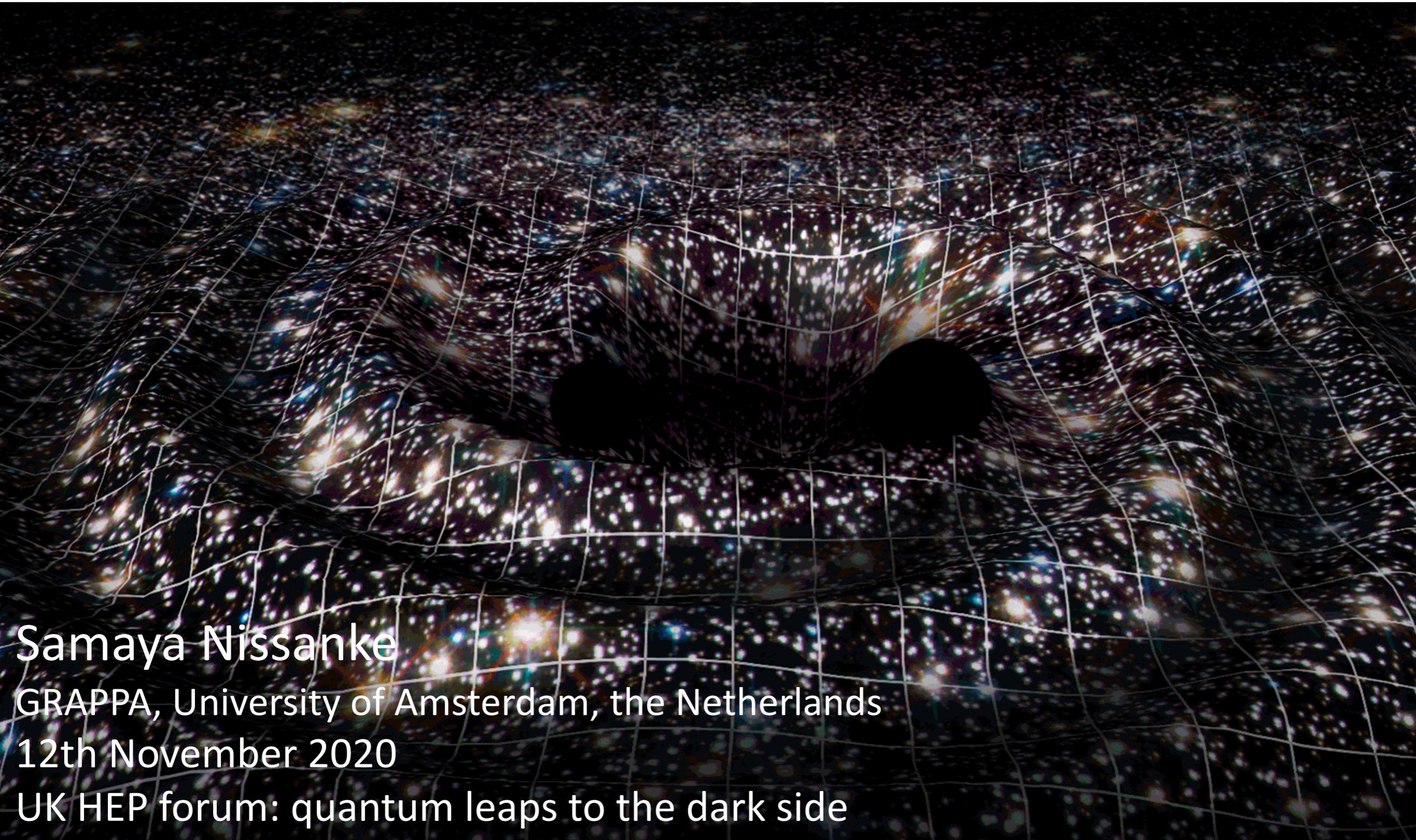


New perspectives onto the Universe in the era of multi-messenger astronomy (MMA)



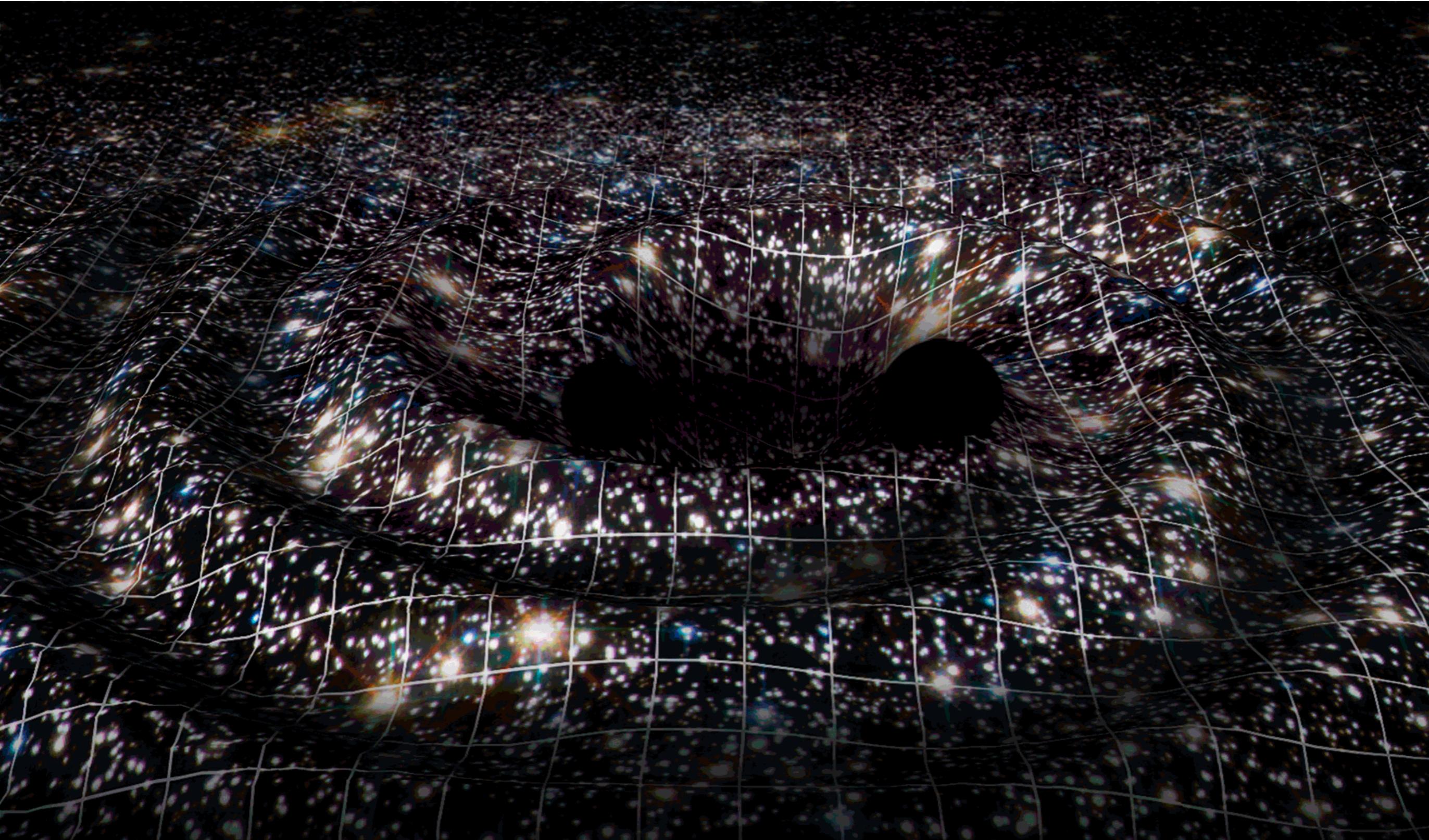
Samaya Nissanke

GRAPPA, University of Amsterdam, the Netherlands

12th November 2020

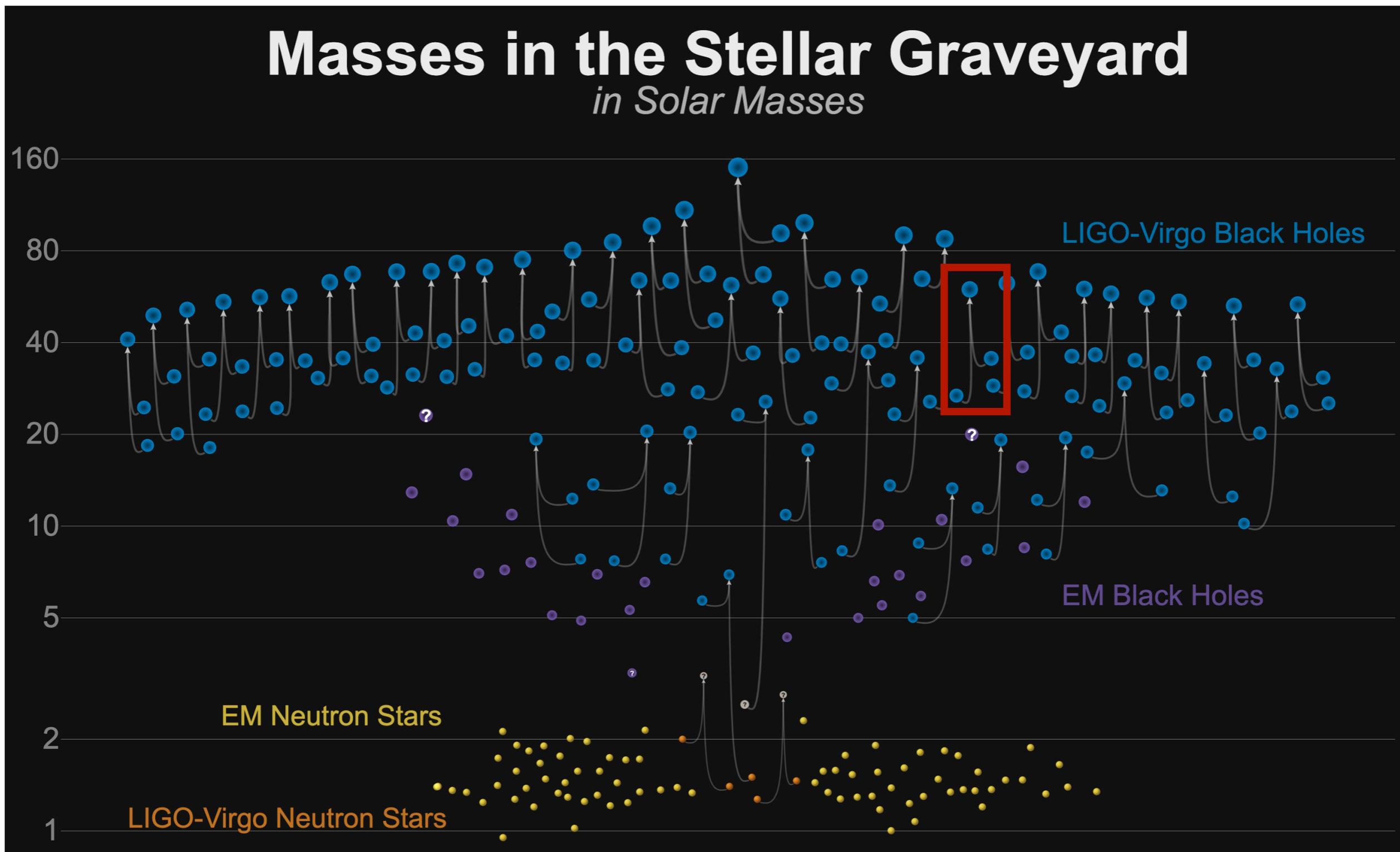
UK HEP forum: quantum leaps to the dark side

New perspectives onto the Universe in the era of multi-messenger astronomy (MMA)



A revolution in the past five years:

gravitational waves (GW), black holes and neutron stars



[1 solar mass or $1 M_{\odot}$ is $\sim 2 \times 10^{30}$ kg]

GWTC-2 plot v1.0

[see LVKC, arXiv 2010.14527 (GWTC 2)
LVC, arXiv 1811.15007 (GWTC-1)]

LIGO-Virgo | Frank Elavsky, Aaron Geller | Northwestern

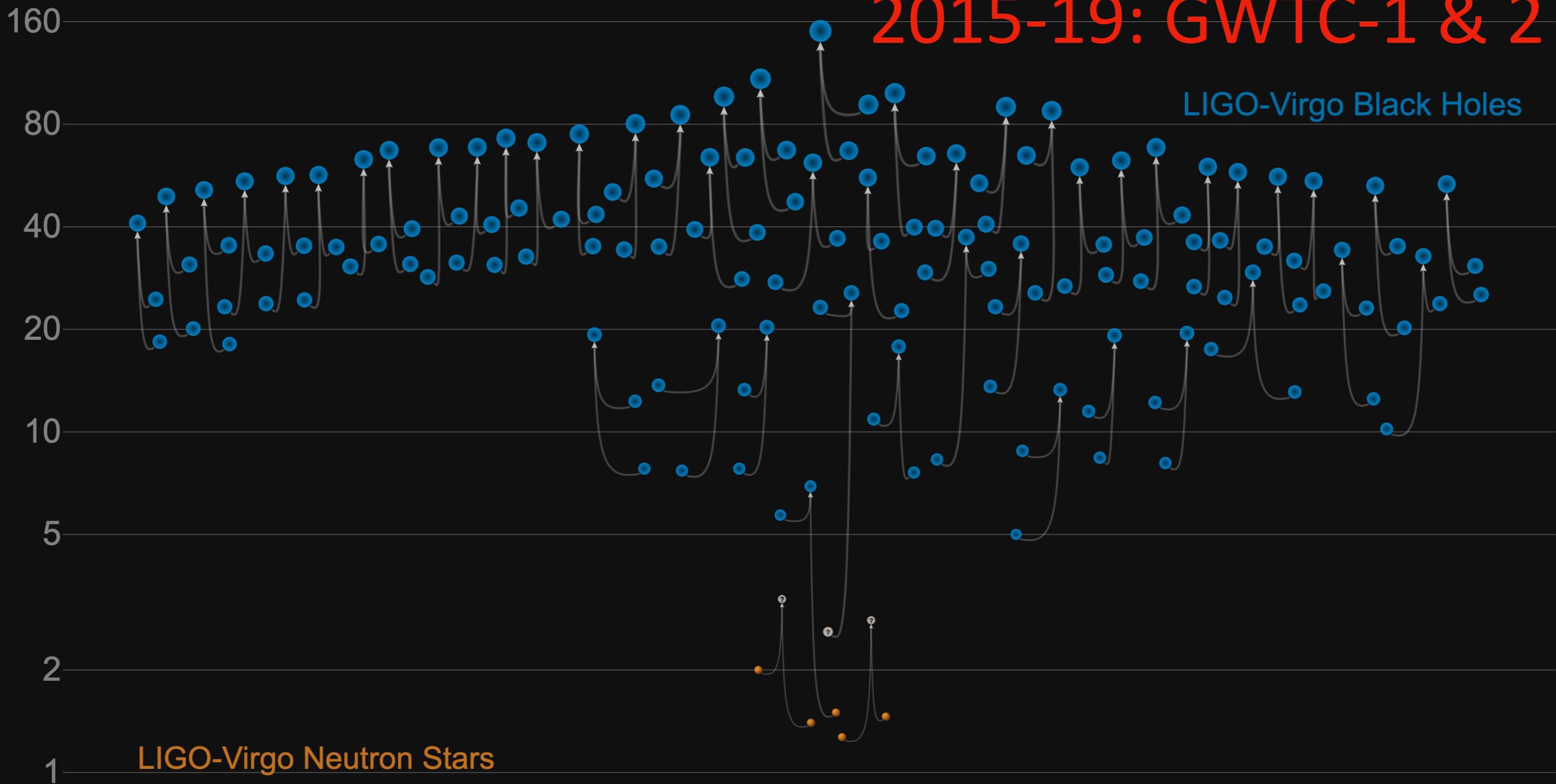
Masses in the Stellar Graveyard

in Solar Masses

2015-19: GWTC-1 & 2

LIGO-Virgo Black Holes

LIGO-Virgo Neutron Stars



[see LVKC, arXiv 2010.14527 (GWTC 2)
LVC, arXiv 1811.15007 (GWTC-1)]

GWTC-2 plot v1.0

LIGO-Virgo | Frank Elavsky, Aaron Geller | Northwestern

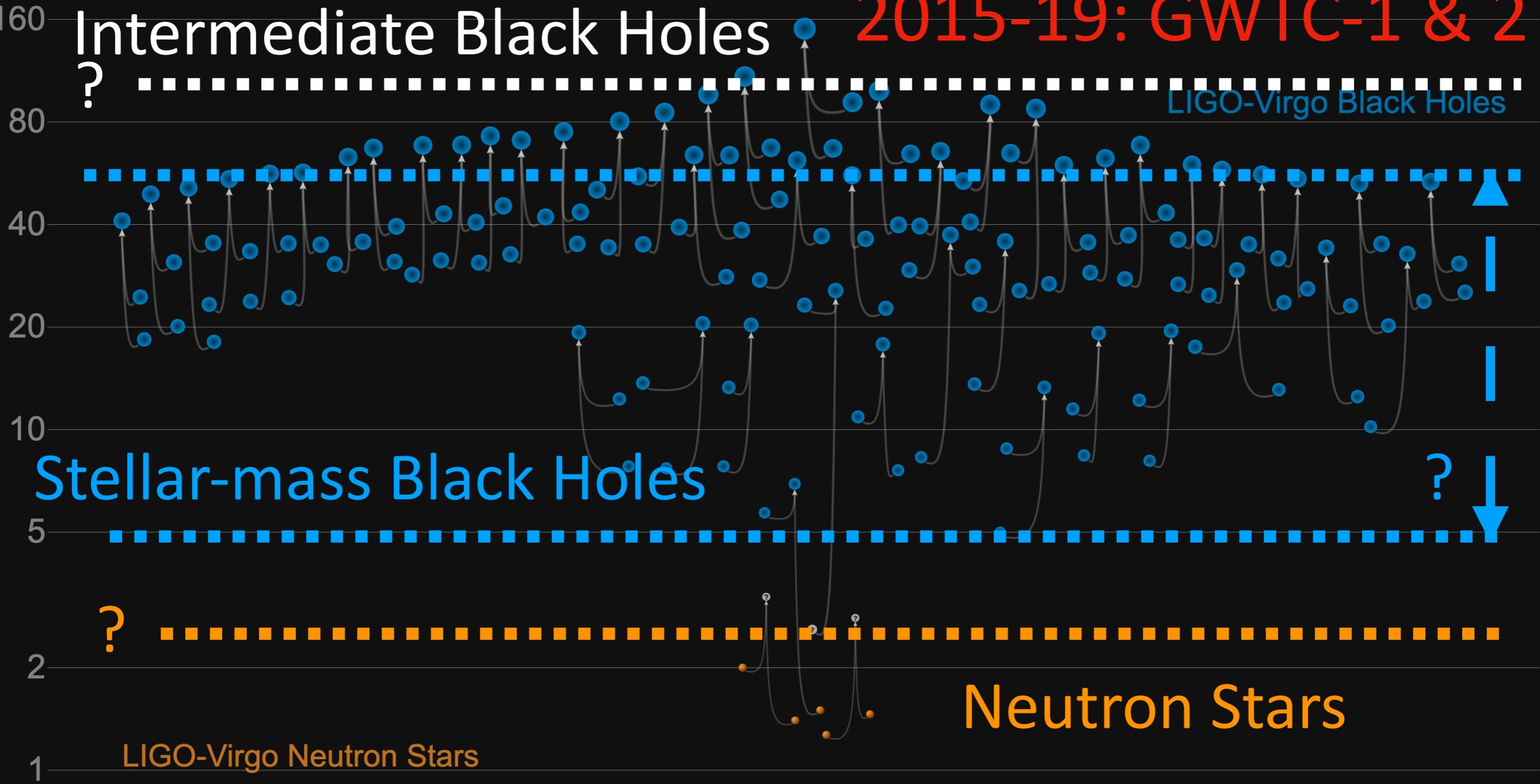
Masses in the Stellar Graveyard

in Solar Masses

2015-19: GWTC-1 & 2

Intermediate Black Holes

?  LIGO-Virgo Black Holes



Stellar-mass Black Holes

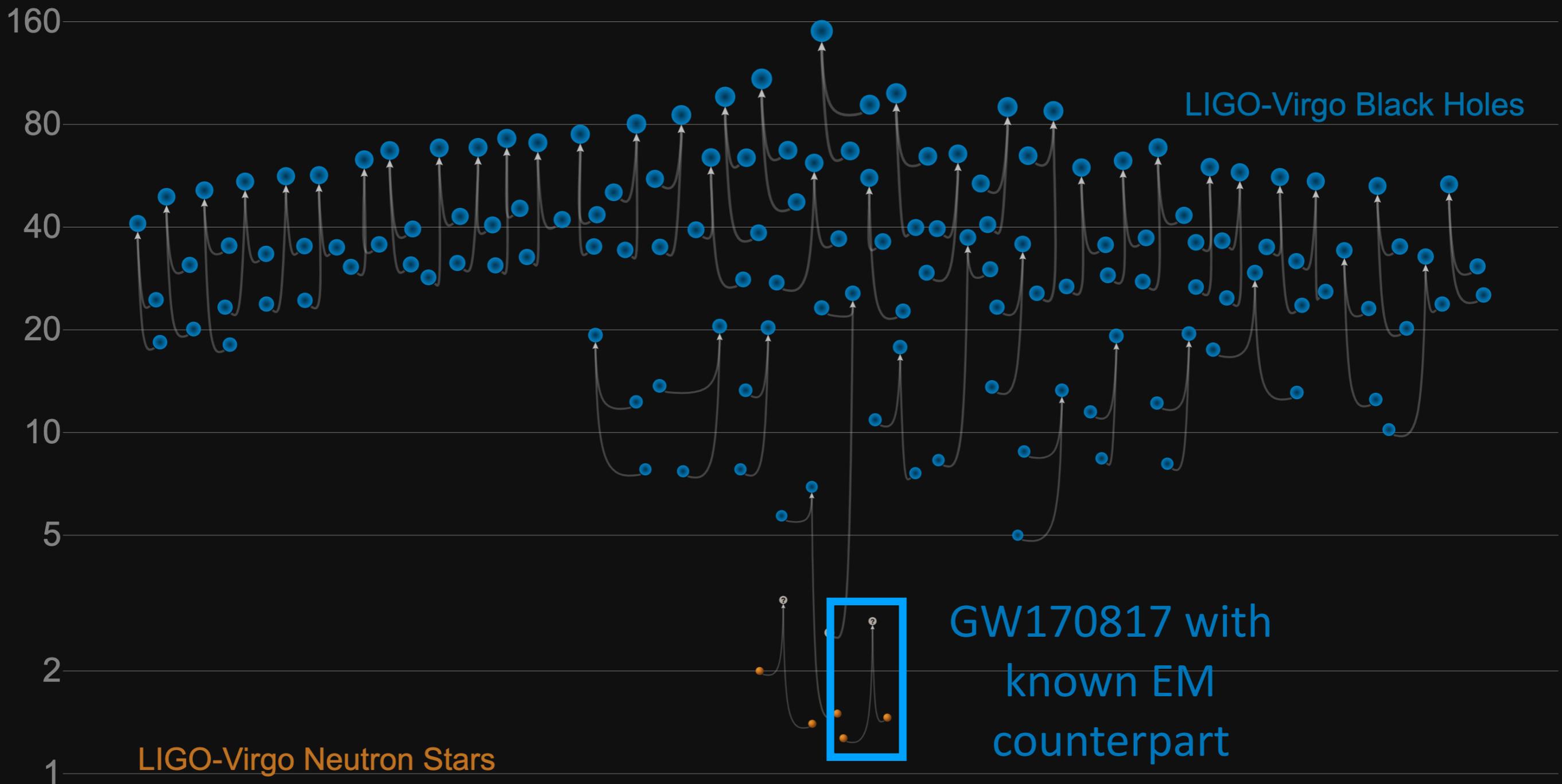
Neutron Stars

LIGO-Virgo Neutron Stars

[see LVKC, arXiv 2010.14527 (GWTC 2)
LVC, arXiv 1811.15007 (GWTC-1)]

