15} \text{ GeV}$$

(assume an upper limit of 10%) However,

$$\xi^{1/2} = 7.9 \times 10^{15} \text{ GeV}$$

Non-minimal Supergravity

We consider the Kähler potential of the form:

$$K = |S|^2 + |\Phi_+|^2 + |\Phi_-|^2 + f_+ \left(\frac{|S|^2}{M_P^2} \right) |\Phi_+|^2 + f_- \left(\frac{|S|^2}{M_P^2} \right) |\Phi_-|^2 + b \frac{|S|^4}{M_P^2}$$

where

$$f_{\pm} \left(\frac{|S|^2}{M_P^2} \right) = c_{\pm} \frac{|S|^2}{M_P^2} = \frac{c_{\pm} s^2}{2}.$$

D-term Inflation

The effective scalar potential for the D-term inflationary model reads

$$V_{\text{eff}}(|S|) = \frac{g^2 \xi^2}{2} \left\{ 1 + \frac{g^2}{16\pi^2} \left[2 \ln\left(z \frac{g^2 \xi}{\Lambda^2}\right) + f_V(z) \right] \right\}.$$

where

$$f_V(z) = (z + 1)^2 \ln\left(1 + \frac{1}{z}\right) + (z - 1)^2 \ln\left(1 - \frac{1}{z}\right)$$

and

$$z \equiv \frac{\lambda^2 s^2}{2g^2 \xi} \exp\left(\frac{s^2}{2} + b \frac{s^4}{4}\right) \frac{1}{(1 + f_+)(1 + f_-)}$$

right-handed sneutrino correction

$$W_\nu = \lambda_\nu \Phi H_u L + \frac{m_\Phi}{2} \Phi^2$$

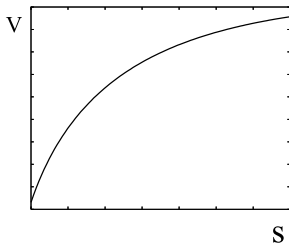
$$K = S^\dagger S + \Phi^\dagger \Phi + \frac{c S^\dagger S \Phi^\dagger \Phi}{M^2}$$

$$\Delta V = -\frac{\kappa |S|^2}{2} + \frac{M_\phi^2 \phi^2}{2},$$

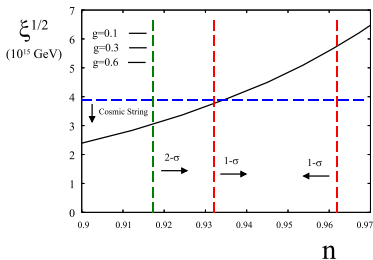
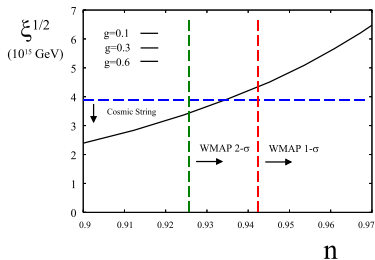
where we have defined κ by

$$\kappa \equiv \frac{(c-1)M_\phi^2|\phi|^2}{M^2}.$$

results

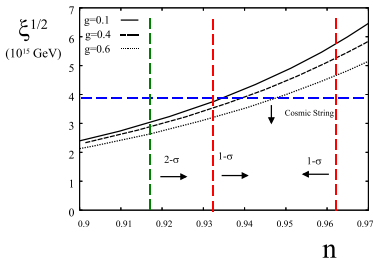
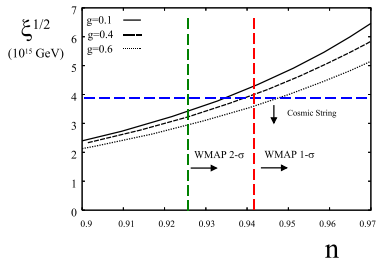


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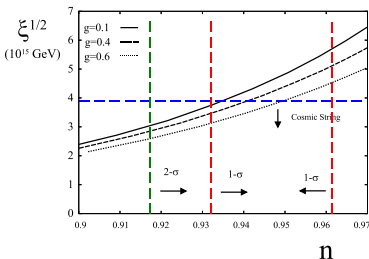
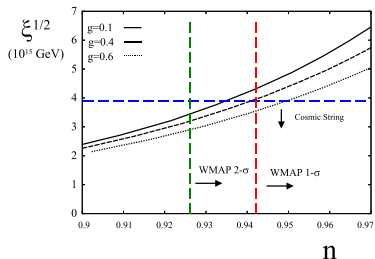
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$$c_+ = c_- = 3.0$$



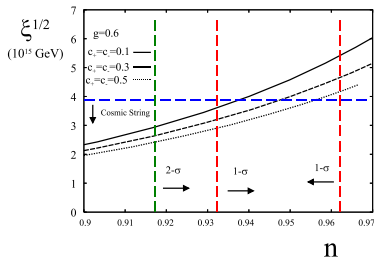
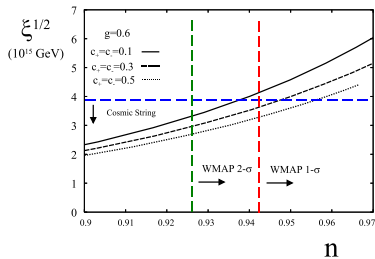
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$$c_+ = 2.0, c_- = 5.0$$

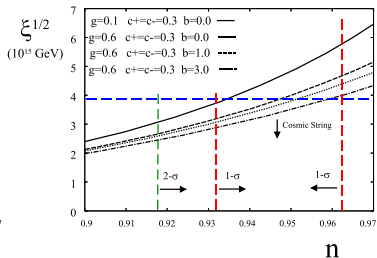
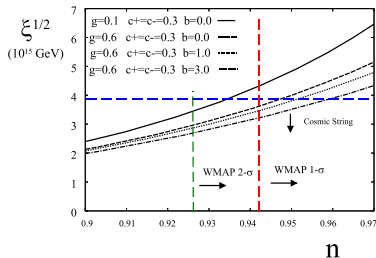


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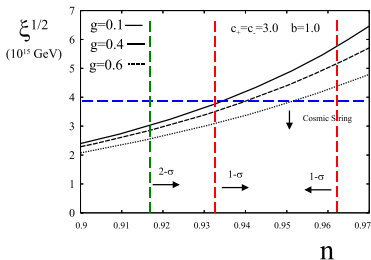
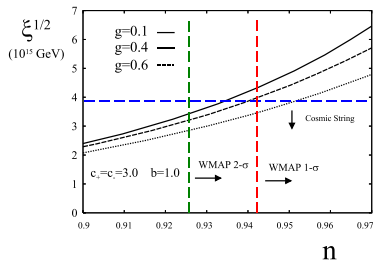
$$g = 0.6$$



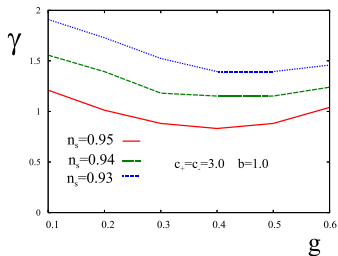
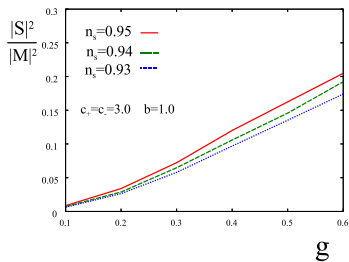
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