Masses in the Stellar Graveyard

in Solar Masses

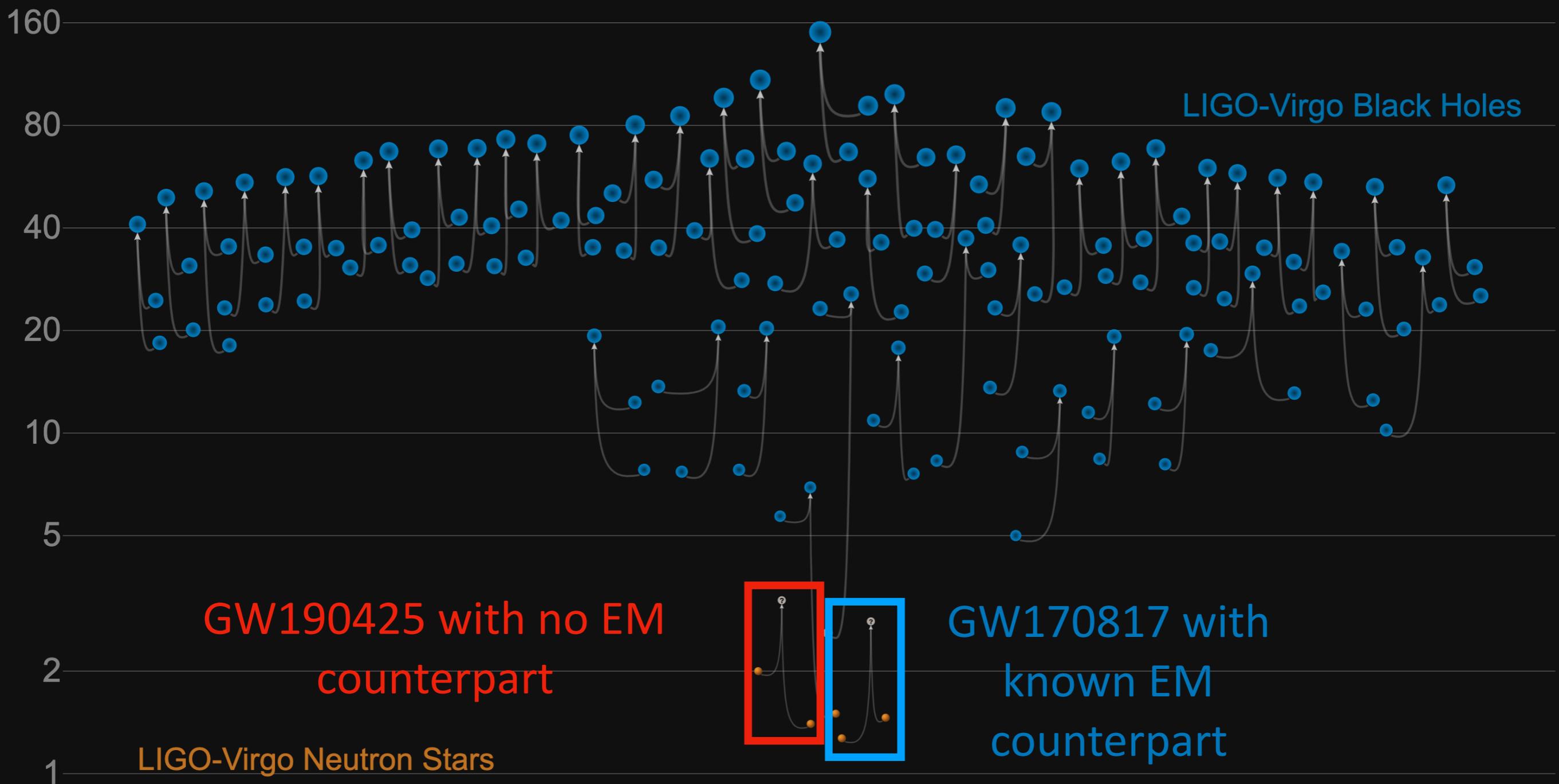


GWTC-2 plot v1.0

LIGO-Virgo | Frank Elavsky, Aaron Geller | Northwestern

Masses in the Stellar Graveyard

in Solar Masses

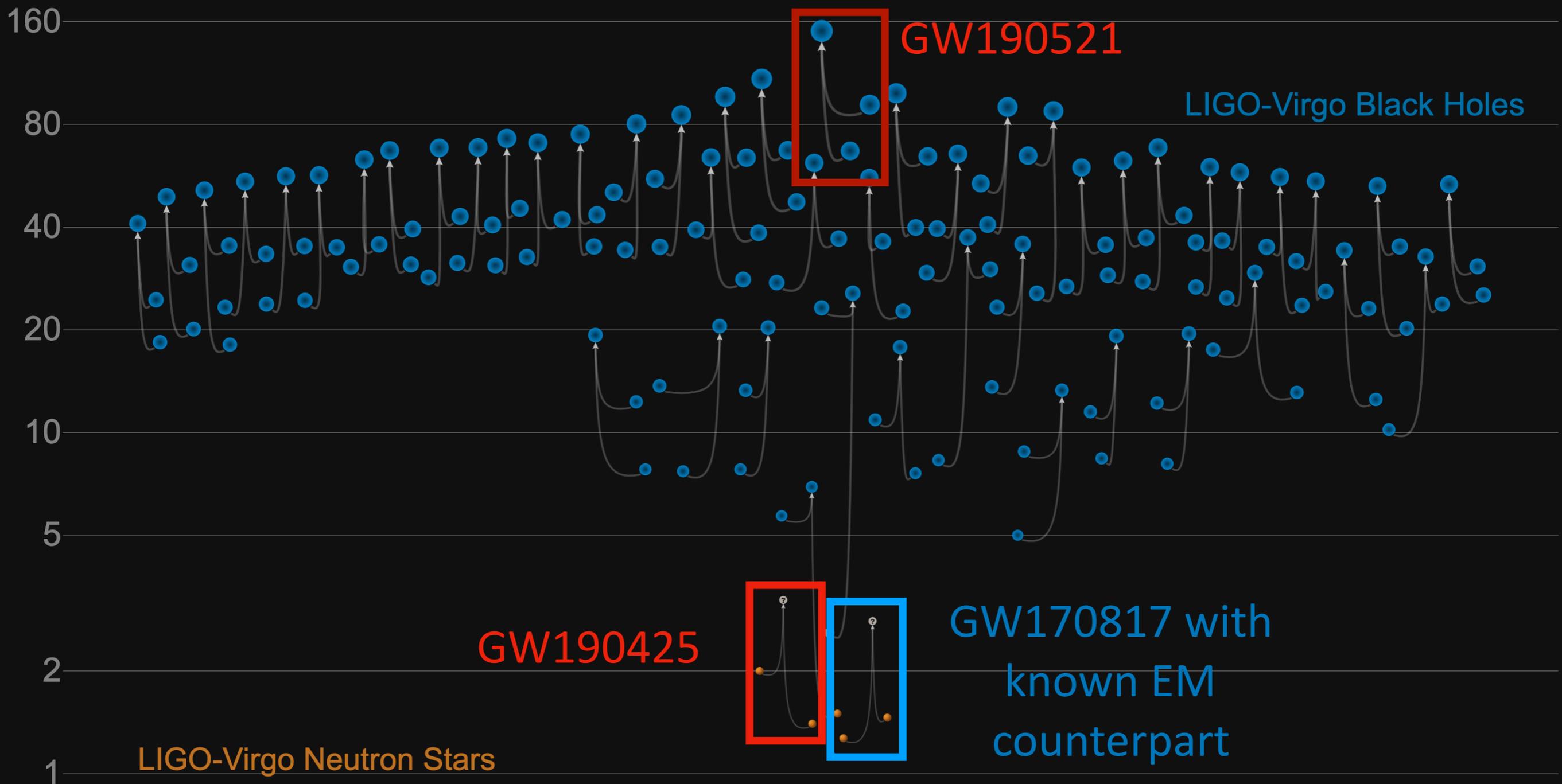


GWTC-2 plot v1.0

LIGO-Virgo | Frank Elavsky, Aaron Geller | Northwestern

Masses in the Stellar Graveyard

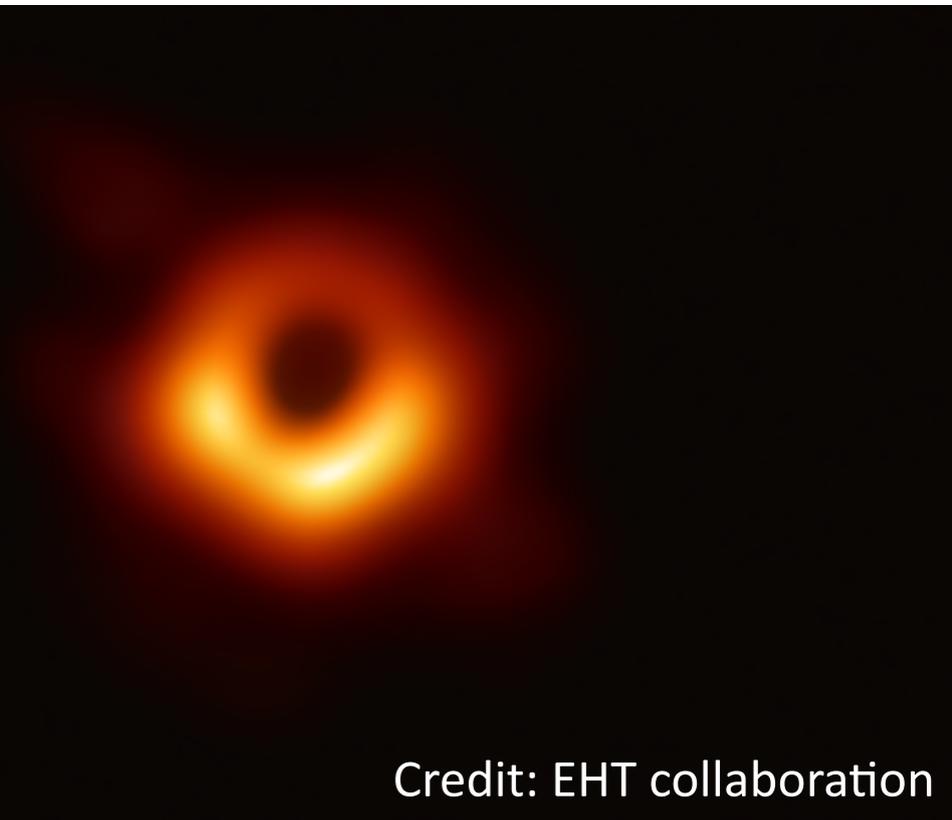
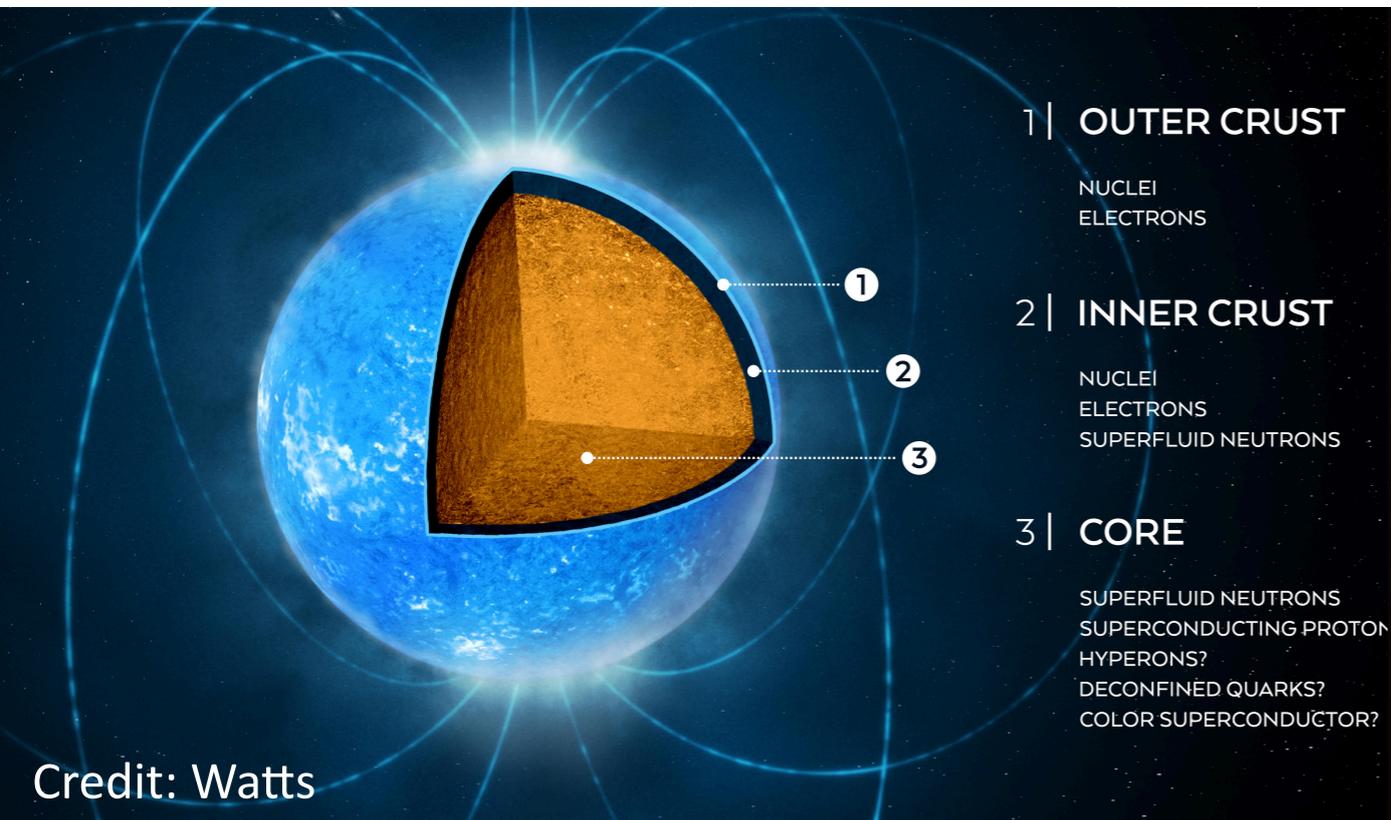
in Solar Masses



GWTC-2 plot v1.0

LIGO-Virgo | Frank Elavsky, Aaron Geller | Northwestern

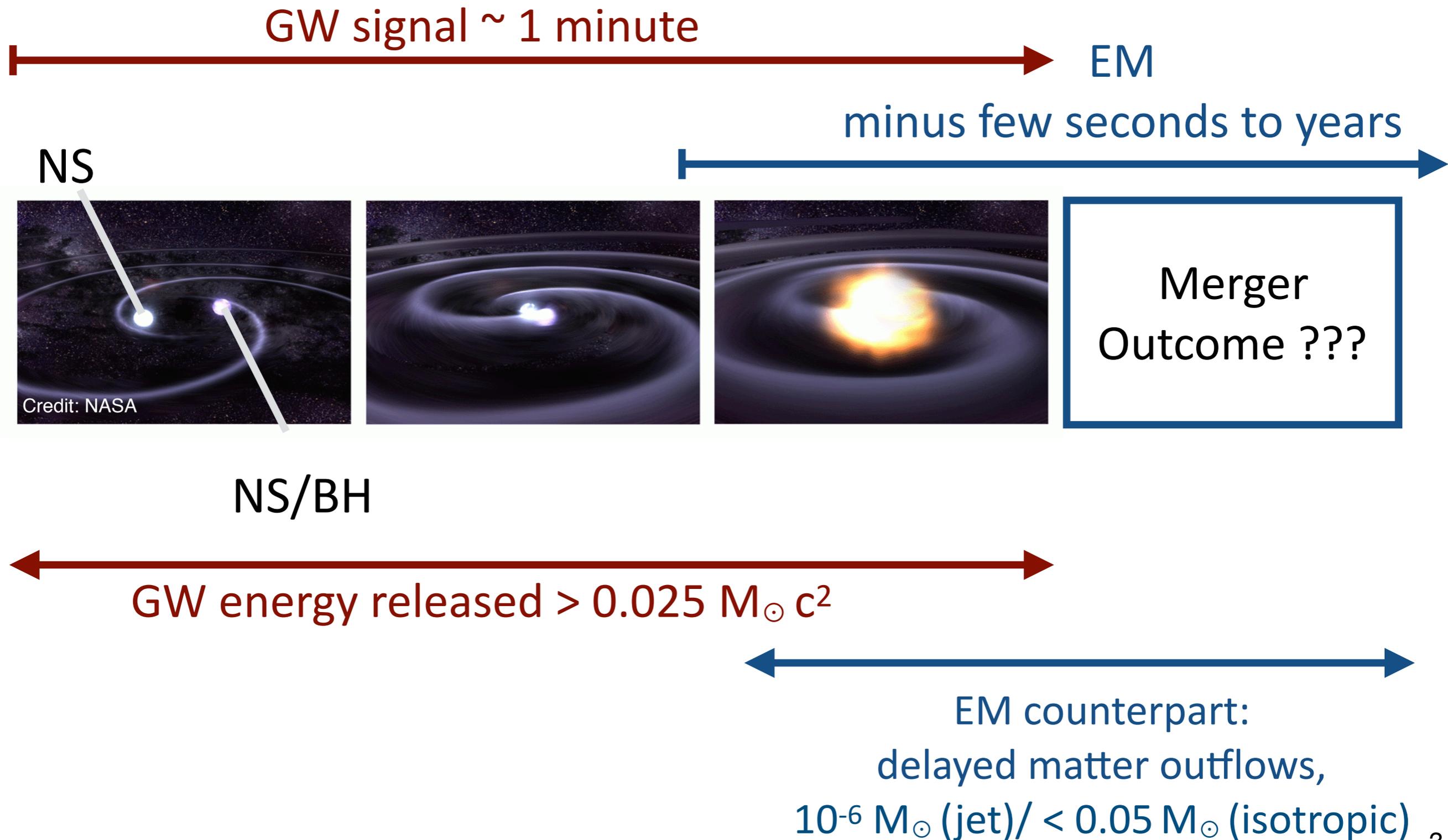
Extreme Matter, Extreme Spacetime Curvature!

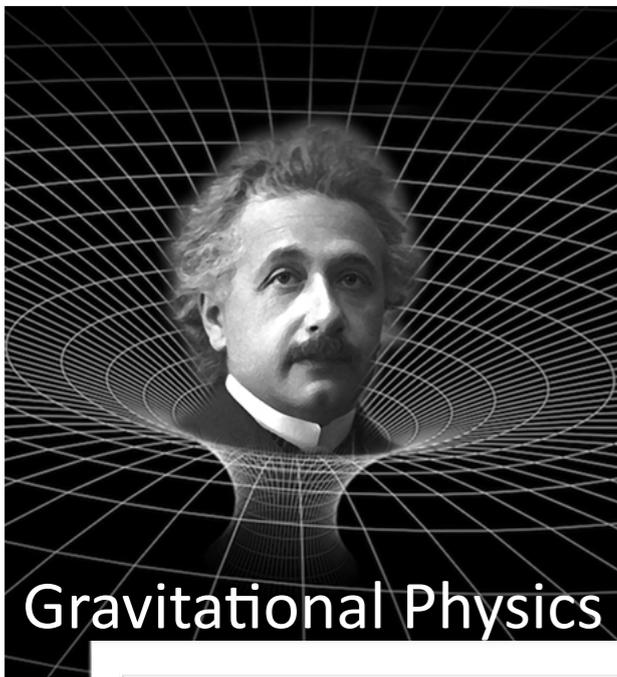


Neutron Stars (NSs) $> 1.2 - 2.7 M_{\odot}$ (?),
~10-15 km radii,
Large magnetic fields,
Extreme matter and exotic equations
of state

Black Holes (BHs): $< 10^{10} M_{\odot}$,
No hair: mass and spin,
Mathematical singularities.

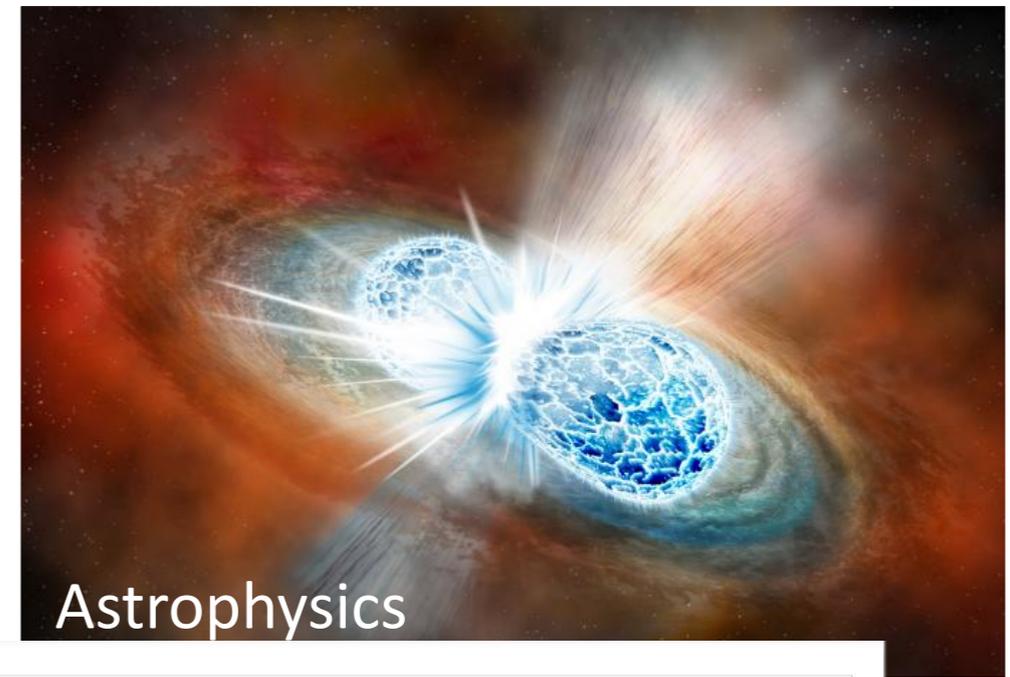
The game changer: First NS Binary Merger detected in GWs & Electromagnetic (EM) radiation!





Gravitational Physics

GW170817: The source of many discoveries



Astrophysics

First Binary Neutron Star detected in Gravitational Waves
First Electromagnetic Counterpart of a GW merger in every waveband!
First Gravitational Wave Standard Siren Hubble Constant Constraint
First Short Gamma Ray Burst - Binary Neutron Star Merger Association
First kilonova discovery and astrophysical sites of r-process heavy elements
First tests of the speed of light and gravity with a GW+EM event ...

Nuclear Physics

| | | | | | | | | | |
|-------|-------|-------|-------|-------|-------|-------|-------|-------|-------|
| 1 H | 2 He | 3 Li | 4 Be | 5 B | 6 C | 7 N | 8 O | 9 F | 10 Ne |
| 11 Na | 12 Mg | 13 Al | 14 Si | 15 P | 16 S | 17 Cl | 18 Ar | 19 K | 20 Ca |
| 21 Sc | 22 Ti | 23 V | 24 Cr | 25 Mn | 26 Fe | 27 Co | 28 Ni | 29 Cu | 30 Zn |
| 31 Ga | 32 Ge | 33 As | 34 Se | 35 Br | 36 Kr | 37 Rb | 38 Sr | 39 Y | 40 Zr |
| 41 Nb | 42 Mo | 43 Tc | 44 Ru | 45 Rh | 46 Pd | 47 Ag | 48 Cd | 49 In | 50 Sn |
| 51 Sb | 52 Te | 53 I | 54 Xe | 55 Cs | 56 Ba | 57 La | 58 Ce | 59 Pr | 60 Nd |
| 61 Pm | 62 Sm | 63 Eu | 64 Gd | 65 Tb | 66 Dy | 67 Ho | 68 Er | 69 Tm | 70 Yb |
| 71 Lu | 72 Hf | 73 Ta | 74 W | 75 Re | 76 Os | 77 Ir | 78 Pt | 79 Au | 80 Hg |
| 81 Tl | 82 Pb | 83 Bi | 84 Po | 85 At | 86 Rn | 87 Fr | 88 Ra | 89 Ac | 90 Th |
| 91 Pa | 92 U | | | | | | | | |

Merging Neutron Stars Exploding Massive Stars Big Bang
Dying Low Mass Stars Exploding White Dwarfs Cosmic Ray Fission

Cosmology

Quantum Fluctuations
Inflation
Afterglow Light Pattern 380,000 yrs.
Dark Ages
1st Stars about 400 million yrs.
Development of Galaxies, Planets, etc.
Big Bang Expansion 13.7 billion years

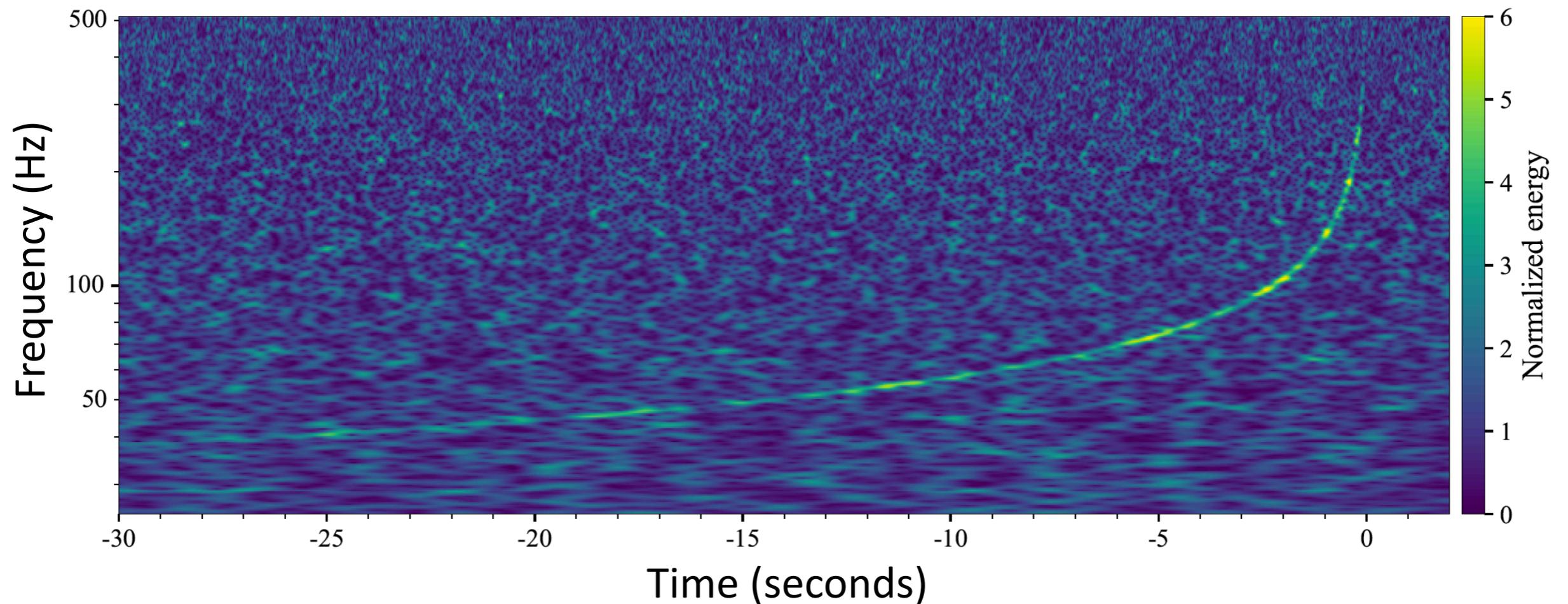
Part I:

What are GWs?

First Measurement of GWs from a Binary Neutron Star Merger

August 17th 2017 at 12:41:04 UTC

[LVC,PRL, 119, 161101 (2017)]



Loudest (SNR ~ 32.2) and longest (~ 100 s) signal so far:
False alarm rate < 1 in 80 000 years

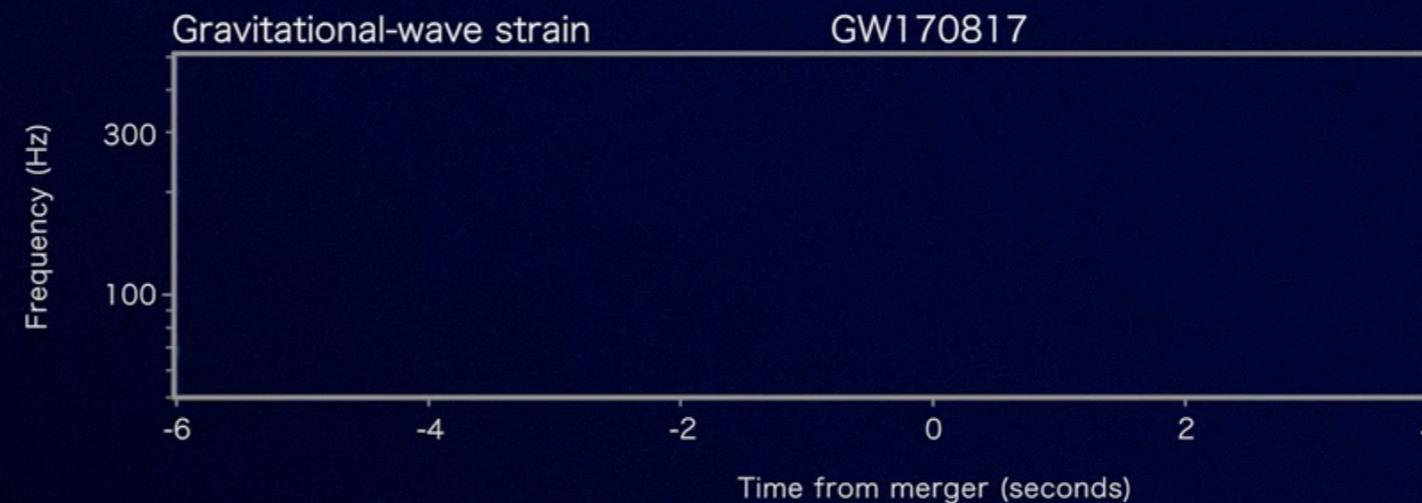
The first two seconds post-merger

[LVC+Fermi+Integral, ApJ Lett, 848:L13, 2017]

Fermi



LIGO



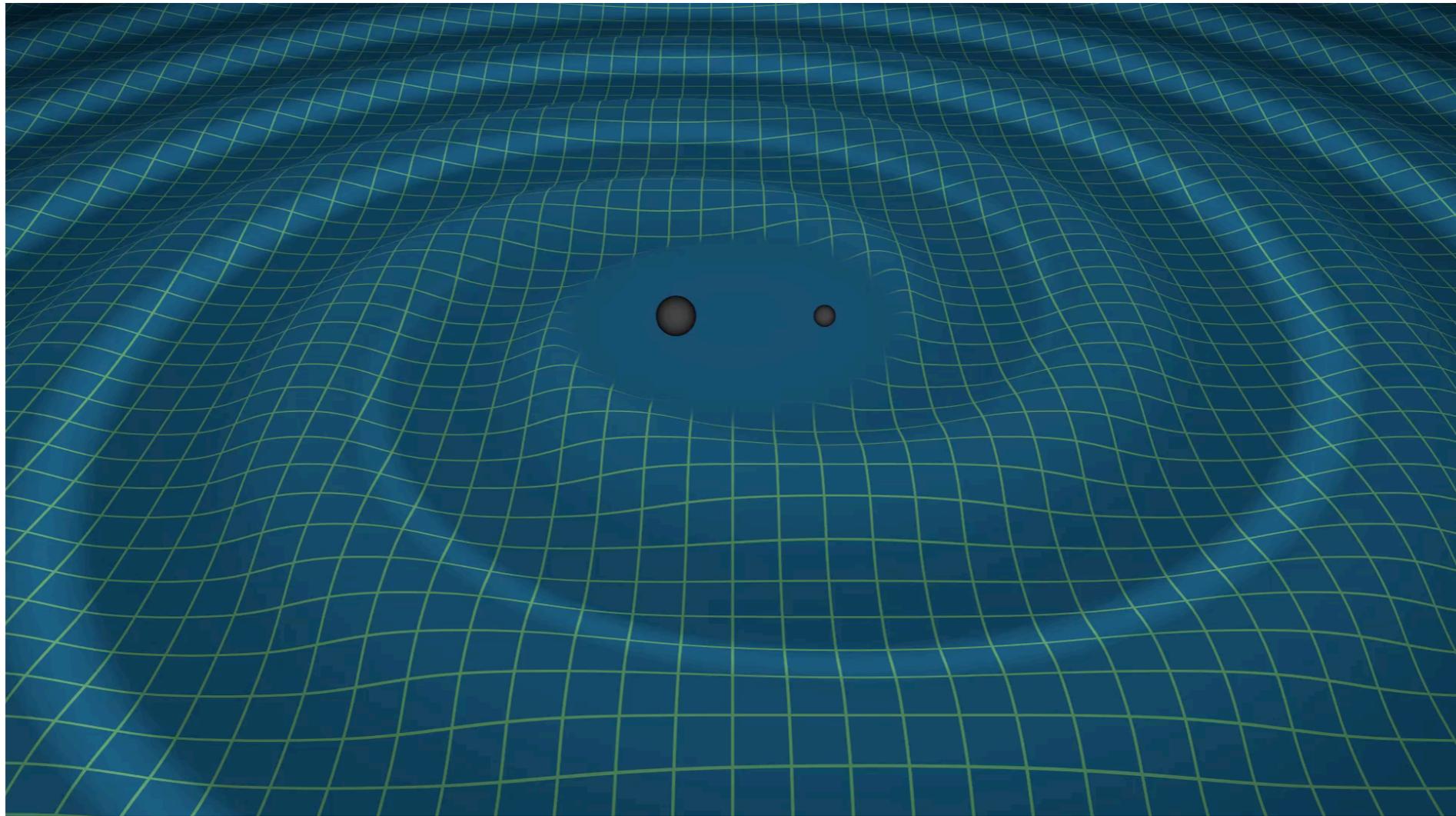
Credit: NASA/LVC

Test of speed of gravity and light and equivalence principle

$$-3 \times 10^{-15} \leq \frac{v_{\text{gw}}}{v_{\text{em}}} - 1 \leq 7 \times 10^{-16}$$

GWs are perturbations in spacetime curvature

Accelerating quadrupole matter sources



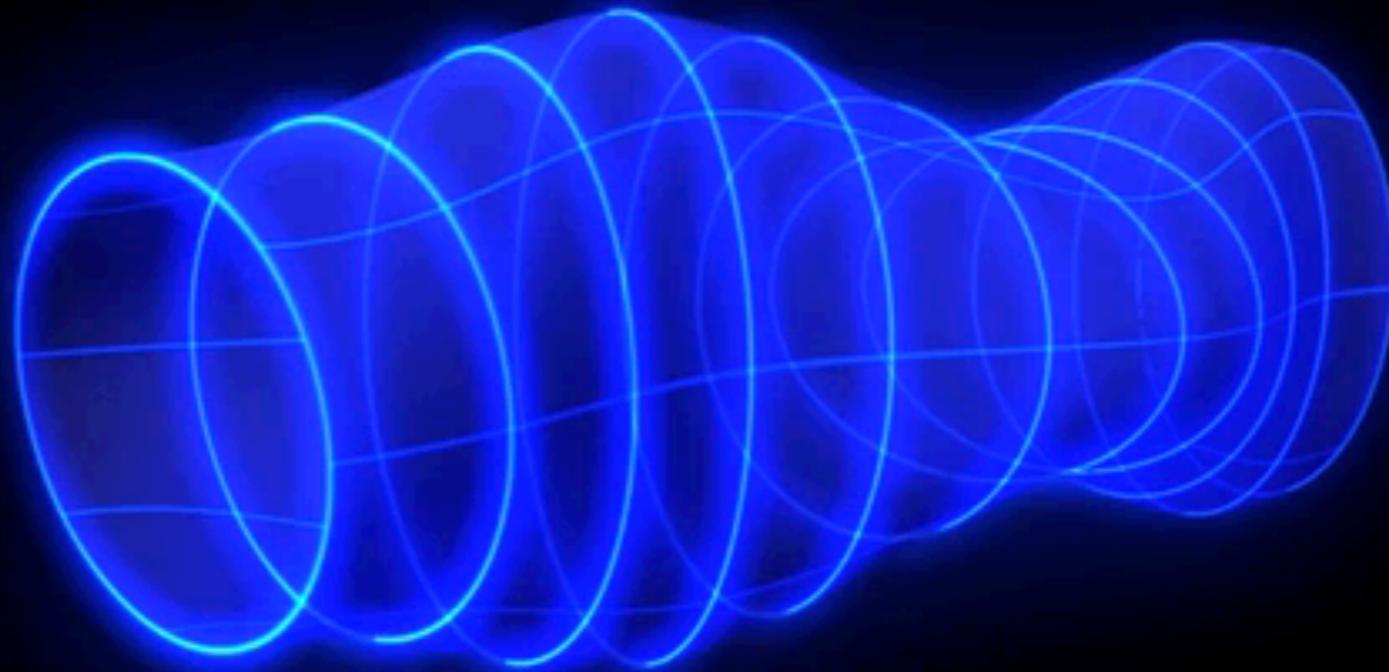
Measurable GW strain $h(t) \sim 1/\text{distance}$

Newtonian Quadrupole formula (1916): $\mathcal{L} = \frac{G}{5c^5} \left\langle \frac{d^3 I_{ij}}{dt^3} \frac{d^3 I_{ij}}{dt^3} \right\rangle + \mathcal{O}\left(\frac{1}{c^7}\right)$

GWs act as transverse tidal fields on matter



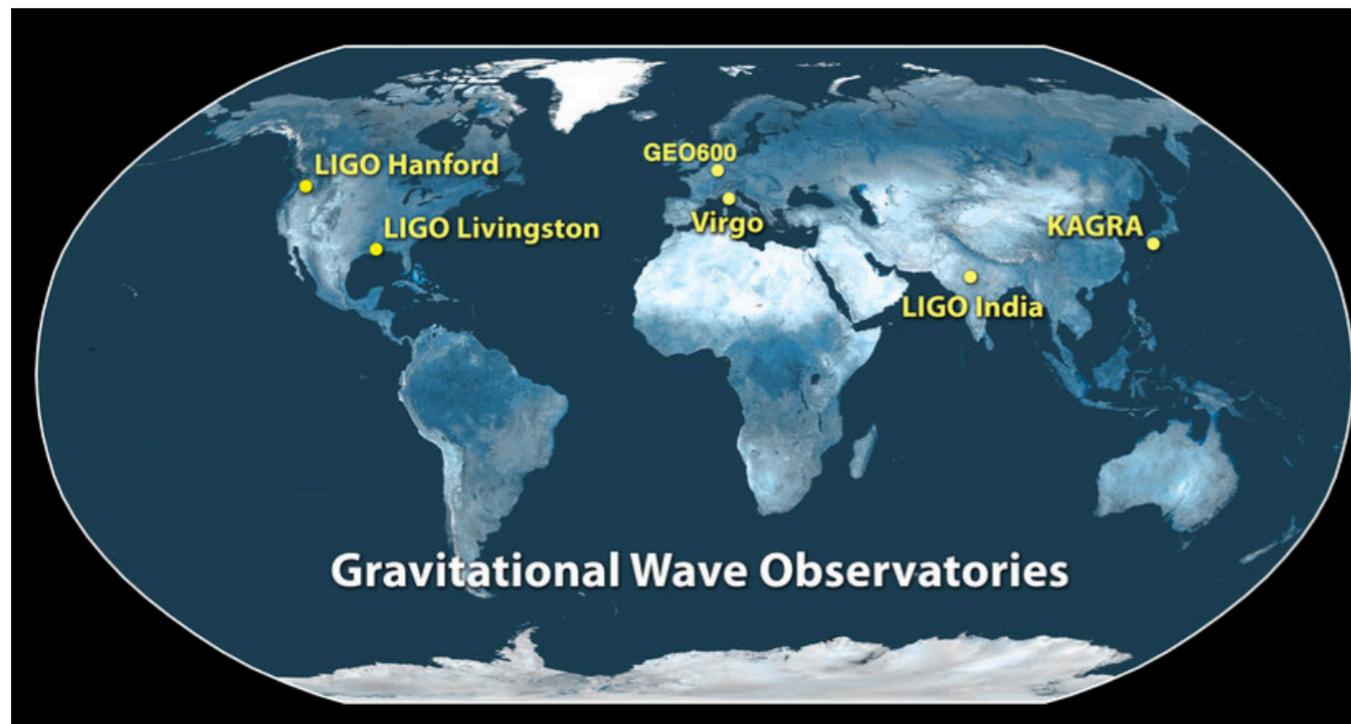
Two polarizations h_+ and h_\times



Coherent, weakly interacting,
bulk dynamic properties of matter

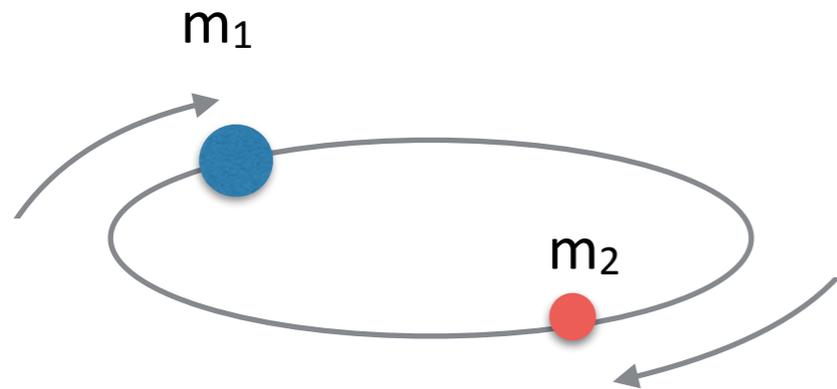
kHz **GW** detectors measure tidal stretching and squeezing of spacetime

$$h \sim \frac{\Delta L}{L} \sim 10^{-21}$$

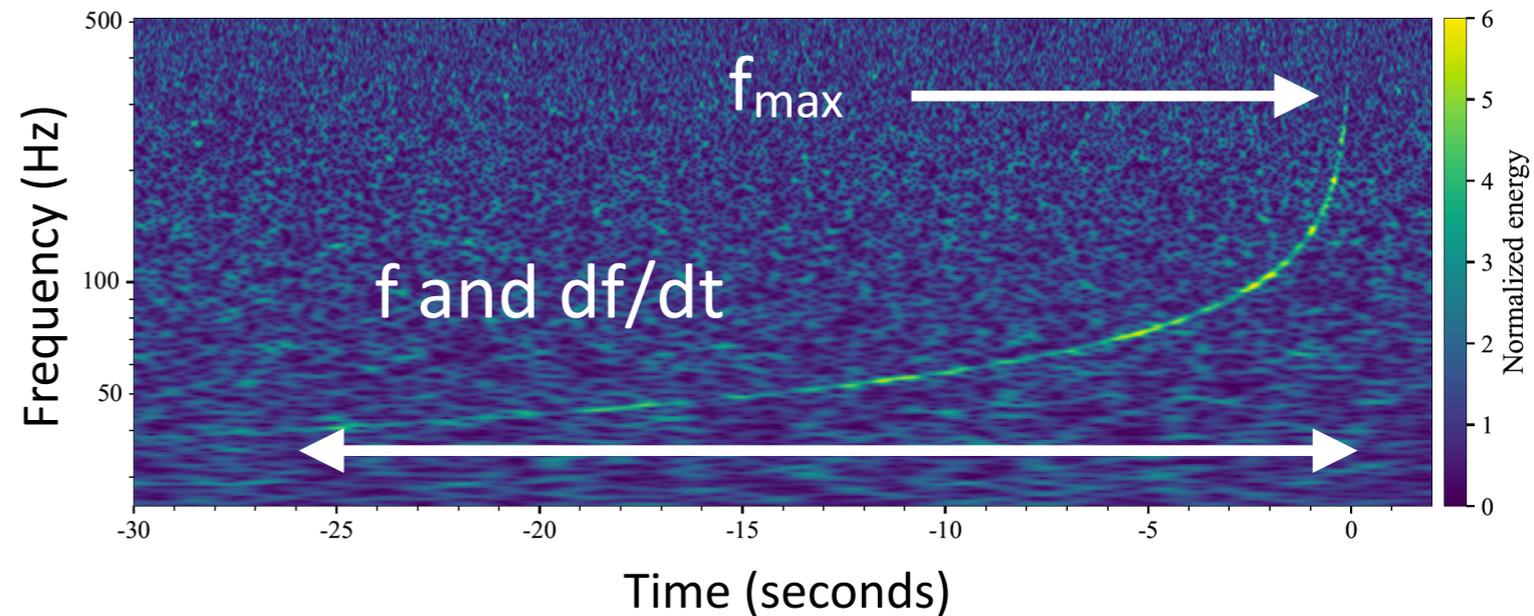


19 - 2048 Hz

Simplest “Newtonian” model explains frequency chirp



$$\left(\frac{dE}{dt}\right)_{\text{rad}} + \left(\frac{dE}{dt}\right)_{\text{orb}} = 0$$



[LVC, PRL 119, 161101 (2017)]

⇒ Frequency chirp:

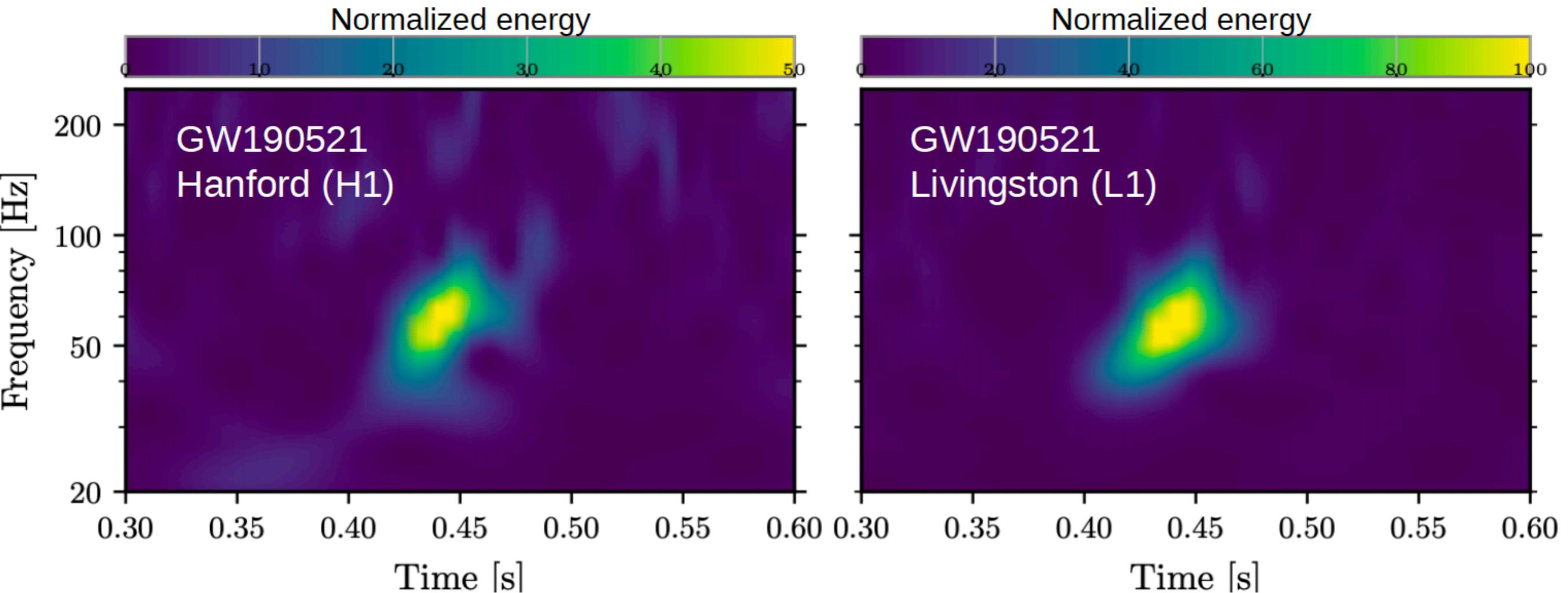
$$\frac{df}{dt} = \frac{96 \pi}{5} \left(\frac{\pi G \mathcal{M}}{c^3} \right)^{5/3} f^{11/3}$$

Chirp mass:

$$\mathcal{M} = \frac{(m_1 m_2)^{3/5}}{(m_1 + m_2)^{1/5}}$$

Most Massive Black Hole Merger seen in GWs

GW190521: May 21st 2019 at 03:02:29 UTC

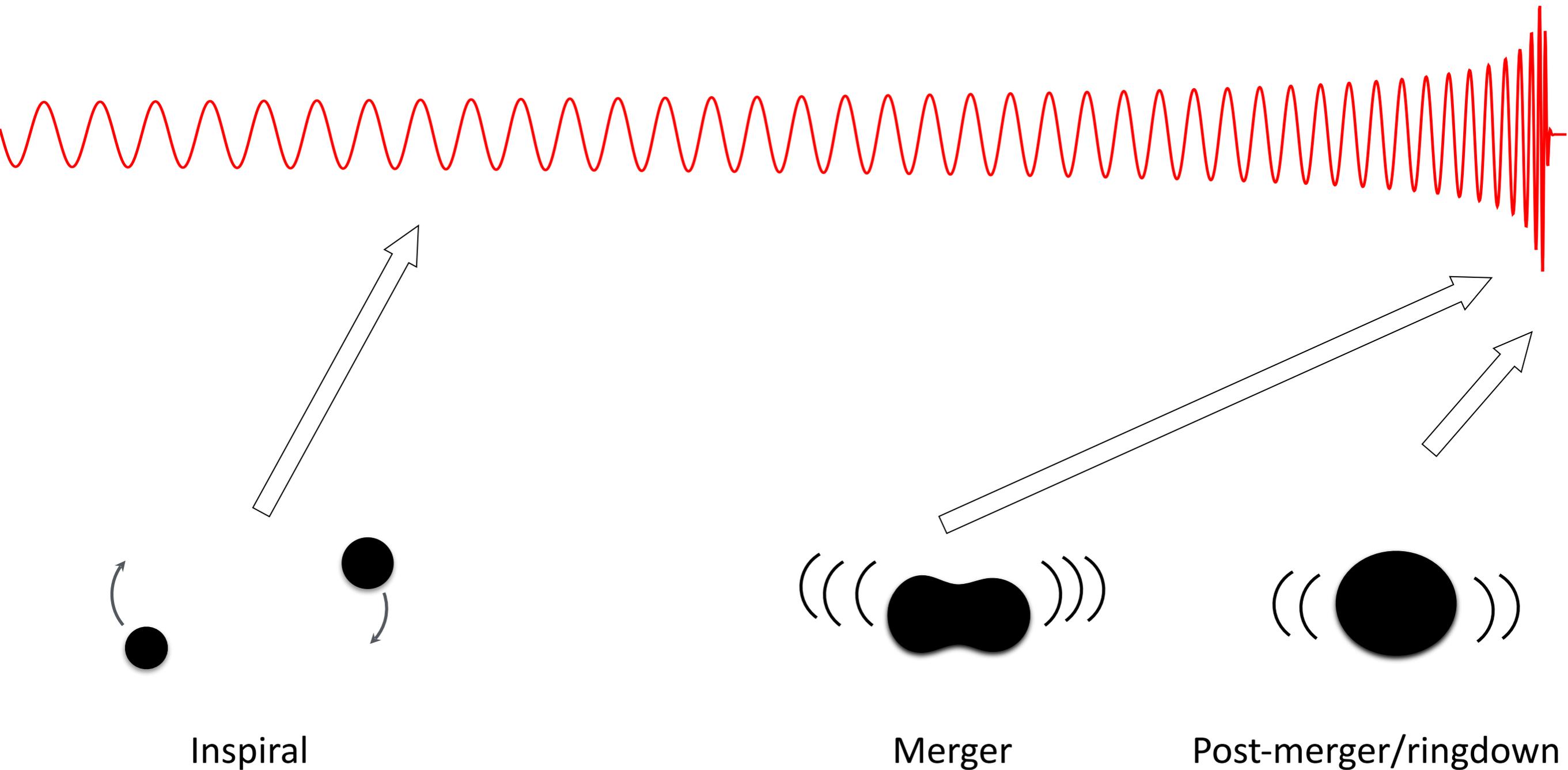


Shortest (~ 0.1 s) signal so far,
False alarm rate < 1 in 5 000 years (Signal-to-noise: 14.7),
Rare Event: 1 every 8 years in Gpc^3

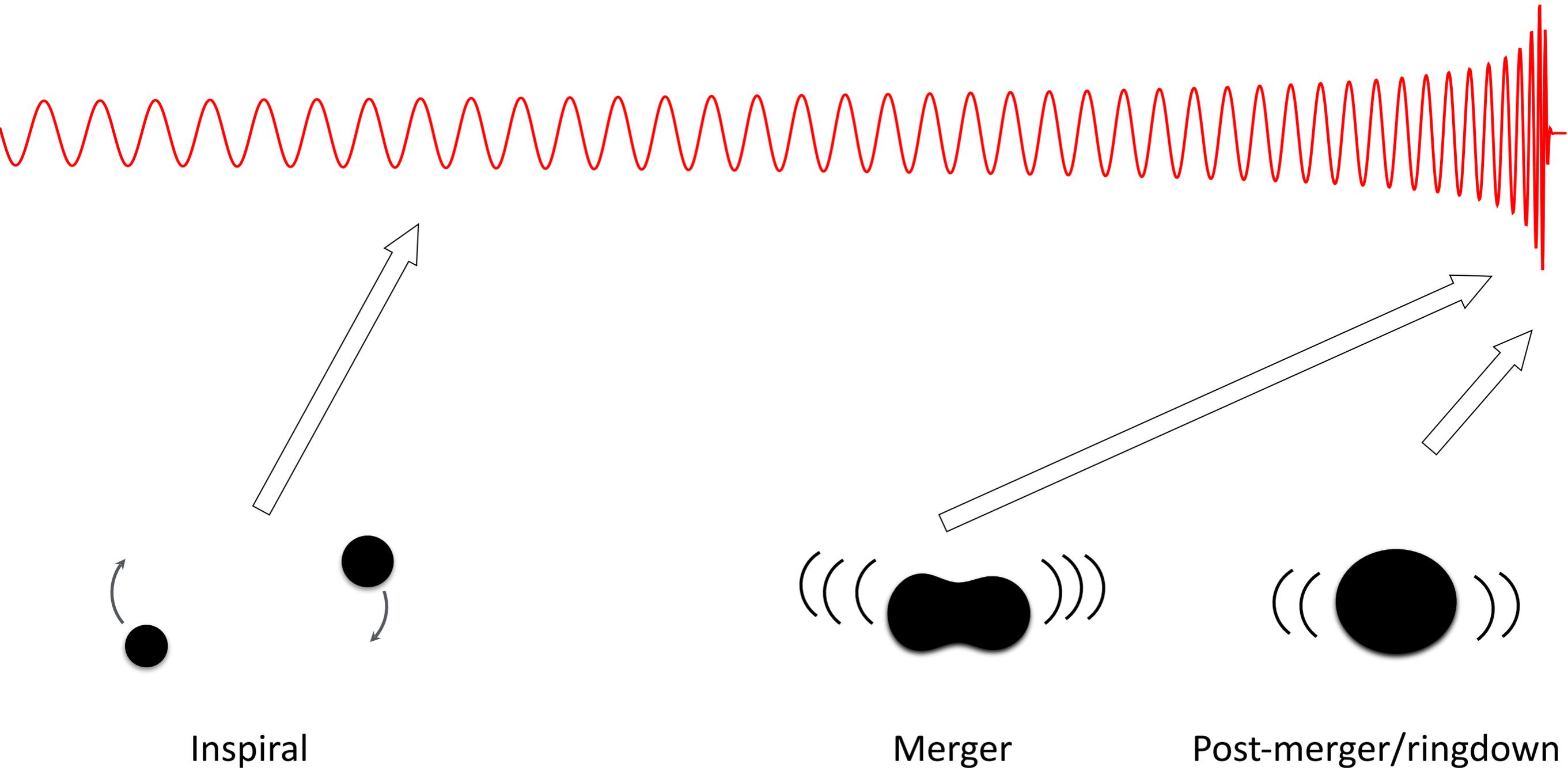
Part II:

Retrieving information
about BHs and NSs from GW

The GW waveform encodes source parameters

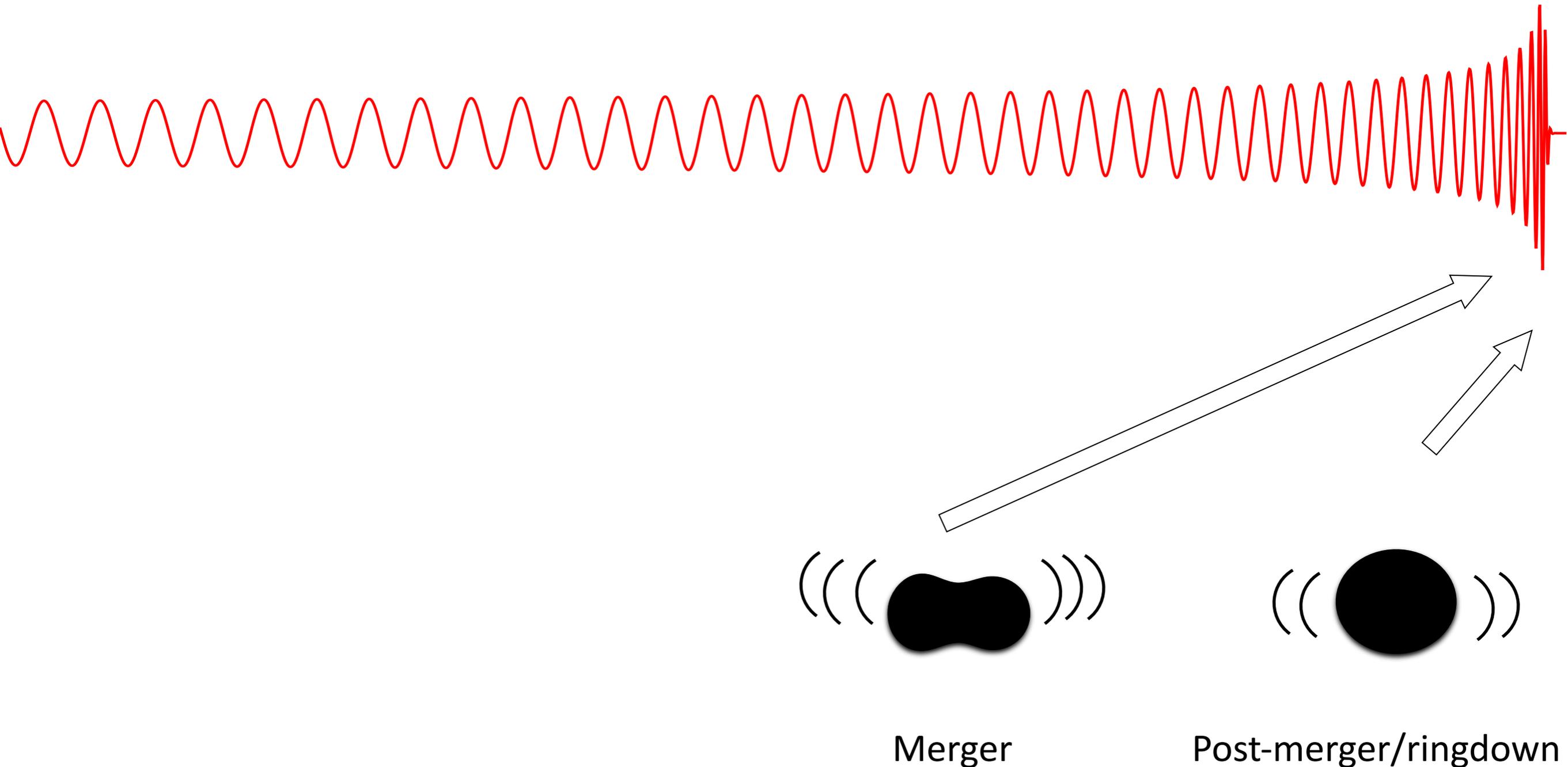


The GW waveform encodes source parameters



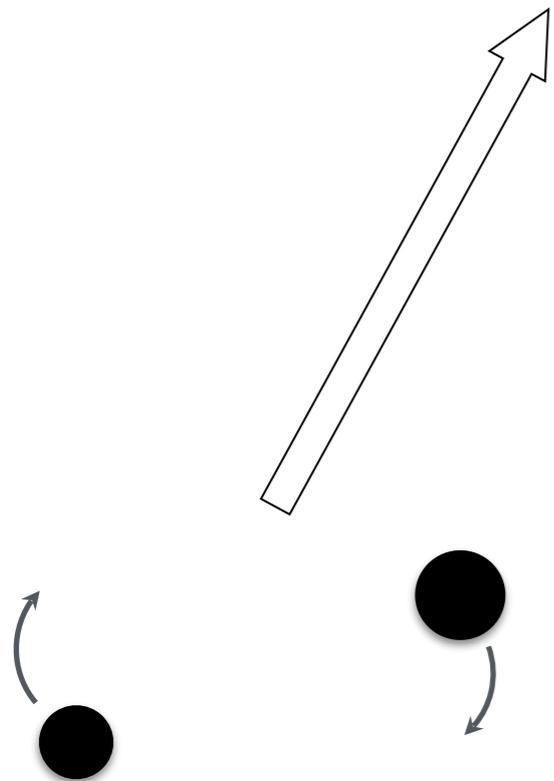
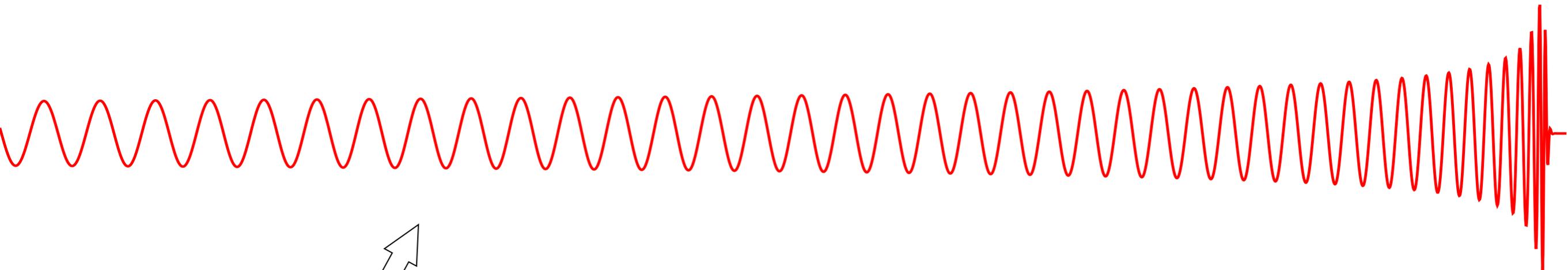
$$\frac{df}{dt} = \frac{96 \pi}{5} \left(\frac{\pi G M}{c^3} \right)^{5/3} f^{11/3}$$

The GW waveform encodes source parameters



~ 10 cycles of merger for ~ 10 M_{\odot} BHs: seconds

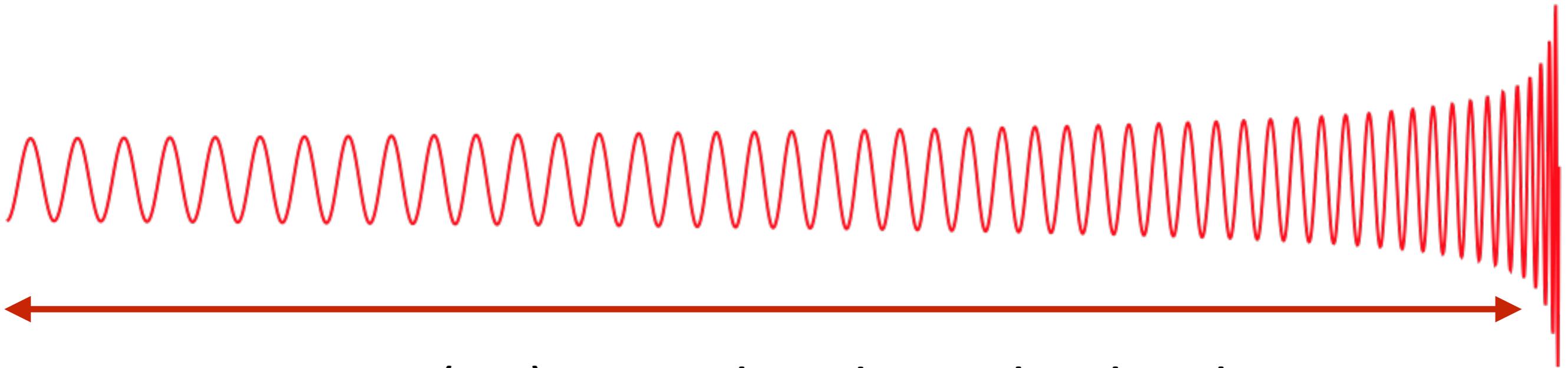
The GW waveform encodes source parameters



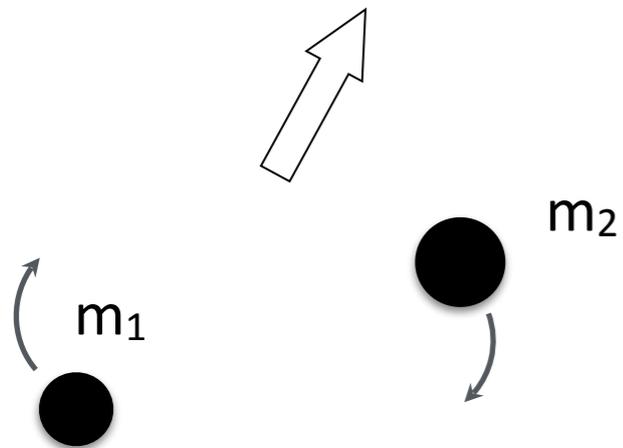
Inspiral

$\sim 10^3$ cycles of the inspiral for few M_{\odot} with LIGO/Virgo: minutes

The GW waveform encodes source parameters

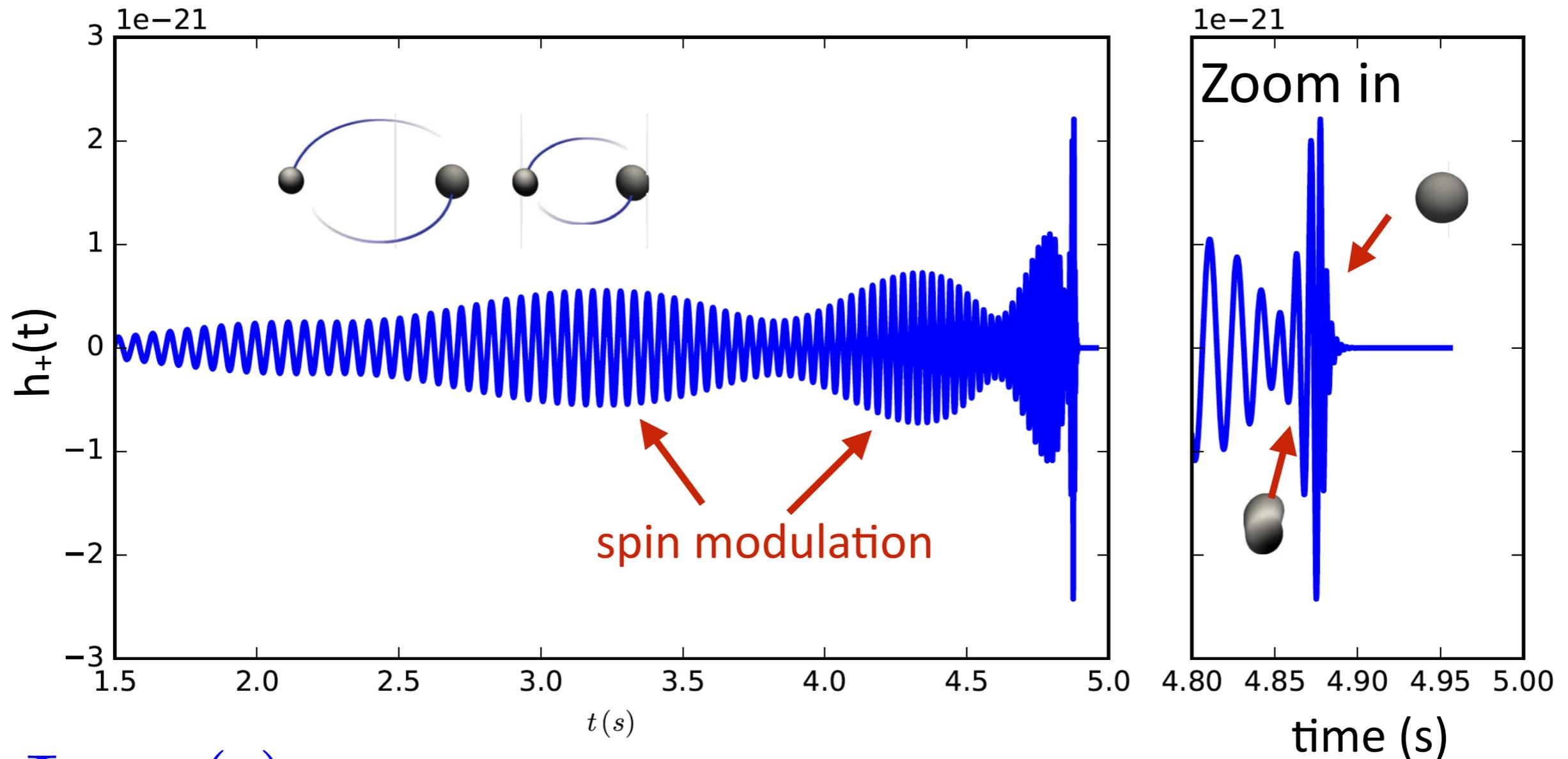


post-Newtonian (PN) Inspiral — driven by the chirp mass



$$1\text{PN} \sim \frac{v^2}{c^2} \sim \frac{Gm}{rc^2} \ll 1$$

The GW waveform encodes source parameters



$$\Phi_{\text{GW}}(t)$$

\Rightarrow chirp mass, reduced mass (1PN), spin-orbit (1.5PN), spin-spin (2PN)...

tidal deformability (5PN)

Extract source information from GWs

$$h_+(t) = \frac{A[\mathcal{M} f(t)]}{D} (1 + \cos^2 \iota) \cos \Phi_{\text{GW}}(t)$$

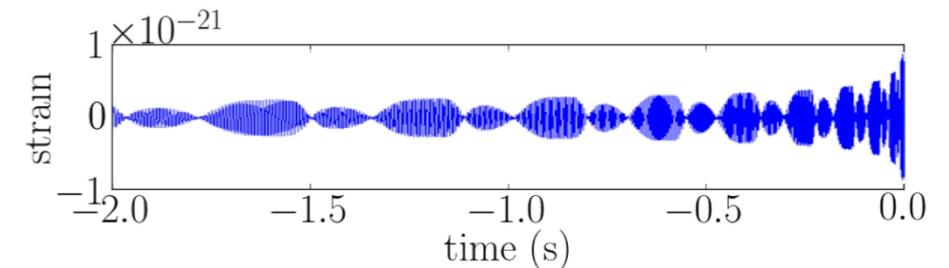
frequency
inclination angle
GW Phase

distance
inclination angle
GW Phase

$h(t)$: 9-16 parameters

- + Redshifted Masses
- + Spins
- + Tidal deformability
- + Geometric properties:
 - Inclination angle
 - Source Position
 - Luminosity distance

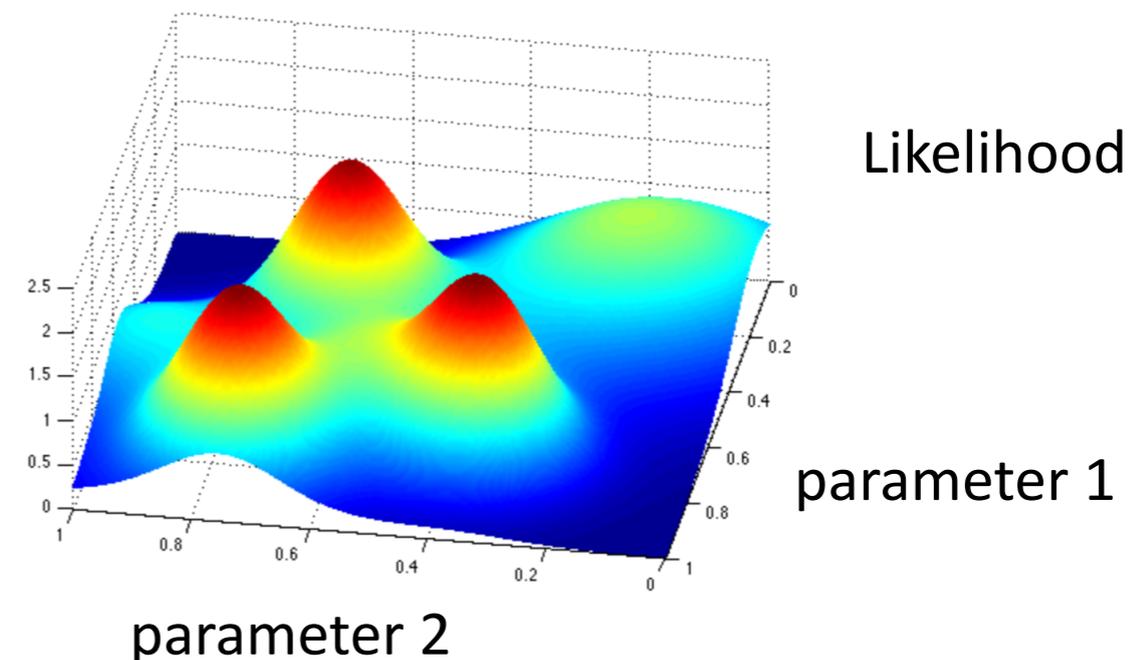
Model $h(t)$



Detector output



Explicitly map out: $p(\theta|s) \propto p(\theta)\mathcal{L}_{\text{total}}(s|\theta)$



using Bayesian Markov Chain Monte Carlo and Nested Sampling Techniques

[see LVC, arXiv: 1602.03840, PRL 116, 241102, 2016]

[see Nissanke et al. 2010, 11, 13a, 13b]

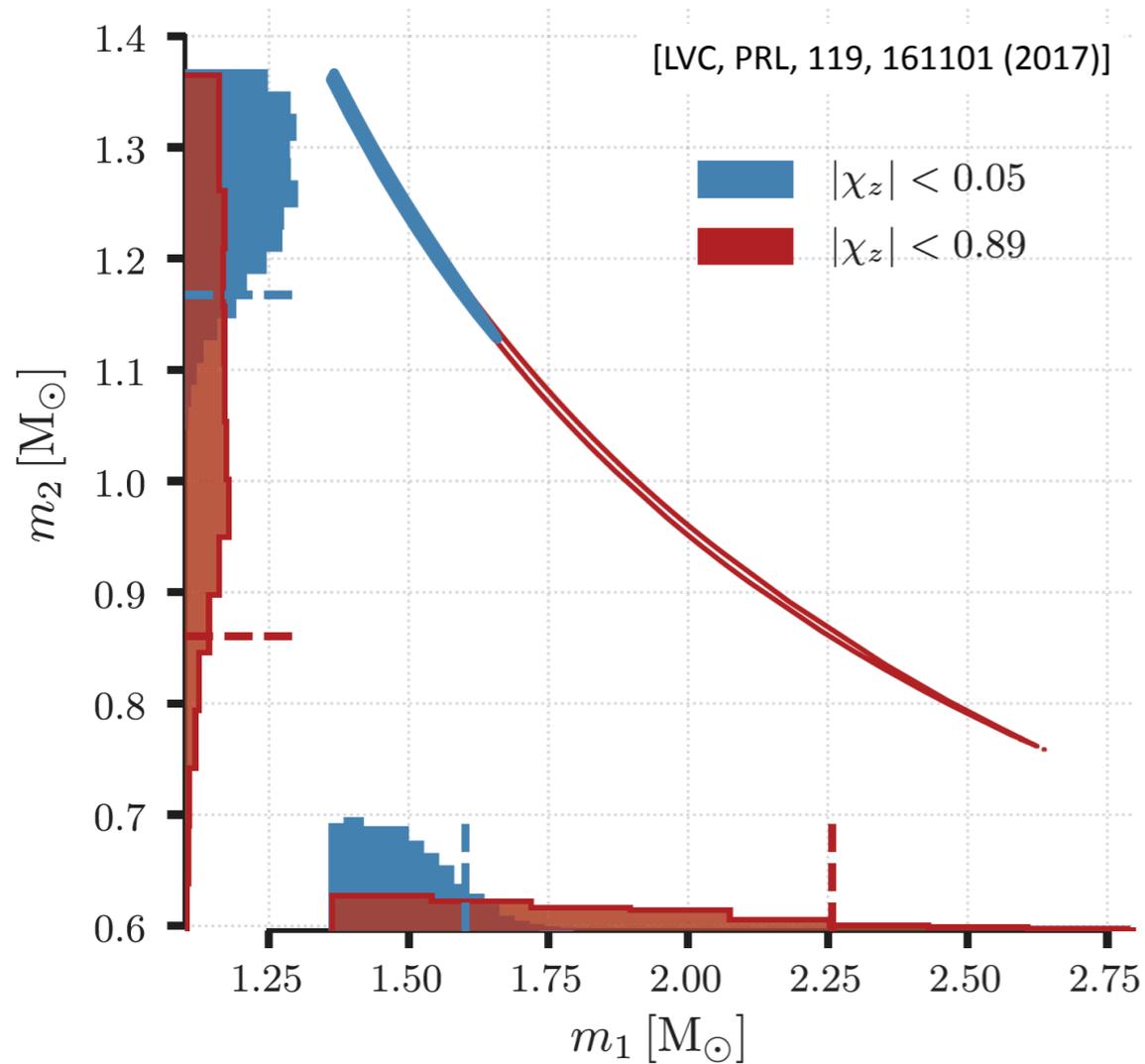
Part III:

GW source parameter inference

Progenitor masses and spins retrieval

GW170817

GW190521

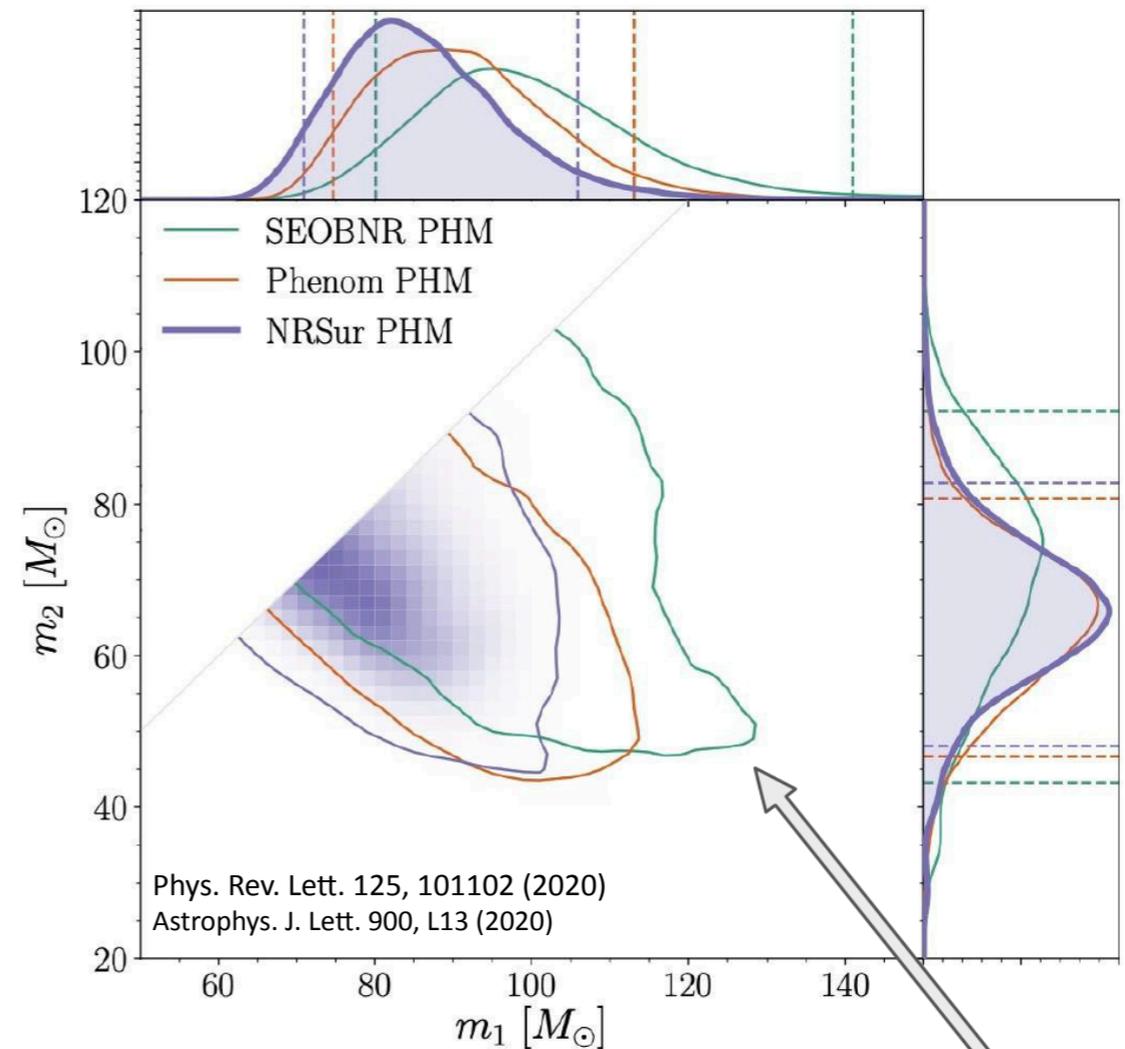


BNS or NS-BH ?

Masses: 1.00 - 1.89 M_\odot (high spins)

1.16 - 1.60 M_\odot (low spins)

Total Mass: $2.73^{+0.04}_{-0.01} M_\odot$



BBH

BH mass 1: $85^{+21}_{-14} M_\odot$

BH mass 2: $66^{+17}_{-19} M_\odot$

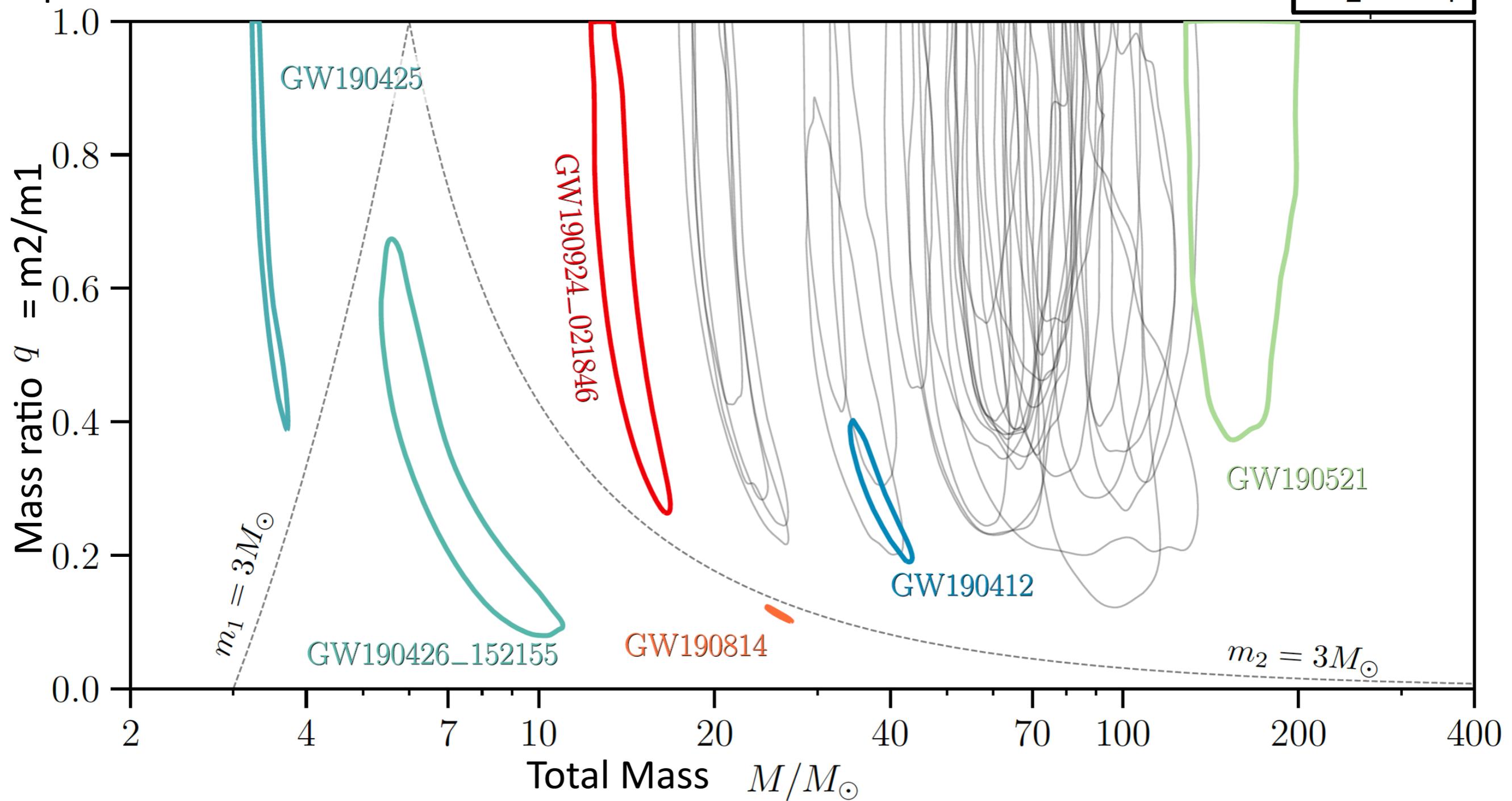
Total Mass: $150^{+29}_{-17} M_\odot$

90% credible regions

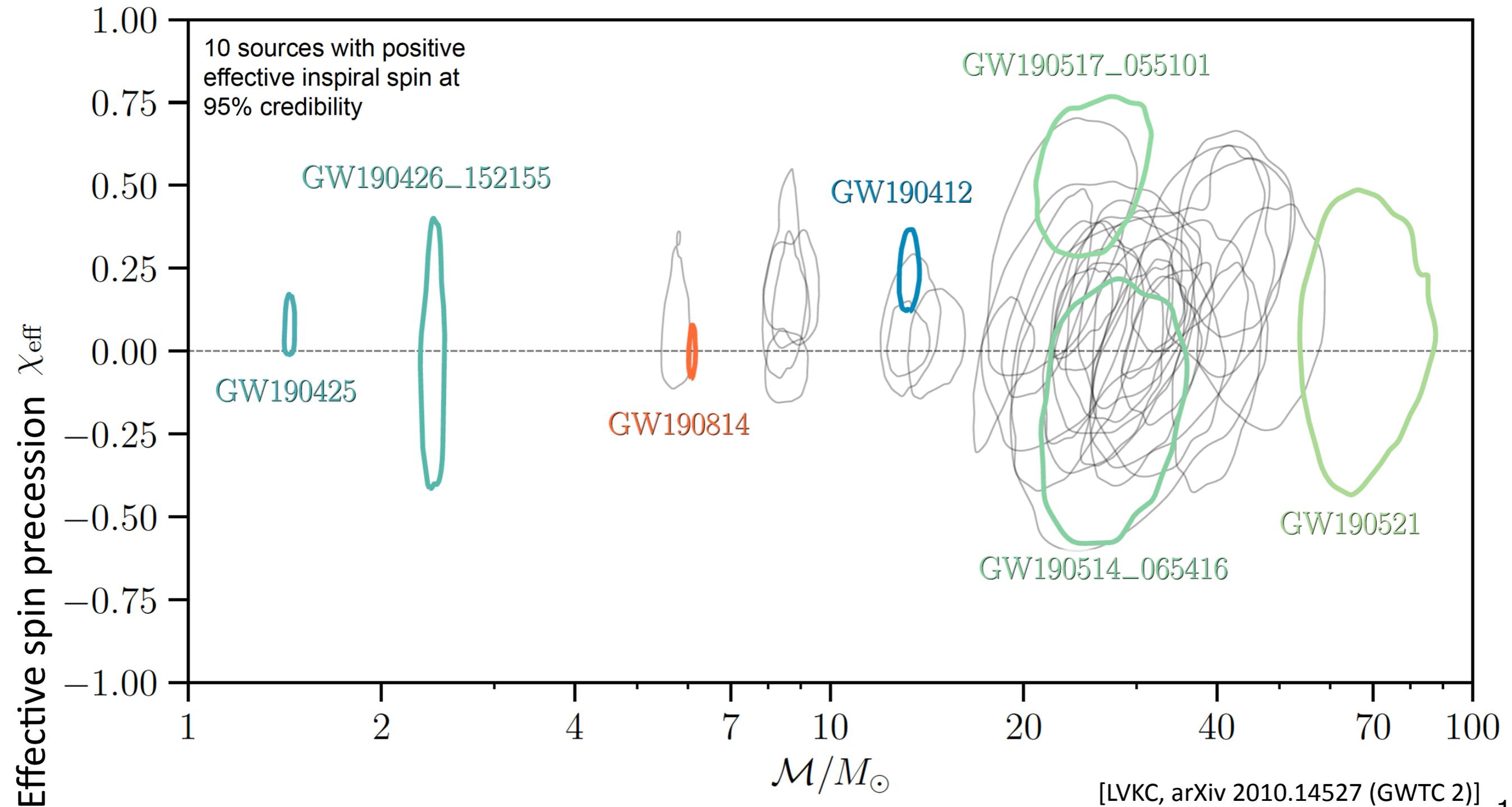
GWTC-2 indicates broad mass distribution also with non-equal masses

Equal Masses

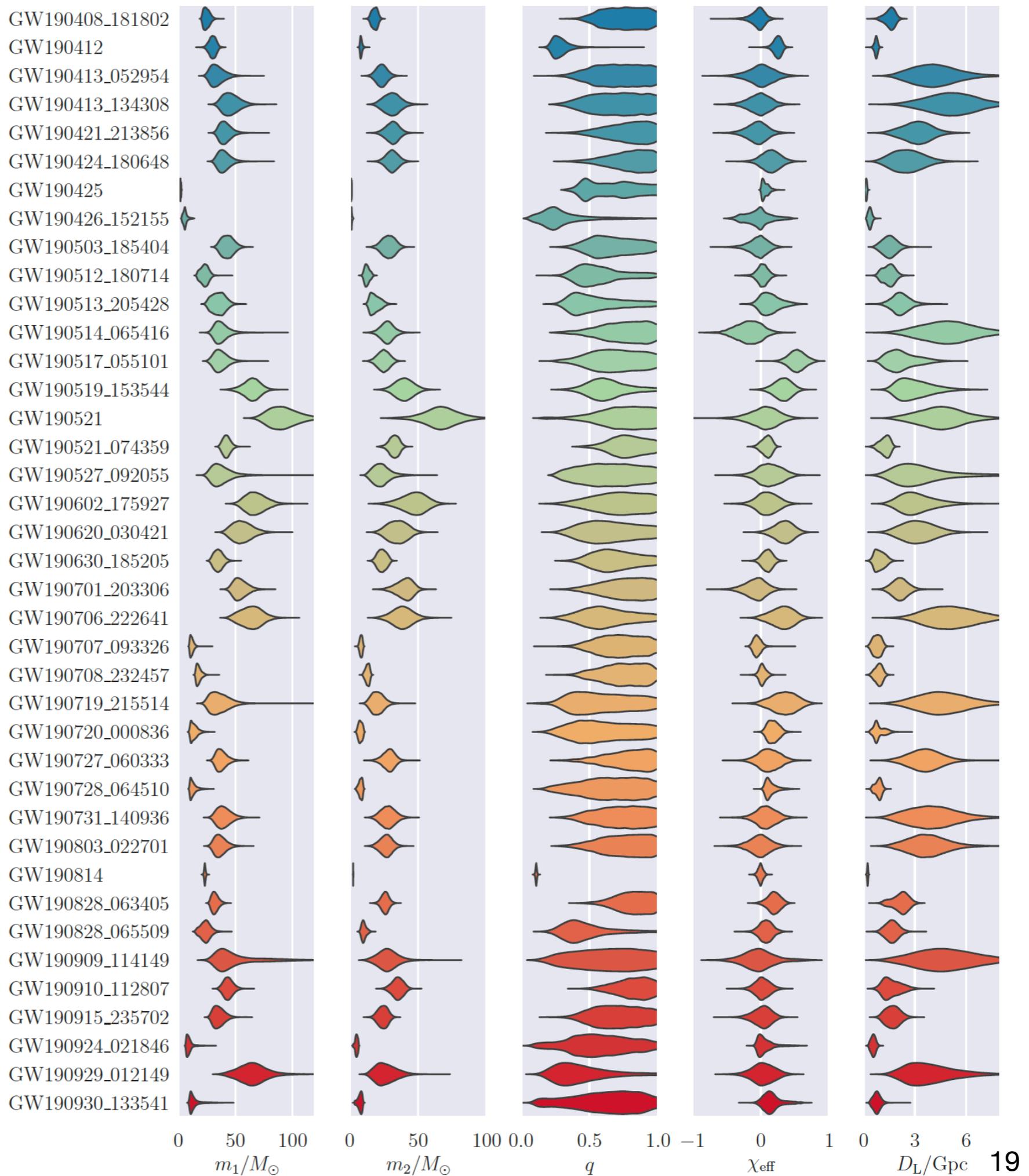
$$m_2 \leq m_1$$



GWTC-2 indicates spin induced precession of orbital plane



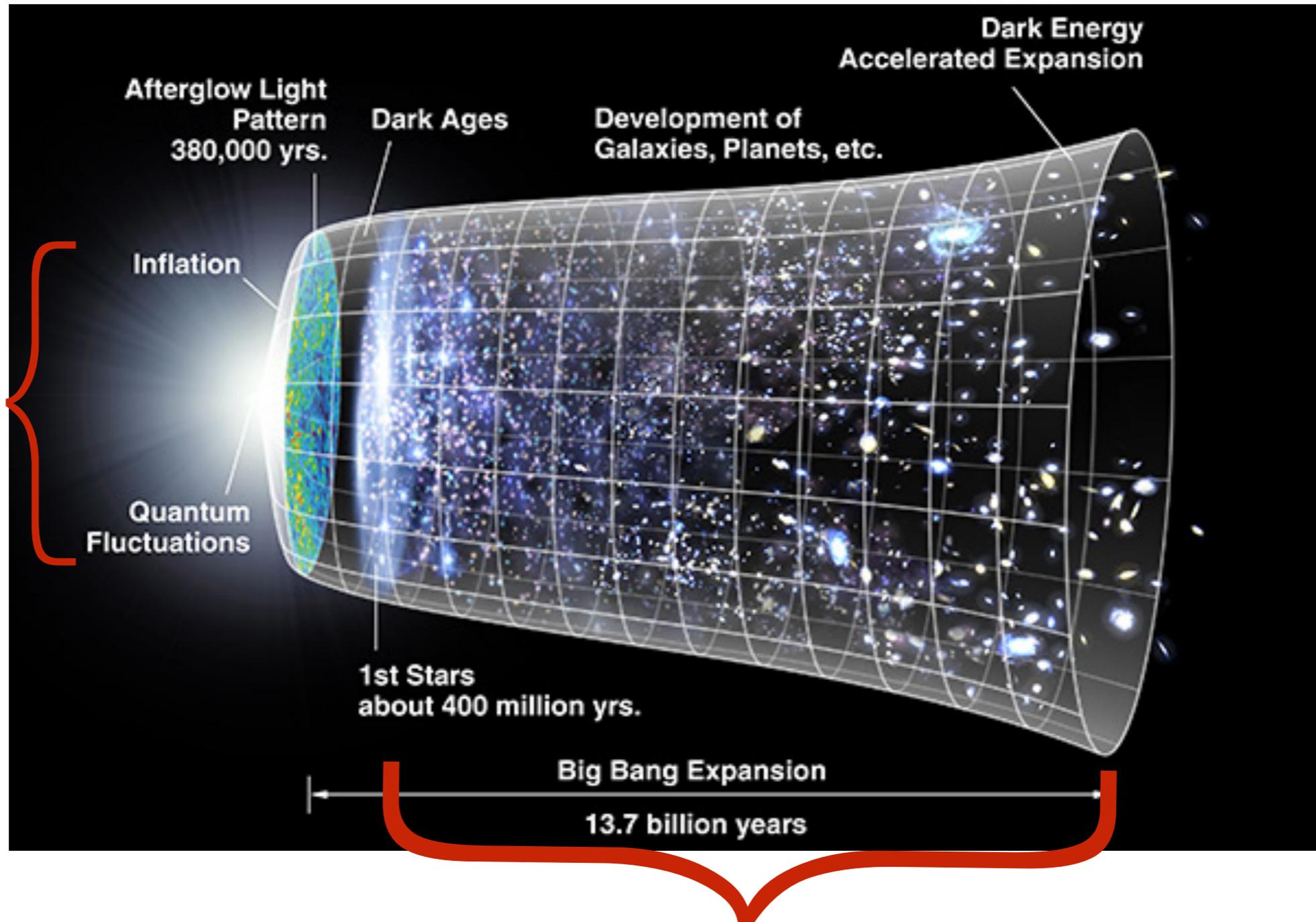
Compact Object Portrait Gallery



[LVKC, arXiv 2010.14527 (GWTC 2)]

Challenge: how, when and where do BBHs form?

Primordial BHs
from density
fluctuations in
early Universe



Pop III: first massive stars (1% of stars in Universe)

Pop II/I: classic field binary evolution (90%)

Pop II/I: rapid rotation

Pop II/I: dynamical formation in globular clusters, stellar clusters

Exotic

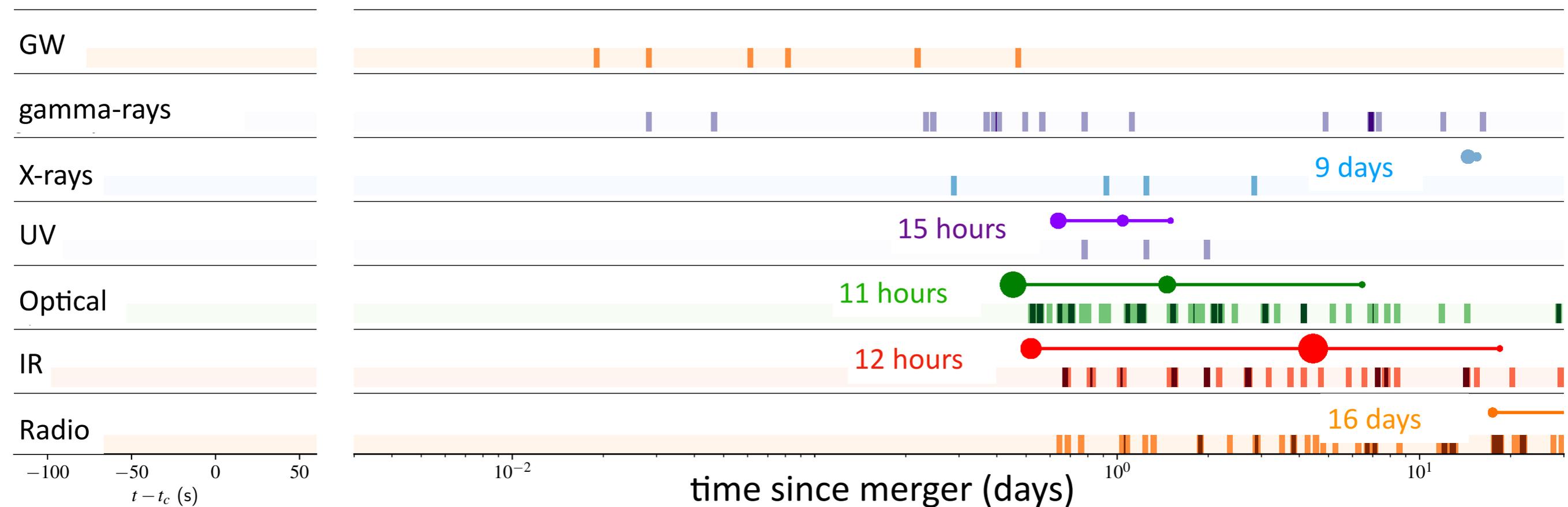
[see LVKC; arXiv: 2010.14533]

Part III:

The Multi-messenger discovery
and cosmology of GW170817

The first month(s) of multi-messenger observations of GW170817

adapted from LIGO, Virgo, EM partners + ApJ 848 L12 (2017) with Nissanke



Global ground and space-based effort:

70+ teams, 100+ instruments, over 3500 co-authors

A Tale of Two+ Matter Outflows \Rightarrow EM counterparts

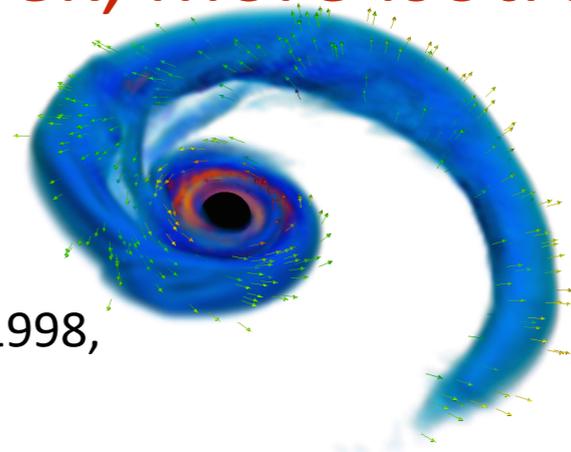
1. Tidal Tails + Disk Winds + Core-bounce Heating \Rightarrow Kilonova: Ultraviolet Optical IR (days-week, more isotropic)

$$M_{ej} \approx 0.01-0.05 M_{\odot}$$

$$E \approx 10^{50} \text{ ergs}$$

$$v \approx 0.1-0.3c$$

[Lattimer and Schramm 1974, Li and Paczynski 1998, Rosswog 1999, Kulkarni 2005, Metzger +.. 2010, Kasen + 2013 ...]



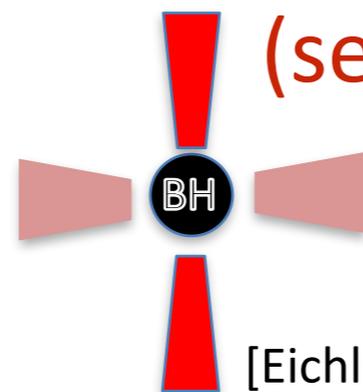
[Foucart et al. 2014]

2. Ultra-relativistic Jet \Rightarrow Short Gamma Ray Burst + afterglow (seconds - months, collimated)

$$M_{ej} \approx 10^{-6} M_{\odot}$$

$$E \approx 10^{49} - 10^{50} \text{ ergs}$$

$$v \approx 0.99c - 0.99995c$$



[Eichler+ 1989, Paczynski 1989,...]

Outflows' expand, cool and are heated by **shocks** or **radioactivity**.

Extract source & environment from EM

$F_{\lambda}(t)$: 15 parameters

- + Energetics and beaming
- + R-process nucleosynthesis
- + Mass ejecta and velocity
- + Environment
- + Redshift, Accurate Position (1")
- + Stellar populations
- + Magnetic field strength
- + Previous binary evolution & mass loss

New field: break degeneracies to measure properties of BHs and NSs

$h(t)$: 9-16 parameters

- + Redshifted Masses (several to tens %)
- + Spins (tens of %)
- + NS radii (tens of %)
- + Geometric properties: (tens of %)
 - Inclination angle
 - Source Position
 - Luminosity distance

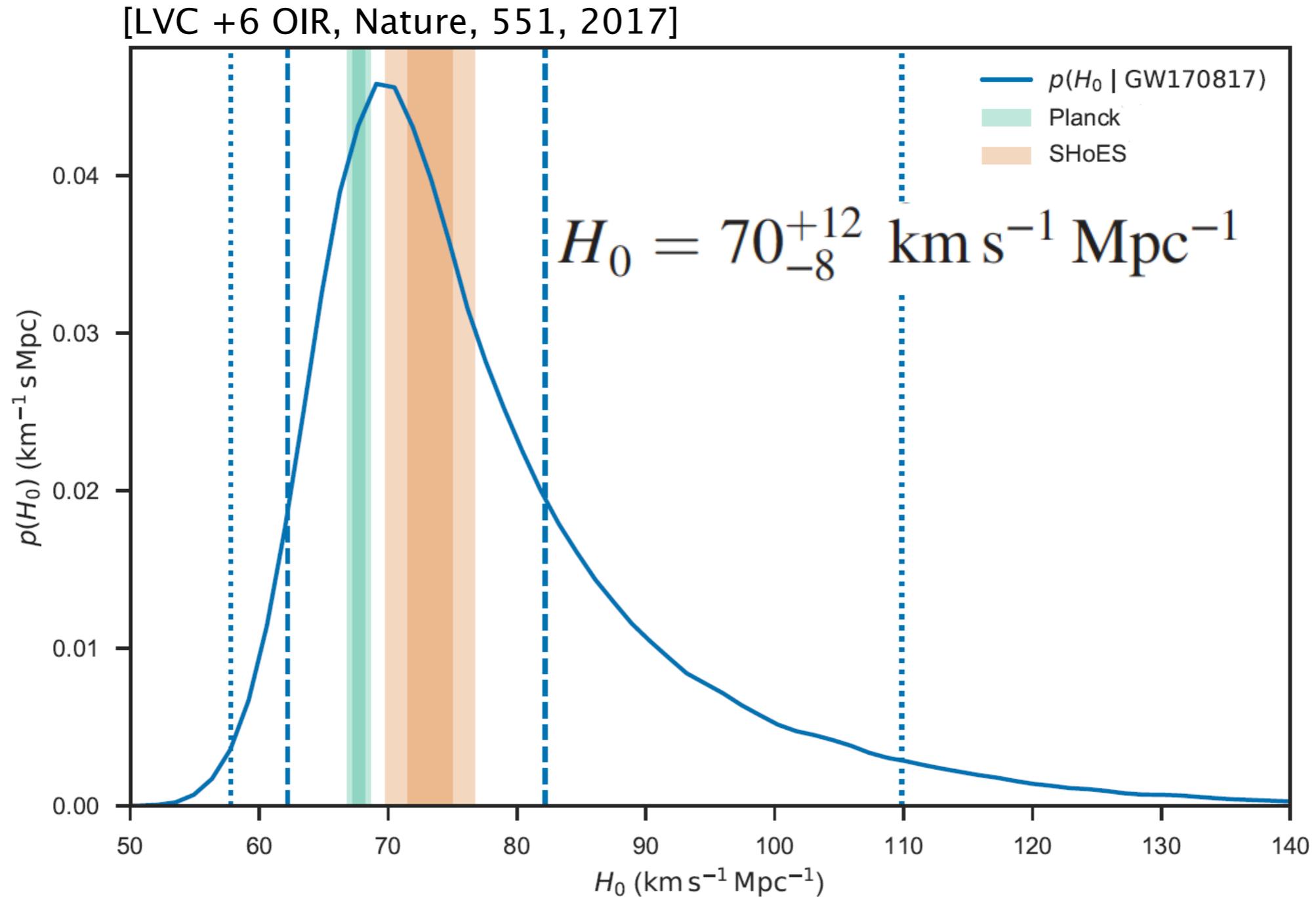
$F_\lambda(t)$: 15 parameters

- + Energetics and beaming
- + R-process nucleosynthesis
- + Mass ejecta and velocity
- + Environment
- + Redshift, Accurate Position (1")
- + Stellar populations
- + Magnetic field strength
- + Previous binary evolution & mass loss

Strong signal binary: Characterisation

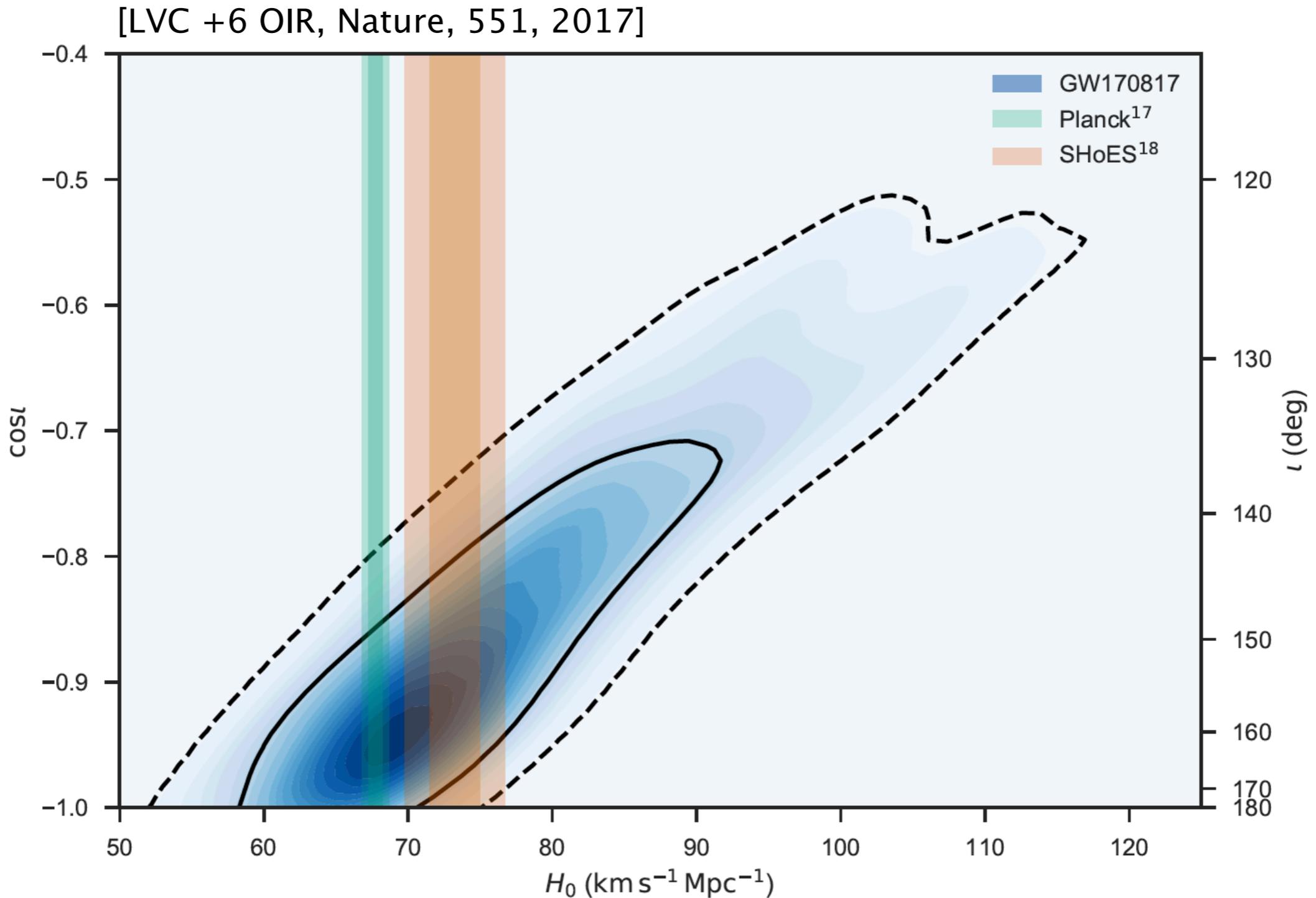
Population: Demographics, ecology and census

GW+EM: Standard siren measurement of Hubble constant



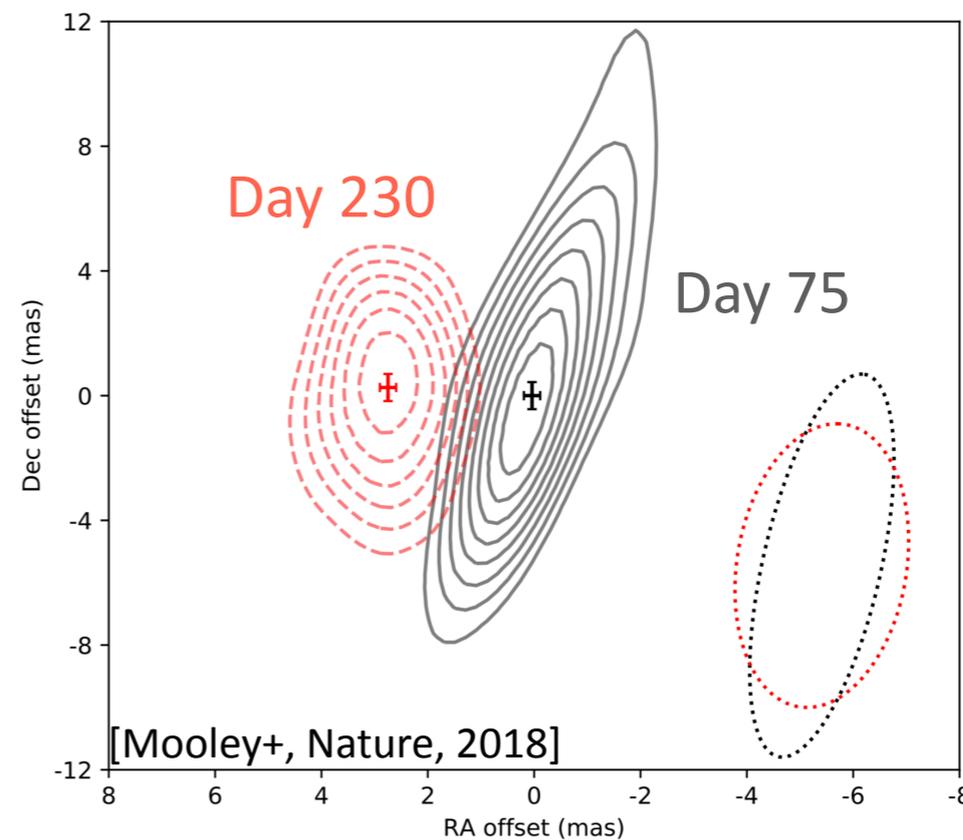
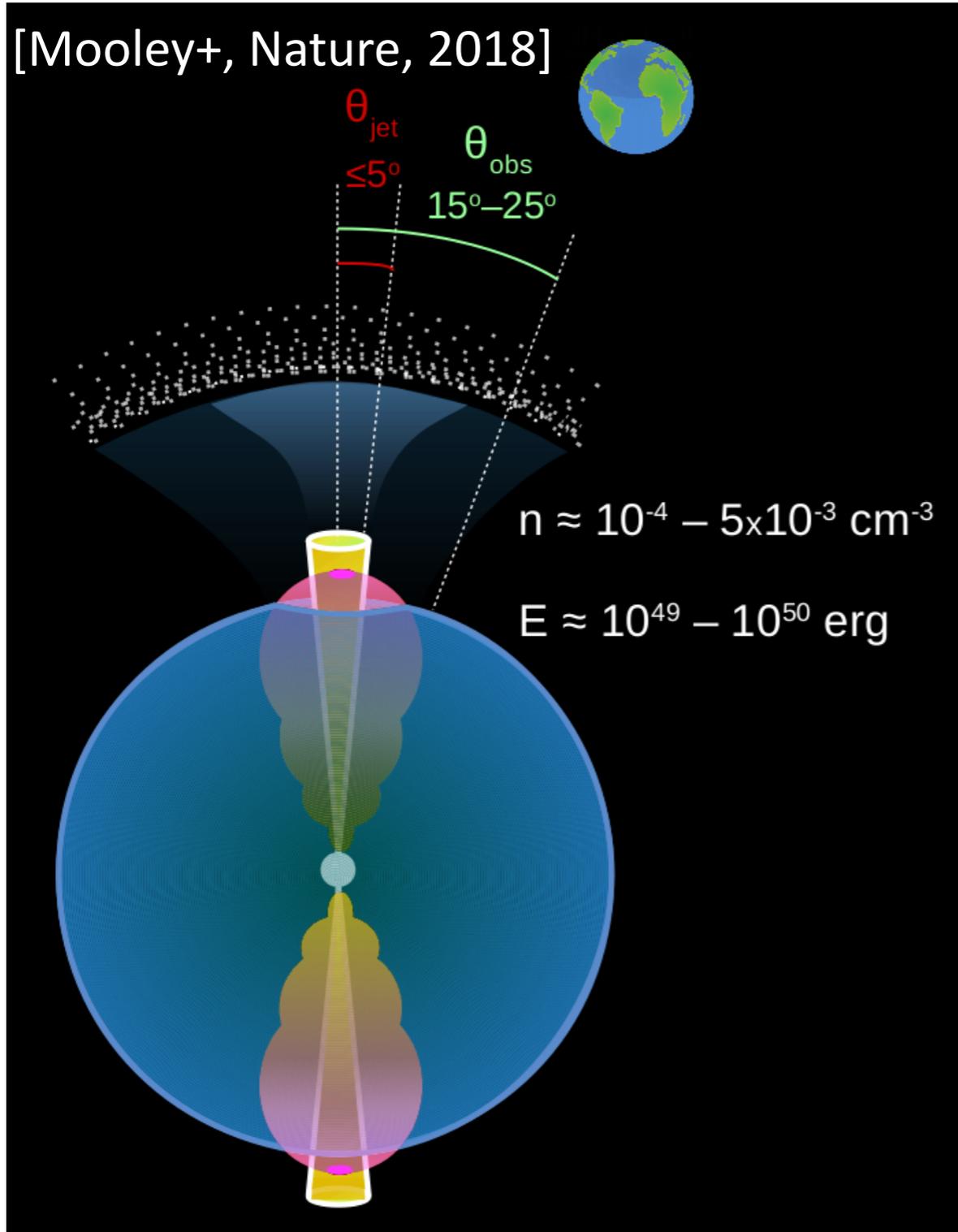
Peculiar velocity error of 150 km/s; Hubble flow velocity of 3017 +/- 166 km s⁻¹

GW challenge: Inclination angle and distance degeneracy



[see SN+10, Guidorzi+ 2017, Finstead +18, Mandel 18]

EM SGRB VLBI \Rightarrow superluminal jet with structure

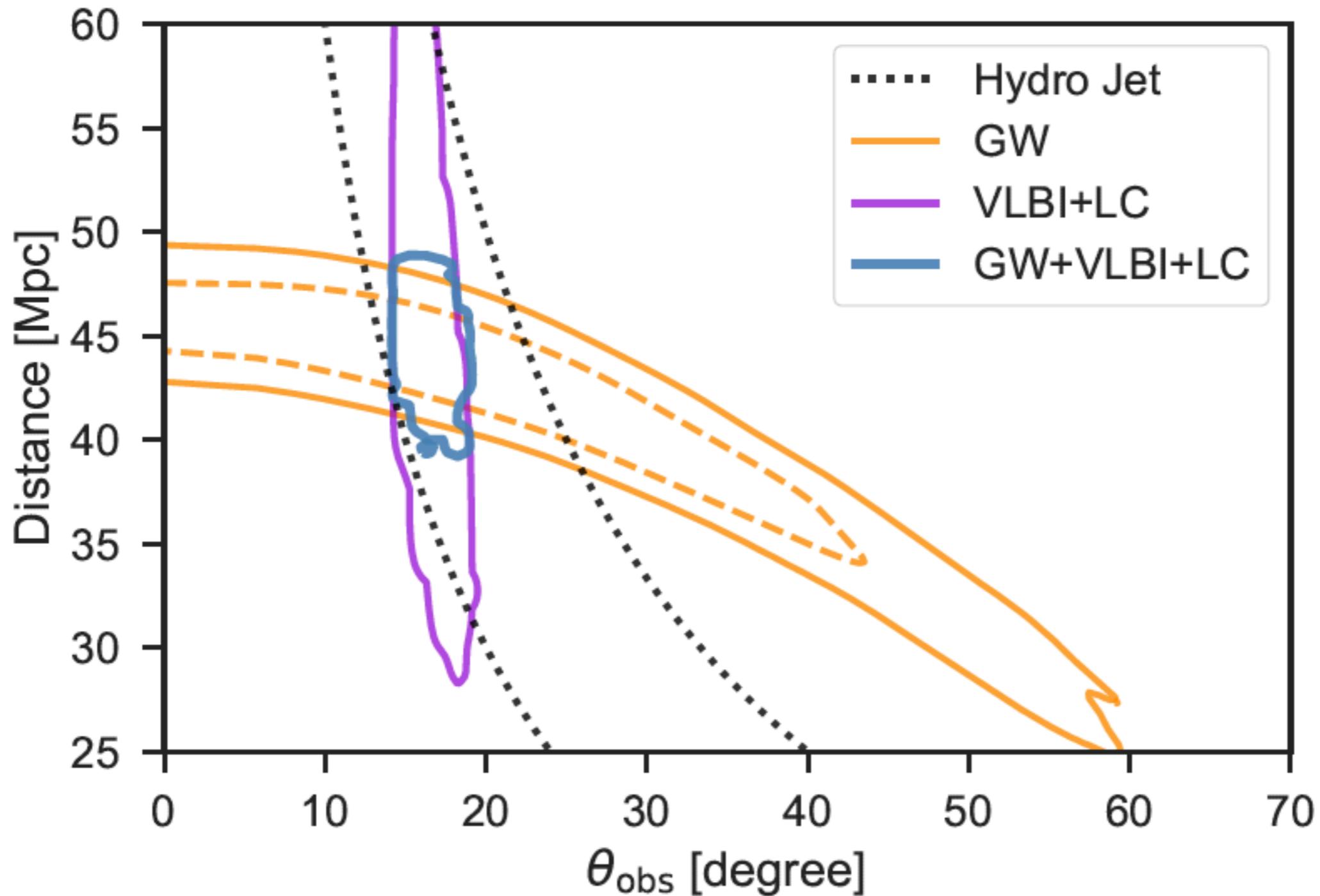


[see also Lazzati 2017,18, Ghirlanda + 2019, Troja+ 2018]

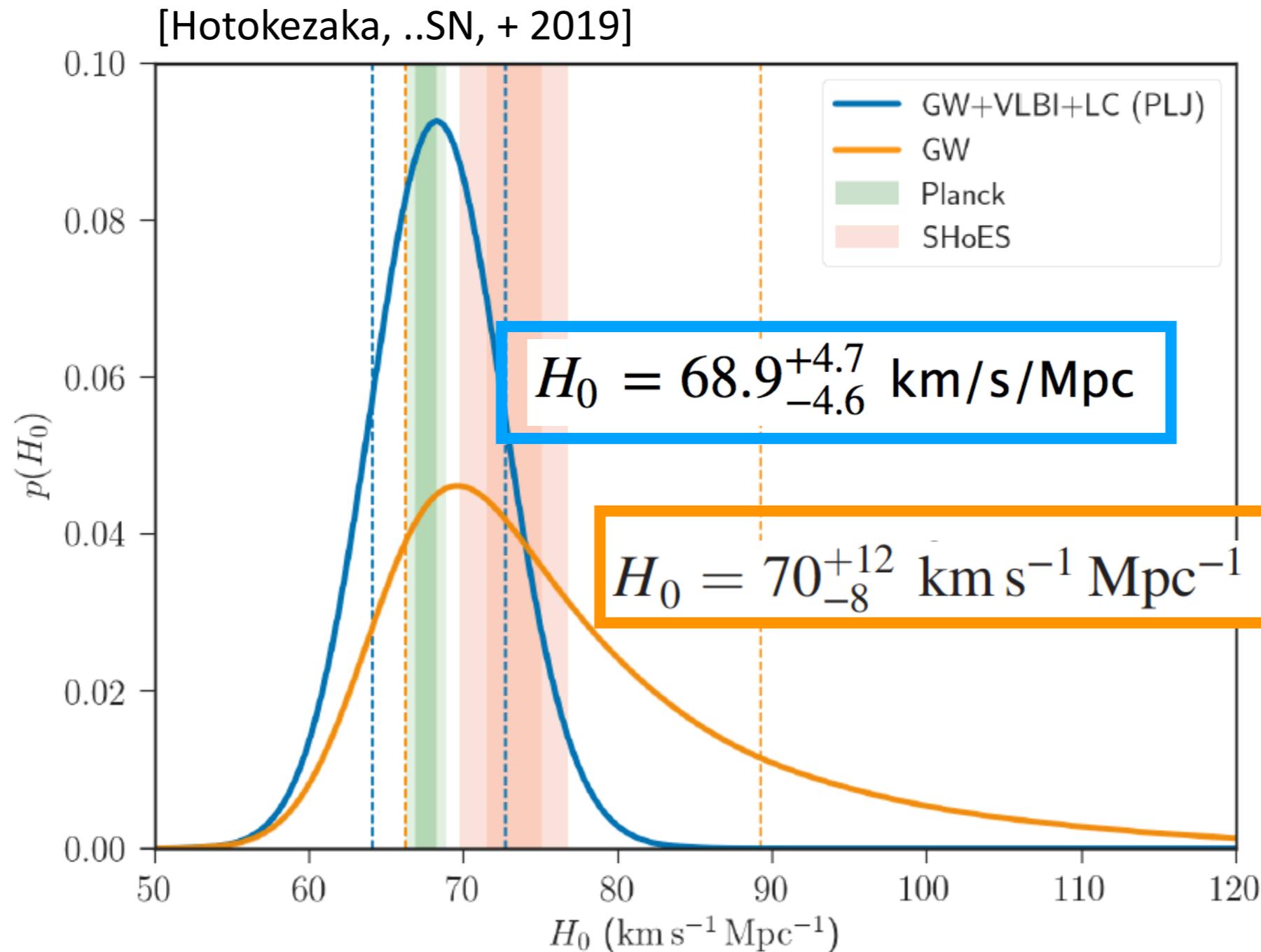
Proper motion of 2.7 mas, $\beta_{\text{app}} \sim 4.7$

GW+ radio: Hubble measurement improves by a factor of 2

[Hotokezaka, ..SN, + 2019, Nature Astro.]



GW+radio: Hubble measurement improves by a factor of 2

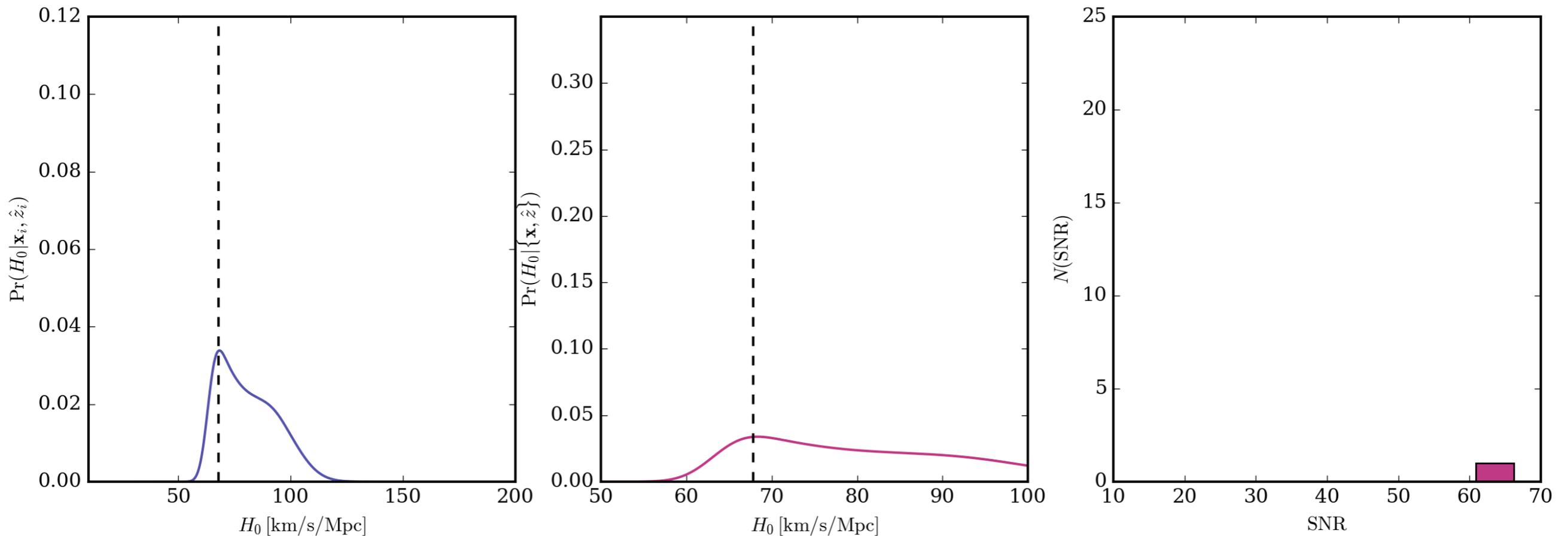


[for improvement of peculiar velocity inference;
Mukherjee, .. SN+ 2019, Davies + 2019, Nicolaou, +2019]

Caveat: astrophysical modelling of jet

GW+EM: importance of populations to potentially resolve “Hubble trouble”

[Feeney, Peiris, Williamson, SN,...,PRL, 2019; Mortlock, ... SN, 2019]



[see also Chen+ 2018, ...]

50 binaries (~8-10 years) to reach a precision of 1.8 % ($1/\sqrt{N}$);
high SNR binaries dominate joint PDF;
important caveat is **assumption of EM**

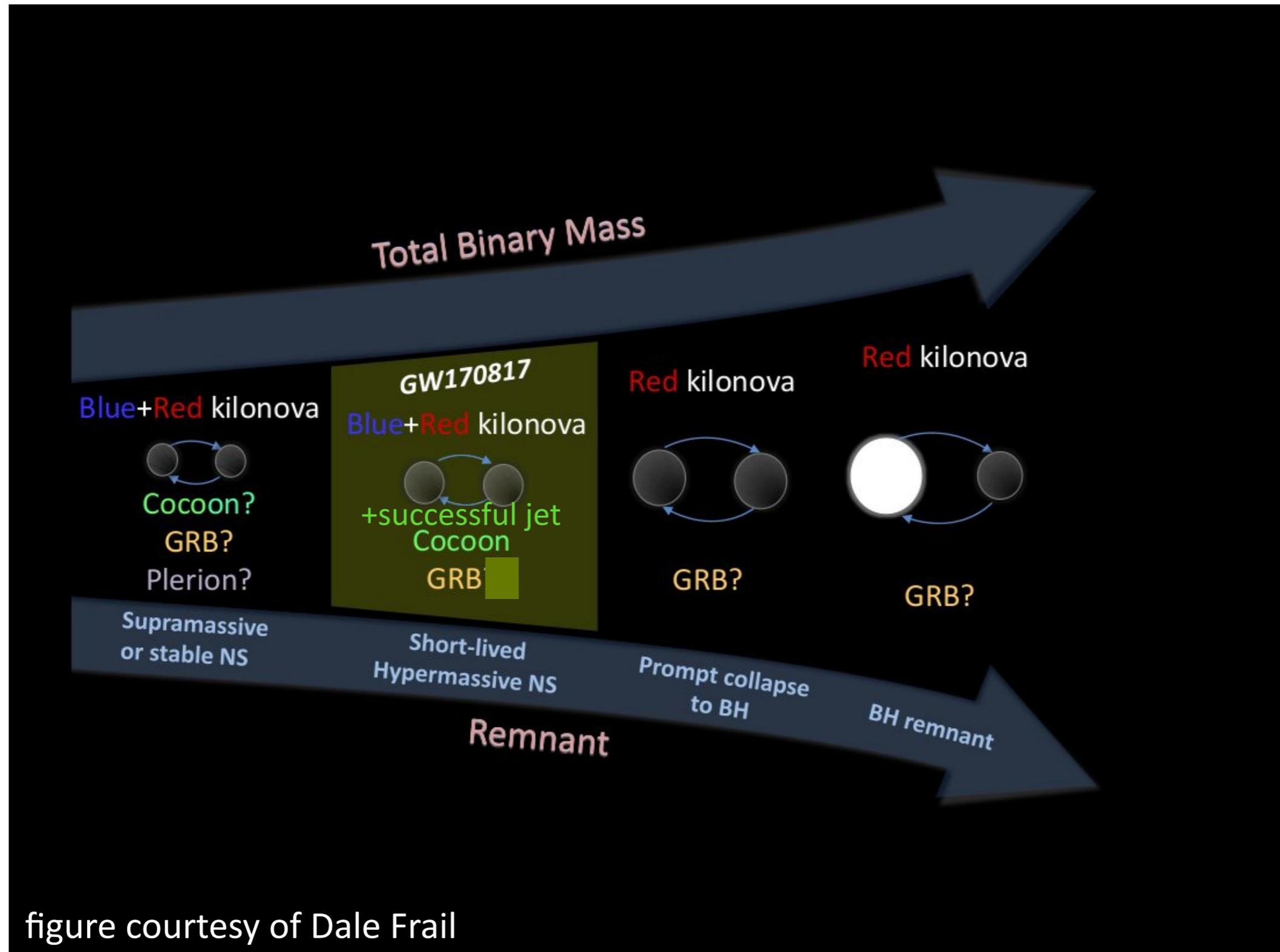
Part IV:

Perspectives

GW170817: follow-up was easy — very close by and bright,
small GW volume

BBH merger rate: $23.9^{+14.9}_{-8.6} \text{ Gpc}^3 \text{ yr}^{-1}$
NS-NS merger rate: $320^{+490}_{-240} \text{ Gpc}^3 \text{ yr}^{-1}$
NS-BH merger rate: $< 610 \text{ Gpc}^3 \text{ yr}^{-1}$

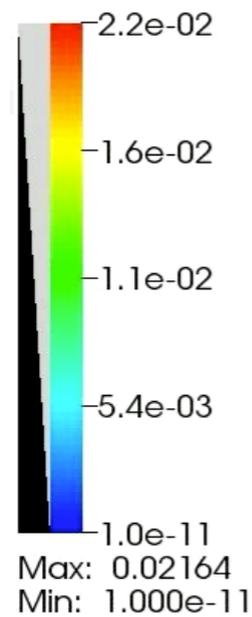
1. GW+EM: Expect a diversity of EM counterparts



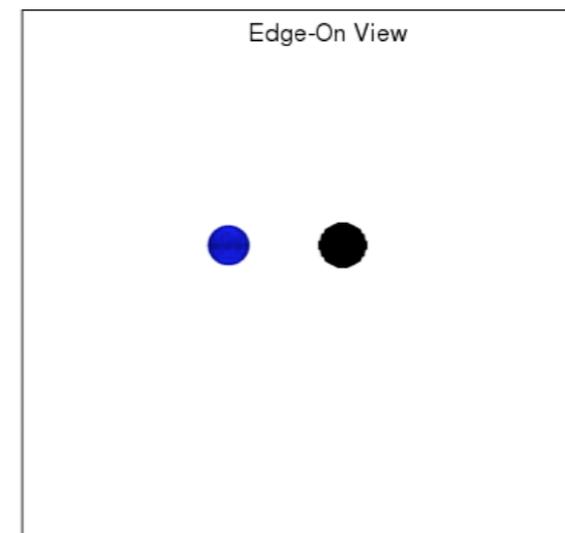
2) GW+EM: New discovery space — NS-BH mergers!

Movie courtesy of Francois Foucart

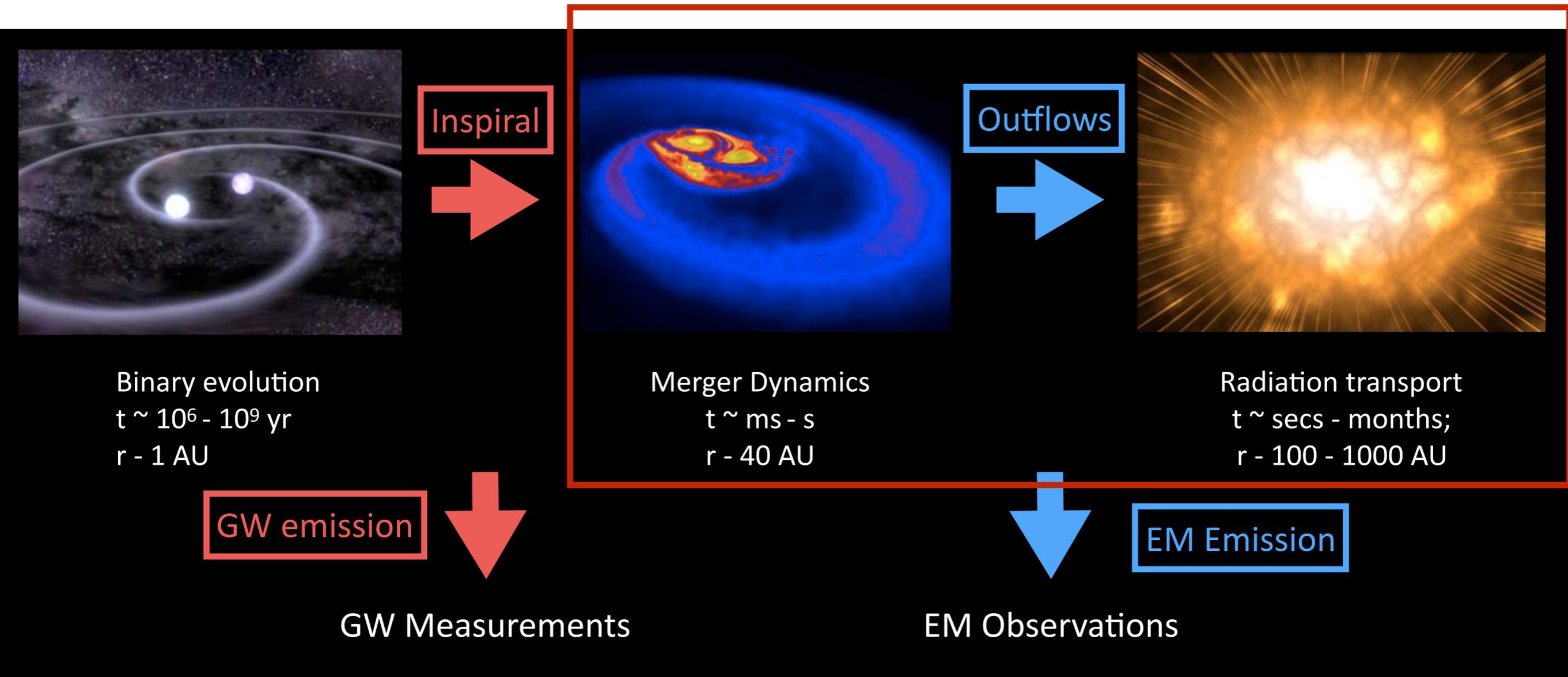
Density



Time=0

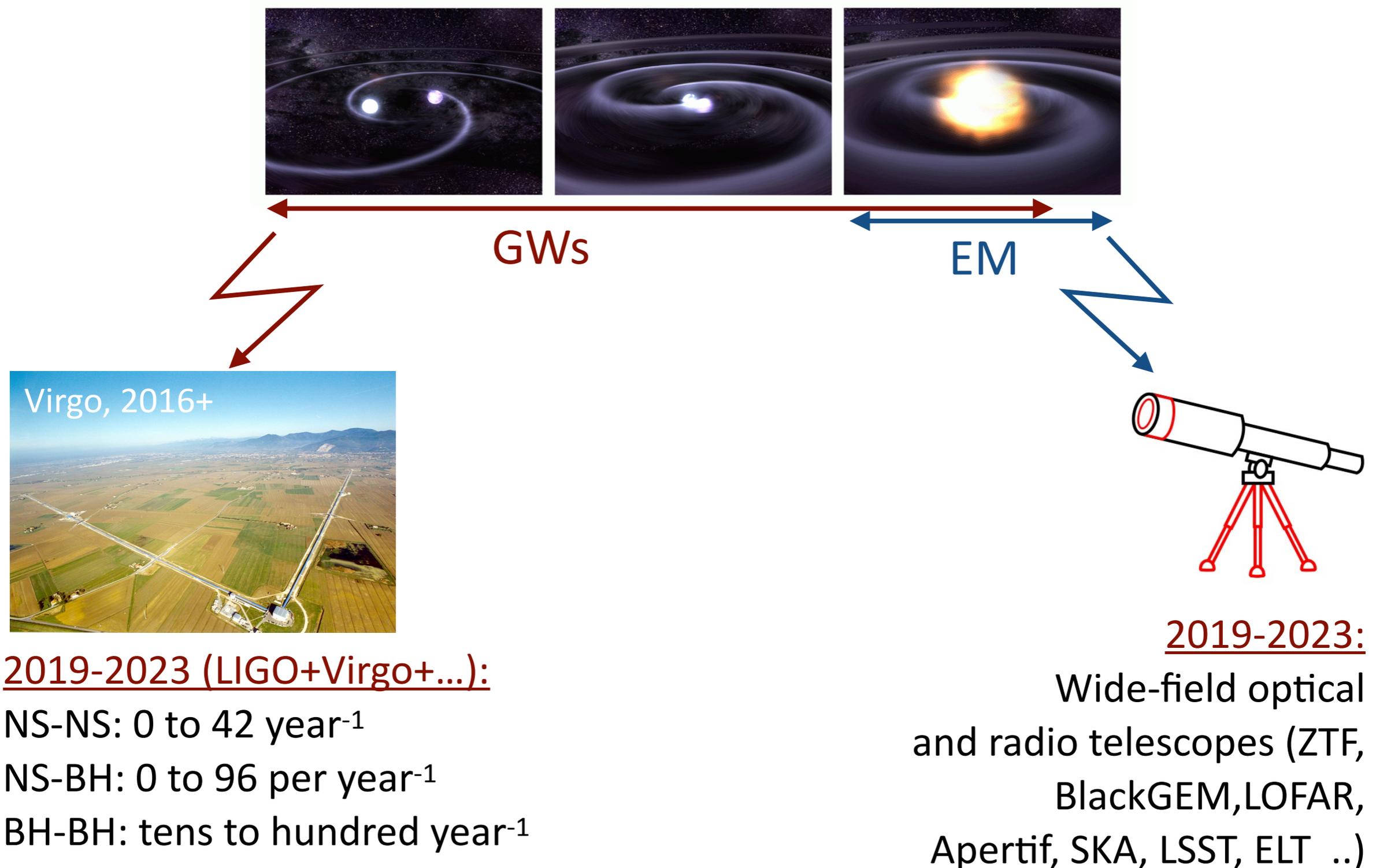


3. **GW+EM** challenge remains to build a coherent model: a key step to joint characterization



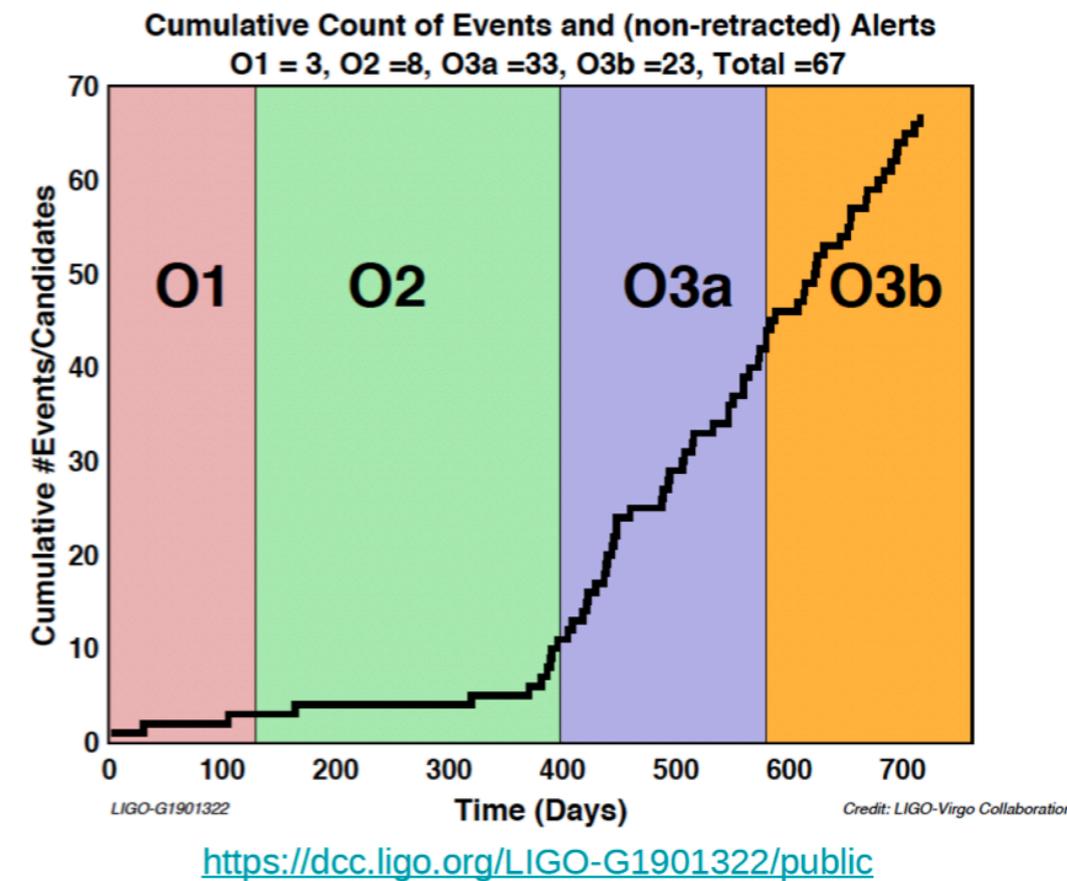
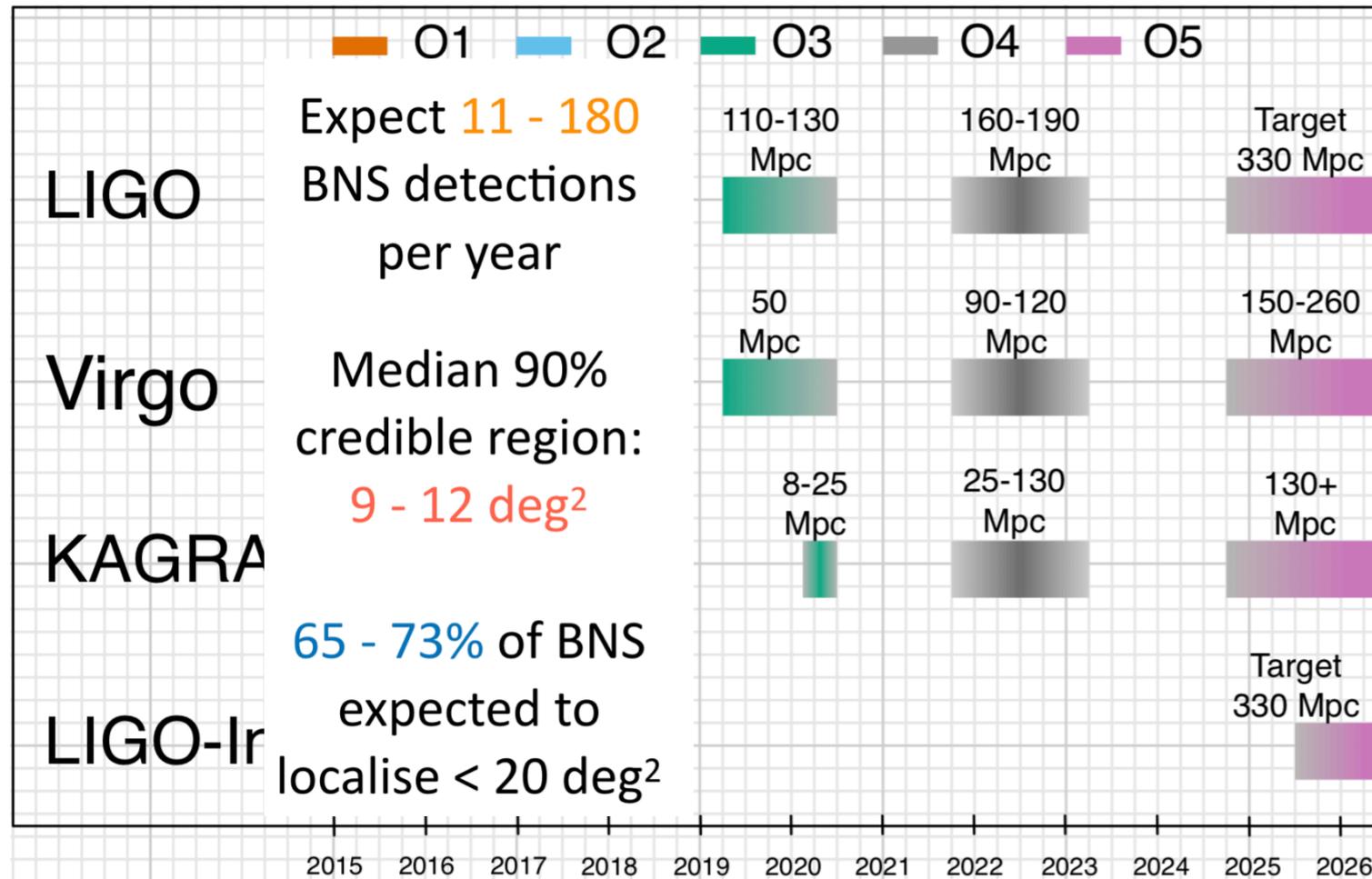
Outflows carry energy from small (10^6 cm) to large distances (10^{14} - 10^{16} cm) for radiation to escape.

4. GW+EM challenge: from individual objects to statistics

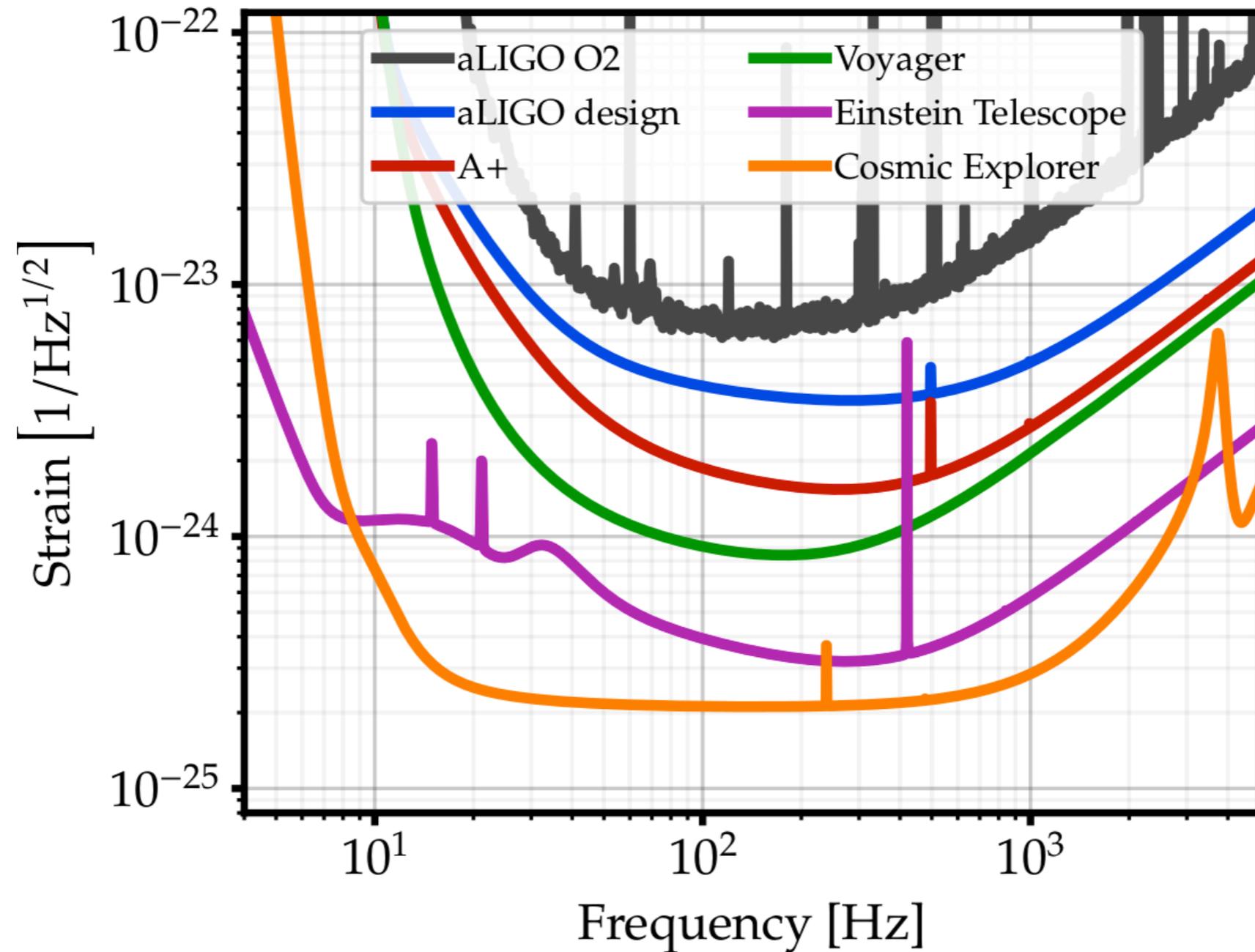


Mid 2020s: small GW areas, depth and rate

[B. P. Abbott et al., Living Rev Relativ, 2019, 1304.0670]

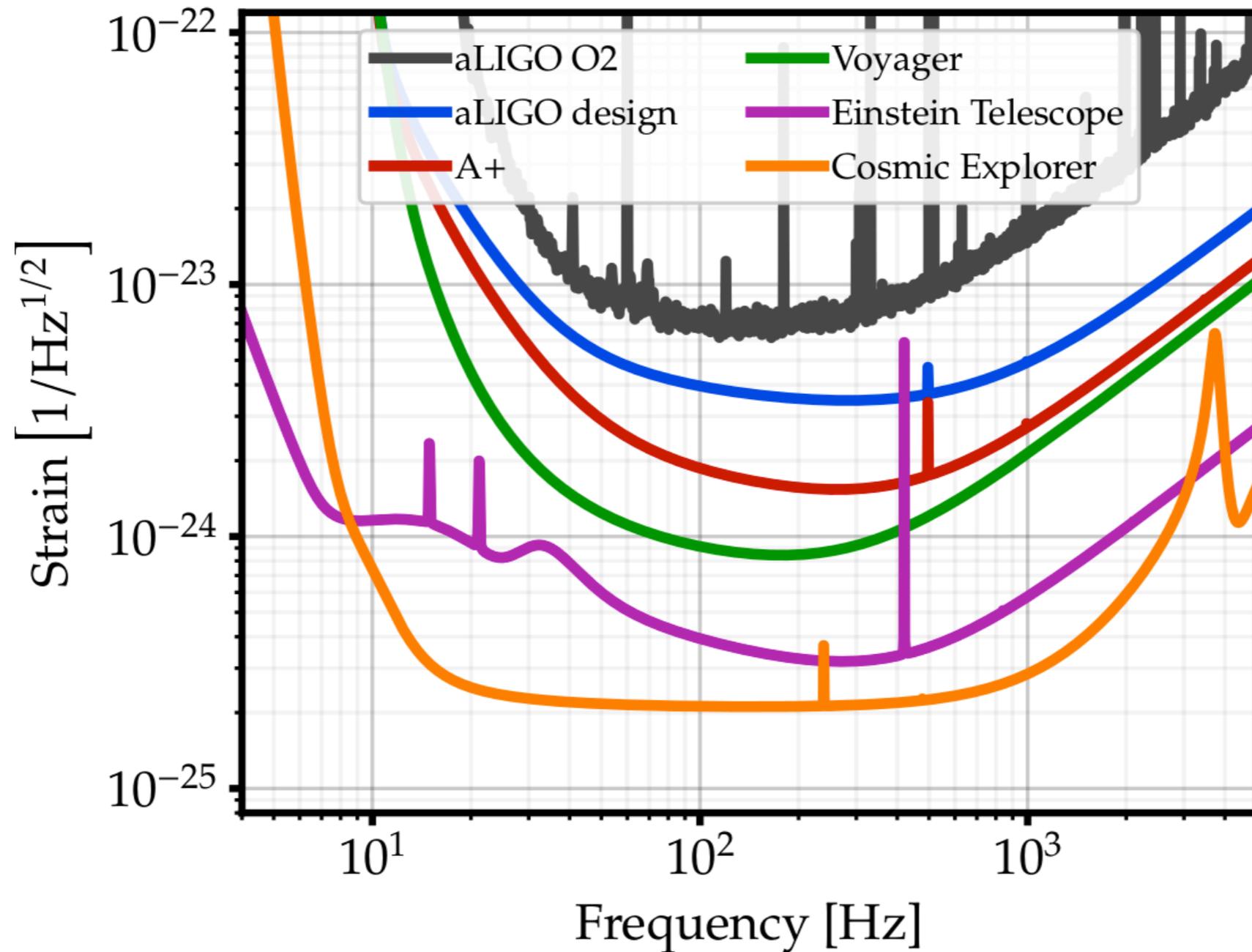


Mid 2020s - 2030s: aLIGO+, Cosmic Explorer



Factor of **2/10** in sensitivity; x **8/1000** in rates

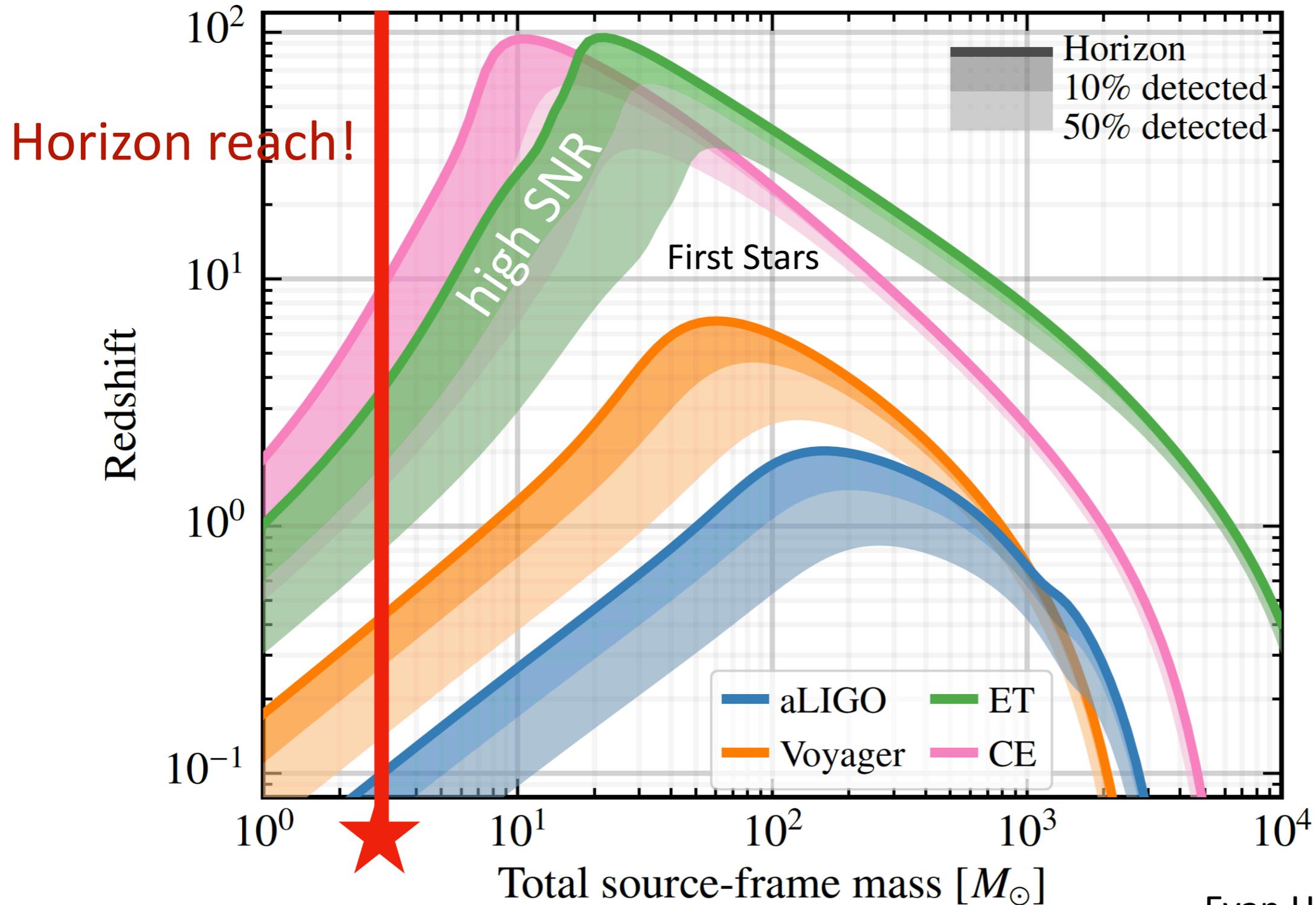
Mid 2020s - 2030s: aLIGO+, Cosmic Explorer



Sensitivity and frequency band increase

2030s: Einstein Telescope and Cosmic Explorer have cosmological reach

MMA co-chairs: Bailes, Kasliwal, Nissanke

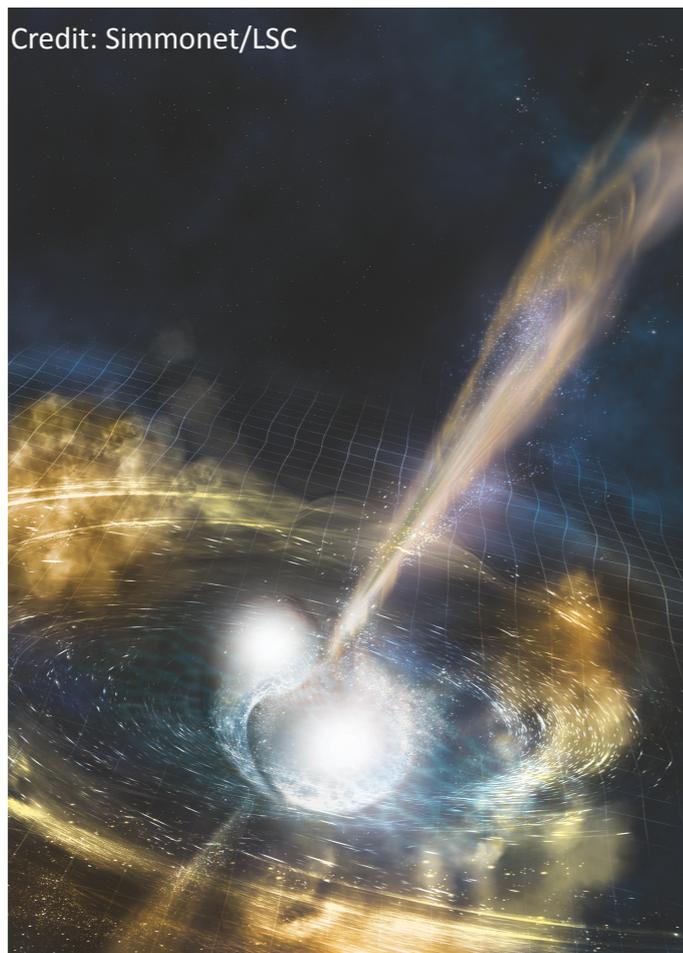


GW170817: $z \sim 10^{-2}$

Evan Hall, MIT

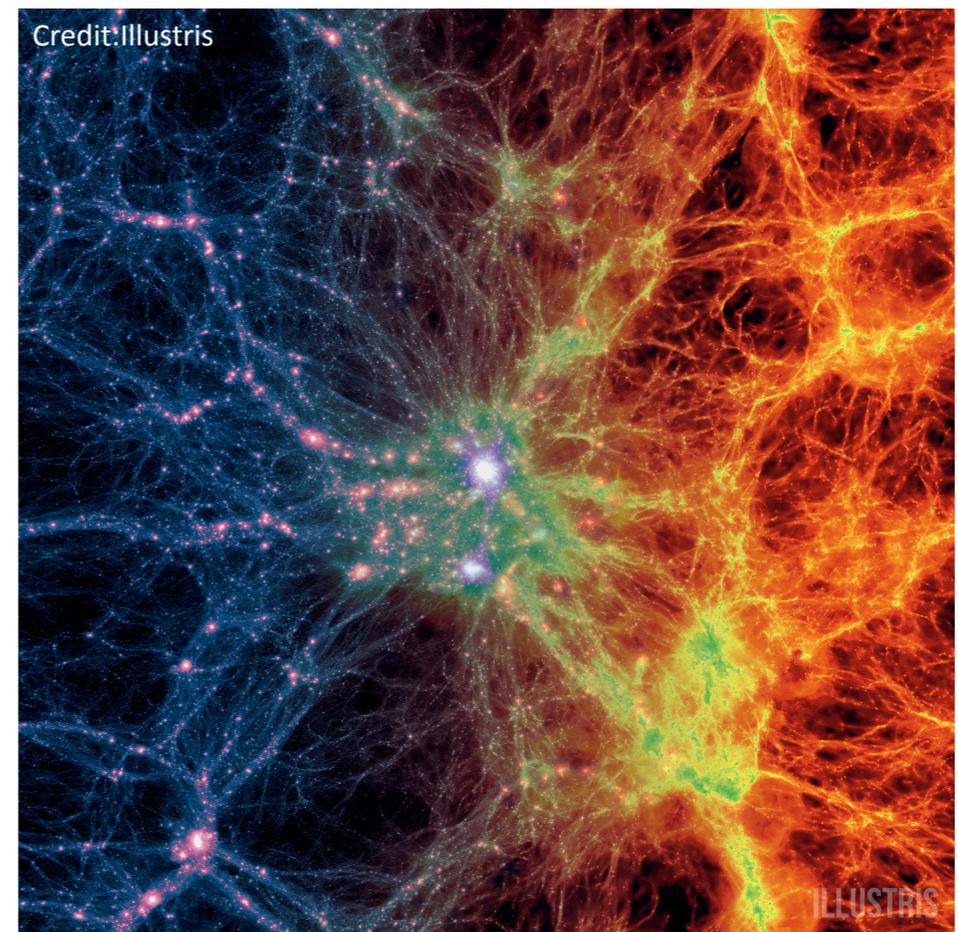
MMO in 2030s is not just EM follow up!

EM follow up of single sources

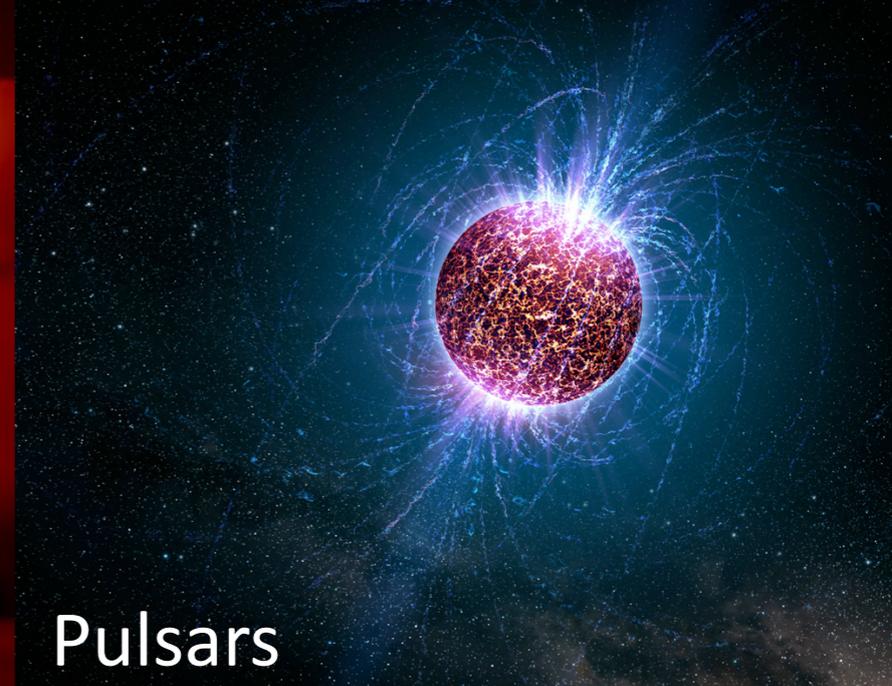
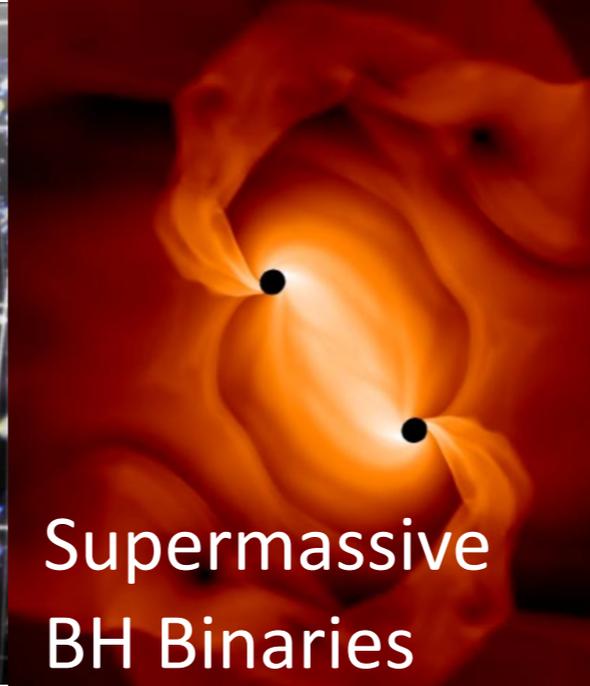
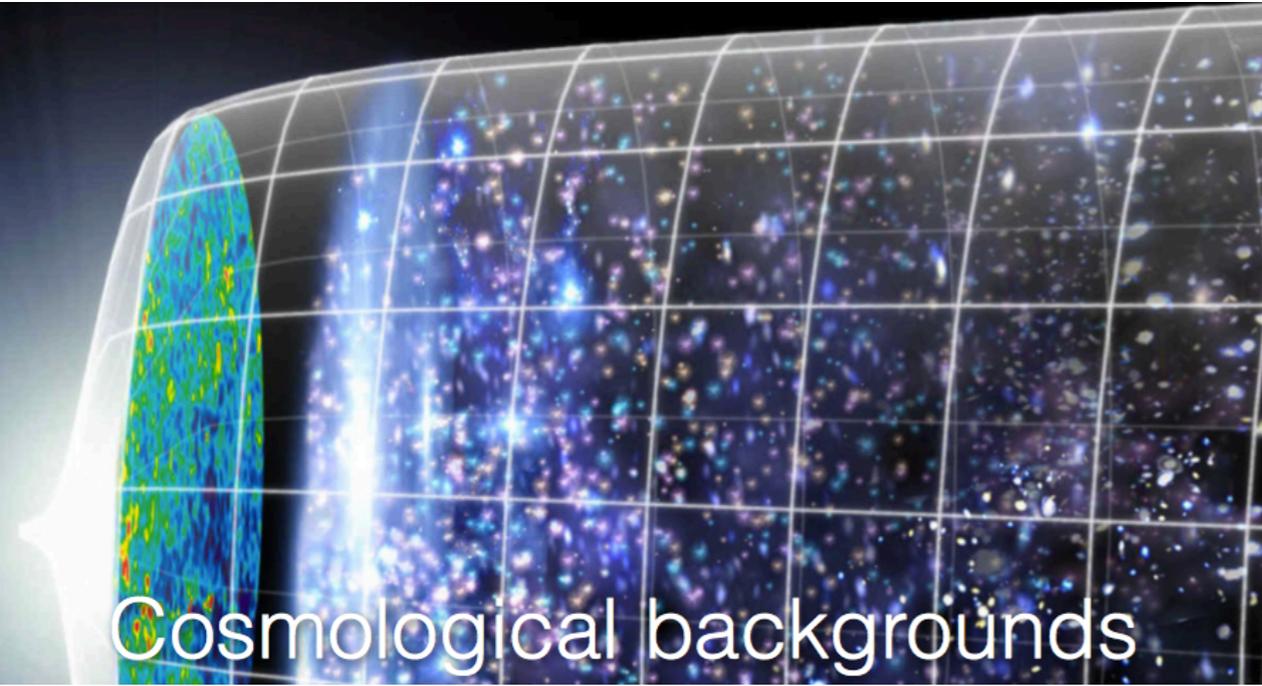


Cherry Pick Loud events
- golden for GW+EM

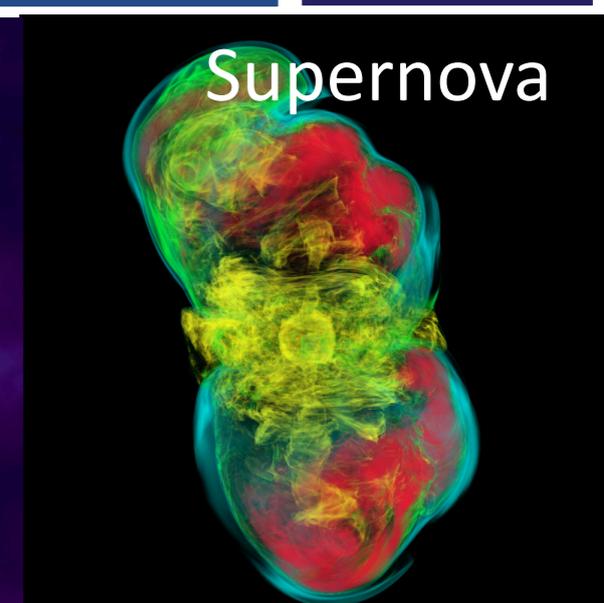
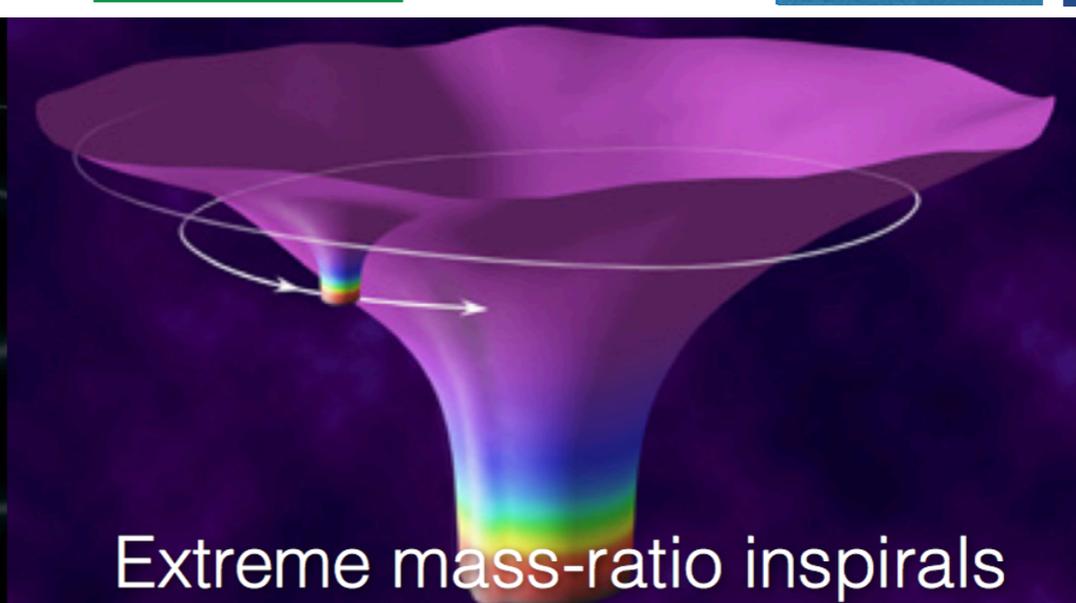
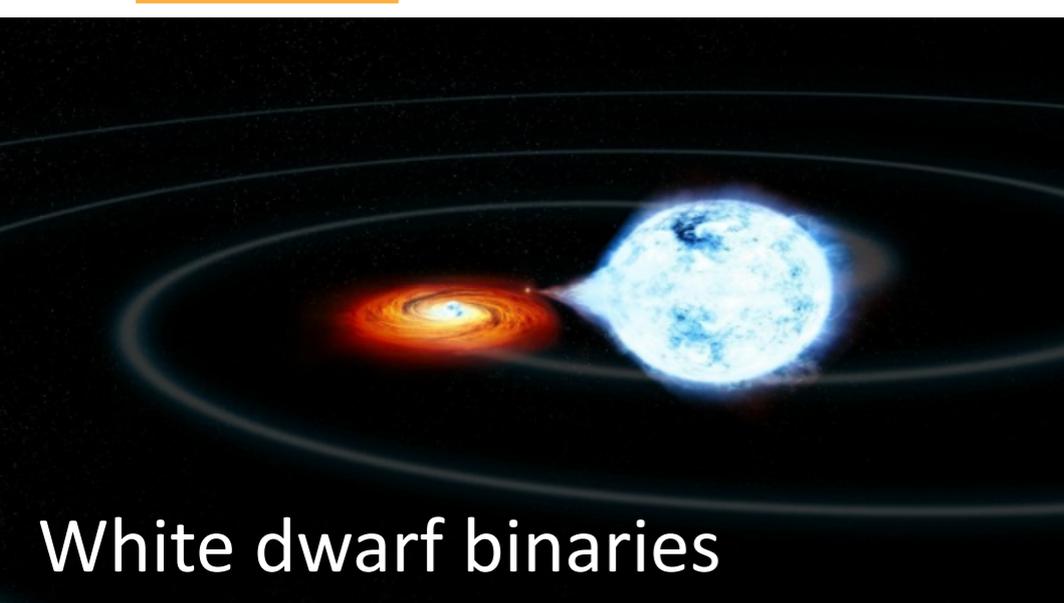
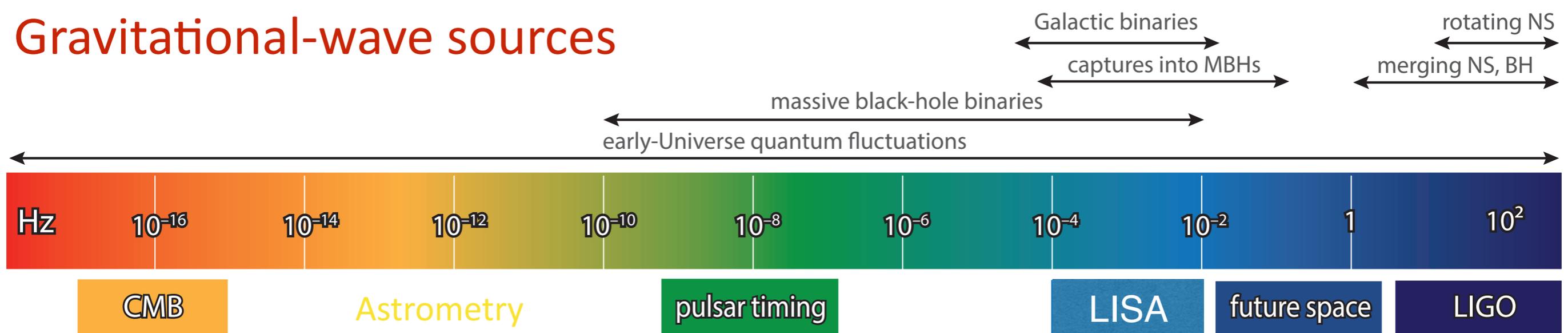
Cross correlating GW and EM source catalogs



Large Scale Structure;
Extragalactic Astronomy



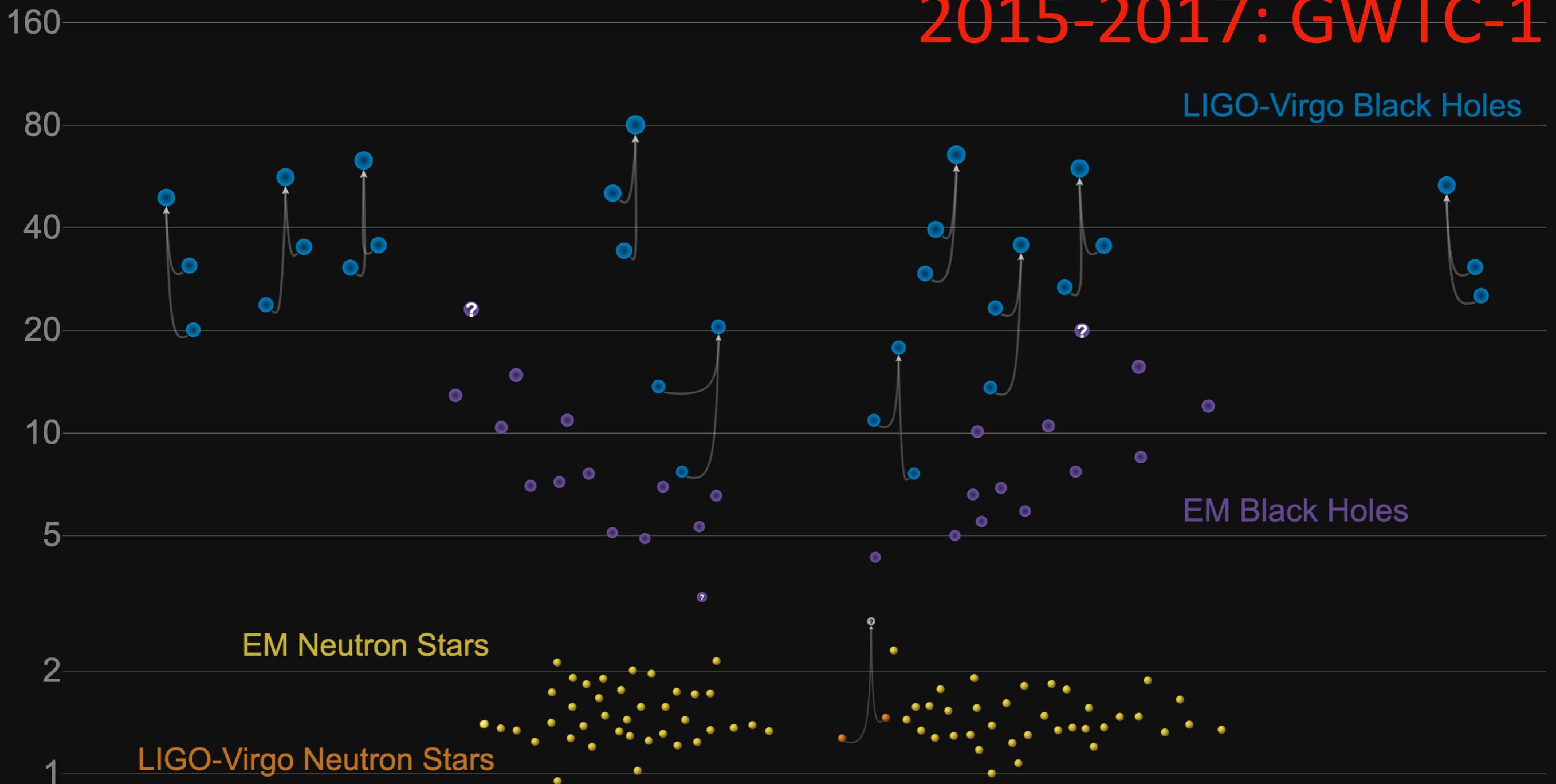
Gravitational-wave sources



Masses in the Stellar Graveyard

in Solar Masses

2015-2017: GWTC-1



GWTC-2 plot v1.0

[see LVC, arXiv 1811.15007 (GWTC-1)]

LIGO-Virgo | Frank Elavsky, Aaron Geller | Northwestern